

The JCMT Gould Belt Survey: first results from SCUBA-2 observations of the Cepheus Flare region

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ABSTRACT

We present observations of the Cepheus Flare obtained as part of the James Clerk Maxwell Telescope (JCMT) Gould Belt Legacy Survey (GBLS) with the SCUBA-2 instrument. We produce a catalogue of sources found by SCUBA-2, and separate these into starless cores and protostars. We determine masses and densities for each of our sources, using source temperatures determined by the *Herschel* Gould Belt Survey. We compare the properties of starless cores in four different molecular clouds: L1147/58, L1172/74, L1251 and L1228. We find that the core mass functions for each region typically show shallower-than-Salpeter behaviour. We find that L1147/58 and L1228 have a high ratio of starless cores to Class II protostars, while L1251 and L1174 have a low ratio, consistent with the latter regions being more active sites of current star formation, while the former are forming stars less actively. We determine that if modelled as thermally supported Bonnor–Ebert spheres, most of our cores have stable configurations accessible to them. We estimate the external pressures on our cores using archival ¹³CO velocity dispersion measurements and find that our cores are typically pressure confined, rather than gravitationally bound. We perform a virial analysis on our cores, and find that they typically cannot be supported against collapse by internal thermal energy alone, due primarily to the measured external pressures. This suggests that the dominant mode of internal support in starless cores in the Cepheus Flare is either non-thermal motions or internal magnetic fields.

Key words: stars: formation – dust, extinction – submillimetre: ISM.

1 INTRODUCTION

The Cepheus Flare region is a collection of star-forming molecular clouds extending to $\sim 10^\circ$ – 20° above the galactic plane at a galactic longitude of $\sim 110^\circ$ (Hubble 1934). Star formation is occurring at several different distances along the line of sight towards the

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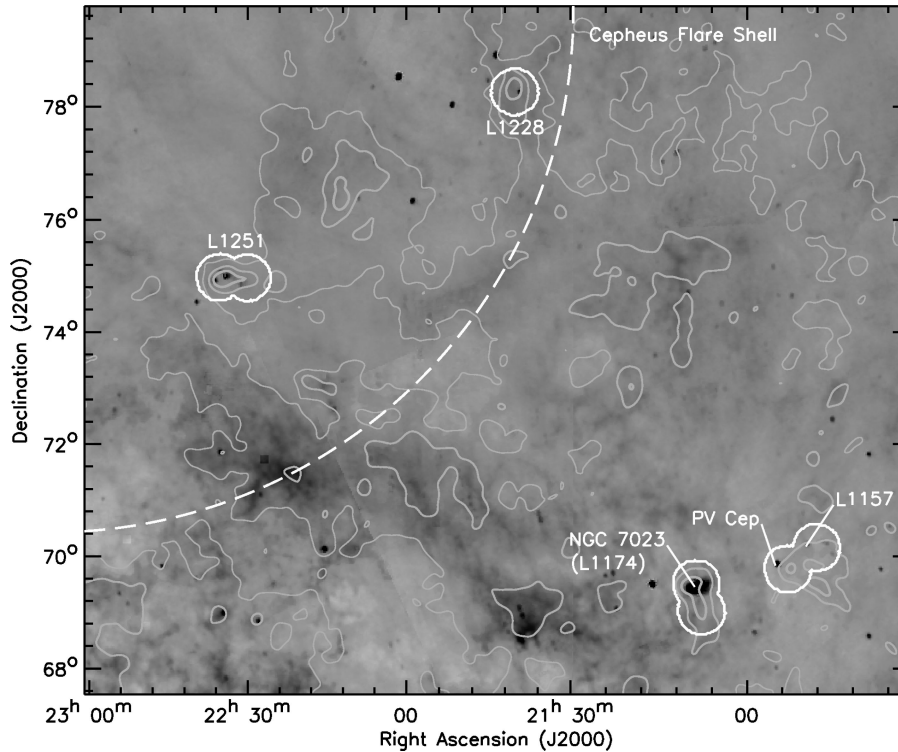


Figure 1. A finding chart of the Cepheus region. The grey-scale image shows *IRAS* 100- μm emission (Miville-Deschênes & Lagache 2005). The grey contours show A_V extinctions of 0.1, 0.5 and 1.0, smoothed with an 8-pixel Gaussian for clarity (Dobashi et al. 2005). The regions enclosed in solid white lines were observed as part of the JCMT GBS (Ward-Thompson et al. 2007). The reflection nebula L1174/NGC 7023 is marked. The L1172 region is immediately to the south of L1174. The variable star PV Cep and the protostar L1157-mm, both in the L1147/58 region, are also marked. The dashed white line shows the approximate position of the Cepheus Flare Shell (K09) – the CFS.

Cepheus Flare: at ~ 160 pc, where star formation is associated with the edge of the Local Bubble; at ~ 300 pc, associated with the Gould Belt; and at ~ 800 pc, associated with the Perseus arm of the Galaxy (Kun, Kiss & Balog 2008, and references therein; Kirk et al. 2009, hereafter K09).

The Gould Belt is a ring of molecular clouds and OB associations ~ 1 kpc in diameter and inclined $\sim 20^\circ$ to the galactic plane (Herschel 1847; Gould 1879). The Gould Belt is considered a ‘laboratory’ for the study of low-mass star formation, as most of the low-mass star-forming regions within 500 pc of the Earth are associated with it. As a result, surveys aimed at mapping substantial fractions of the Gould Belt have been undertaken using the JCMT (Ward-Thompson et al. 2007), the *Herschel Space Observatory* (André et al. 2010), and the *Spitzer Space Telescope* (Evans et al. 2009).

In this paper, we present SCUBA-2 observations of the intermediate-distance material in Cepheus associated with the Gould Belt. These data were taken as part of the James Clerk Maxwell Telescope (JCMT) Gould Belt Legacy Survey (GBLS; Ward-Thompson et al. 2007). There are five dark cloud complexes in the Cepheus Flare, which are associated with the Gould Belt: L1147/48/52/55/57/58, L1172/74, L1247/51, L1228 and L1241 (Lynds 1962). We present SCUBA-2 data for all or part of each of these regions, with the exception of L1241.

The Cepheus Flare is a sparsely filled region in which star formation appears to be proceeding in a variety of different environments. *IRAS* 100 μm observations of the Cepheus Flare (Miville-Deschênes & Lagache 2005) are shown in Fig. 1, with contours of A_V extinction overlaid (Dobashi et al. 2005). The regions of the highest visual

extinction are not distributed evenly across the Cepheus Flare, but instead are principally located on its north-eastern and south-western sides. In addition, Cepheus has a central region of relatively low extinction ($A_V < 3$; Dobashi et al. 2005) in which little star formation is occurring, although there is not a complete lack of molecular gas or young stars here (Tachihara et al. 2005). K09 found that YSOs in the Cepheus Flare are typically found in small, isolated groups, with a much higher fraction of distributed YSOs (the fraction of YSOs not associated with a group) than is typical: 41 per cent of YSOs in Cepheus are distributed, compared to an average of ~ 10 per cent across clouds observed as part of the *Spitzer* c2d survey (Evans et al. 2009).

The Cepheus Flare is defined by the interaction of a collection of shells with the local ISM, of which the most significant to the evolution of the region appears to be the Cepheus Flare Shell (CFS; Grenier et al. 1989; Olano, Meschin & Niemela 2006), an expanding supernova bubble with a radius ~ 9.5 , whose centre is located to the east of the Cepheus Flare at galactic coordinates $l \sim 120^\circ$, $b \sim 17^\circ$. The approximate position of the CFS is marked on Fig. 1. The shell divides the north-eastern and south-western star-forming regions. Olano et al. (2006) suggest that star formation in the eastern regions of the Cepheus Flare has been triggered by the passage of the CFS. K09 note that the current position of the CFS is consistent with that of L1228, and suggest that star formation in this region is being enhanced by the interaction with the shell. A possible geometry of the clouds associated with the CFS is proposed by Kun et al. (2008). In this geometry, the various intermediate-distance dark clouds are located approximately on the current surface of the CFS. As the CFS has an approximate radius of ~ 50 pc and is located at a distance

Table 1. Cepheus regions observed as part of the JCMT GBLs, with approximate central positions in equatorial and galactic coordinates listed.

Region	R.A. (J2000) (hours:min)	Dec. (J2000) (deg:arcmin)	l ($^{\circ}$)	b ($^{\circ}$)	Distance (pc)	Distance reference
L1147/58	21:02	+68:00	104.0	14.1	325 ± 13	Straizys et al. (1992)
L1172/74	20:41	+67:52	102.6	15.6	288 ± 25	Straizys et al. (1992)
L1251	22:34	+75:14	114.4	14.7	300^{+50}_{-10}	Kun et al. (2008)
L1228	20:58	+77:38	111.7	20.2	200^{+100}_{-10}	Kun et al. (2008)

of ~ 300 pc from the Earth (Olano et al. 2006), there are significant differences, both fractional and absolute, between the distances of the various dark clouds associated with the CFS, despite those dark clouds appearing along very similar lines of sight (see Table 1 for distances).

In this study, we identify, and investigate the properties of, starless cores in the Cepheus Flare. We investigate the cores' stability against collapse and the relative importance of gravity and external surface pressure in their confinement. Previous analysis of GBS data of the Ophiuchus molecular cloud (an intermediate-mass star-forming region forming stars in a clustered manner; e.g. Wilking, Gagné & Allen 2008) has suggested that dense starless cores in that region are typically confined by external surface pressure rather than self-gravity (Pattle et al. 2015). We here investigate whether starless cores in the various different environments found in the Cepheus Flare behave in a similar manner.

This paper is laid out as follows. In Section 2, we discuss the observations and data reduction. In Section 3, we discuss source extraction and characterization, and present our catalogue of sources. In Section 4, we discuss the properties of the starless cores in our catalogue. In Section 5, we discuss the counting statistics of starless and protostellar sources in Cepheus. In Section 6, we assess the stability of our cores using the Bonnor–Ebert (BE) criterion. In Section 7, we discuss the energy balance in the starless cores in our catalogue, and put an upper limit on the degree to which the cores are virially bound. In Section 8, we summarize our conclusions.

2 OBSERVATIONS

The SCUBA-2 (Holland et al. 2013) observations used here form part of the JCMT GBLs (Ward-Thompson et al. 2007). Continuum observations at $850 \mu\text{m}$ and $450 \mu\text{m}$ were made using fully sampled 30 arcmin diameter circular regions (PONG1800 mapping mode; Kackley et al. 2010) at resolutions of 14.1 arcsec and 9.6 arcsec, respectively. The Cepheus Flare was observed with SCUBA-2 in 41 observations taken between 2012 March 30 and 2014 October 24. The L1174 field was observed four times in very dry (Grade 1; $\tau_{225\text{GHz}} < 0.05$) weather. The remaining fields were each observed six times in dry (Grade 2; $0.05 < \tau_{225\text{GHz}} < 0.08$) weather, except for one field, L1147/58 East (containing the star PV Cep, discussed next), which was observed seven times. Larger regions were mosaicked with overlapping scans. Four final output maps were produced, the central co-ordinates of which are listed in Table 1.

The data were reduced using an iterative map-making technique (*makemap* in SMURF; Chapin et al. 2013), and gridded to 3 arcsec pixels at $850 \mu\text{m}$ and 2 arcsec pixels at $450 \mu\text{m}$, as part of the Legacy Release 1 (LR1) GBLs data set (see Mairs et al. 2015). The iterations were halted when the map pixels, on average, changed by < 0.1 per cent of the estimated map rms. The initial reductions of each individual scan were co-added to form a mosaic from which a mask based on signal-to-noise ratio was produced for each region.

The final mosaic was produced from a second reduction using this mask to define areas of emission. Detection of emission structure and calibration accuracy are robust within the masked regions, and are uncertain outside of the masked region.

A spatial filter of 10 arcmin is used in the reduction, which means that flux recovery is robust for sources with a Gaussian FWHM less than 2.5 arcmin. Sources between 2.5 arcmin and 7.5 arcmin in size will be detected, but both the flux and the size are underestimated because the Fourier components with scales greater than 5 arcmin are removed by the filtering process. Detection of sources larger than 7.5 arcmin is dependent on the mask used for reduction. The mask introduces further spatial filtering, as after all but the final iteration of the map-maker, all emission outside the region enclosed by the mask is suppressed. Therefore, the recovery of extended structure outside the masked regions is limited.

The data are calibrated in mJy arcsec^{-2} , using aperture flux conversion factors (FCFs) of 2.34 and $4.71 \text{ Jy pW}^{-1} \text{ arcsec}^{-2}$ at $850 \mu\text{m}$ and $450 \mu\text{m}$, respectively, derived from average values of JCMT calibrators (Dempsey et al. 2013). The estimated 1σ errors on the FCFs are $0.08 \text{ Jy pW}^{-1} \text{ arcsec}^{-2}$ and $0.50 \text{ Jy pW}^{-1} \text{ arcsec}^{-2}$ at $850 \mu\text{m}$ and $450 \mu\text{m}$, respectively. The PONG scan pattern leads to lower noise levels in the map centre and overlap regions, while data reduction and emission artefacts can lead to small variations in the noise level over the whole map.

The SCUBA-2 $850\text{-}\mu\text{m}$ data of Cepheus are shown in Figs 2 (L1172/74), 3 (L1147/58), 4 (L1251) and 5 (L1228). The sources we extract from the data are marked as coloured ellipses: light green in L1174, dark green in L1172, red in L1147/58, blue in L1251 and purple in L1228. This colour coding is continued throughout this paper.

The emission measured in the $850\text{-}\mu\text{m}$ filter on SCUBA-2 can be contaminated by the CO $J = 3 \rightarrow 2$ transition (Drabek et al. 2012) which, with a rest wavelength of $867.6 \mu\text{m}$, is covered by the SCUBA-2 $850\text{-}\mu\text{m}$ filter that has a half-power bandwidth of $85 \mu\text{m}$ (Holland et al. 2013). The only regions in the map which are expected to be substantially CO-contaminated are local to the PV Cep and L1157-mm protostars (discussed in Section 3.1), with which there are strong outflows associated (the CO contribution from the outflow associated with L1157-mm is clearly visible as extensions north and south of the source in Figs 3 and 6, below). However, as can be seen in Fig. 3, both PV Cep and L1157-mm are isolated objects, and CO emission from their outflows is unlikely to affect the fluxes measured for any of the other sources in the field.

Table 2 lists the 1σ RMS noise levels in each of the regions observed, measured on the default LR1 pixel widths of 2 arcsec at $450 \mu\text{m}$ and 3 arcsec at $850 \mu\text{m}$. The $450\text{-}\mu\text{m}$ RMS noise levels vary somewhat between different regions observed in the same weather band. This is due to the differing $450\text{-}\mu\text{m}$ sensitivity across Band 2 weather conditions. The $850\text{-}\mu\text{m}$ RMS noise is the highest in L1174, despite this region having been observed in the best weather, due to the presence of the NGC 7023 reflection nebula (see Section 3.1). The bright, extended emission from NGC 7023 makes

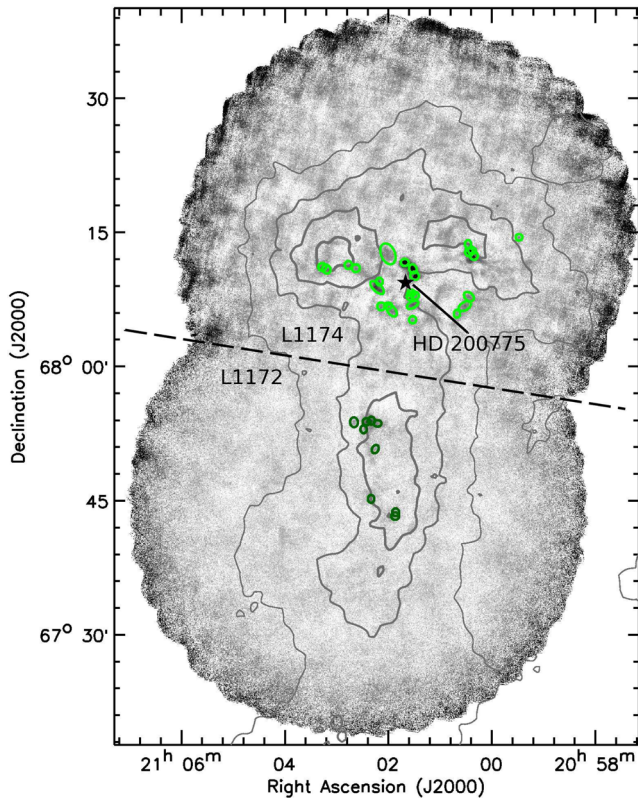


Figure 2. SCUBA-2 850- μm observations of the L1172 (south) and L1174 (north) regions, with sources marked. The ellipses show twice the FWHM size of each source. Contours of A_V of 0.5, 1.0, 1.5 and 2.0 magnitudes are shown for reference (Dobashi et al. 2005). The dashed line marks approximately the boundary between the L1172 and L1174 regions. The position of the Herbig Ae/Be star HD 200775 is marked.

it more difficult for the data reduction process to converge on a solution.

The 450- μm and 850- μm SCUBA-2 data presented in this paper are available at: <http://dx.doi.org/10.11570/16.0002>.

3 RESULTS

3.1 Cepheus Flare region

The Cepheus Flare consists of several distinct areas of high column density, each of which is at a different distance and likely to have a different star formation history. Thus, we consider each separately in the following analysis, and summarize their properties here.

L1172/L1174 is a site of clustered star formation. The dark cloud L1174, shown in the northern part of Fig. 2, is coincident with the NGC 7023 reflection nebula, also known as the Iris Nebula (Herschel 1802). The nebula is illuminated by the Herbig Ae/Be star HD 200775 (R.A. (J2000) = $21^{\text{h}}01^{\text{m}}39^{\text{s}}.920$, Dec. (J2000) = $+68^{\circ}09'47''.76$; van Leeuwen 2007) of spectral classification B2Ve (Guetter 1968). The position of HD 200775 is marked on Fig. 2, although HD 200775 itself is not visible in the SCUBA-2 data.

L1172 lies to the south of L1174, and is also shown in Fig. 2. It is forming stars much less actively than the neighbouring L1174.

L1147/L1158 contains the Lynds dark nebulae L1147, L1148, L1152, L1155, L1157, and L1158 (Lynds 1962). This region is considered to be the least affected by the CFS, and to be forming

stars with a low efficiency (K09). Only L1147, L1152, and L1155 were observed with SCUBA-2. All of the emission seen in the western area shown in Fig. 3 is associated with L1152, except for the bright protostar L1157-mm and its associated outflow (Kun et al. 2008), which are discussed below. All of the emission in the eastern region of Fig. 3 is associated with L1155, with the exception of the bright point source in the north-east, the star PV Cep (Li et al. 1994; discussed below).

Both L1152 and L1155 appear relatively quiescent (K09). There is some evidence that L1155 may be undergoing external heating: Nutter, Stamatellos & Ward-Thompson (2009) found evidence for a ~ 2 K temperature gradient across one of the cores in the region, L1155C, which they ascribed to the effects of the nearby A6V star BD+67 1263 (marked on Fig. 3).

The SCUBA-2 field contains two bright PMS stars: PV Cep (R.A. (J2000) = $20^{\text{h}}45^{\text{m}}53^{\text{s}}.943$, Dec. (J2000) = $+67^{\circ}57'38''.66$; Skrutskie et al. 2016) and L1157-mm (R.A. (J2000) = $20^{\text{h}}39^{\text{m}}06^{\text{s}}.2$, Dec. (J2000) = $+68^{\circ}02'15''$; K09). PV Cep is a highly variable (Kun et al. 2009) A5 Herbig Ae/Be star (Li et al. 1994), with which an extended outflow is associated (Reipurth, Bally & Devine 1997). PV Cep has a high westerly proper motion of ~ 20 km s^{-1} , and is likely to have escaped from the NGC 7023 cluster, which is discussed below (Goodman & Arce 2004). L1157-mm is a Class 0 protostar with an extremely strong molecular outflow (Chini et al. 2001). The outflow is highly visible in the 850- μm SCUBA-2 observations, and can be seen in Figs 3 and 6.

L1251, shown in Fig. 4, consists of three submillimetre-bright regions, the western L1251A, the central L1251C and the eastern L1251E (Sato et al. 1994), surrounded by a network of filaments. L1251 appears to be actively forming stars; in particular, there is a small group of young stars, L1251B, embedded within the L1251E region (Sato et al. 1994; Lee et al. 2007). K09 suggest that star formation in L1251 may have been triggered or enhanced by the passage of the CFS ~ 4 Myr ago.

L1228, shown in Fig. 5, is a small cloud that is likely to be located on the near side of the CFS, unlike the other clouds discussed here (Kun et al. 2008). L1228 runs $\sim 3^{\circ}$ along an approximately north-south axis. As can be seen from the extinction contours in Fig. 1, only the central part of L1228 was observed by the JCMT GBLS. K09 note that L1228 is at a location consistent with the current position of the CFS, and suggest that star formation here may be in the process of being enhanced by interaction with the shell.

Enlargements of the regions of significant 850- μm emission within each of the areas observed with SCUBA-2 are shown in Fig. 6.

3.2 Source extraction

We identified sources in the SCUBA-2 850- μm data using CSAR (Cardiff Source-extraction AlgoRithm; Kirk et al. 2013). CSAR is a dendrogram-based source-finding algorithm that was run in its non-hierarchical mode. CSAR identifies a source based on a peak in the emission map and assigns neighbouring pixels to that source if those pixels are above an assigned signal-to-noise criterion, and continues to do so until the contour level at which the source becomes confused with its neighbours is reached.

We gridded each of the SCUBA-2 850- μm maps on to 6-arcsec pixels before performing the source extraction. The LR1 default pixel size is 3 arcsec at 850 μm . However, the beam noise resulting from this oversampling of the data prevented CSAR from finding closed contours around extended low-surface-brightness sources.

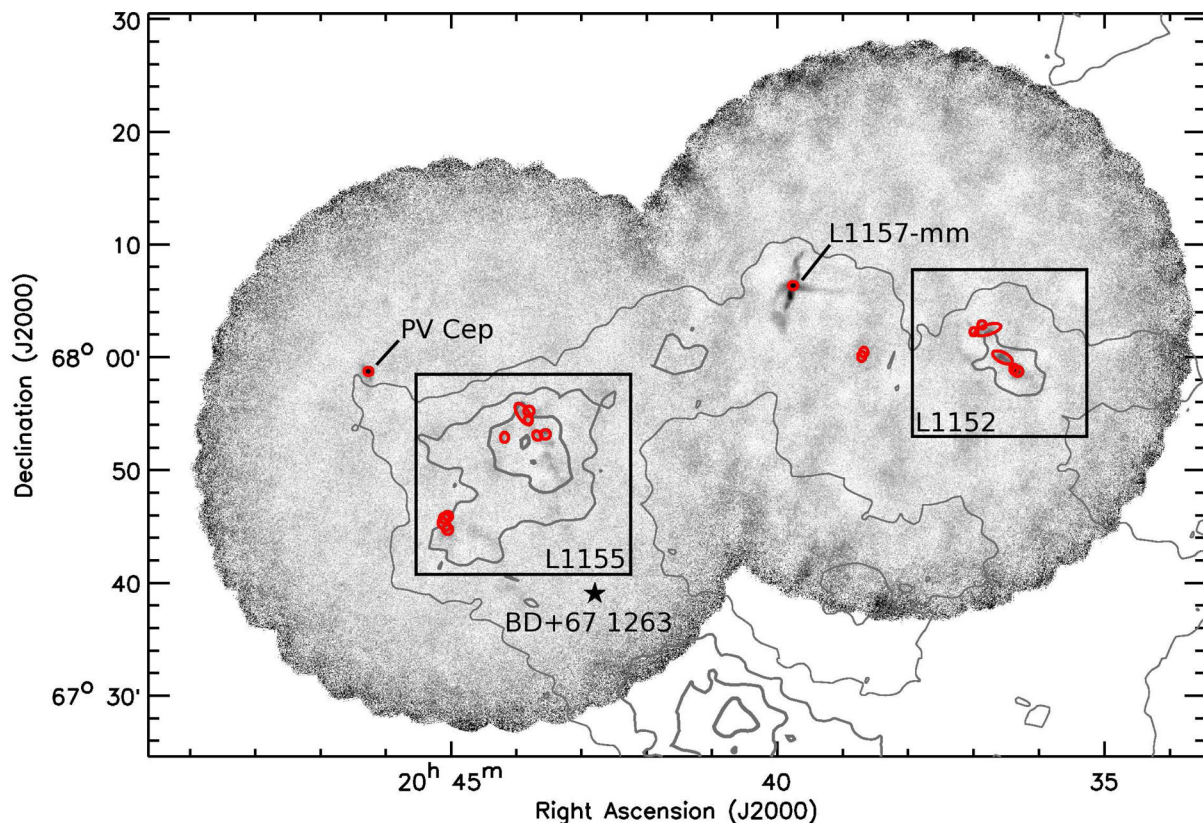


Figure 3. SCUBA-2 850- μm observations of the L1147/L1158 region, with sources marked. The ellipses show twice the FWHM size of each source. Contours of A_V of 0.5, 1.0, 1.5 and 2.0 mag are shown for reference (Dobashi et al. 2005). The protostars PV Cep and L1157-mm are labelled, and the position of the A6V star BD+67 1263 is marked. The boxes mark approximately the extent of the L1152 and L1155 regions.

