

Using late-time optical and near-infrared spectra to constrain Type Ia supernova explosion properties

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ABSTRACT

The late-time spectra of Type Ia supernovae (SNe Ia) are powerful probes of the underlying physics of their explosions. We investigate the late-time optical and near-infrared spectra of seven SNe Ia obtained at the VLT with XShooter at >200 d after explosion. At these epochs, the inner Fe-rich ejecta can be studied. We use a line-fitting analysis to determine the relative line fluxes, velocity shifts, and line widths of prominent features contributing to the spectra ([Fe II], [Ni II], and [Co III]). By focusing on [Fe II] and [Ni II] emission lines in the ~ 7000 – 7500 Å region of the spectrum, we find that the ratio of stable [Ni II] to mainly radioactively produced [Fe II] for most SNe Ia in the sample is consistent with Chandrasekhar-mass delayed-detonation explosion models, as well as sub-Chandrasekhar mass explosions that have metallicity values above solar. The mean measured Ni/Fe abundance of our sample is consistent with the solar value. The more highly ionized [Co III] emission lines are found to be more centrally located in the ejecta and have broader lines than the [Fe II] and [Ni II] features. Our analysis also strengthens previous results that SNe Ia with higher Si II velocities at maximum light preferentially display blueshifted [Fe II] 7155 Å lines at late times. Our combined results lead us to speculate that the majority of normal SN Ia explosions produce ejecta distributions that deviate significantly from spherical symmetry.

Key words: line: profiles – techniques: spectroscopic – supernovae: general.

1 INTRODUCTION

Type Ia supernovae (SNe Ia) are the violent luminous deaths of carbon–oxygen white dwarfs in binary systems. Despite being precision probes of the cosmological parameters (e.g. Betoule et al. 2014), we still do not understand the mechanisms that lead to their explosions. Proposed explosion scenarios are generally split into

two broad classes depending on the properties of their companion stars at the time of explosion. These are the ‘double-degenerate’ scenario (Iben & Tutukov 1984; Webbink 1984), where the companion star is another white dwarf and the ‘single-degenerate’ scenario (Whelan & Iben 1973), where the companion star is a main sequence or giant star. Alternatively, the competing scenarios can be framed in terms of the primary white dwarf mass at the time of explosion: those that explode near the Chandrasekhar mass limit, ‘ M_{ch} explosions’ or those that explode significantly below the Chandrasekhar mass limit, ‘sub- M_{ch} explosions’. Historically, M_{ch}

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explosions were thought to be responsible for the bulk of normal SNe Ia. However, sub- M_{ch} explosions have been the focus of a number of recent studies. These sub- M_{ch} models include both double-detonation scenarios (Nomoto 1982; Livne & Glasner 1990; Bildsten et al. 2007; Fink et al. 2010; Shen & Bildsten 2014), where a thin layer of He on the surface of the white dwarf ignites to cause a subsequent core detonation, as well as double white-dwarf mergers where the individual masses can be well below the Chandrasekhar mass (Yoon, Podsiadlowski & Rosswog 2007; Dan et al. 2011; Pakmor et al. 2012).

Observational evidence in favour (and disfavour) of both M_{ch} and sub- M_{ch} explosions has been put forward, with the current evidence suggesting that there is more than one channel for producing normal SNe Ia. This evidence comes from a number of probes, not limited to companion star interaction at early times, signatures of circumstellar material, presence of H and He features stripped from a companion star, studies of SN remnants nucleosynthetic yields (e.g. Mattila et al. 2005; Leonard 2007; Sternberg et al. 2011; Maguire et al. 2013; Seitenzahl et al. 2013b; Scalzo et al. 2014; Olling et al. 2015; Marion et al. 2016; Martínez-Rodríguez et al. 2017).

A key distinguishing prediction between M_{ch} and sub- M_{ch} explosions is the different ratios of stable to radioactive isotopes of Fe-group elements produced. This ratio depends on the central density of the white dwarf at the time of explosion (Höflich et al. 2004). Steadily accreting M_{ch} models predict significantly higher central densities ($\sim 10^9 \text{ g cm}^{-3}$) than sub- M_{ch} models where the central densities are expected to be much lower ($\sim 10^7 \text{ g cm}^{-3}$; Rosswog et al. 2009; Pakmor et al. 2012; Shen et al. 2018). Higher central densities manifest themselves as larger ratios of stable (e.g. ^{54}Fe , ^{58}Ni) to radioactively produced (^{56}Fe) isotopes due to increased electron capture in these high-density environments.

Investigating the central density requires spectra at late times when the inner regions of the ejecta are visible. This occurs at phases of $\gtrsim 200$ d after peak magnitude when the outer ejecta have become transparent and we can see the inner Fe-group-dominated regions. Detailed modelling has been performed to provide a comparison with a number of nearby SNe Ia (e.g. Axelrod 1980; Ruiz-Lapuente & Lucy 1992; Spyromilio et al. 1992; Liu, Jeffery & Schultz 1997; Mazzali et al. 2007, 2015; Fransson & Jerkstrand 2015; Botyánszki & Kasen 2017; Wilk, Hillier & Dessart 2018). However, there are many differences between the models that depend on the inputs to the code, as well as in the details of the radiative transfer codes themselves. A good match to the shape and strength of the Ni and Fe complex at $\sim 7300 \text{ \AA}$ is difficult to achieve, which results in different implied masses of neutron-rich stable Ni material. Some predicted the presence of significant quantities of stable Ni for SN 2011fe (Mazzali et al. 2015), while others suggested for the same SN that stable Ni is not required (Botyánszki & Kasen 2017). Therefore, we have adopted an agnostic approach to determining the abundance ratios using empirical, model-independent methods. However, we stress that radiative transfer modelling of SNe Ia at late times is vital for providing detailed constraints on ejecta abundances and understanding SN Ia explosion physics.

The late-time spectra of SNe Ia have also provided information on potential asymmetries in the ejecta (Motohara et al. 2006; Gerardy et al. 2007; Maeda et al. 2010b). A link between maximum light Si II velocity gradients and the wavelength shifts of late-time emission lines was identified, with SNe Ia with high Si II velocities having preferentially redshifted late-time emission lines (Maeda et al. 2010a). It was suggested that variations in the observed properties