Source extraction was performed on the low-variance regions of the maps, where the variance, as measured in the variance array, was very low, $<0.005 \text{ (mJy arcsec}^{-2}\text{)}^2$. The criteria chosen for a robustly detected source were a peak flux density $F_{\nu}^{\text{peak}} \geq 5\sigma$ and a minimum of a 1σ drop in flux density between adjacent sources (i.e. a local minimum in flux density at least 1σ less than peak value of the fainter of the two sources), where σ is the RMS noise level of the data. We adopted 1σ values of $0.041 \text{ mJy arcsec}^{-2}$ in L1174, and $0.028 \text{ mJy arcsec}^{-2}$ elsewhere on 6-arcsec pixels at $850 \mu\text{m}$.

We identified 27 sources in L1147/58, 26 sources in L1174, 9 sources in L1172, 42 sources in L1251 and 20 sources in L1228. Of the 27 sources in L1147/58, 7 were rejected due to their association with the L1157-mm outflow and hence likely to be artefacts resulting from CO contamination in the SCUBA-2 850- μm data. Rejecting these left us with 20 reliable sources in L1147/58. There were no sources in other regions that we considered likely to be CO artefacts.

The sources we identified in each cloud are shown in detail on Fig. 6, and on Figs 3–5 for reference. Due to the significant overlap between some of the sources, we fitted each source using a multiple-Gaussian fitting routine. This model, which utilizes the fitting routine *mpfit* (Markwardt 2009), is described in detail by Pattle et al. (2015). The fitting routine models the flux density of sources in crowded regions by fitting a two-dimensional Gaussian and an inclined-plane background to each of a set of associated sources simultaneously. Sources are considered to be neighbours if they are separated by less than twice the FWHM of the larger source. Groups to be fitted simultaneously are defined such that each source in a group is a neighbour to at least one other source in the group,

and no source has any neighbours outside of the group. The source positions and sizes determined using CSAR were supplied as initial input to the fitting routine. The Gaussian fitting routine was constrained such that for each source, the x and y coordinates of the source could vary no more than 6 arcsec from their initial position, the source semimajor and semiminor axes could not vary by more than 10 per cent of their initial values, and the source position angle could vary by no more than 5° . The total flux of the source was constrained to be positive.

It should be noted that while the Gaussian model is a popular and widely used choice of model for characterizing the properties of starless cores (e.g. Ward-Thompson et al. 1994; Hirota, Ito & Yamamoto 2002; Enoch et al. 2008; Gómez et al. 2014; Pattle et al. 2015), the underlying geometry of a starless core is unlikely to obey a Gaussian distribution, instead typically showing a flat central plateau and power-law wings (e.g. Alves, Lada & Lada 2001), which may be characterized using a BE geometry (Ebert 1955; Bonnor 1956) or a Plummer-like geometry (Plummer 1911; Whitworth et al. 1996). However, the Gaussian model remains a very useful tool for characterizing the properties of ensembles of starless cores, due to its analytic tractability. Gaussian fits may underestimate core size (Terebey, Chandler & Andre 1993), typically fitting the central plateau of the core and underestimating the extent of the wings. However, two arguments mitigate against the effect of this on our core sample. First, if we were significantly underestimating the size of our cores, then we would expect to see positive annuli of unfitted flux in the residuals of our Gaussian fits, which is not the case. Secondly, it can be shown that for Gaussian and Plummer-like distributions with the same total mass and central

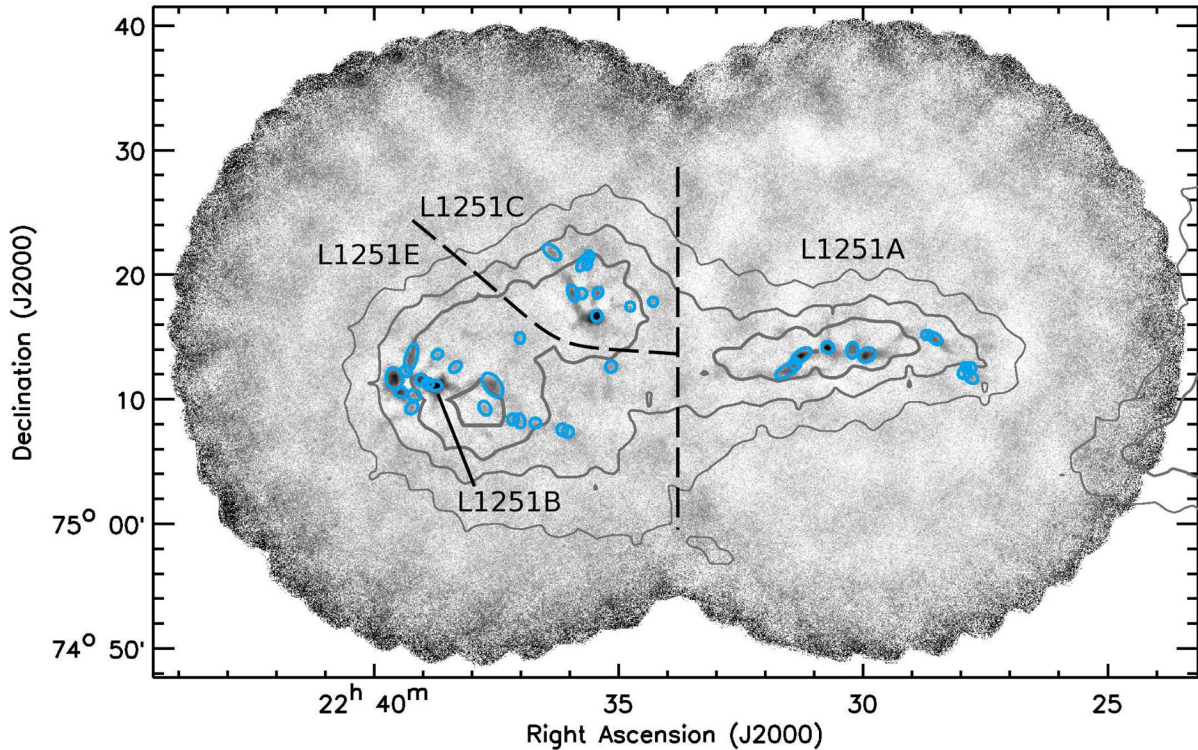


Figure 4. SCUBA-2 850- μm observations of the L1251 region, with sources marked. The ellipses show twice the FWHM size of each source. Contours of A_V of 0.5, 1.0, 1.5 and 2.0 mag are shown for reference (Dobashi et al. 2005). The dashed lines mark approximately the boundaries between the L1251A, L1251C and L1251E regions. The protostellar cluster L1251B is labelled.

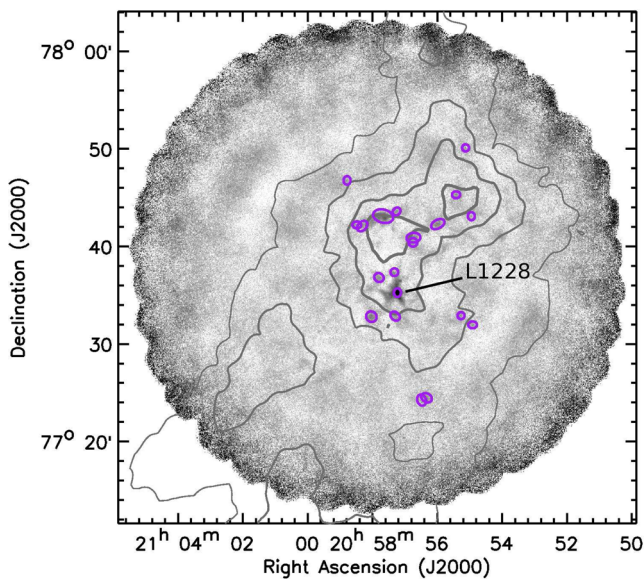


Figure 5. SCUBA-2 850- μm observations of the L1228 region, with sources marked. The ellipses show twice the FWHM size of each source. Contours of A_V of 0.5, 1.0, 1.5 and 2.0 mag are shown for reference (Dobashi et al. 2005). The protostar L1228 is labelled.

density, the characteristic sizes of the two distributions are very similar, $R_{\text{Plummer}} = 1.17 R_{\text{Gaussian}}$, where R_{Plummer} is the characteristic size of the Plummer-like distribution and R_{Gaussian} is the Gaussian width (assuming a power-law index for the Plummer-like distribution of 4; see Pattle 2016 for derivations of the masses of the two

distributions). This suggests that we are unlikely to be significantly underestimating the size of our cores using a Gaussian distribution.

In this analysis we are concerned with the ensemble properties of starless cores in the Cepheus molecular cloud, and so require an approximate size and mass estimate for each core, which can be usefully provided by a Gaussian fit to the data. Future detailed analyses of the interior structure of starless cores using SCUBA-2 data will require more sophisticated modelling of core geometries.

For each of our sources, Table 3 lists the position, angular size, orientation, peak and total flux densities, signal-to-noise ratio at 450 μm , classification as starless or protostellar and the region in which the source is located. For the 850- μm flux densities, both the modelled values and the values determined from aperture photometry are listed. For the 450- μm flux densities, only values determined from aperture photometry are listed. The aperture photometry measurements were made using elliptical apertures with major and minor axis diameters of twice the FWHM values listed in Table 3, and as shown in Figs 3–6.

Prior to aperture photometry measurements being made, the 450- μm data were convolved to match the resolution of the 850- μm data using a convolution kernel constructed as described by Pattle et al. (2015), following the method proposed by Aniano et al. (2011). The convolution kernel used was constructed using the SCUBA-2 450- μm and 850- μm beam models given by Dempsey et al. (2013). However, the peak 450- μm flux densities, and the 450- μm signal-to-noise ratios, were determined from the original, non-convolved map.

We emphasize that due to the significant overlap between many of the sources (see Figs 3–5), there will be double-counting of pixels in many of the flux densities determined from aperture photometry, and the flux density values determined from aperture photometry are

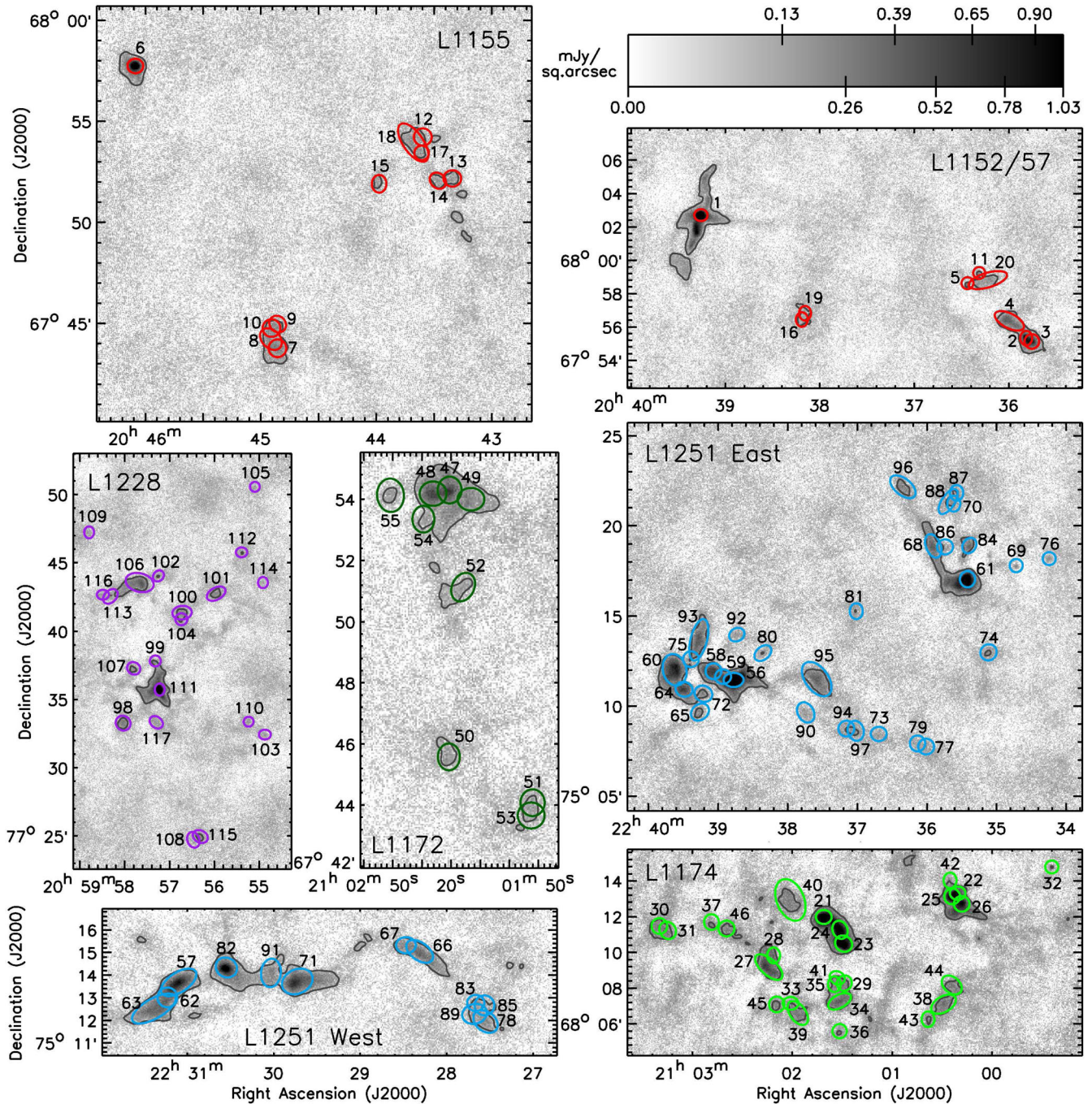


Figure 6. SCUBA-2 850- μm observations of regions of significant emission. Sources extracted in this work are numbered as in Table 3 and colour-coded by region: red – L1147/58; light green – L1172; dark green – L1174; blue – L1251; purple – L1228. The data are shown in square-root scaling.

likely to be overestimates of the amount of emission associated with a source. The aperture-photometry-determined peak flux densities are those of the brightest pixel in the source aperture, and so may be identical for overlapping sources. The modelling-determined peak flux densities are the best-fitting peak flux densities assuming the sources obey Gaussian distributions.

It can be seen in Table 3 that the aperture-photometry-determined 850- μm flux densities are typically ~ 30 per cent higher than the model 850- μm flux densities in isolated (non-overlapping) sources. This is due to the inclined-plane background that is fitted to the measured emission along with the Gaussian source model.

Note that the 450- μm and 850- μm aperture-photometry-determined flux densities do not have the SCUBA-2 aperture photometry corrections discussed by Dempsey et al. (2013) applied to them. The SCUBA-2 aperture photometry corrections are determined for point sources, and account for flux in the secondary beam of the JCMT not enclosed by a small aperture (the JCMT’s secondary beam has a FWHM of 25 arcsec at 450 μm and 48 arcsec at 850 μm ; see Dempsey et al. 2013). We do not use these aperture photometry corrections in this work, as their applicability to either extended sources or non-circular apertures is not certain. Furthermore, for aperture diameters from 25 to 50 arcsec (i.e. the

Table 2. The mean 1σ RMS noise levels in each of the regions observed, measured on the default LR1 pixel sizes of 2 arcsec at 450 μm and 3 arcsec at 850 μm .

Region	450- μm RMS mJy sqa^{-1}	850- μm RMS mJy sqa^{-1}
L1174	1.03 ± 0.08	0.069 ± 0.006
L1172	2.16 ± 0.16	0.056 ± 0.004
L1155	2.44 ± 0.11	0.056 ± 0.004
L1157	2.20 ± 0.09	0.055 ± 0.005
L1251 E	1.77 ± 0.09	0.059 ± 0.003
L1251 W	0.93 ± 0.04	0.054 ± 0.005
L1228	0.87 ± 0.05	0.059 ± 0.007

vast majority of our sources), the 450- μm and 850- μm aperture photometry corrections are identical, while for sources larger than 50 arcsec, the difference between the 450- μm and 850- μm corrections is very small, typically ~ 1 per cent (Dempsey et al. 2013). As we are using the aperture-photometry-determined flux densities only as a ratio quantity (see Section 3.4), use of the aperture photometry corrections (or otherwise) should not affect our results. However, as aperture-photometry-corrected flux densities may be useful for other purposes, we direct the reader to Dempsey et al. (2013) for further information.

In the analysis that follows, we use the best-fitting model 850- μm total flux densities in order to determine source masses. The ratio of the 450- μm and 850- μm aperture-photometry-determined total flux densities is used to determine source temperatures, for those sources with a peak 450- μm signal-to-noise ratio ≥ 3 – see Section 3.4.

3.3 Source characterization

Of the 117 sources in our Cepheus Flare catalogue, 23 were associated with at least one protostar in the K09 *Spitzer* catalogue (the K09 catalogue lists 143 protostellar sources and covers all of the regions observed with SCUBA-2). Protostar associations are listed in Table 4, along with the K09 source with which they are associated, the evolutionary class of that source (as determined from the infrared spectral index, α_{IR} by K09), and alternative identifications. It should be noted that due to the ~ 300 pc distances to the Cepheus Flare clouds, a single SCUBA-2 source in Cepheus may be associated with more than one protostellar object. In particular, source 56 contains six embedded sources, the L1251B group.

The K09 *Spitzer* catalogue is the only systematic protostar catalogue produced from *Spitzer* observations of Cepheus to date. We compared the K09 results to a more limited recent study by Dunham et al. (2013), who revise the classification of a number of protostars detected by the *Spitzer* c2d (Evans et al. 2009) and Gould Belt (P.I. L. Allen; see e.g. K09) surveys. Dunham et al. (2013) extend the methods developed by Evans et al. (2009) for correcting protostellar fluxes and luminosities for extinction, providing corrected classifications for *Spitzer*-detected protostars associated with at least one submillimetre detection at wavelengths $\geq 350\mu\text{m}$. Dunham et al. (2013) include 20 protostars in Cepheus in their sample, all of which are included in the K09 catalogue. The Dunham et al. (2013) extinction corrections alter the α_{IR} classification of two of the 20 stars that they consider in Cepheus, both of which we detect with SCUBA-2: Source 48 (K09 Source 50), which is reclassified from Flat to Class II, and Source 91 (K09 Source 66), which is reclassified from Class I to Flat. Source 111 (K09 source 9) also moves from

Class I to the Class I/Flat boundary. As these extinction-corrected classifications are available for only a subset of the *Spitzer* sources in Cepheus, and as only a small minority of the source classifications are changed by the correction for extinction, we continue to use the classifications given in K09 throughout this work. This is in order to use a self-consistent set of source classifications.

Temperatures for each of our sources were supplied by Di Francesco et al. (in preparation). These temperatures were determined from SED fitting to the 160–500- μm *Herschel* observations taken towards the Cepheus Flare as part of the *Herschel* Gould Belt Survey (GBS) (André et al. 2010). The *Herschel* data were fitted by Di Francesco et al. (in preparation) using the model

$$F_{\nu} = \frac{MB_{\nu}(T)\kappa_{\nu}}{D^2}, \quad (1)$$

where F_{ν} is the measured flux density, $B_{\nu}(T)$ is the Planck function, M is the source mass, D is the distance to the source and the Beckwith et al. (1990) parametrization of dust opacity, $\kappa_{\nu} = 0.1(\nu/10^{12}\text{Hz})^{\beta} \text{cm}^2 \text{g}^{-1}$ is used assuming a dust emissivity index $\beta = 2.0$. We use this model for dust opacity and this value of dust emissivity index throughout the rest of this work, in order to combine the *Herschel* data with our own in a self-consistent manner. This model of dust properties, adopted by the *Herschel* GBS, is described in detail by e.g. Könyves et al. (2015).

We note that combined SCUBA-2 and *Herschel* observations have demonstrated variations in β towards star-forming regions in the range $\beta = 1.6$ –2.0 (Sadavoy et al. 2013) and $\beta = 1.0$ –2.7 (Chen et al. 2016), with lower values of β typically observed towards protostellar cores. Sadavoy et al. (2013) found $\beta \approx 2.0$ towards filaments and moderately dense material, suggesting that $\beta = 2.0$ is a representative value for the starless cores in our sample, but may be less appropriate for the protostellar sources that we observe.

The SED fitting process is described in detail by Könyves et al. (2015). It must be emphasized that the only quantity derived from the *Herschel* data that we use is the source temperature. We discuss our own determinations of source masses – from their SCUBA-2 850 μm flux densities – below. All of our sources were observed as part of the *Herschel* GBS. However, the sources on the western edge of L1152 are on the very edge of the *Herschel* field, and hence their temperatures may be less reliable than those in other parts of the region. Temperatures of cores without embedded sources are typically in the range 9–15 K, except in the NGC 7023 (L1174) region, where temperatures of up to ~ 50 K are measured.

Source masses were determined using the Hildebrand (1983) formulation

$$M = \frac{F_{\nu}^{\text{total}}(850\mu\text{m})D^2}{\kappa_{\nu(850\mu\text{m})}B_{\nu(850\mu\text{m})}(T)}, \quad (2)$$

where $F_{\nu}(850\mu\text{m})$ is the best-fitting model flux density at 850 μm , D is the source distance as listed in Table 1, $B_{\nu(850\mu\text{m})}(T)$ is the Planck function and $\kappa_{\nu(850\mu\text{m})}$ is the dust mass opacity as parametrized by Beckwith et al. (1990), where β is again taken to be 2.0. Note that equation (2) is functionally identical to equation (1). However, we determine the masses of our sources using our model SCUBA-2 850- μm flux densities and *Herschel*-determined temperatures, whereas equation (1) was used to determine best-fitting source temperatures by fitting the flux densities measured in four *Herschel* wavebands to each pixel in the *Herschel* observations. We use the mean fitted temperature in the pixels enclosed by the source apertures shown on Fig. 6. Detection of a SCUBA-2 source does not necessarily mean that there is a *Herschel* source at the same position.

Table 3. Sources identified in SCUBA-2 850- μm emission by CSAR and characterized using multiple-Gaussian fitting in the Cepheus Flare region. FWHMs are as measured, without deconvolution. Position angles are measured east of north, and listed for elliptical sources only. ‘Model’ peak and total flux density values are the results of the multiple-Gaussian fitting process, and are used in subsequent mass calculations. ‘Photometry’ peak and total flux density values are determined from aperture photometry, using the source sizes shown in Fig. 6, and hence flux density will be double-counted in some of these measurements. ‘Photometry’ measurements are used in the subsequent calculation of flux-ratio-determined temperature. 450- μm signal to noise is measured on peak value. Sources marked with ‘*’ overlap significantly with at least one other source, listed in the final column. See text for details.

Source index	RA (J2000)	Dec (J2000)	FWHM (arcsec)	Angle ($^{\circ}$)	Model				Photometry		450 μm S/N	Type	Region	Overlaps
					$F_{\nu(850)}^{\text{peak}}$ (mJy/sqa)	$F_{\nu(850)}^{\text{total}}$ (Jy)	$F_{\nu(850)}^{\text{peak}}$ (mJy/sqa)	$F_{\nu(450)}^{\text{peak}}$ (mJy/sqa)	$F_{\nu(850)}^{\text{total}}$ (Jy)	$F_{\nu(450)}^{\text{total}}$ (Jy)				
1	20:39:05.28	68:02:20.40	21.6 \times 24.0	100.1	3.07	1.80	5.28	45.4	2.31	11.42	20.6	P	47/58	–
2*	20:35:45.11	67:53:02.40	21.6 \times 26.4	5.4	0.83	0.54	1.23	10.8	0.98	5.72	4.9	P	47/58	3
3*	20:35:41.76	67:52:48.00	26.4 \times 26.4	–	0.78	0.62	1.17	10.8	1.14	5.88	4.9	C	47/58	2
4	20:35:54.72	67:54:10.80	57.8 \times 26.4	152.0	0.39	0.68	0.49	8.7	0.91	4.49	4.0	C	47/58	–
5*	20:36:18.96	67:56:42.00	21.6 \times 21.6	–	0.10	0.05	0.40	7.7	0.08	0.87	3.5	P	47/58	20
6	20:45:53.28	67:57:39.60	23.4 \times 21.6	170.2	1.66	0.95	2.87	24.4	1.27	7.87	11.1	P	47/58	–
7*	20:44:48.48	67:43:12.00	26.4 \times 26.4	–	0.16	0.13	0.35	8.0	0.42	2.49	3.6	C	47/58	8
8*	20:44:51.60	67:43:40.80	37.3 \times 26.4	125.0	0.14	0.15	0.35	8.0	0.51	2.73	3.6	C	47/58	7,10
9*	20:44:47.52	67:44:24.00	23.1 \times 26.4	43.0	0.12	0.08	0.31	7.4	0.26	0.79	3.4	C	47/58	10
10*	20:44:50.88	67:44:13.20	26.4 \times 26.4	–	0.16	0.12	0.32	7.4	0.35	1.40	3.4	C	47/58	8,9
11*	20:36:10.80	67:57:14.40	21.6 \times 21.6	–	0.16	0.08	0.30	6.2	0.11	0.52	2.8	P	47/58	20
12*	20:43:24.48	67:53:09.60	26.4 \times 25.7	170.0	0.07	0.06	0.32	7.2	0.24	0.77	3.3	C	47/58	18
13*	20:43:10.56	67:51:00.00	26.4 \times 24.3	10.0	0.10	0.07	0.27	7.3	0.23	0.39	3.3	C	47/58	14
14*	20:43:18.24	67:50:56.40	21.6 \times 26.4	37.5	0.10	0.06	0.27	7.9	0.2	1.25	3.6	C	47/58	13
15	20:43:49.20	67:51:00.00	21.6 \times 26.4	173.0	0.09	0.06	0.27	7.6	0.16	0.75	3.4	C	47/58	–
16*	20:38:06.96	67:55:30.00	26.4 \times 21.6	80.0	0.06	0.04	0.26	6.1	0.19	0.28	2.8	C	47/58	19
17*	20:43:25.68	67:52:22.80	21.6 \times 21.6	177.0	0.08	0.04	0.30	8.2	0.17	0.94	3.7	C	47/58	18
18*	20:43:29.76	67:52:55.20	66.4 \times 29.8	121.7	0.13	0.30	0.32	8.2	0.68	3.36	3.7	C	47/58	12,17
19*	20:38:04.57	67:55:51.60	21.6 \times 26.4	0.0	0.04	0.02	0.26	6.3	0.19	0.44	2.9	C	47/58	16
20*	20:36:05.76	67:56:45.60	72.2 \times 26.4	19.6	0.20	0.43	0.29	7.6	0.48	2.35	3.4	C	47/58	5,11
21	21:01:40.81	68:12:03.60	26.4 \times 23.7	10.0	1.33	0.94	2.03	18.4	1.94	16.91	18.4	C	L1174	–
22*	21:00:19.68	68:13:22.80	22.8 \times 26.4	100.0	0.76	0.52	1.98	14.6	1.47	8.07	14.6	P	L1174	25,26
23*	21:01:28.80	68:10:33.60	29.8 \times 26.4	147.2	1.61	1.43	1.77	19.0	2.14	18.08	19.0	P	L1174	24
24*	21:01:30.96	68:11:20.40	31.2 \times 25	112.8	1.47	1.30	1.61	17.3	2.20	19.26	17.3	P	L1174	23
25*	21:00:23.04	68:13:12.00	26.4 \times 26.4	–	1.17	0.93	1.98	14.6	1.72	10.83	14.6	P	L1174	22,26
26*	21:00:17.28	68:12:46.80	26.4 \times 26.4	–	0.99	0.78	1.12	7.2	1.36	7.33	7.2	C	L1174	22,25
27*	21:02:13.92	68:09:14.40	57.2 \times 26.4	127.9	0.62	1.05	0.70	7.3	1.48	11.64	7.3	C	L1174	28
28*	21:02:11.04	68:09:54.00	21.6 \times 26.4	10.0	0.20	0.13	0.47	5.6	0.46	3.58	5.6	C	L1174	27
29*	21:01:28.32	68:08:20.40	26.4 \times 24.6	71.8	0.24	0.18	0.41	4.8	0.44	2.71	4.8	C	L1174	35,41
30*	21:03:20.16	68:11:31.20	26.4 \times 26.4	–	0.12	0.10	0.52	3.7	0.50	1.96	3.7	C	L1174	31
31*	21:03:15.12	68:11:16.80	30.8 \times 26.4	115.8	0.17	0.15	0.52	3.2	0.55	1.49	3.2	C	L1174	30
32	20:59:22.56	68:14:49.20	22.3 \times 21.6	10.0	0.18	0.10	0.39	3.2	0.15	0.22	3.2	P	L1174	–
33*	21:02:00.72	68:07:12.00	26.4 \times 21.6	172.8	0.10	0.06	0.36	3.4	0.29	0.68	3.4	C	L1174	39
34	21:01:31.20	68:07:19.20	42.3 \times 21.9	24.4	0.26	0.27	0.48	5.2	0.73	4.69	5.2	C	L1174	–
35*	21:01:34.32	68:08:16.80	21.6 \times 26.4	0.0	0.07	0.04	0.40	4.2	0.34	1.69	4.2	P	L1174	29,41
36	21:01:31.20	68:05:38.40	24 \times 24	–	0.22	0.15	0.35	3.8	0.25	1.44	3.8	C	L1174	–
37	21:02:48.72	68:11:45.60	24.9 \times 26.4	10.0	0.10	0.08	0.34	2.7	0.30	0.81	2.7	C	L1174	–
38	21:00:28.56	68:07:08.40	45.5 \times 26.4	40.3	0.25	0.34	0.43	4.7	0.78	3.55	4.7	C	L1174	–
39*	21:01:56.39	68:06:39.60	46.1 \times 26.4	136.3	0.22	0.30	0.40	3.1	0.74	2.37	3.1	C	L1174	33
40	21:02:00.96	68:13:01.20	73.6 \times 46	122.7	0.22	0.86	0.41	4.3	1.61	6.08	4.3	C	L1174	–
41*	21:01:32.64	68:08:38.40	26.4 \times 21.6	161.8	0.23	0.15	0.41	4.2	0.33	1.96	4.2	P	L1174	29,35
42	21:00:24.25	68:14:06.00	26.4 \times 21.6	95.6	0.10	0.07	0.37	3.9	0.26	0.38	3.9	C	L1174	–
43	21:00:37.92	68:06:18.00	21.6 \times 26.4	1.4	0.14	0.09	0.31	4.1	0.27	0.74	4.1	C	L1174	–
44	21:00:23.52	68:08:13.20	38 \times 26.4	148.7	0.18	0.20	0.40	4.1	0.57	2.61	4.1	C	L1174	–
45	21:02:09.12	68:07:08.40	25.6 \times 26.4	170.0	0.15	0.11	0.39	3.6	0.36	1.46	3.6	C	L1174	–
46	21:02:39.60	68:11:24.00	27 \times 26.4	175.2	0.17	0.14	0.38	3.5	0.41	0.70	3.5	C	L1174	–
47*	21:02:20.64	67:54:21.60	23.5 \times 26.4	177.6	0.41	0.29	0.72	8.4	0.70	3.72	3.8	P	L1172	48
48*	21:02:26.40	67:54:14.40	26.4 \times 24	170.0	0.38	0.28	0.65	7.1	0.65	3.27	3.2	P	L1172	47
49	21:02:13.20	67:54:03.60	22.3 \times 26.4	80.0	0.10	0.07	0.36	5.9	0.35	1.49	2.7	C	L1172	–
50	21:02:20.64	67:45:36.00	21.6 \times 26.4	170.0	0.09	0.06	0.27	8.0	0.18	1.14	3.7	C	L1172	–
51*	21:01:51.60	67:44:06.00	23.8 \times 26.4	170.0	0.08	0.05	0.24	7.5	0.20	0.72	3.4	C	L1172	53
52	21:02:15.84	67:51:10.80	29.5 \times 21.6	53.4	0.10	0.07	0.26	6.1	0.21	1.00	2.8	C	L1172	–
53*	21:01:52.08	67:43:40.80	26.4 \times 25.7	10.0	0.08	0.06	0.24	7.5	0.21	0.61	3.4	C	L1172	51
54	21:02:29.76	67:53:24.00	21.6 \times 26.4	170.0	0.07	0.05	0.24	5.6	0.16	0.31	2.5	C	L1172	–
55	21:02:41.28	67:54:10.80	33 \times 26.4	84.7	0.17	0.17	0.26	5.3	0.19	0.55	2.4	C	L1172	–
56*	22:38:47.04	75:11:31.20	33.4 \times 23.8	9.9	3.54	3.19	4.22	34.9	3.67	23.33	19.7	P	L1251	58,59

Table 3 – *continued*

Source index	RA (J2000)	Dec (J2000)	FWHM (arcsec)	Angle (°)	Model		Photometry				450 μ m S/N	Type	Region	Overlaps
					$F_{\nu(850)}^{\text{peak}}$ (mJy/sqa)	$F_{\nu(850)}^{\text{total}}$ (Jy)	$F_{\nu(850)}^{\text{peak}}$ (mJy/sqa)	$F_{\nu(450)}^{\text{peak}}$ (mJy/sqa)	$F_{\nu(850)}^{\text{total}}$ (Jy)	$F_{\nu(450)}^{\text{total}}$ (Jy)				
57*	22:31:04.32	75:13:37.20	53.7 × 26.4	27.2	0.92	1.48	1.02	7.2	1.57	7.71	7.7	P	L1251	62,63
58*	22:39:04.56	75:12:00.00	26.4 × 26.4	–	0.76	0.60	0.70	8.4	0.98	3.47	4.8	C	L1251	59
59*	22:38:56.16	75:11:42.00	26.4 × 24.3	177.6	0.65	0.47	3.50	28.9	1.08	5.37	16.3	C	L1251	56,58
60*	22:39:38.40	75:12:03.60	53.3 × 39.0	93.2	0.80	1.89	0.92	6.6	2.51	7.00	3.7	C	L1251	64
61	22:35:22.56	75:17:06.00	27.1 × 25.3	80.0	1.89	1.47	2.81	23.0	2.18	14.62	13.0	P	L1251	–
62*	22:31:12.48	75:12:57.60	26.4 × 26.4	–	0.34	0.27	0.69	4.5	0.66	3.50	4.8	C	L1251	57,63
63*	22:31:22.08	75:12:28.80	65.7 × 26.4	19.9	0.41	0.81	0.45	4.5	1.03	6.41	4.8	C	L1251	57,62
64*	22:39:30.00	75:10:58.80	28.7 × 24.4	158.9	0.58	0.46	0.73	6.4	0.80	2.61	3.6	C	L1251	60
65	22:39:16.08	75:09:43.20	32.3 × 25.6	49.5	0.19	0.18	0.35	5.1	0.31	0.51	2.9	C	L1251	–
66*	22:28:15.36	75:14:38.40	40.6 × 23.4	146.6	0.42	0.45	0.45	3.8	0.58	2.40	4.1	C	L1251	67
67*	22:28:24.72	75:14:56.40	26.4 × 21.6	170.0	0.15	0.10	0.39	3.5	0.30	1.36	3.8	C	L1251	66
68*	22:35:52.32	75:18:57.60	41.0 × 24.5	109.8	0.39	0.45	0.51	8.3	0.66	3.78	4.7	C	L1251	86
69	22:34:39.84	75:17:49.20	21.6 × 21.6	–	0.14	0.08	0.27	4.2	0.08	0.74	2.4	P	L1251	–
70*	22:35:34.08	75:21:18.00	21.6 × 26.4	172.3	0.11	0.07	0.31	5.0	0.24	0.78	2.8	C	L1251	87,88
71	22:29:41.52	75:13:30.00	42.6 × 32.1	15.0	0.66	1.02	0.75	5.7	1.32	7.10	6.1	C	L1251	–
72	22:39:13.20	75:10:44.40	30.4 × 26.4	163.5	0.18	0.17	0.36	6.4	0.33	1.08	3.6	C	L1251	–
73	22:36:40.80	75:08:31.20	26.4 × 23.8	10.0	0.08	0.06	0.26	7.0	0.21	1.07	4.0	C	L1251	–
74	22:35:04.80	75:13:01.20	27.7 × 26.4	29.7	0.14	0.12	0.27	5.2	0.29	2.19	2.9	C	L1251	–
75*	22:39:24.24	75:12:39.60	26.4 × 26.4	–	0.15	0.12	0.32	6.1	0.28	-0.17	3.4	C	L1251	–
76	22:34:10.80	75:18:10.80	21.6 × 21.6	–	0.14	0.07	0.26	4.7	0.12	0.31	2.7	P	L1251	93
77*	22:35:59.76	75:07:48.00	26.4 × 26.4	–	0.09	0.07	0.25	4.9	0.27	0.30	2.8	C	L1251	–
78*	22:27:31.44	75:11:24.00	36.1 × 25.8	143.7	0.16	0.17	0.31	2.9	0.36	1.62	3.1	C	L1251	79
79*	22:36:07.20	75:07:58.80	25.7 × 26.4	0.0	0.08	0.06	0.24	5.5	0.23	0.06	3.1	C	L1251	83,85,89
80	22:38:21.60	75:13:01.20	31.8 × 21.6	39.1	0.16	0.13	0.26	6.0	0.18	1.32	3.4	C	L1251	77
81	22:37:00.00	75:15:21.60	26.4 × 21.6	96.8	0.09	0.06	0.28	4.9	0.21	0.24	2.8	C	L1251	–
82	22:30:30.72	75:14:13.20	29.2 × 25.8	146.8	1.04	0.88	1.41	12.8	1.13	5.79	13.7	P	L1251	–
83*	22:27:37.69	75:12:14.40	23.1 × 21.6	10.0	0.10	0.06	0.31	2.9	0.18	0.90	3.1	C	L1251	–
84	22:35:20.64	75:18:57.60	27.3 × 21.6	65.4	0.32	0.21	0.42	6.7	0.33	2.29	3.8	P	L1251	78,85,89
85*	22:27:31.68	75:12:07.20	26.4 × 26.4	8.4	0.06	0.05	0.31	2.9	0.25	1.07	3.1	C	L1251	–
86*	22:35:42.00	75:18:54.00	26.4 × 24.9	10.0	0.14	0.10	0.35	6.6	0.25	1.69	3.8	C	L1251	78,83,89
87*	22:35:31.44	75:21:54.00	22.1 × 26.4	10.0	0.09	0.06	0.28	4.7	0.25	0.92	2.7	C	L1251	68
88*	22:35:38.88	75:21:25.20	47.6 × 21.6	64.1	0.08	0.09	0.31	5.3	0.42	1.63	3.0	C	L1251	70,88
89*	22:27:38.87	75:11:45.60	33.3 × 26.4	3.3	0.10	0.10	0.31	3.3	0.31	1.69	3.5	C	L1251	70,87
90	22:37:44.16	75:09:43.20	35.8 × 26.4	129.3	0.16	0.17	0.27	5.5	0.28	0.20	3.1	C	L1251	78,83,85
91	22:29:59.76	75:13:55.20	38.1 × 26.4	73.9	0.20	0.23	0.35	4.3	0.49	3.50	4.6	P	L1251	–
92	22:38:44.40	75:14:02.40	26.4 × 21.6	18.2	0.13	0.08	0.26	6.7	0.18	0.73	3.8	C	L1251	–
93*	22:39:17.52	75:13:44.40	71.2 × 27.5	77.1	0.42	0.94	0.50	7.5	0.85	0.66	4.2	C	L1251	75
94*	22:37:08.88	75:08:49.20	26.4 × 26.4	–	0.10	0.08	0.28	5.5	0.30	1.13	3.1	C	L1251	97
95	22:37:34.57	75:11:34.80	65.6 × 38.3	134.0	0.38	1.07	0.46	6.4	1.50	5.33	3.6	C	L1251	–
96	22:36:18.72	75:22:15.60	50.3 × 27.2	130.8	0.22	0.34	0.35	5.3	0.56	0.30	3.0	C	L1251	–
97*	22:37:00.71	75:08:42.00	33.7 × 26.4	114.4	0.11	0.11	0.27	5.7	0.37	1.76	3.2	C	L1251	94
98	20:58:02.16	77:33:18.00	33.7 × 31.2	126.0	0.24	0.28	0.34	3.1	0.52	2.00	3.5	C	L1228	–
99	20:57:18.24	77:37:51.60	24.0 × 24.0	–	0.20	0.13	0.27	3.9	0.23	1.32	4.4	C	L1228	–
100*	20:56:41.28	77:41:24.00	42.7 × 30.0	15.8	0.18	0.26	0.30	3.5	0.60	2.16	3.9	C	L1228	104
101	20:55:54.24	77:42:46.80	44.5 × 26.4	18.2	0.19	0.25	0.31	3.2	0.48	1.86	3.6	C	L1228	–
102	20:57:13.68	77:44:06.00	26.4 × 21.6	35.0	0.11	0.07	0.26	3.5	0.20	1.00	3.9	C	L1228	–
103	20:54:49.44	77:32:24.00	26.4 × 21.6	170.0	0.12	0.07	0.23	2.3	0.15	0.36	2.6	C	L1228	–
104*	20:56:42.24	77:40:55.20	26.4 × 26.4	–	0.12	0.09	0.29	3.5	0.32	1.11	3.9	C	L1228	100
105	20:54:59.04	77:50:34.80	21.6 × 21.6	–	0.04	0.02	0.24	3.1	0.14	0.52	3.5	C	L1228	–
106	20:57:39.60	77:43:37.20	62.8 × 40.7	175.2	0.36	1.03	0.53	4.4	1.33	5.85	4.9	C	L1228	–
107	20:57:47.76	77:37:19.20	30.0 × 26.5	147.4	0.17	0.15	0.30	2.5	0.31	0.64	2.8	C	L1228	–
108*	20:56:27.37	77:24:43.20	35.2 × 26.4	118.4	0.07	0.07	0.26	2.8	0.29	0.29	3.1	C	L1228	115
109	20:58:49.68	77:47:16.80	26.2 × 21.6	100.0	0.07	0.04	0.21	2.7	0.19	0.30	3.0	C	L1228	–
110	20:55:11.28	77:33:21.60	21.8 × 21.6	55.0	0.09	0.05	0.24	2.7	0.16	0.28	3.0	C	L1228	–
111	20:57:12.24	77:35:45.60	26.4 × 21.8	100.5	2.20	1.43	3.65	29.0	2.03	9.00	32.2	P	L1228	–
112	20:55:18.48	77:45:46.80	25.1 × 21.6	0.5	0.06	0.04	0.31	3.3	0.21	1.04	3.7	C	L1228	–
113*	20:58:19.92	77:42:36.00	37.4 × 25.9	35.0	0.11	0.12	0.25	3.2	0.32	1.42	3.5	C	L1228	116
114	20:54:49.44	77:43:33.60	21.6 × 24.6	173.7	0.08	0.05	0.22	2.7	0.13	0.33	3.0	C	L1228	–
115*	20:56:18.00	77:24:57.60	34.6 × 26.4	142.4	0.16	0.17	0.28	2.6	0.38	0.51	2.9	C	L1228	108
116*	20:58:30.24	77:42:43.20	26.4 × 21.6	3.0	0.11	0.07	0.25	3.2	0.19	1.01	3.5	C	L1228	113
117	20:57:17.05	77:33:21.60	31.9 × 24.1	129.7	0.13	0.11	0.26	2.6	0.28	0.75	2.9	C	L1228	–

Table 4. The protostellar sources in our catalogue, with their identification and evolutionary class from K09, and alternative identifications. With the exception of L1157-mm and stars with an identification of the form XX Cep, alternative identifications are given in the following order of preference: *IRAS* Point or Faint Source Catalogs (*IRAS* – Beichman et al. 1988; Moshir, Kopman & Conrow 1992), 2MASS All-Sky Catalog of Point Sources (2MASS – Skrutskie et al. 2016), *Spitzer* Gould Belt Survey (SSTgbs – K09). For the L1251B cluster, designations from Lee et al. (2006) are also given. K09 identifications given in brackets indicate an offset between our source central coordinates and the coordinates of the K09 source greater than the JCMT 850- μm beam size, but less than the radius of the source as listed in Table 3.

Source ID	K09 ID	K09 Class	Alternative ID
1	134	I	L1157-mm
2	1	I	<i>IRAS</i> 20353+6742
5	3	II	<i>IRAS</i> 20359+6475
6	135	I	PV Cep
11	2	II	2MASS J2036+1165+6757093
22	17	I	SSTgbs J2100207+6813169
22	(100)	F	SSTgbs J2100224+6813042
23	27	II	2MASS J21012637+6810385
23	137	II	SSTgbs J2101271+6810380
24	34	I	2MASS J21013280+6811204
25	18	I	SSTgbs J2100221+6812585
25	(100)	F	SSTgbs J2100224+6813042
32	15	II	FT Cep
35	(104)	F	PW Cep
41	(104)	F	PW Cep
47	49	I	<i>IRAS</i> 21017+6742
48	50	F	SSTgbs J2102273+6754186
48	(53)	II	2MASS J21022993+6754083
56	89	I	2MASS J22384282+7511369; L1251B IRS 4
56	90	I	SSTgbs J2238469+7511337; L1251B IRS 1
56	92	I	2MASS J22385287+7511235; L1251B IRS 2
56	107	III	<i>IRAS</i> 22376+7455; L1251B IRS 3
56	108	III	SSTgbs J2238440+7511266; L1251B IRS 5
56	109	II	2MASS J22384807+7511488; L1251B IRS 6
57	68	II	SSTgbs J2231056+7513372
61	143	I	<i>IRAS</i> 22343+7501
69	69	F	2MASS J22344051+7517444
76	142	F	<i>IRAS</i> 22331+7502
82	67	I	SSTgbs J2230318+7514094
84	(76)	II	2MASS J22351668+7518471
91	66	I	<i>IRAS</i> 22290+7458
111	9	F	<i>IRAS</i> 20582+7724; L1228

Mean source volume densities were determined using the equation

$$n(\text{H}_2) = \frac{M}{\mu m_{\text{H}}} \frac{1}{\frac{4}{3}\pi R^3}, \quad (3)$$

where R is the equivalent deconvolved mean FWHM of the source. The equivalent deconvolved mean FWHM was taken to be the geometric mean of the best-fitting major and minor FWHMs, with the JCMT 850 μm effective beam FWHM (14.1 arcsec) subtracted in quadrature. The mean molecular weight μ was taken to be 2.86, assuming that the gas is ~ 70 per cent H_2 by mass (Kirk et al. 2013). We give densities in terms of H_2 rather than of total density of particles as the density ranges traced by different molecular species are typically expressed in terms of H_2 number density (see e.g. Di Francesco et al. 2007). The range of densities traced by isotopologues of CO is relevant to our determination of core stability in Section 7, and so we express particle number density in terms of $n(\text{H}_2)$ throughout this work. Assuming a typical mean particle mass of 2.3 amu, our H_2 number densities can be converted to total gas particle number densities by multiplication by a factor of 1.24.

Mean source column densities were determined using the equation

$$N(\text{H}_2) = \frac{M}{\mu m_{\text{H}}} \frac{1}{\pi R^2}, \quad (4)$$

with symbols defined as above.

The derived properties of our sources: temperature, mass, column density, volume density and deconvolved FWHM, are listed in Table 5. For the protostellar sources in our catalogue, the temperatures, and hence the masses, determined from the dust emission are those of the protostellar envelopes, and not of the protostars themselves (see e.g. Pattle et al. 2015). The modified blackbody model used to fit temperatures is applicable only to envelope-dominated sources; the temperatures and masses determined for the Class II and III protostars in our catalogue (11 sources, listed in Table 4) may not be representative.

3.4 SCUBA-2-derived temperatures and masses

We derived temperatures for our sources from the ratio of the total SCUBA-2 450- μm and 850- μm flux densities measured along the

Table 5. Properties of the sources. See text for discussion.

Source index	T_{Herschel} (K)	$T_{\text{SCUBA-2}}$ (K)	$M(T_{\text{Herschel}})$ (M_{\odot})	$M(T_{\text{SCUBA-2}})$ (M_{\odot})	$N(\text{H}_2)$ ($\times 10^{21} \text{ cm}^{-2}$)	$n(\text{H}_2)$ ($\times 10^4 \text{ cm}^{-3}$)	Deconv. FWHM (pc)
1	14.8 ± 0.2	10.7 ± 1.7	2.43 ± 0.07	4.36 ± 1.35	42.66 ± 1.30	36.81 ± 1.12	0.028
2	12.4 ± 0.3	–	0.99 ± 0.05	–	14.91 ± 0.80	11.93 ± 0.64	0.030
3	11.5 ± 0.2	–	1.31 ± 0.06	–	14.77 ± 0.66	10.21 ± 0.45	0.035
4	10.6 ± 0.1	–	1.70 ± 0.07	–	7.17 ± 0.31	3.03 ± 0.13	0.057
5	13.1 ± 0.2	–	0.08 ± 0.01	–	1.78 ± 0.30	1.68 ± 0.28	0.026
6	16.2 ± 0.2	13.5 ± 2.8	1.11 ± 0.03	1.51 ± 0.55	20.19 ± 0.58	17.75 ± 0.51	0.028
7	11.7 ± 0.0	–	0.26 ± 0.02	–	2.93 ± 0.19	2.02 ± 0.13	0.035
8	11.5 ± 0.1	–	0.32 ± 0.02	–	2.29 ± 0.16	1.26 ± 0.09	0.044
9	11.7 ± 0.1	–	0.17 ± 0.02	–	2.31 ± 0.25	1.76 ± 0.19	0.032
10	11.5 ± 0.1	–	0.26 ± 0.02	–	2.93 ± 0.21	2.02 ± 0.14	0.035
11	12.3 ± 0.0	–	0.15 ± 0.01	–	3.21 ± 0.30	3.03 ± 0.28	0.026
12	11.5 ± 0.1	–	0.12 ± 0.02	–	1.36 ± 0.19	0.96 ± 0.14	0.035
13	11.5 ± 0.0	–	0.15 ± 0.02	–	1.94 ± 0.21	1.42 ± 0.15	0.033
14	11.6 ± 0.0	–	0.13 ± 0.02	–	1.95 ± 0.24	1.56 ± 0.19	0.030
15	12.4 ± 0.0	–	0.11 ± 0.01	–	1.64 ± 0.20	1.32 ± 0.16	0.030
16	13.3 ± 0.1	–	0.06 ± 0.01	–	0.94 ± 0.18	0.75 ± 0.15	0.030
17	11.5 ± 0.1	–	0.09 ± 0.02	–	1.79 ± 0.33	1.68 ± 0.31	0.026
18	11.3 ± 0.1	–	0.65 ± 0.03	–	2.05 ± 0.11	0.75 ± 0.04	0.066
19	13.2 ± 0.0	–	0.04 ± 0.01	–	0.57 ± 0.18	0.45 ± 0.14	0.030
20	11.5 ± 0.1	–	0.91 ± 0.04	–	2.99 ± 0.14	1.12 ± 0.05	0.065
21	16.7 ± 0.5	23.3 ± 9.6	0.82 ± 0.05	0.50 ± 0.29	13.68 ± 0.86	11.52 ± 0.72	0.029
22	13.1 ± 0.1	11.9 ± 2.1	0.68 ± 0.02	0.81 ± 0.27	12.00 ± 0.39	10.40 ± 0.34	0.028
23	27.6 ± 1.5	21.8 ± 8.2	0.60 ± 0.05	0.83 ± 0.45	7.31 ± 0.58	5.25 ± 0.42	0.034
24	19.2 ± 1.1	23.5 ± 9.8	0.91 ± 0.08	0.68 ± 0.40	11.17 ± 1.04	8.07 ± 0.75	0.034
25	13.2 ± 0.1	13.7 ± 2.9	1.19 ± 0.03	1.12 ± 0.41	17.06 ± 0.41	13.30 ± 0.32	0.031
26	11.9 ± 0.3	11.6 ± 2.0	1.21 ± 0.07	1.27 ± 0.42	17.40 ± 0.94	13.57 ± 0.73	0.031
27	20.3 ± 1.4	18.9 ± 6.0	0.68 ± 0.08	0.75 ± 0.36	3.70 ± 0.42	1.78 ± 0.20	0.051
28	21.8 ± 0.6	18.7 ± 6.0	0.08 ± 0.01	0.09 ± 0.05	1.45 ± 0.17	1.31 ± 0.15	0.027
29	26.9 ± 1.1	–	0.08 ± 0.01	–	1.22 ± 0.14	1.00 ± 0.11	0.030
30	15.5 ± 0.2	–	0.09 ± 0.01	–	1.36 ± 0.17	1.06 ± 0.13	0.031
31	15.1 ± 0.2	–	0.16 ± 0.02	–	1.81 ± 0.18	1.27 ± 0.13	0.035
32	14.1 ± 0.1	–	0.11 ± 0.01	–	2.87 ± 0.33	2.97 ± 0.34	0.024
33	25.3 ± 0.3	–	0.03 ± 0.01	–	0.57 ± 0.10	0.51 ± 0.09	0.027
34	19.5 ± 0.4	14.0 ± 3.1	0.18 ± 0.02	0.32 ± 0.12	1.81 ± 0.15	1.17 ± 0.10	0.038
35	28.4 ± 1.2	–	0.02 ± 0.01	–	0.34 ± 0.10	0.31 ± 0.09	0.027
36	18.2 ± 0.4	–	0.11 ± 0.01	–	2.09 ± 0.22	1.87 ± 0.20	0.027
37	15.5 ± 0.0	–	0.07 ± 0.01	–	1.16 ± 0.16	0.94 ± 0.13	0.030
38	19.5 ± 0.7	–	0.23 ± 0.02	–	1.67 ± 0.16	0.92 ± 0.09	0.044
39	24.2 ± 0.6	–	0.15 ± 0.01	–	1.07 ± 0.09	0.58 ± 0.05	0.045
40	18.8 ± 1.3	–	0.62 ± 0.08	–	1.38 ± 0.18	0.43 ± 0.06	0.079
41	31.0 ± 1.1	–	0.05 ± 0.01	–	1.01 ± 0.12	0.91 ± 0.11	0.027
42	13.2 ± 0.1	–	0.08 ± 0.01	–	1.61 ± 0.27	1.45 ± 0.25	0.027
43	21.2 ± 0.3	–	0.05 ± 0.01	–	1.02 ± 0.14	0.92 ± 0.13	0.027
44	18.9 ± 0.5	–	0.15 ± 0.01	–	1.30 ± 0.13	0.80 ± 0.08	0.040
45	23.6 ± 0.5	–	0.06 ± 0.01	–	0.88 ± 0.11	0.70 ± 0.09	0.031
46	16.3 ± 0.3	–	0.12 ± 0.01	–	1.69 ± 0.17	1.29 ± 0.13	0.032
47	12.0 ± 0.0	–	0.45 ± 0.01	–	7.56 ± 0.21	6.40 ± 0.18	0.029
48	12.1 ± 0.1	–	0.42 ± 0.02	–	6.81 ± 0.28	5.68 ± 0.23	0.029
49	11.8 ± 0.1	–	0.10 ± 0.01	–	1.91 ± 0.23	1.69 ± 0.20	0.028
50	12.9 ± 0.0	–	0.07 ± 0.01	–	1.42 ± 0.19	1.29 ± 0.17	0.027
51	12.7 ± 0.1	–	0.08 ± 0.01	–	1.27 ± 0.18	1.06 ± 0.15	0.029
52	12.5 ± 0.0	–	0.10 ± 0.01	–	1.61 ± 0.19	1.34 ± 0.15	0.029
53	12.6 ± 0.1	–	0.08 ± 0.01	–	1.20 ± 0.16	0.95 ± 0.13	0.031
54	12.3 ± 0.1	–	0.07 ± 0.01	–	1.33 ± 0.22	1.20 ± 0.20	0.027
55	12.5 ± 0.1	–	0.24 ± 0.01	–	2.58 ± 0.15	1.73 ± 0.10	0.036
56	14.9 ± 0.3	13.9 ± 3.0	3.61 ± 0.14	4.07 ± 1.51	39.85 ± 1.52	27.26 ± 1.04	0.036
57	11.0 ± 0.1	10.7 ± 1.6	2.89 ± 0.09	3.09 ± 0.95	15.59 ± 0.48	7.46 ± 0.23	0.051
58	11.1 ± 0.1	–	1.16 ± 0.03	–	15.37 ± 0.42	11.51 ± 0.31	0.032
59	13.1 ± 0.5	10.8 ± 1.7	0.66 ± 0.06	0.96 ± 0.30	9.83 ± 0.84	7.81 ± 0.67	0.031
60	10.3 ± 0.2	–	4.22 ± 0.21	–	14.75 ± 0.72	5.68 ± 0.28	0.063
61	18.4 ± 0.1	14.9 ± 3.5	1.19 ± 0.02	1.67 ± 0.66	16.04 ± 0.27	12.15 ± 0.21	0.032
62	10.8 ± 0.1	–	0.55 ± 0.03	–	7.30 ± 0.33	5.46 ± 0.25	0.032
63	11.0 ± 0.2	–	1.57 ± 0.07	–	6.72 ± 0.29	2.86 ± 0.12	0.057

Table 5 – continued

Source index	T_{Herschel} (K)	$T_{\text{SCUBA-2}}$ (K)	$M(T_{\text{Herschel}})$ (M_{\odot})	$M(T_{\text{SCUBA-2}})$ (M_{\odot})	$N(\text{H}_2)$ ($\times 10^{21} \text{ cm}^{-2}$)	$n(\text{H}_2)$ ($\times 10^4 \text{ cm}^{-3}$)	Deconv. FWHM (pc)
64	10.2 ± 0.1	–	1.05 ± 0.04	–	13.82 ± 0.46	10.31 ± 0.35	0.033
65	11.4 ± 0.1	–	0.33 ± 0.02	–	3.45 ± 0.25	2.30 ± 0.16	0.036
66	11.1 ± 0.1	–	0.87 ± 0.03	–	7.58 ± 0.26	4.62 ± 0.16	0.040
67	11.4 ± 0.1	–	0.18 ± 0.01	–	3.18 ± 0.26	2.76 ± 0.23	0.028
68	11.4 ± 0.1	–	0.81 ± 0.03	–	6.63 ± 0.23	3.91 ± 0.14	0.041
69	13.0 ± 0.1	–	0.11 ± 0.01	–	2.64 ± 0.27	2.70 ± 0.28	0.024
70	12.4 ± 0.1	–	0.11 ± 0.01	–	2.03 ± 0.22	1.76 ± 0.19	0.028
71	10.7 ± 0.1	11.6 ± 2.0	2.09 ± 0.07	1.80 ± 0.59	11.78 ± 0.39	5.75 ± 0.19	0.050
72	11.1 ± 0.1	–	0.32 ± 0.02	–	3.46 ± 0.24	2.36 ± 0.16	0.036
73	12.5 ± 0.0	–	0.09 ± 0.01	–	1.35 ± 0.17	1.09 ± 0.14	0.030
74	12.0 ± 0.0	–	0.19 ± 0.01	–	2.36 ± 0.16	1.71 ± 0.12	0.034
75	11.1 ± 0.0	–	0.23 ± 0.02	–	3.09 ± 0.20	2.31 ± 0.15	0.032
76	14.4 ± 0.1	–	0.09 ± 0.01	–	2.12 ± 0.23	2.17 ± 0.24	0.024
77	13.0 ± 0.0	–	0.10 ± 0.01	–	1.30 ± 0.14	0.97 ± 0.10	0.032
78	12.1 ± 0.1	–	0.28 ± 0.02	–	2.50 ± 0.18	1.54 ± 0.11	0.039
79	12.9 ± 0.0	–	0.09 ± 0.01	–	1.18 ± 0.15	0.90 ± 0.11	0.032
80	12.8 ± 0.0	–	0.18 ± 0.01	–	2.47 ± 0.17	1.87 ± 0.13	0.032
81	13.5 ± 0.0	–	0.08 ± 0.01	–	1.38 ± 0.18	1.20 ± 0.16	0.028
82	11.9 ± 0.1	11.1 ± 1.8	1.49 ± 0.05	1.70 ± 0.54	17.66 ± 0.56	12.53 ± 0.39	0.034
83	11.7 ± 0.0	–	0.10 ± 0.01	–	2.15 ± 0.28	2.07 ± 0.27	0.025
84	12.2 ± 0.1	–	0.34 ± 0.02	–	5.74 ± 0.26	4.85 ± 0.22	0.029
85	11.8 ± 0.1	–	0.08 ± 0.01	–	1.06 ± 0.18	0.79 ± 0.13	0.032
86	12.0 ± 0.1	–	0.17 ± 0.01	–	2.41 ± 0.20	1.88 ± 0.16	0.031
87	12.6 ± 0.1	–	0.09 ± 0.01	–	1.48 ± 0.20	1.26 ± 0.17	0.029
88	12.3 ± 0.1	–	0.15 ± 0.02	–	1.16 ± 0.13	0.67 ± 0.08	0.042
89	11.9 ± 0.1	–	0.17 ± 0.02	–	1.66 ± 0.18	1.07 ± 0.12	0.038
90	11.8 ± 0.1	–	0.29 ± 0.02	–	2.57 ± 0.15	1.57 ± 0.09	0.040
91	11.9 ± 0.2	–	0.38 ± 0.03	–	3.10 ± 0.22	1.82 ± 0.13	0.041
92	12.5 ± 0.0	–	0.13 ± 0.01	–	2.24 ± 0.21	1.94 ± 0.18	0.028
93	10.8 ± 0.1	–	1.90 ± 0.07	–	7.13 ± 0.27	2.84 ± 0.11	0.061
94	12.1 ± 0.0	–	0.13 ± 0.01	–	1.66 ± 0.16	1.24 ± 0.12	0.032
95	11.4 ± 0.1	–	1.96 ± 0.06	–	5.57 ± 0.17	1.93 ± 0.06	0.070
96	12.4 ± 0.1	–	0.54 ± 0.02	–	3.02 ± 0.11	1.48 ± 0.05	0.050
97	12.3 ± 0.1	–	0.17 ± 0.01	–	1.65 ± 0.14	1.05 ± 0.09	0.038
98	11.7 ± 0.1	–	0.22 ± 0.01	–	3.79 ± 0.16	3.25 ± 0.14	0.028
99	11.8 ± 0.1	–	0.10 ± 0.01	–	3.85 ± 0.25	4.97 ± 0.32	0.019
100	11.6 ± 0.1	–	0.20 ± 0.01	–	2.76 ± 0.14	2.11 ± 0.11	0.032
101	12.4 ± 0.0	–	0.17 ± 0.01	–	2.58 ± 0.12	2.07 ± 0.10	0.030
102	12.2 ± 0.1	–	0.05 ± 0.01	–	2.09 ± 0.25	2.71 ± 0.32	0.019
103	13.1 ± 0.1	–	0.05 ± 0.01	–	1.88 ± 0.20	2.44 ± 0.26	0.019
104	11.8 ± 0.1	–	0.07 ± 0.01	–	2.12 ± 0.19	2.37 ± 0.21	0.022
105	13.2 ± 0.0	–	0.01 ± 0.00	–	0.66 ± 0.25	1.01 ± 0.38	0.016
106	11.3 ± 0.2	–	0.86 ± 0.04	–	5.39 ± 0.23	2.78 ± 0.12	0.047
107	11.7 ± 0.1	–	0.12 ± 0.01	–	2.91 ± 0.19	2.99 ± 0.19	0.024
108	13.3 ± 0.1	–	0.04 ± 0.01	–	0.89 ± 0.13	0.82 ± 0.12	0.026
109	13.7 ± 0.1	–	0.02 ± 0.00	–	1.00 ± 0.18	1.31 ± 0.23	0.019
110	13.3 ± 0.1	–	0.03 ± 0.00	–	1.70 ± 0.26	2.58 ± 0.39	0.016
111	12.5 ± 0.1	9.8 ± 1.4	0.98 ± 0.02	1.57 ± 0.46	38.47 ± 0.64	49.75 ± 0.83	0.019
112	12.5 ± 0.0	–	0.03 ± 0.00	–	1.17 ± 0.21	1.58 ± 0.29	0.018
113	12.3 ± 0.1	–	0.08 ± 0.01	–	1.57 ± 0.13	1.42 ± 0.12	0.027
114	13.0 ± 0.0	–	0.03 ± 0.00	–	1.32 ± 0.21	1.81 ± 0.28	0.018
115	13.0 ± 0.0	–	0.11 ± 0.01	–	2.20 ± 0.11	2.06 ± 0.11	0.026
116	12.5 ± 0.1	–	0.05 ± 0.01	–	1.90 ± 0.22	2.47 ± 0.28	0.019
117	11.8 ± 0.0	–	0.09 ± 0.01	–	2.22 ± 0.17	2.33 ± 0.17	0.023

line of sight for each source. The temperature of a source can be determined from measurements of flux densities F_{ν_1} and F_{ν_2} at frequencies, respectively, ν_1 and ν_2 using the relation

$$\frac{F_{\nu_1}}{F_{\nu_2}} = \left(\frac{\nu_1}{\nu_2}\right)^{3+\beta} \frac{e^{\frac{h\nu_2}{k_B T}} - 1}{e^{\frac{h\nu_1}{k_B T}} - 1}, \quad (5)$$

where T is the source temperature and β is the dust opacity index, as before (see e.g. Buckle et al. 2015). Assuming $\beta = 2.0$, then in the case of the ratio of SCUBA-2 450- μm and 850- μm fluxes this relation becomes

$$\frac{F_{\nu(450\mu\text{m})}}{F_{\nu(850\mu\text{m})}} = 24.05 \times \frac{e^{16.96/T} - 1}{e^{32.02/T} - 1}. \quad (6)$$

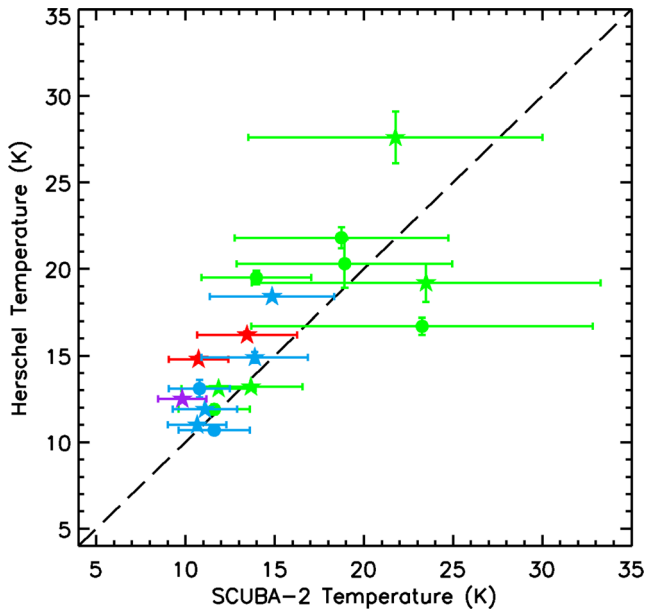


Figure 7. Comparison of SCUBA-2- and *Herschel*-derived (Di Francesco et al., in preparation) temperatures. Red sources lie in L1147/58, light green sources in L1174, blue sources in L1251 and purple sources in L1228. Error bars show 1σ uncertainties as listed in Table 5. The dashed line is the 1:1 line.

This equation can be solved numerically for source temperature T . This analysis presumes that the $450\text{-}\mu\text{m}$ and $850\text{-}\mu\text{m}$ emission traces the same dust population, and that the line-of-sight temperature variation is minimal (see e.g. Shetty et al. 2009). A detailed analysis of dust temperatures determined from SCUBA-2 $450\text{-}\mu\text{m}$ and $850\text{-}\mu\text{m}$ observations of the W40 region was performed by Rumble et al. (2016). We choose $\beta = 2.0$ for consistency with the *Herschel*-derived temperature measurements, as described above.

As discussed in Section 3.2, $450\text{-}\mu\text{m}$ aperture photometry measurements were made using data that were convolved to match the resolution of the $850\text{-}\mu\text{m}$ data using a convolution kernel constructed as described by Pattle et al. (2015), using the SCUBA-2 $450\text{-}\mu\text{m}$ and $850\text{-}\mu\text{m}$ beam models given by Dempsey et al. (2013).

We derived temperatures for those of our sources with a detection $\geq 5\sigma$ at SCUBA-2 $450\text{-}\mu\text{m}$. These temperatures are listed in Table 5. Fig. 7 compares the temperatures derived using SCUBA-2 and *Herschel* data, showing that SCUBA-2-derived and *Herschel*-derived source temperatures (Di Francesco et al., in preparation) are typically in agreement, albeit with large uncertainties on the SCUBA-2-derived temperatures.

This analysis suggests that determining source temperatures using only the ratio of SCUBA-2 $450\text{-}\mu\text{m}$ and $850\text{-}\mu\text{m}$ data will produce reliable results in low-temperature cores. This is as expected, as equation (5) is insensitive to temperature in the Rayleigh–Jeans (RJ) limit ($h\nu/k_{\text{B}}T \ll 1$). The $450\text{-}\mu\text{m}$ data point will fall on the RJ tail of the spectral energy distribution if $T \gg 32\text{ K}$, while the $850\text{-}\mu\text{m}$ data point will fall on the RJ tail if $T \gg 17\text{ K}$. It can be seen in Fig. 7 that the uncertainties on our SCUBA-2-derived temperatures increase substantially when $T > 20\text{ K}$, due to the decreasing sensitivity of the flux density ratio to temperature as source temperature increases.

Fig. 7 shows that there is a slight tendency for source temperatures determined from *Herschel* measurements to be higher than those determined from SCUBA-2 measurements. While this behaviour is not statistically significant, this is consistent with the

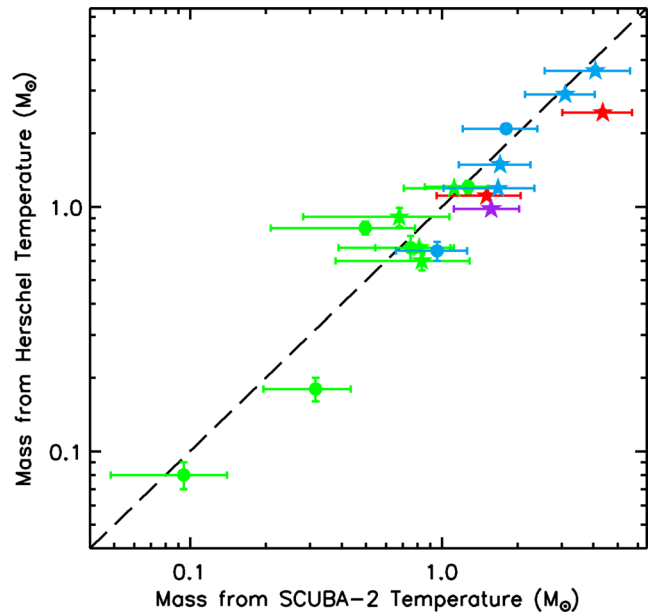


Figure 8. Comparison of source masses determined using SCUBA-2- and *Herschel*-derived (Di Francesco et al., in preparation) temperatures. Red sources lie in L1147/58, light green sources in L1174, blue sources in L1251 and purple sources in L1228. Error bars show 1σ uncertainties as listed in Table 5. The dashed line is the 1:1 line.

shorter wavelength *Herschel* observations being sensitive to emission from warmer material than the longer wavelength SCUBA-2 observations.

We calculated the masses of each of our sources using equation (2), our SCUBA-2-derived temperatures and our best-fitting model $850\text{-}\mu\text{m}$ flux densities. Source masses determined from SCUBA-2 temperatures are listed in Table 5. Fig. 8 compares the masses derived using SCUBA-2-based and *Herschel*-based source temperatures (Di Francesco et al., in preparation), and shows that the two measures of mass are generally in agreement. There is a slight tendency to SCUBA-2-temperature masses to be higher than *Herschel*-temperature masses. This is a result of the tendency for *Herschel*-derived temperatures to be higher than those derived from SCUBA-2 data.

Thus, we conclude that there is reasonable agreement between SCUBA-2-derived and *Herschel*-derived source temperatures for our sources, and that SCUBA-2-derived source temperatures will be accurate when there is a good ($>5\sigma$) source detection at $450\text{-}\mu\text{m}$, and when neither the $450\text{-}\mu\text{m}$ nor the $850\text{-}\mu\text{m}$ data point falls on the Rayleigh–Jeans tail of the spectral energy distribution. For a detailed comparison of core properties determined from SCUBA-2 and *Herschel* data, see Ward-Thompson et al. (2016).

4 DISCUSSION OF DERIVED PROPERTIES

The masses and sizes of our sources are shown in Fig. 9. Our sources typically occupy the upper part of the mass/size plane, in which prestellar cores are expected to lie (Simpson, Nutter & Ward-Thompson 2008), being overdense relative to transient, unbound structure (c.f. André et al. 2010). The grey band on Fig. 9 shows the region in which transient, gravitationally unbound CO clumps are expected to lie (Elmegreen & Falgarone 1996).

The temperatures and volume densities of our sources are shown in Fig. 10. It can be seen that with the exception of sources in

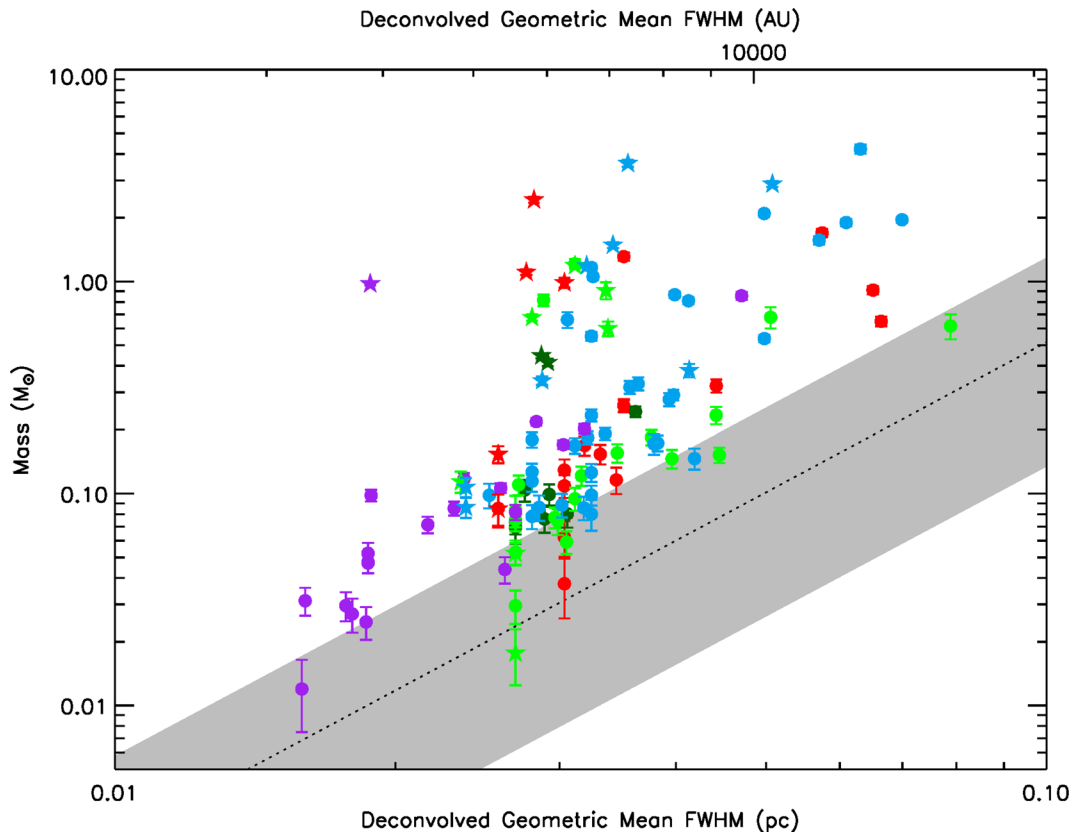


Figure 9. Comparison of source mass (determined using *Herschel*-derived temperatures) and source size (geometric mean of modelled major and minor FWHMs) for the sources in Cepheus. Circles represent starless cores; stars represent cores with embedded protostars. Red sources lie in L1147/58, light green sources in L1174, dark green sources in L1172, blue sources in L1251 and purple sources in L1228.

L1174 – the reflection nebula – the cores in our sample have a narrow range of temperatures ($\sim 9\text{--}15$ K).

In order to determine a mass function for each set of starless cores in our sample, we analysed the cumulative distribution functions of starless core masses for each region in Cepheus, using the maximum likelihood (ML) estimator for an infinite power-law distribution (Koen 2006; Maschberger & Kroupa 2009). Throughout the following discussion, we assume that the masses of cores can be modelled by a power-law function,

$$\xi(M)dM \propto M^{-\alpha}dM, \quad (7)$$

where $\xi(M)dM$ is the number of cores in the mass range M to $M + dM$.

The empirical cumulative distribution function \hat{F} is given, for the i th data point in our sample, by

$$\hat{F}(X_i) \equiv \frac{i}{n+1}, \quad (8)$$

where n is the number of data points X . The ML estimator for the exponent α of an infinite power-law distribution is

$$\alpha_{ml} = 1 + \frac{n}{\left(\sum_{i=1}^n \ln(X_i)\right) - n \ln(\min(X))}. \quad (9)$$

The unbiased maximum likelihood (UML) estimator, α_{uml} is then

$$\alpha_{uml} = 1 + \frac{n-1}{n}(\alpha_{ml} - 1). \quad (10)$$

Uncertainties were estimated by performing a set of Monte Carlo experiments, drawing a set of data points randomly from our dis-

tribution of masses, from which α_{ml} was recalculated. The error quoted is the standard deviation of the distribution of α_{uml} which results from this procedure.

In this analysis, we consider only starless cores, excluding all sources with embedded objects. This is in order to construct cumulative mass distributions comparable to the core mass function (CMF).

The cumulative mass distribution functions for each region of Cepheus except L1172 are shown in Fig. 11, while the cumulative mass distribution function for all of the starless cores in our sample combined is shown in Fig. 12. The ML estimator mass functions for each region are listed in Table 6. L1172 is excluded from this analysis as the region contains only seven starless cores, too few to accurately constrain the power-law index of the region's core mass function.

As can be seen from Fig. 11 and Table 6, the core mass function in each region in the Cepheus Flare other than L1172 can be characterized by a power law above a mass of $0.05 M_{\odot}$. The L1147/58, L1174 and L1251 α_{uml} values are similar, and show a high-mass CMF slope of $\alpha_{uml} = 1.8 \pm 0.2$, 2.0 ± 0.2 and 1.8 ± 0.1 , respectively. The L1228 region, however, has a high-mass CMF slope of $\alpha_{uml, L1228} = 2.3 \pm 0.3$. Whether this difference in CMF slope is indicative of a difference in behaviour between L1228 and the remainder of the sample, or merely of the small sample sizes in each region, is difficult to determine.

We combined the cores from each individual region in order to overcome the problem of small number statistics. The cumulative mass distribution for all of the starless cores we detect in Cepheus

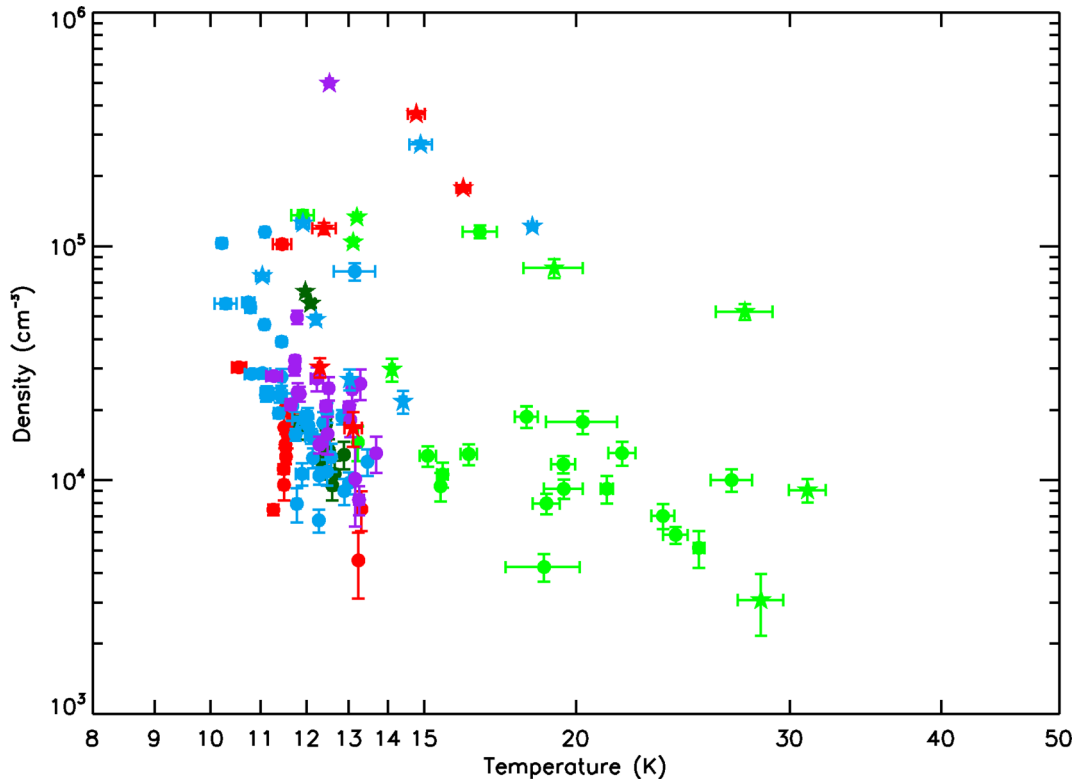


Figure 10. Comparison of temperature and density for the sources in Cepheus. Colour and symbol coding is as in Fig. 9 (red – L1147/58; light green – L1172; dark green – L1174; blue – L1251; purple – L1228).

(including those in L1172) is shown in Fig. 12. There appears to be a break in starless core masses between 0.3 and $0.5 M_{\odot}$, with no starless cores being detected in this mass range. Determining α_{uml} over the mass range $M > 0.08 M_{\odot}$ gives a power-law index of 1.9 ± 0.1 , with starless cores in the mass range 0.08 – $0.3 M_{\odot}$ conforming well to a power-law distribution (see left panel of Fig. 12). Determining α_{uml} for the high-mass cores only ($M > 0.5 M_{\odot}$) gives a steeper power-law index, of 2.6 ± 0.3 . Whether these high-mass starless cores represent a different population is not clear.

Both Chabrier (2003) and Kroupa (2001) predict a power-law index of 2.3 for the high-mass end of the stellar initial mass function (IMF), consistent with the Salpeter (1955) high-mass IMF. Previously, a number of authors have suggested a link between the stellar IMF and the CMF (e.g. Motte, André & Neri 1998; Nutter & Ward-Thompson 2007). In Cepheus, we see a high-mass CMF slope of 2.6 ± 0.3 (when $M > 0.5 M_{\odot}$), consistent with the Kroupa–Chabrier–Salpeter value.

The break in core masses can be seen in Fig. 9, for both starless cores (as discussed in this section) and for cores with embedded protostars. Inspection of Fig. 9 further shows that these most massive cores have a higher average radius than the rest of the population. This might suggest that the more massive cores without embedded sources are a separate population of starless ‘clumps’; objects that might be expected to fragment to form multiple starless cores. The lower mass population of starless clumps might, due to their large radii and low masses and temperatures, be below the detectability limit of SCUBA-2 (Ward-Thompson et al. 2016). However, whether these highest mass objects are in fact a separate population is by no means certain.

5 COUNTING STATISTICS

We compared the number of starless cores in our sample with the number of embedded (Class I and Flat) and Class II sources detected by K09 in the same area, in order to make a crude estimate of the relative level of star formation activity in the different regions of the Cepheus Flare. The absolute number counts are shown in Fig. 13, while the counts normalized to the number of Class II sources in the region are shown in Fig. 14.

Fig. 13 shows that in absolute terms, L1251 contains the highest number of both starless cores and embedded sources, and the second highest number of Class II sources. L1174 contains the highest number of Class II sources; a natural result for a region in which clustered star formation has been ongoing for some time (Kun et al. 2008). L1174 has the second highest number of embedded sources after L1251, and the joint second-highest number of starless cores, along with L1228. L1228, L1147/L1158 and L1172 have low number counts of both embedded and Class II sources. This shows that the sites of ongoing active star formation, L1251 and L1174, have the highest absolute number of sources in almost all categories, while the regions of less-active star formation generally have lower numbers of starless cores as well as of embedded sources. However, in order to determine the evolutionary status of each region, the ratio of starless cores to embedded sources must be considered.

Fig. 14 shows the number of sources of each type in each region, normalized to the number of Class II sources. Again, a difference in behaviour can be seen between the less-active regions, L1147/58 and L1228, and the active regions L1174 and L1251. In the less-active regions, there is a high ratio of starless cores to Class II sources: $\sim 3.8:1$ in L1148/L1157, and $\sim 2.7:1$ in L1228. However,

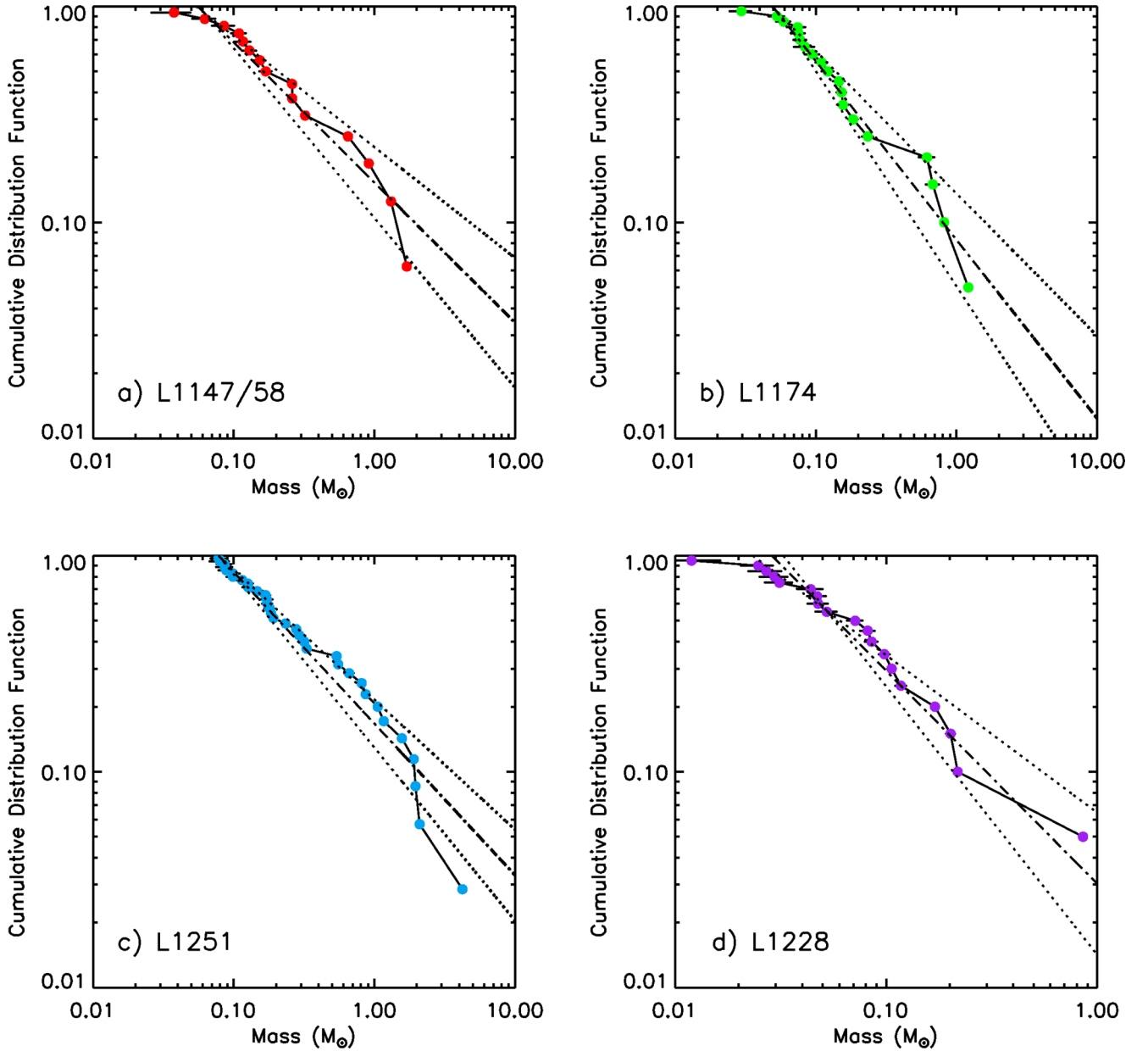


Figure 11. Cumulative probability plots by region. Colour coding is as in Fig. 9 (red – L1147/58; light green – L1172; blue – L1251; purple – L1228).

in the active star-forming regions, this ratio is much lower: $\sim 1.4:1$ in L1251, while in L1174 Class II sources outnumber starless cores, with a ratio $\sim 0.7:1$. L1172 shows an intermediate behaviour, with a ratio $\sim 1.8:1$. However, the low counting statistics in all classes in L1172 make any interpretation of this result difficult.

Since we do not know whether the star formation rate (SFR) in Cepheus has been constant over time, we consider two scenarios: a constant SFR over a very long time or a relatively short and finite burst of star formation.

In the first scenario, we can interpret those regions with a lower ratio of starless cores to Class II sources (L1251 and L1174) as having a high (but constant) SFR, i.e. converting gas into stars efficiently. We interpret regions with a higher ratio of starless cores to Class II sources (L1147/58 and L1228) as having a lower (but still constant) SFR, i.e. these regions are forming stars less efficiently.

In the second scenario, we can interpret those regions with a higher ratio of starless cores to Class II sources as being at an earlier evolutionary stage than those with a lower starless core to Class II ratio, i.e. the regions with a high ratio have thus far converted only a small amount of their reservoir of available material into stars, while the regions with a low ratio have significantly depleted their local reservoir of dense gas.

The ratios that we observe are likely to result from a combination of these effects. We can attempt to determine which effect is more likely to dominate for each region by considering what we know of their star formation histories and the current or historical influences on them. However, all of the following interpretation must be used with care, as our absolute number counts of cores and Class II sources may not be large enough to put our conclusions on a strongly statistically significant footing.

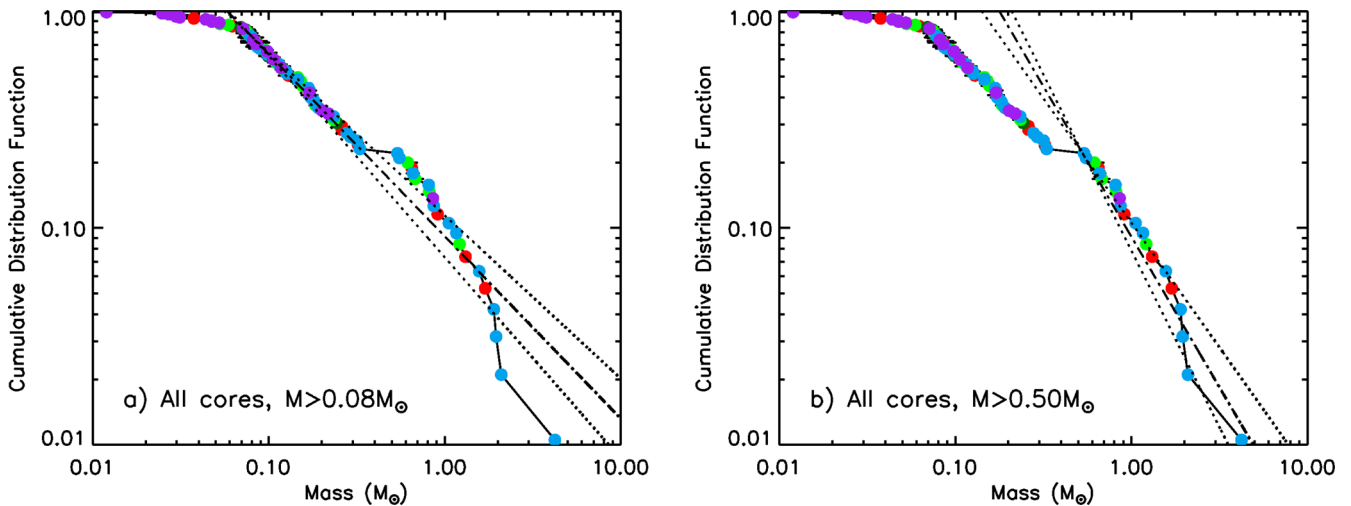


Figure 12. Cumulative probability plots for the starless cores in Cepheus. Left-hand panel: power-law distribution for cores with masses $>0.08 M_{\odot}$. Right-hand panel: power-law distribution for cores with masses $>0.50 M_{\odot}$. Colour coding is as in Fig. 9 (red – L1147/58; light green – L1172; dark green – L1174; blue – L1251; purple – L1228).

Table 6. ML estimator power-law indices for cores in Cepheus.

Region	α_{uml}	Mass range
L1147/L1158	1.8 ± 0.2	$>0.05 M_{\odot}$
L1174	2.0 ± 0.2	$>0.05 M_{\odot}$
L1251	1.8 ± 0.1	$>0.05 M_{\odot}$
L1228	2.3 ± 0.3	$>0.05 M_{\odot}$
All	1.9 ± 0.1	$>0.08 M_{\odot}$
All	2.6 ± 0.3	$>0.5 M_{\odot}$

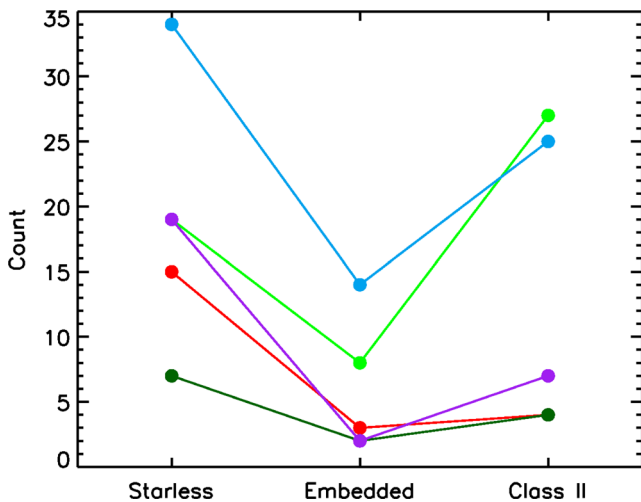


Figure 13. Regional statistics of absolute number of starless, embedded and Class II sources in Cepheus. Colour coding is as in Fig. 9 (red – L1147/58; light green – L1172; dark green – L1174; blue – L1251; purple – L1228).

Star formation in L1251 may have been triggered or enhanced by passage of the Cepheus Flare Shell (CFS) through the region ~ 4 Myr ago, while L1228 may currently be interacting with the CFS (see Section 3.1, above). This is consistent with the high starless-core-to-Class-II ratio in these regions resulting from the second scenario described above, with the low core-to-Class-II ratio in L1251 indicating that the region is significantly more evolved than

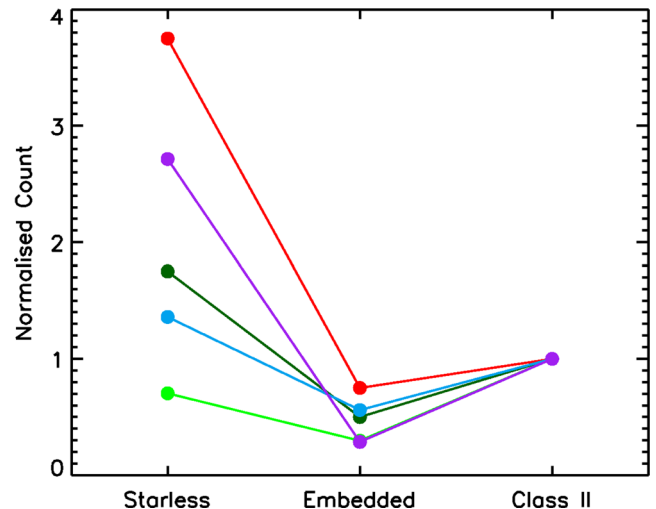


Figure 14. Regional statistics of starless, embedded and Class II sources in Cepheus, normalized to the number of Class II sources. Colour coding is as in Fig. 9 (red – L1147/58; light green – L1172; dark green – L1174; blue – L1251; purple – L1228).

L1228, in which star formation has only recently been triggered or enhanced.

L1147/58, with a high core-to-Class-II ratio, shows no significant signs of recent external influence (see Section 3.1, above), suggesting that here the first scenario might be more likely, and star formation is an ongoing, inefficient process. The clustered star formation in L1174 (low core-to-Class-II ratio) is more difficult to interpret; star formation has been ongoing in this region for some time, suggesting that L1174 might be running out of gas to convert into stars, thus favouring the first scenario.

6 BE STABILITY ANALYSIS

The BE model of a starless core (Ebert 1955; Bonnor 1956) is frequently used as a measure of the stability of starless cores (e.g. Alves et al. 2001). The BE model treats a core as an isothermal, self-gravitating, polytropic sphere bounded by external pressure. For a

given core temperature and external pressure, there is a maximum mass at which the core can be stable against gravitational collapse. The critically stable BE mass is frequently used as a proxy for virial mass (e.g. Könyves et al. 2015). In the following analysis, we consider the stability of the starless cores in Cepheus against gravitational collapse according to the BE model, under the assumption that our cores can be accurately characterized as BE spheres. As discussed above, detailed modelling of core geometries is beyond the scope of this study, and so we cannot definitively state whether our cores have morphologies consistent with the BE model.

6.1 Choice of bounding radius

As discussed above, the BE model treats cores as being bound by external pressure. We have hitherto modelled our cores as having Gaussian density distributions, without a defined edge radius.

We define our edge radius r_e as

$$r_e = \alpha \sqrt{2 \ln \left(\frac{\rho_0}{\rho_e} \right)}, \quad (11)$$

where ρ_0 is the central density of the core and ρ_e is the density at the pressure-confined edge, assuming that the core obeys a Gaussian density distribution at all radii smaller than the edge radius.

The central density ρ_0 can be determined from the modelled mass M and Gaussian width α of each core (listed in Table 5):

$$\rho_0 = \frac{M}{(2\pi\alpha^2)^{3/2}}. \quad (12)$$

We here consider two different bounding densities,

$$\rho_e = \mu m_{\text{H}} \times 10^4 \text{ cm}^{-3}, \quad (13)$$

and

$$\rho_e = \mu m_{\text{H}} \times 10^3 \text{ cm}^{-3}, \quad (14)$$

as representing a physically plausible range of densities at which our cores might be bound, and for consistency with our analysis of external pressure based on measurements of ^{13}CO linewidths in Section 7, below. These choices of bounding density are consistent with our measurements: Fig. 10 shows that the mean density of our cores is $\gtrsim 10^4 \text{ cm}^{-3}$ in almost all cases (the exceptions being warm, low-column-density cores in L1174), and in all of our cores, the central density inferred using equation (12) is $> 10^4 \text{ cm}^{-3}$ (see Table 7).

6.2 Critically stable BE sphere

The mass at which a BE sphere at temperature T , with sound speed $c_s(T)$, and bounded by external pressure P_{EXT} , is critically stable against gravitational collapse is given by

$$M_{\text{BE,crit}} = 1.18 \frac{c_s(T)^4}{P_{\text{EXT}}^{1/2} G^{3/2}}. \quad (15)$$

This can alternatively be expressed in terms of the critically stable BE radius $R_{\text{BE,crit}}$,

$$M_{\text{BE,crit}} = 2.4 \frac{c_s^2}{G} R_{\text{BE,crit}} = 2.4 \frac{k_{\text{B}} T}{\mu m_{\text{H}} G} R_{\text{BE,crit}}. \quad (16)$$

In an attempt to determine whether our starless cores are likely to be virially bound, we determined their BE critically stable masses (Ebert 1955; Bonnor 1956), under the assumption that $R_{\text{BE,crit}} = r_e$. The critically stable BE masses of our cores are listed in Table 7, and

are plotted against our observed core masses in Fig. 15. Fig. 15 suggests that the majority of our cores have stable, pressure-confined (i.e. non-critical) BE solutions. Our choice of bounding density does not significantly affect which of our cores have stable BE solutions.

7 ENERGY BALANCE AND STABILITY

We attempted to assess the energy balance of the starless cores in the Cepheus molecular cloud and to determine the applicability of our BE analysis by estimating the external pressure on our cores using measurements presented by Yonekura et al. (1997) (hereafter Y97). Y97 conducted a large-scale $^{13}\text{CO } J=1 \rightarrow 0$ survey of the Cepheus Flare region using two 4-m telescopes at Nagoya University. Their observations had a resolution of 2.4 arcmin. Each of the regions in our survey is entirely covered by a different, single, Y97 source: Y97 Source 8 for L1147/58, Y97 Source 14 for L1172 and L1174, Y97 Source 79 for L1251 and Y97 Source 66 for L1228. Thus, we can estimate only a single value for external pressure in each region, which we then assume is representative for all of the cores within that region.

Previous studies of starless cores suggested that external gas pressure might be instrumental in confining dense cores in at least some cases (Maruta et al. 2010; Pattle et al. 2015). We apply the method used by Pattle et al. (2015) to estimate the external pressure on starless cores in the Ophiuchus molecular cloud to our sample of starless cores.

We consider the gas pressure in material traced by ^{13}CO to be the external pressure acting on our starless cores, since CO is expected to trace the outer layers, or envelopes, of starless cores (Di Francesco et al. 2007). Higher density tracers such as N_2H^+ are expected to trace the denser inner material of the cores themselves.

We estimate the external pressure P_{EXT} in each region from the linewidths measured by Y97 using the ideal gas law

$$P_{\text{EXT}} \approx \rho_{^{13}\text{CO}} \langle \sigma_{\text{gas},^{13}\text{CO}}^2 \rangle. \quad (17)$$

We consider two different models of the gas pressure in material traced by ^{13}CO . In the first instance, we assume that our cores are bounded at the maximum gas density traced by ^{13}CO , $\rho_{^{13}\text{CO}} \sim 10^4 \text{ cm}^{-3}$ (Di Francesco et al. 2007). We also consider the case in which our cores are bounded at the ‘typical’ density of gas traced by ^{13}CO , $\rho_{^{13}\text{CO}} \sim 10^3 \text{ cm}^{-3}$ (Di Francesco et al. 2007). Hereafter, we refer to these models as ‘high-bounding-density’ and ‘low-bounding-density’ respectively.

The measured mean density of material within a molecular cloud depends strongly on the volume over which it is being assessed. Using the values of total mass and surface area for Cepheus listed by Dunham et al. (2015), $2610 \pm 170 M_{\odot}$ and 38 pc^2 , respectively, we find, assuming a spherical geometry, $\langle n(\text{H}_2) \rangle = 210 \text{ cm}^{-3}$. This value is determined over all areas mapped by *Spitzer* with $A_V \gtrsim 3$. However, K09 list a total cloud mass of $1003 M_{\odot}$ over a total area of 0.4 pc^2 , considering only areas with $A_V > 5$. Again assuming a spherical geometry, this implies a mean volume density $\langle n(\text{H}_2) \rangle = 7.5 \times 10^4 \text{ cm}^{-3}$. Moreover, the sizes and masses listed for individual clumps by K09 suggest densities $\sim 10^5 \text{ cm}^{-3}$ in at least some star-forming clumps. Where in this range of densities our ‘bounding density’ – i.e. the *minimum density of material associated with a potentially star-forming core* – is likely to fall is not immediately clear. However, the apparent threshold for star formation of $A_V \sim 7$ (or higher) in local clouds (e.g. Molinari et al. 2014, and references therein) suggests that star-forming cores are likely to exist within

Table 7. Data relating to starless cores' virial stability: (1) core ID, (2) modelled peak density, (3) gas velocity dispersion determined from low-resolution Y97 ^{13}CO measurements, (4–7) modelled bounding radius, external pressure, BE critically stable mass, and ratio of observed to BE-critical mass, for a bounding density of $n(\text{H}_2) = 10^3 \text{ cm}^{-3}$, (8–11) as 4–7, for a bounding density of $n(\text{H}_2) = 10^4 \text{ cm}^{-3}$.

Source ID	n_0 ($\times 10^4 \text{ cm}^{-3}$)	$\sigma_{^{13}\text{CO, gas}}$ (kms^{-1})	$n(\text{H}_2) = 10^3 \text{ cm}^{-3}$				$n(\text{H}_2) = 10^4 \text{ cm}^{-3}$			
			$\frac{r_{^{13}\text{CO}}}{\text{FWHM}}$	$P_{\text{EXT}}/k_{\text{B}}$ (Kcm^{-3})	M_{BE} (M_{\odot})	$\frac{M}{M_{\text{BE}}}$	$\frac{r_{^{13}\text{CO}}}{\text{FWHM}}$	$P_{\text{EXT}}/k_{\text{B}}$ (Kcm^{-3})	M_{BE} (M_{\odot})	$\frac{M}{M_{\text{BE}}}$
3	35.45 ± 1.58	0.4	1.46	0.6	0.969 ± 0.017	1.35 ± 0.08	1.13	6.2	0.768 ± 0.013	1.71 ± 0.10
4	22.23 ± 0.63	0.4	1.40	0.6	1.307 ± 0.018	1.30 ± 0.08	1.06	6.2	0.957 ± 0.013	1.77 ± 0.10
7	5.82 ± 0.98	0.4	1.21	0.6	0.847 ± 0.003	0.31 ± 0.02	0.80	6.2	0.596 ± 0.002	0.44 ± 0.03
8	5.87 ± 0.70	0.4	1.21	0.6	0.998 ± 0.008	0.32 ± 0.03	0.80	6.2	0.659 ± 0.005	0.49 ± 0.04
9	61.65 ± 1.77	0.4	1.52	0.6	0.759 ± 0.008	0.22 ± 0.03	1.22	6.2	0.525 ± 0.006	0.32 ± 0.04
10	7.03 ± 0.45	0.4	1.24	0.6	0.838 ± 0.006	0.31 ± 0.02	0.84	6.2	0.590 ± 0.004	0.44 ± 0.03
12	6.10 ± 0.65	0.4	1.22	0.6	0.750 ± 0.005	0.15 ± 0.02	0.81	6.2	0.471 ± 0.003	0.25 ± 0.04
13	7.02 ± 0.50	0.4	1.24	0.6	0.760 ± 0.003	0.20 ± 0.02	0.84	6.2	0.511 ± 0.002	0.30 ± 0.03
14	10.52 ± 0.98	0.4	1.30	0.6	0.706 ± 0.003	0.18 ± 0.02	0.92	6.2	0.481 ± 0.002	0.27 ± 0.03
15	3.32 ± 0.48	0.4	1.12	0.6	0.743 ± 0.002	0.15 ± 0.02	0.66	6.2	0.494 ± 0.001	0.22 ± 0.03
16	4.47 ± 0.60	0.4	1.17	0.6	0.740 ± 0.003	0.08 ± 0.02	0.73	6.2	0.441 ± 0.002	0.14 ± 0.03
17	4.94 ± 0.52	0.4	1.19	0.6	0.601 ± 0.003	0.14 ± 0.03	0.76	6.2	0.413 ± 0.002	0.21 ± 0.04
18	3.69 ± 0.52	0.4	1.14	0.6	1.368 ± 0.006	0.47 ± 0.03	0.69	6.2	0.814 ± 0.004	0.80 ± 0.04
19	5.42 ± 0.67	0.4	1.20	0.6	0.681 ± 0.002	0.06 ± 0.02	0.78	6.2	0.340 ± 0.001	0.11 ± 0.03
20	4.57 ± 0.57	0.4	1.17	0.6	1.442 ± 0.013	0.63 ± 0.04	0.74	6.2	0.934 ± 0.008	0.98 ± 0.05
21	4.66 ± 0.53	0.7	1.18	1.6	1.168 ± 0.038	0.70 ± 0.07	0.74	16.0	0.930 ± 0.030	0.88 ± 0.08
26	3.30 ± 0.45	0.7	1.12	1.6	0.914 ± 0.020	1.33 ± 0.10	0.66	16.0	0.733 ± 0.016	1.66 ± 0.12
27	3.88 ± 0.18	0.7	1.15	1.6	2.087 ± 0.141	0.33 ± 0.08	0.70	16.0	1.446 ± 0.098	0.47 ± 0.10
28	4.18 ± 0.68	0.7	1.16	1.6	1.154 ± 0.030	0.07 ± 0.01	0.72	16.0	0.767 ± 0.020	0.10 ± 0.01
29	6.02 ± 0.35	0.7	1.22	1.6	1.515 ± 0.059	0.05 ± 0.01	0.80	16.0	0.961 ± 0.038	0.08 ± 0.01
30	40.02 ± 2.51	0.7	1.47	1.6	0.927 ± 0.009	0.10 ± 0.01	1.15	16.0	0.594 ± 0.006	0.16 ± 0.02
31	36.13 ± 1.19	0.7	1.46	1.6	1.024 ± 0.016	0.15 ± 0.02	1.14	16.0	0.677 ± 0.011	0.23 ± 0.03
33	28.02 ± 2.62	0.7	1.43	1.6	1.175 ± 0.012	0.03 ± 0.00	1.10	16.0	0.619 ± 0.006	0.05 ± 0.01
34	46.19 ± 1.11	0.7	1.49	1.6	1.425 ± 0.032	0.13 ± 0.01	1.18	16.0	0.930 ± 0.021	0.20 ± 0.02
36	6.17 ± 0.71	0.7	1.22	1.6	1.012 ± 0.022	0.11 ± 0.01	0.81	16.0	0.705 ± 0.015	0.16 ± 0.02
37	4.54 ± 0.53	0.7	1.17	1.6	0.872 ± 0.001	0.09 ± 0.01	0.74	16.0	0.547 ± 0.001	0.14 ± 0.02
38	3.48 ± 0.39	0.7	1.13	1.6	1.623 ± 0.059	0.14 ± 0.02	0.67	16.0	1.012 ± 0.037	0.23 ± 0.03
39	3.68 ± 0.46	0.7	1.14	1.6	1.897 ± 0.044	0.08 ± 0.01	0.69	16.0	1.048 ± 0.024	0.15 ± 0.02
40	4.42 ± 0.44	0.7	1.17	1.6	2.489 ± 0.175	0.25 ± 0.08	0.73	16.0	1.199 ± 0.084	0.51 ± 0.11
42	1.78 ± 0.32	0.7	1.02	1.6	0.708 ± 0.004	0.12 ± 0.02	0.46	16.0	0.478 ± 0.003	0.17 ± 0.03
43	4.06 ± 0.34	0.7	1.16	1.6	1.073 ± 0.013	0.05 ± 0.01	0.71	16.0	0.669 ± 0.008	0.08 ± 0.01
44	1.07 ± 0.31	0.7	0.92	1.6	1.381 ± 0.035	0.11 ± 0.01	0.15	16.0	0.834 ± 0.021	0.18 ± 0.02
45	6.49 ± 0.68	0.7	1.23	1.6	1.304 ± 0.029	0.05 ± 0.01	0.82	16.0	0.763 ± 0.017	0.08 ± 0.01
46	3.27 ± 0.46	0.7	1.12	1.6	1.015 ± 0.016	0.12 ± 0.01	0.65	16.0	0.673 ± 0.010	0.18 ± 0.02
49	1.48 ± 0.20	0.7	0.99	1.6	0.661 ± 0.004	0.16 ± 0.02	0.38	16.0	0.455 ± 0.003	0.23 ± 0.03
50	3.15 ± 0.37	0.7	1.12	1.6	0.680 ± 0.001	0.11 ± 0.01	0.64	16.0	0.451 ± 0.001	0.16 ± 0.02
51	5.04 ± 0.86	0.7	1.19	1.6	0.701 ± 0.004	0.11 ± 0.02	0.76	16.0	0.450 ± 0.002	0.17 ± 0.02
52	3.19 ± 0.44	0.7	1.12	1.6	0.722 ± 0.003	0.14 ± 0.02	0.65	16.0	0.481 ± 0.002	0.21 ± 0.02
53	2.76 ± 0.28	0.7	1.09	1.6	0.727 ± 0.005	0.11 ± 0.02	0.61	16.0	0.457 ± 0.003	0.18 ± 0.02
54	2.44 ± 0.30	0.7	1.07	1.6	0.647 ± 0.005	0.11 ± 0.02	0.57	16.0	0.424 ± 0.003	0.16 ± 0.03
55	4.49 ± 0.46	0.7	1.17	1.6	0.916 ± 0.005	0.27 ± 0.02	0.74	16.0	0.633 ± 0.004	0.38 ± 0.02
58	39.98 ± 1.08	0.8	1.47	2.3	0.875 ± 0.007	1.33 ± 0.04	1.15	22.5	0.697 ± 0.005	1.67 ± 0.05
59	27.11 ± 2.31	0.8	1.42	2.3	0.947 ± 0.037	0.70 ± 0.09	1.09	22.5	0.740 ± 0.029	0.89 ± 0.10
60	19.73 ± 0.97	0.8	1.38	2.3	1.488 ± 0.031	2.84 ± 0.23	1.04	22.5	1.141 ± 0.024	3.70 ± 0.27
62	18.98 ± 0.87	0.8	1.38	2.3	0.799 ± 0.008	0.69 ± 0.04	1.03	22.5	0.611 ± 0.006	0.90 ± 0.05
63	9.94 ± 0.43	0.8	1.29	2.3	1.349 ± 0.020	1.16 ± 0.07	0.91	22.5	0.983 ± 0.015	1.60 ± 0.09
64	35.80 ± 1.20	0.8	1.46	2.3	0.802 ± 0.007	1.32 ± 0.05	1.14	22.5	0.636 ± 0.005	1.66 ± 0.06
65	8.00 ± 0.57	0.8	1.26	2.3	0.872 ± 0.011	0.38 ± 0.03	0.87	22.5	0.622 ± 0.008	0.53 ± 0.04
66	16.04 ± 0.55	0.8	1.35	2.3	0.993 ± 0.007	0.87 ± 0.04	1.00	22.5	0.751 ± 0.006	1.15 ± 0.05
67	9.58 ± 0.79	0.8	1.28	2.3	0.685 ± 0.004	0.26 ± 0.02	0.90	22.5	0.497 ± 0.003	0.36 ± 0.03
68	13.57 ± 0.48	0.8	1.33	2.3	1.044 ± 0.008	0.78 ± 0.03	0.97	22.5	0.780 ± 0.006	1.04 ± 0.04
70	6.12 ± 0.67	0.8	1.22	2.3	0.707 ± 0.005	0.16 ± 0.02	0.81	22.5	0.489 ± 0.003	0.23 ± 0.03
71	19.98 ± 0.67	0.8	1.38	2.3	1.225 ± 0.015	1.71 ± 0.08	1.04	22.5	0.940 ± 0.011	2.23 ± 0.10
72	8.18 ± 0.56	0.8	1.26	2.3	0.835 ± 0.010	0.38 ± 0.03	0.87	22.5	0.597 ± 0.007	0.53 ± 0.04
73	3.78 ± 0.48	0.8	1.14	2.3	0.723 ± 0.001	0.12 ± 0.02	0.69	22.5	0.466 ± 0.001	0.19 ± 0.02
74	5.94 ± 0.40	0.8	1.21	2.3	0.820 ± 0.002	0.23 ± 0.02	0.80	22.5	0.566 ± 0.001	0.34 ± 0.02
75	8.03 ± 0.53	0.8	1.26	2.3	0.756 ± 0.003	0.31 ± 0.02	0.87	22.5	0.540 ± 0.002	0.43 ± 0.03
77	3.38 ± 0.36	0.8	1.13	2.3	0.800 ± 0.002	0.12 ± 0.01	0.66	22.5	0.505 ± 0.001	0.20 ± 0.02
78	5.36 ± 0.38	0.8	1.20	2.3	0.955 ± 0.011	0.29 ± 0.02	0.78	22.5	0.650 ± 0.008	0.43 ± 0.03

Table 7 – continued

Source ID	n_0 ($\times 10^4 \text{ cm}^{-3}$)	$\sigma_{13\text{CO, gas}}$ (kms^{-1})	$\frac{r_{13\text{CO}}}{\text{FWHM}}$	$n(\text{H}_2) = 10^3 \text{ cm}^{-3}$			$n(\text{H}_2) = 10^4 \text{ cm}^{-3}$			$\frac{M}{M_{\text{BE}}}$
				$P_{\text{EXT}}/k_{\text{B}}$ (Kcm^{-3})	M_{BE} (M_{\odot})	$\frac{M}{M_{\text{BE}}}$	$\frac{r_{13\text{CO}}}{\text{FWHM}}$	$P_{\text{EXT}}/k_{\text{B}}$ (Kcm^{-3})	M_{BE} (M_{\odot})	
79	3.12 ± 0.40	0.8	1.11	2.3	0.769 ± 0.002	0.11 ± 0.01	0.64	22.5	0.478 ± 0.002	0.18 ± 0.02
80	6.48 ± 0.45	0.8	1.23	2.3	0.846 ± 0.003	0.22 ± 0.02	0.82	22.5	0.590 ± 0.002	0.31 ± 0.02
81	4.17 ± 0.54	0.8	1.16	2.3	0.734 ± 0.003	0.11 ± 0.01	0.72	22.5	0.481 ± 0.002	0.16 ± 0.02
83	7.20 ± 0.95	0.8	1.24	2.3	0.611 ± 0.002	0.16 ± 0.02	0.84	22.5	0.431 ± 0.001	0.23 ± 0.03
85	2.75 ± 0.46	0.8	1.09	2.3	0.705 ± 0.005	0.11 ± 0.02	0.60	22.5	0.425 ± 0.003	0.19 ± 0.03
86	6.54 ± 0.55	0.8	1.23	2.3	0.768 ± 0.006	0.22 ± 0.02	0.82	22.5	0.536 ± 0.004	0.31 ± 0.03
87	4.37 ± 0.60	0.8	1.17	2.3	0.701 ± 0.004	0.12 ± 0.02	0.73	22.5	0.463 ± 0.003	0.19 ± 0.03
88	2.34 ± 0.26	0.8	1.07	2.3	0.926 ± 0.004	0.16 ± 0.02	0.55	22.5	0.536 ± 0.003	0.27 ± 0.03
89	3.70 ± 0.40	0.8	1.14	2.3	0.864 ± 0.010	0.20 ± 0.02	0.69	22.5	0.555 ± 0.007	0.31 ± 0.04
90	5.45 ± 0.32	0.8	1.20	2.3	0.939 ± 0.004	0.31 ± 0.02	0.78	22.5	0.640 ± 0.003	0.46 ± 0.03
92	6.75 ± 0.64	0.8	1.23	2.3	0.720 ± 0.003	0.18 ± 0.02	0.83	22.5	0.504 ± 0.002	0.25 ± 0.02
93	9.87 ± 0.38	0.8	1.29	2.3	1.412 ± 0.018	1.35 ± 0.08	0.91	22.5	1.028 ± 0.013	1.85 ± 0.10
94	4.32 ± 0.42	0.8	1.17	2.3	0.770 ± 0.002	0.16 ± 0.02	0.73	22.5	0.508 ± 0.001	0.25 ± 0.02
95	6.72 ± 0.21	0.8	1.23	2.3	1.638 ± 0.015	1.19 ± 0.05	0.83	22.5	1.146 ± 0.011	1.71 ± 0.07
96	5.13 ± 0.19	0.8	1.19	2.3	1.232 ± 0.005	0.44 ± 0.02	0.77	22.5	0.833 ± 0.004	0.65 ± 0.03
97	3.64 ± 0.31	0.8	1.14	2.3	0.899 ± 0.004	0.19 ± 0.02	0.68	22.5	0.575 ± 0.003	0.30 ± 0.03
98	11.28 ± 0.48	0.7	1.31	1.6	0.722 ± 0.005	0.30 ± 0.01	0.93	16.0	0.532 ± 0.003	0.41 ± 0.02
99	17.27 ± 1.11	0.7	1.36	1.6	0.502 ± 0.002	0.20 ± 0.01	1.01	16.0	0.382 ± 0.002	0.26 ± 0.02
100	7.31 ± 0.38	0.7	1.24	1.6	0.771 ± 0.005	0.26 ± 0.02	0.85	16.0	0.545 ± 0.004	0.37 ± 0.02
101	7.20 ± 0.33	0.7	1.24	1.6	0.781 ± 0.003	0.22 ± 0.01	0.84	16.0	0.551 ± 0.002	0.31 ± 0.01
102	9.42 ± 1.11	0.7	1.28	1.6	0.488 ± 0.006	0.11 ± 0.01	0.90	16.0	0.353 ± 0.004	0.15 ± 0.02
103	8.47 ± 0.91	0.7	1.27	1.6	0.516 ± 0.003	0.09 ± 0.01	0.88	16.0	0.370 ± 0.003	0.13 ± 0.01
104	8.25 ± 0.74	0.7	1.26	1.6	0.538 ± 0.004	0.13 ± 0.01	0.87	16.0	0.385 ± 0.003	0.19 ± 0.02
105	3.52 ± 1.32	0.7	1.13	1.6	0.397 ± 0.000	0.03 ± 0.01	0.67	16.0	0.253 ± 0.000	0.05 ± 0.02
106	9.66 ± 0.41	0.7	1.28	1.6	1.135 ± 0.017	0.75 ± 0.05	0.90	16.0	0.825 ± 0.013	1.04 ± 0.06
107	10.38 ± 0.67	0.7	1.29	1.6	0.599 ± 0.005	0.20 ± 0.01	0.92	16.0	0.438 ± 0.003	0.27 ± 0.02
108	2.86 ± 0.41	0.7	1.10	1.6	0.644 ± 0.006	0.07 ± 0.01	0.62	16.0	0.392 ± 0.003	0.11 ± 0.02
109	4.53 ± 0.81	0.7	1.17	1.6	0.500 ± 0.003	0.05 ± 0.01	0.74	16.0	0.332 ± 0.002	0.07 ± 0.01
110	8.96 ± 1.35	0.7	1.27	1.6	0.451 ± 0.003	0.07 ± 0.01	0.89	16.0	0.325 ± 0.002	0.10 ± 0.01
112	5.47 ± 1.00	0.7	1.20	1.6	0.451 ± 0.001	0.06 ± 0.01	0.78	16.0	0.308 ± 0.000	0.09 ± 0.02
113	4.93 ± 0.41	0.7	1.19	1.6	0.657 ± 0.003	0.12 ± 0.01	0.76	16.0	0.442 ± 0.002	0.19 ± 0.02
114	6.29 ± 0.98	0.7	1.22	1.6	0.471 ± 0.001	0.06 ± 0.01	0.81	16.0	0.327 ± 0.001	0.09 ± 0.01
115	7.14 ± 0.37	0.7	1.24	1.6	0.699 ± 0.001	0.15 ± 0.01	0.84	16.0	0.492 ± 0.001	0.22 ± 0.01
116	8.57 ± 0.99	0.7	1.27	1.6	0.494 ± 0.003	0.10 ± 0.01	0.88	16.0	0.355 ± 0.003	0.13 ± 0.02
117	8.11 ± 0.60	0.7	1.26	1.6	0.575 ± 0.001	0.15 ± 0.01	0.87	16.0	0.411 ± 0.001	0.21 ± 0.02

regions with densities significantly higher than the mean value in the cloud.

As discussed above, we typically find mean core densities $\sim 10^4 \text{ cm}^{-3}$ in our sample, and infer peak core densities $> 10^4 \text{ cm}^{-3}$ but generally $< 10^5 \text{ cm}^{-3}$ (with some exceptions in the densest cores). Hence, we assume that the bounding densities of our cores cannot significantly exceed 10^4 cm^{-3} , and are likely to be lower. Assuming that potentially star-forming cores exist within regions of density higher than the background cloud average, we consider 10^3 – 10^4 cm^{-3} to be representative of the range of densities at which our cores are likely to be bound.

Y97 find the highest ^{13}CO linewidth in L1251, the lowest in L1147/L1158, and the same, intermediate, value in L1172, L1174 and L1228. It is possible that there are, locally, higher external pressures within these regions than are captured by the low-resolution Y97 measurements.

In the following analysis, we treat both the thermal and non-thermal components of the velocity dispersion in ^{13}CO as representing a hydrostatic pressure on our cores – i.e. we are treating the non-thermal component of the velocity dispersion as a modification to the sound speed in the gas (the microturbulent assumption; Chandrasekhar 1951a,b). Whilst on the majority of size scales in

molecular clouds, this has been demonstrated to be an invalid assumption, it has been shown that in both the compressible and incompressible cases, turbulence can provide support against cloud collapse (and hence, conversely, can provide an ‘inward’ pressure promoting collapse) on scales smaller than the thermal Jeans wavelength in the cloud (Mac Low & Klessen 2004, and references therein). For typical conditions in our cores, $T \sim 15 \text{ K}$ and $n(\text{H}_2) \sim 10^4 \text{ cm}^{-3}$, the thermal Jeans wavelength ($\lambda_J = c_s \sqrt{\pi/(G\rho)}$, where c_s is sound speed and ρ is gas density; Jeans 1928) in our cores is $\lambda_J \sim 0.2 \text{ pc}$, an order of magnitude larger than the size scale of our cores (see Tables 5 and 7). Hence, the assumption that the non-thermal component of the velocity dispersion can be treated as a hydrostatic pressure is justifiable in the case of our cores.

7.1 Virial analysis of cores in Cepheus

We performed a virial analysis on our cores, in order to test their stability against collapse, and hence to test the validity of the BE analysis above. We estimate the terms in the virial equation, in order to determine the stability of our cores. Throughout the following analysis, we assume that our cores obey a Gaussian core

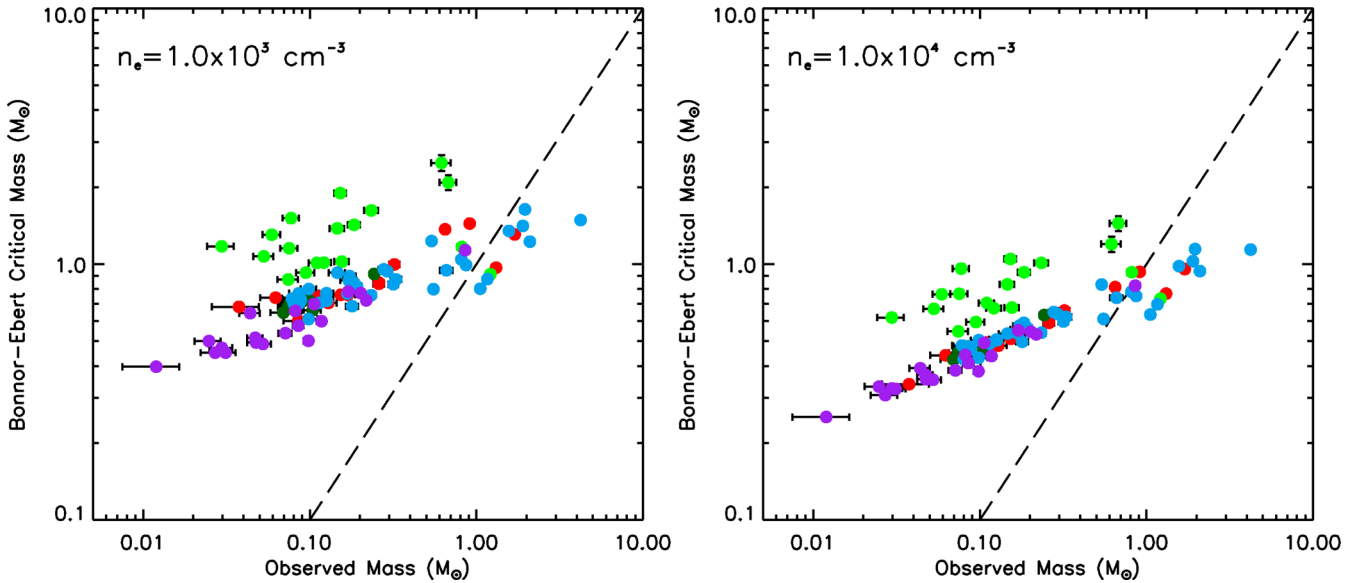


Figure 15. BE stability plot for starless cores in Cepheus, with observed mass plotted against the BE critically stable mass. Left-hand panel: bounding density $n_e = 10^3 \text{ cm}^{-3}$; right-hand panel: bounding density $n_e = 10^4 \text{ cm}^{-3}$. Cores to the right of the dashed line have no stable BE configuration. Colour coding is as in Fig. 9 (red – L1147/58; light green – L1172; dark green – L1174; blue – L1251; purple – L1228).

density profile, consistent with the core fitting process discussed in Section 3.2.

We use the virial equation in the form

$$\frac{1}{2} \ddot{\mathcal{I}} = \Omega_G + 2\Omega_K + \Omega_P, \quad (18)$$

where $\ddot{\mathcal{I}}$ is the second derivative of the moment of inertia \mathcal{I} , Ω_G is the gravitational potential energy of the core, Ω_K is the internal thermal energy of the core and Ω_P is the external pressure energy of the core. We do not include the internal magnetic energy in the following analysis. A core with $\ddot{\mathcal{I}} > 0$ is virially unbound and dispersing, a core with $\ddot{\mathcal{I}} < 0$ is virially bound and collapsing and a core with $\ddot{\mathcal{I}} = 0$ is in virial equilibrium with its surroundings.

We determined external pressure energies for our cores using the equation

$$\Omega_P = -4\pi P_{\text{EXT}} r_e^3. \quad (19)$$

We assume that ^{13}CO traces material adjacent to our cores, and again assume that the cores are confined by external pressure at a density of either 10^3 cm^{-3} or 10^4 cm^{-3} , in both cases at the radius r_e defined in Section 6.1 (see equations 11 and 14). These bounding densities are chosen as representing a range of densities from a typical gas density traced by ^{13}CO ($\sim 10^3 \text{ cm}^{-3}$) to the gas density at which ^{13}CO ceases to be an effective tracer ($\sim 10^4 \text{ cm}^{-3}$; Di Francesco et al. 2007), and as being a physically plausible density at which our cores might be bounded (see discussion in Section 6.1, above).

We determined gravitational potential energies for each of our cores using the equation for the gravitational potential energy of a spherically symmetrical Gaussian density distribution truncated at a radius r_e :

$$\Omega_G = -16\pi^2 G \rho_0^2 \alpha^5 \left[\frac{\sqrt{\pi}}{4} \operatorname{erf}\left(\frac{r_e}{\alpha}\right) - \sqrt{\frac{\pi}{2}} e^{-\frac{1}{2}\left(\frac{r_e}{\alpha}\right)^2} \operatorname{erf}\left(\frac{r_e}{\alpha\sqrt{2}}\right) + \frac{1}{2} \frac{r_e}{\alpha} e^{-\left(\frac{r_e}{\alpha}\right)^2} \right]. \quad (20)$$

where ρ_0 is the modelled central density of the core and α is the modelled Gaussian width of the core – see Pattle (2016) for a derivation of this result.

We were able to put a lower limit on the internal energy of each of our cores by estimating the thermal kinetic energy of the core,

$$\Omega_{\kappa, \tau} = \frac{3}{2} M c_s^2 = \frac{3}{2} M \frac{k_B T}{\mu m_H}. \quad (21)$$

Pattle et al. (2015) found that a significant fraction of the internal kinetic energy of starless cores in the Ophiuchus molecular cloud is non-thermal. Unless starless cores in Cepheus are substantially dissimilar to those in Ophiuchus, there is likely to be a substantial non-thermal component to the internal energy of our cores. Hence the values given by equation (21) are a lower limit on the true value of Ω_{κ} . The values of Ω_G , Ω_P , Ω_{κ} and the virial parameter $\frac{1}{2} \ddot{\mathcal{I}}$ that we determine are listed in Table 8.

7.2 Virial stability of cores in Cepheus

Our best estimate of the virial plane for Cepheus is shown in Fig. 16. The virial plane diagram was introduced by Pattle et al. (2015) as a means of assessing the virial stability and mode of confinement of starless cores. The abscissa shows the virial ratio, $-(\Omega_G + \Omega_P)/2\Omega_{\kappa}$. Cores to the right of the vertical dashed line in the virial plane are virially bound, while cores to the left of this line are virially unbound. The ordinate shows the ratio of gravitational potential energy to external pressure energy, Ω_G/Ω_P . The dominant mode of confinement of cores below the horizontal dashed line is external pressure, while the dominant mode of confinement of cores above this line is self-gravity.

It must be stressed that the values shown for the virial ratio $-(\Omega_G + \Omega_P)/2\Omega_{\kappa}$ are upper limits (for the assumed bounding density); Fig. 16 shows the greatest extent to which our cores could be virially bound.

Fig. 16 suggests that in the high-bounding-density case, our cores are not thermally supported: in the absence of non-thermal internal energy and/or an internal magnetic field, $-(\Omega_G + \Omega_P) > 2\Omega_{\kappa}$ in all

Table 8. Terms in the virial equation for cores in Cepheus: (1) core ID, (2) thermal internal energy (3–5) gravitational potential energy, external pressure energy from ^{13}CO measurements and the virial parameter, for a bounding density of $n(\text{H}_2) = 10^3 \text{ cm}^{-3}$, (6–8) as 3–5 for a bounding density of $n(\text{H}_2) = 10^4 \text{ cm}^{-3}$.

Source ID	Ω_{k}	$n(\text{H}_2)=10^3 \text{ cm}^{-3}$			$n(\text{H}_2)=10^4 \text{ cm}^{-3}$		
		Ω_{G}	Ω_{p}	$\frac{1}{2}\ddot{\mathcal{I}}$ ($\times 10^{41}$ erg)	Ω_{G}	Ω_{p}	$\frac{1}{2}\ddot{\mathcal{I}}$
3	13.0	-28.04	-4.3	-6.4	-30.08	-20.2	-24.4
4	15.4	-29.27	-13.1	-11.5	-36.49	-47.1	-52.7
7	2.6	-1.14	-2.6	+1.5	-1.56	-8.2	-4.5
8	3.2	-1.42	-4.4	+0.6	-2.23	-10.7	-6.5
9	1.7	-0.53	-1.9	+1.0	-0.75	-5.4	-2.8
10	2.6	-1.14	-2.6	+1.4	-1.55	-8.2	-4.5
12	1.2	-0.24	-1.9	+0.2	-0.42	-3.7	-1.8
13	1.5	-0.43	-1.9	+0.7	-0.65	-5.1	-2.7
14	1.3	-0.33	-1.5	+0.7	-0.48	-4.2	-2.1
15	1.2	-0.24	-1.4	+0.7	-0.37	-3.6	-1.6
16	0.7	-0.08	-1.1	+0.2	-0.15	-1.8	-0.5
17	0.8	-0.17	-1.0	+0.6	-0.24	-2.8	-1.3
18	6.3	-4.00	-11.9	-3.3	-7.51	-18.9	-13.8
19	0.4	-0.03	-0.9	-0.1	-0.06	-0.6	+0.2
20	9.0	-7.78	-13.3	-3.0	-12.75	-30.0	-24.7
21	11.7	-13.24	-6.2	+4.0	-14.10	-30.0	-20.6
26	12.5	-27.02	-8.2	-10.2	-28.50	-40.4	-43.9
27	11.9	-5.42	-19.1	-0.8	-7.67	-56.0	-39.9
28	1.4	-0.13	-2.6	+0.1	-0.20	-6.4	-3.8
29	1.8	-0.12	-3.1	+0.4	-0.21	-6.4	-3.0
30	1.3	-0.18	-3.7	-1.3	-0.29	-7.9	-5.7
31	2.0	-0.42	-5.4	-1.8	-0.67	-13.2	-9.8
33	0.6	-0.02	-1.7	-0.4	-0.04	-1.5	-0.3
34	3.1	-0.55	-6.7	-1.1	-0.89	-15.6	-10.3
36	1.7	-0.27	-3.0	+0.2	-0.37	-9.0	-5.9
37	1.0	-0.11	-3.1	-1.2	-0.20	-6.1	-4.3
38	3.9	-0.77	-9.8	-2.7	-1.35	-19.0	-12.5
39	3.2	-0.34	-8.1	-2.1	-0.68	-9.3	-3.6
40	10.0	-3.28	-38.2	-21.5	-6.68	-21.1	-7.7
42	1.0	-0.16	-2.7	-0.9	-0.23	-7.1	-5.4
43	1.0	-0.06	-2.2	-0.3	-0.11	-4.3	-2.5
44	2.4	-0.34	-6.6	-2.2	-0.62	-11.2	-7.1
45	1.2	-0.07	-2.9	-0.5	-0.14	-4.2	-2.0
46	1.7	-0.28	-4.2	-1.0	-0.44	-10.3	-7.3
49	1.1	-0.23	-3.0	-1.1	-0.34	-8.7	-6.9
50	0.8	-0.12	-2.5	-1.0	-0.19	-6.3	-4.8
51	0.8	-0.12	-2.9	-1.4	-0.20	-6.4	-4.9
52	1.1	-0.20	-3.3	-1.4	-0.31	-8.4	-6.6
53	0.9	-0.13	-3.3	-1.7	-0.23	-6.6	-5.1
54	0.7	-0.11	-2.5	-1.1	-0.17	-5.9	-4.6
55	2.6	-0.97	-6.9	-2.7	-1.39	-20.1	-16.3
58	11.1	-23.88	-12.5	-14.1	-25.42	-60.3	-63.4
59	7.5	-8.19	-9.4	-2.7	-8.97	-42.7	-36.7
60	37.5	-162.68	-76.0	-163.7	-184.08	-322.0	-431.1
62	5.1	-5.42	-10.2	-5.4	-6.16	-43.0	-38.9
63	14.9	-25.21	-45.5	-40.8	-31.79	-160.5	-162.4
64	9.3	-19.56	-12.3	-13.3	-20.97	-58.3	-60.7
65	3.2	-1.75	-11.0	-6.3	-2.32	-36.1	-31.9
66	8.3	-10.89	-18.1	-12.4	-12.66	-73.0	-69.1
67	1.8	-0.67	-5.3	-2.5	-0.86	-18.6	-15.9
68	8.0	-9.26	-19.0	-12.3	-11.03	-73.6	-68.6
70	1.2	-0.28	-4.6	-2.4	-0.40	-13.3	-11.3
71	19.4	-50.83	-37.4	-49.4	-57.43	-158.8	-177.4
72	3.0	-1.66	-10.5	-6.0	-2.18	-34.5	-30.6
73	1.0	-0.16	-4.7	-3.0	-0.26	-10.5	-8.9
74	2.0	-0.65	-7.8	-4.5	-0.93	-22.4	-19.4
75	2.2	-0.99	-7.8	-4.3	-1.31	-25.6	-22.4
77	1.1	-0.18	-5.6	-3.6	-0.32	-11.4	-9.6

Table 8 – *continued*

Source ID	Ω_{k}	$n(\text{H}_2)=10^3 \text{ cm}^{-3}$			$n(\text{H}_2)=10^4 \text{ cm}^{-3}$		
		Ω_{G}	Ω_{p}	$\frac{1}{2}\ddot{\mathcal{I}}$ ($\times 10^{41}$ erg)	Ω_{G}	Ω_{p}	$\frac{1}{2}\ddot{\mathcal{I}}$
78	2.9	-1.18	-12.0	-7.4	-1.74	-33.0	-28.9
79	1.0	-0.14	-5.1	-3.4	-0.25	-9.8	-8.1
80	2.0	-0.62	-7.1	-3.6	-0.87	-21.2	-18.0
81	0.9	-0.13	-3.9	-2.3	-0.21	-9.3	-7.7
83	1.0	-0.23	-3.5	-1.8	-0.31	-11.1	-9.4
85	0.8	-0.12	-5.1	-3.6	-0.23	-8.7	-7.3
86	1.7	-0.54	-6.4	-3.5	-0.75	-19.4	-16.6
87	0.9	-0.16	-4.2	-2.5	-0.25	-10.3	-8.7
88	1.6	-0.33	-10.2	-7.5	-0.63	-14.3	-11.9
89	1.8	-0.48	-9.3	-6.3	-0.79	-20.3	-17.6
90	3.0	-1.28	-12.5	-7.9	-1.88	-34.5	-30.5
92	1.4	-0.34	-4.7	-2.4	-0.47	-14.5	-12.2
93	17.7	-34.67	-55.4	-54.6	-43.79	-195.0	-203.4
94	1.3	-0.30	-6.2	-3.9	-0.47	-15.1	-12.9
95	19.2	-32.43	-73.5	-67.5	-44.88	-223.8	-230.3
96	5.8	-3.49	-24.0	-15.9	-5.22	-64.1	-57.8
97	1.8	-0.48	-9.5	-6.3	-0.81	-20.4	-17.5
98	2.2	-0.98	-4.1	-0.7	-1.20	-15.1	-11.9
99	1.0	-0.30	-1.4	+0.3	-0.34	-5.7	-4.0
100	2.0	-0.76	-5.1	-1.8	-1.02	-16.1	-13.0
101	1.8	-0.57	-4.3	-1.2	-0.77	-13.6	-10.7
102	0.6	-0.09	-1.1	-0.1	-0.11	-3.9	-2.9
103	0.5	-0.07	-1.1	-0.1	-0.09	-3.6	-2.6
104	0.7	-0.14	-1.7	-0.3	-0.18	-5.5	-4.2
105	0.1	-0.01	-0.5	-0.2	-0.01	-1.0	-0.7
106	8.3	-9.13	-18.0	-10.4	-11.58	-62.8	-57.7
107	1.2	-0.34	-2.3	-0.3	-0.43	-8.4	-6.4
108	0.5	-0.05	-2.0	-1.0	-0.08	-3.4	-2.5
109	0.3	-0.02	-0.8	-0.3	-0.03	-2.1	-1.6
110	0.4	-0.04	-0.7	-0.0	-0.05	-2.3	-1.7
112	0.3	-0.02	-0.8	-0.3	-0.04	-2.3	-1.7
113	0.9	-0.15	-2.6	-1.1	-0.23	-6.9	-5.4
114	0.3	-0.03	-0.8	-0.2	-0.04	-2.4	-1.8
115	1.2	-0.26	-2.7	-0.6	-0.35	-8.5	-6.5
116	0.5	-0.07	-1.1	-0.1	-0.09	-3.6	-2.7
117	0.9	-0.19	-2.0	-0.5	-0.24	-6.6	-5.1

but one case (core 19, which is marginally unbound). If the cores' dominant support mechanism were internal *thermal* motions, then all but two of our cores (core 19, and core 3, discussed below) would be simultaneously undergoing pressure-driven collapse.

The physical picture in the low-bounding-density case is less clear than in the high-density case. In this case, the cores are less strongly bound and less pressure-confined than in the high-density case. Fig. 16 shows that our cores remain typically virially bound and pressure-confined in the low-bounding-density-case, but that a significant fraction of the cores are in or near virial equilibrium, particularly those in L1147/58, L1174 and L1228. Many of the cores in L1147/58 are marginally unbound in this analysis, along with four cores in L1174 and one core in L1228. All of the cores in L1251 and L1172 remain virially bound. In this case, there are seven gravitationally and virially bound cores in the sample, and one core that is gravitationally dominated but marginally virially unbound.

In either case, these results suggest that a significant fraction of our cores are simultaneously undergoing pressure-driven collapse. As this scenario is unlikely, we hypothesize three possible alternative scenarios.

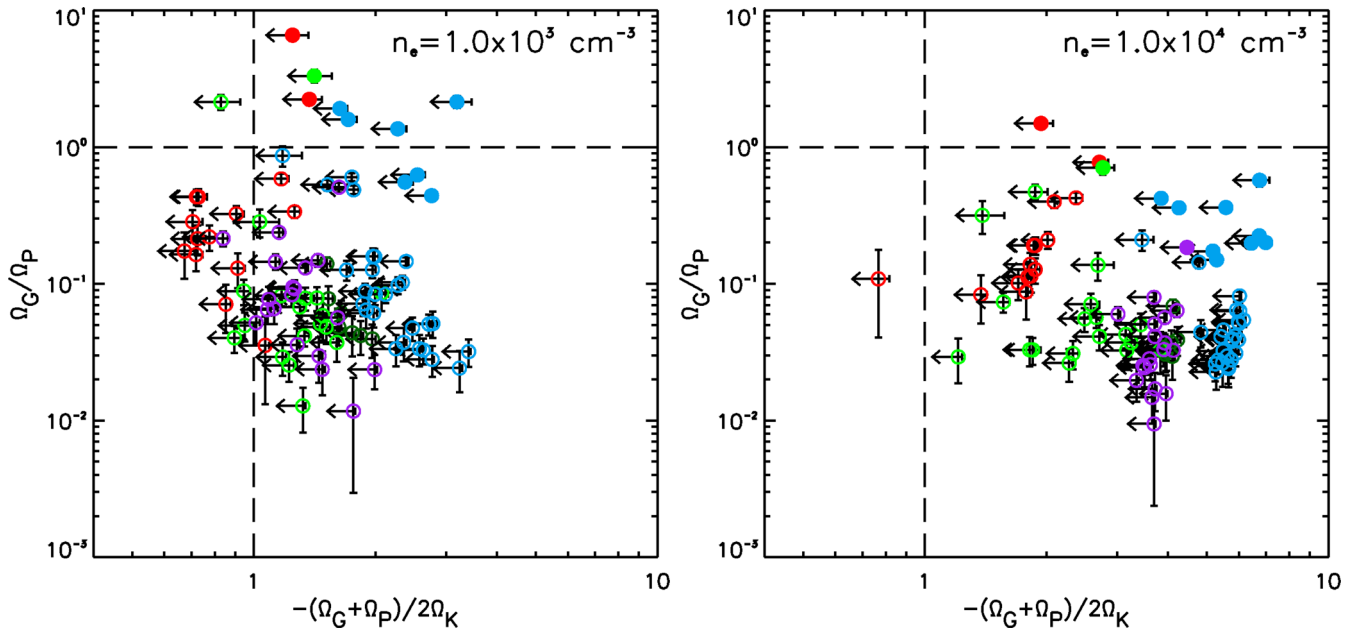


Figure 16. The virial plane for our starless cores. Left-hand panel: bounding density $n_e = 10^3 \text{ cm}^{-3}$; right-hand panel: bounding density $n_e = 10^4 \text{ cm}^{-3}$. Virial stability is plotted on the x -axis. The ratio of gravitational potential energy to external pressure energy is plotted on the y -axis. The vertical dashed line indicates the line of virial stability, with the right-hand side of the plot being bound and the left-hand side being unbound. The horizontal dashed line marks equipartition between external pressure energy and gravitational potential energy; cores above the line are gravitationally bound, while cores below the line are pressure-confined. Closed circles indicate cores with a mass greater than their BE-critical mass. Colour coding as in Fig. 9 (red – L1147/58; light green – L1172; dark green – L1174; blue – L1251; purple – L1228).

(1) The Y97 measurements overestimate the external pressure on our cores. If the low-resolution, and hence large-scale, Y97 ^{13}CO measurements do not correspond accurately to the gas immediately surrounding our dense cores, the linewidths measured by Y97 will not accurately represent the pressure confining the core. As turbulent motions are expected to dissipate on small scales (e.g. Larson 1981; Solomon et al. 1987), velocities measured with the Y97 beam size of 2.4 arcmin (corresponding to 12 pc at a distance of 300 pc) may not be representative of the velocities of material surrounding the sub-parsec-scale cores we consider here.

(2) The non-thermal motions of the gas surrounding the core do not create the effect of a hydrostatic pressure on the core, or do not do so in such a manner that the measured linewidth accurately represents the pressure on the core caused by non-thermal gas motions (see e.g. Mac Low & Klessen 2004).

(3) The dominant mechanism of core support in Cepheus has not been accounted for in our virial analysis. In this scenario, cores are predominantly supported by some combination of internal non-thermal motions and/or internal magnetic field. Pattle et al. (2015) found that the majority of the starless cores in the highest-column-density regions of the Ophiuchus molecular cloud were typically in approximate virial equilibrium with their surroundings, and marginally pressure dominated, with the majority of support against collapse provided by non-thermal internal motions. Fig. 16 is consistent with starless cores in Cepheus behaving in a similar way to those in Ophiuchus, presuming that there is sufficient internal support from non-thermal internal motions and/or internal magnetic fields to bring the cores into approximate virial equilibrium with their surroundings.

None of these hypotheses are contradictory, and all may contribute to the apparent over-estimation of the degree to which our cores are pressure-confined.

Values of the virial ratio are marked as upper limits on Fig. 16, as we can identify the information missing from our determination of the cores' virial ratios: an estimate of the cores' non-thermal and magnetic internal energies. If either of hypotheses (1) and (2) above are valid, then the values of the confinement ratio Ω_G/Ω_P shown on Fig. 16 are lower limits on the true values. However, we do not mark them as such on Fig. 16, as we do not know which of our hypotheses best explain the measured values of external pressure.

The high-density analysis suggests that there is one gravitationally bound prestellar core amongst our sample: core 3 in L1147/58, for which $\Omega_G > \Omega_P$ and $-(\Omega_G + \Omega_P) > 2\Omega_K$. This core is among the 13 cores predicted to be gravitationally unstable by the BE criterion. It is worth noting that the majority of the cores that are only mildly pressure dominated are unstable according to the BE criterion. If the measured values of the confinement ratio are in fact lower limits on the true values, many of these cores may prove to be gravitationally bound.

The low-density analysis suggests that there are seven gravitationally bound prestellar cores amongst our sample, all of which are unstable according to the BE criterion. Three of the cores that are mildly pressure dominated are unstable according to the BE criterion.

The BE stability criterion is in better agreement with our virial analysis in the low-density case than in the high-density case. If the BE criterion is in fact an accurate measure of the stability of our cores, it suggests that the low-density model ($n_e = 10^3 \text{ cm}^{-3}$) is a more accurate description of the energy balance of our cores than the high-density model.

That the low-density model is more likely to be accurate is also supported by the values of the virial ratio shown in Fig. 16: in the low-density case, a significant fraction of the cores have virial ratios consistent with their being in or near virial equilibrium with their surroundings, as might be expected in reasonably long-lived

star-forming regions (all of the Cepheus clouds have been forming stars for long enough to form at least some Class II protostars; see K09).

In the high-density case, almost all of the cores appear to be strongly pressure-confined and virially bound, suggesting that they are all effectively imploding under pressure – an unlikely situation for cores in star-forming regions that appear to have been forming stars continuously for a significant length of time, particularly those showing no signs of recent external influence. Hence, we conclude that of our two models, a density of 10^3 cm^{-3} is the more likely to be representative of the true density at which our cores are confined by pressure from the surrounding molecular cloud.

7.3 Resolving the virial balance of cores in Cepheus

The minimum additional information that is required in order to determine which of our cores are in fact virially bound is a measure of the cores’ internal linewidths, i.e. observations of the cores in an optically thin dense gas tracer such as C^{18}O or nitrogen-bearing tracers (e.g. NH_3 , N_2H^+). This would allow determination of the level of core support from non-thermal internal motions.

Ideally, a measure of the magnetic field strength in the cores is also required, to determine whether magnetic fields play a significant rôle in core support in Cepheus. This might be achieved using a wide-field polarimeter such as POL-2 on the JCMT (Friberg et al. 2016; Ward-Thompson et al., in preparation).

Our estimates of the external surface pressure on the cores could be improved with higher resolution observations of the Cepheus Flare clouds in medium-density tracers such as ^{13}CO . Measurements of the linewidth of gas surrounding the cores taken with an instrument with angular resolution comparable to the angular size of the cores (e.g. using HARP-B on the JCMT; Buckle et al. 2010) would allow us to exclude hypothesis (1), above.

8 SUMMARY

In this paper, we have extracted sources from the SCUBA-2 data of the L1147/L1158, L1172/L1174, L1251 and L1228 regions of the Cepheus Flare. We have characterized our sources using their 850- μm flux densities and temperatures supplied by the *Herschel* GBS (Di Francesco et al., in preparation). We have compared the properties of cores in the different Cepheus Flare regions in order to determine the mode of star formation proceeding in each region. We have determined the relative importance of gravity and external pressure in confining our cores, and have determined an upper limit on the degree to which our cores are virially bound.

We identified 117 sources across the Cepheus Flare region using the CSAR source extraction algorithm, of which 23 were associated with a protostar in the K09 *Spitzer* catalogue. Of our 117 sources, 20 were located in L1147/L1158, 26 in L1174, 9 sources in L1172, 42 in L1251 and 20 in L1228. We determined the best-fitting flux densities of our sources using the multiple-Gaussian fitting algorithm described by Pattle et al. (2015).

We determined masses for each of our sources using our best-fitting flux densities and temperatures supplied by the *Herschel* GBS (Di Francesco et al., in preparation). We found that our cores typically lie in the ‘prestellar’ part of the mass/size plane. Our cores typically have temperatures in the range ~ 9 –15 K, with the exception of cores associated with the L1174 reflection nebula that have temperatures up to ~ 30 K.

We determined source temperatures from the ratio of SCUBA-2 450- μm and 850- μm flux densities, for those of our sources with a

detection $\geq 5\sigma$ at 450 μm . We found that temperatures determined from *Herschel* and SCUBA-2 data were generally in agreement for our sources. We found a slight tendency for *Herschel*-derived temperatures to be higher than SCUBA-2-derived temperatures, consistent with *Herschel* observations sampling slightly warmer material. Source masses derived from SCUBA-2 temperatures are correspondingly slightly higher than those derived from *Herschel* temperatures. We concluded that the SCUBA-2 flux density ratio is a reliable determinant of a source’s temperature when neither the 450- μm nor the 850- μm data point is on the Rayleigh–Jeans tail of a source’s spectral energy distribution – i.e. when $T \lesssim 20$ K.

We analysed the cumulative distribution functions of core masses for each region in Cepheus, using the ML estimator for an infinite power-law distribution, and found that the core mass function in each region shows a sub-Salpeter power-law behaviour, with the exception of L1228 that has a power-law index consistent with the Salpeter IMF. Determining the power-law index over all cores, we found a sub-Salpeter value of $\alpha = 1.88 \pm 0.09$ over the mass range $M > 0.08 M_{\odot}$. For the highest mass cores, we found a CMF power-law index $\alpha = 2.61 \pm 0.27$ over the mass range $M > 0.5 M_{\odot}$ (again determined over all cores), marginally consistent with the Salpeter IMF.

We compared the number of starless cores detected in each region with the numbers of embedded and Class II sources found by K09. We found that L1147/L1158 and L1228 have a high ratio of starless cores to Class II sources, while L1251 and L1174 have a low ratio. This is consistent with L1174 and L1251 being active sites of star formation, while L1147/L1158 and L1228 form stars in a less-active mode.

We determined the BE critically stable masses of our cores, and found that the BE model predicts that most of our cores have stable BE solutions accessible to them.

We performed a virial analysis on our cores, determining the external pressure on our cores using ^{13}CO velocity dispersion measurements determined by Y97. We found that, assuming a bounding density for our cores of 10^4 cm^{-3} , all but one of our cores are virially bound and there is only one gravitationally bound prestellar core among our sample, with the rest of the cores being confined by external pressure. However, we found that if we assume a bounding density of 10^3 cm^{-3} , seven of our cores are gravitationally bound, and the majority of the cores are approximately virialized or mildly pressure-confined. We concluded that, if the Y97 measurements are representative of the conditions in the gas confining our cores, our cores typically cannot be supported by internal thermal energy alone. In the absence of non-thermal internal motions or an internal magnetic field, a significant fraction of our cores would be significantly out of virial equilibrium and collapsing under pressure. We hence hypothesize that cores in the Cepheus molecular cloud may not typically be thermally supported.

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REFERENCES

- Alves J. F., Lada C. J., Lada E. A., 2001, *Nature*, 409, 159
- André P., Belloche A., Motte F., Peretto N., 2007, *A&A*, 472, 519
- André P. et al., 2010, *A&A*, 518, L102
- Aniano G., Draine B. T., Gordon K. D., Sandstrom K., 2011, *PASP*, 123, 1218
- Beckwith S. V. W., Sargent A. I., Chini R. S., Guesten R., 1990, *AJ*, 99, 924
- Beichman C. A., Neugebauer G., Habing H. J., Clegg P. E., Chester T. J., eds, 1988, *Infrared Astronomical Satellite (IRAS) Catalogs and Atlases*. Vol. 1: Explanatory Supplement. Available at: irsa.ipac.caltech.edu/IRASdocs/exp_sup/
- Bonnor W. B., 1956, *MNRAS*, 116, 351
- Buckle J. V. et al., 2010, *MNRAS*, 401, 204
- Buckle J. V. et al., 2015, *MNRAS*, 449, 2472
- Chabrier G., 2003, *PASP*, 115, 763
- Chandrasekhar S., 1951a, *Proc. R. Soc. Ser. A*, 210, 18
- Chandrasekhar S., 1951b, *Proc. R. Soc. Ser. A*, 210, 26
- Chapin E. L., Berry D. S., Gibb A. G., Jenness T., Scott D., Tilanus R. P. J., Economou F., Holland W. S., 2013, *MNRAS*, 430, 2545
- Chen M. C.-Y. et al., 2016, *ApJ*, 826, 95
- Chini R., Ward-Thompson D., Kirk J. M., Nielbock M., Reipurth B., Sievers A., 2001, *A&A*, 369, 155
- Currie M. J., Berry D. S., Jenness T., Gibb A. G., Bell G. S., Draper P. W., 2014, in Manset N., Forshay P., eds, *ASP Conf. Ser. Vol. 485, Astronomical Data Analysis Software and Systems XXIII*. Astron. Soc. Pac., San Francisco, p. 391
- Dempsey J. T. et al., 2013, *MNRAS*, 430, 2534
- Di Francesco J., Evans N. J. I., Caselli P., Myers P. C., Shirley Y., Aikawa Y., Tafalla M., 2007, in Reipurth B., Jewitt D., Tucson K. K., eds, *Protostars and Planets V*. Univ. Arizona Press, Tucson, AZ, p. 17
- Dobashi K., Uehara H., Kandori R., Sakurai T., Kaiden M., Umemoto T., Sato F., 2005, *PASJ*, 57, 1
- Drabek E. et al., 2012, *MNRAS*, 426, 23
- Dunham M. M. et al., 2013, *AJ*, 145, 94
- Dunham M. M. et al., 2015, *ApJS*, 220, 11
- Ebert R., 1955, *Z. Astrophys.*, 37, 217
- Elmegreen B. G., Falgarone E., 1996, *ApJ*, 471, 816
- Enoch M. L., Evans N. J. I., Sargent A. I., Glenn J., Rosolowsky E., Myers P., 2008, *ApJ*, 684, 1240
- Evans N. J., II, et al., 2009, *ApJS*, 181, 321
- Friberg P., Bastien P., Berry D., Savini G., Graves S. F., Pattle K., 2016, *Proc. SPIE*, 9914, 991403
- Gómez L., Wyrowski F., Schuller F., Menten K. M., Ballesteros-Paredes J., 2014, *A&A*, 561, A148
- Goodman A. A., Arce H. G., 2004, *ApJ*, 608, 831
- Gould B. G., 1879, *Resultados del Obs. Nac. Argentino*, 1, 0
- Grenier I. A., Lebrun F., Arnaud M., Dame T. M., Thaddeus P., 1989, *ApJ*, 347
- Guetter H. H., 1968, *PASP*, 80, 197
- Herschel W., 1802, *R. Soc. Phil. Trans. Ser. I*, 92, 477
- Herschel Sir J. F. W., 1847, *Results of astronomical observations made during the years 1834, 5, 6, 7, 8, at the Cape of Good Hope*. Smith, Elder and Co., London
- Hildebrand R. H., 1983, *QJRAS*, 24, 267
- Hirota T., Ito T., Yamamoto S., 2002, *ApJ*, 565, 359
- Holland W. S. et al., 2013, *MNRAS*, 430, 2513
- Hubble E., 1934, *ApJ*, 79, 8
- Jeans J. H., 1928, *Astronomy and Cosmogony*. Cambridge Univ. Press, Cambridge
- Kackley R., Scott D., Chapin E., Friberg P., 2010, *Proc. SPIE*, 7740, 77401Z
- Kirk J. M. et al., 2009, *ApJS*, 185, 198 (K09)
- Kirk J. M. et al., 2013, *MNRAS*, 432, 1424
- Koen C., 2006, *MNRAS*, 365, 590
- Könyves V. et al., 2015, *A&A*, 584, A91
- Kroupa P., 2001, *MNRAS*, 322, 231
- Kun M., Kiss Z. T., Balog Z., 2008, in Reipurth B., ed., *Handbook of Star Forming Regions*, Vol. I. Astron. Soc. Pac., San Francisco, p. 136
- Kun M., Balog Z., Kenyon S. J., Mamajek E. E., Gutermuth R. A., 2009, *ApJS*, 185, 451
- Larson R. B., 1981, *MNRAS*, 194, 809
- Lee J.-E. et al., 2006, *ApJ*, 648, 491
- Lee J.-E., Di Francesco J., Bourke T. L., Evans N. J., II, Wu J., 2007, *ApJ*, 671, 1748
- Li W., Evans N. J., II, Harvey P. M., Colome C., 1994, *ApJ*, 433, 199
- Lynds B. T., 1962, *ApJS*, 7, 1
- Mac Low M.-M., Klessen R. S., 2004, *Rev. Mod. Phys.*, 76, 125
- Mairs S. et al., 2015, *MNRAS*, 454, 2557
- Markwardt C. B., 2009, in Bohlender D. A., Durand D., Dowler P., eds, *ASP Conf. Ser. Vol. 411, Astronomical Data Analysis Software and Systems XVIII*. Astron. Soc. Pac., San Francisco, p. 251
- Maruta H., Nakamura F., Nishi R., Ikeda N., Kitamura Y., 2010, *ApJ*, 714, 680
- Maschberger T., Kroupa P., 2009, *MNRAS*, 395, 931
- Miville-Deschênes M.-A., Lagache G., 2005, *ApJS*, 157, 302
- Molinari S. et al., 2014, in Beuther H., Klessen R., Dullemond C., Henning T., eds, *Protostars and Planets VI*. Univ. Arizona Press, Tucson, AZ, p. 125
- Moshir M., Kopman G., Conrow T. A. O., 1992, *IRAS Faint Source Survey, Explanatory Supplement Version 2*. Jet Propulsion Laboratory, Pasadena
- Motte F., André P., Neri R., 1998, *A&A*, 336, 150
- Nutter D., Ward-Thompson D., 2007, *MNRAS*, 374, 1413
- Nutter D., Stamatellos D., Ward-Thompson D., 2009, *MNRAS*, 396, 1851
- Olano C. A., Meschin P. I., Niemela V. S., 2006, *MNRAS*, 369, 867
- Pattle K., 2016, *MNRAS*, 459, 2651
- Pattle K. et al., 2015, *MNRAS*, 450, 1094
- Plummer H. C., 1911, *MNRAS*, 71, 460
- Reipurth B., Bally J., Devine D., 1997, *AJ*, 114, 2708
- Rumble D. et al., 2016, *MNRAS*, 460, 4150
- Sadavoy S. I. et al., 2013, *ApJ*, 767, 126
- Salpeter E. E., 1955, *ApJ*, 121, 161
- Sato F., Mizuno A., Nagahama T., Onishi T., Yonekura Y., Fukui Y., 1994, *ApJ*, 435, 279
- Shetty R., Kauffmann J., Schnee S., Goodman A. A., Ercolano B., 2009, *ApJ*, 696, 2234
- Simpson R. J., Nutter D., Ward-Thompson D., 2008, *MNRAS*, 391, 205
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Solomon P. M., Rivolo A. R., Barrett J., Yahil A., 1987, *ApJ*, 319, 730
- Straizys V., Cernis K., Kazlauskas A., Meistas E., 1992, *Balt. Astron.*, 1, 149
- Tachihara K., Neuhauser R., Kun M., Fukui Y., 2005, *A&A*, 437, 919
- Terebey S., Chandler C. J., André P., 1993, *ApJ*, 414, 759
- van Leeuwen F., 2007, *A&A*, 474, 653
- Ward-Thompson D., Scott P. F., Hills R. E., André P., 1994, *MNRAS*, 268, 276
- Ward-Thompson D. et al., 2007, *PASP*, 119, 855
- Ward-Thompson D. et al., 2016, *MNRAS*, 463, 1008

- Whitworth A. P., Bhattal A. S., Francis N., Watkins S. J., 1996, *MNRAS*, 283, 1061
- Wilking B. A., Gagné M., Allen L. E., 2008, in Reipurth B., ed., *Handbook of Star Forming Regions*, Vol. II. *Astron. Soc. Pac.*, San Francisco, p. 351
- Yonekura Y., Dobashi K., Mizuno A., Ogawa H., Fukui Y., 1997, *ApJS*, 110, 21 (Y97)
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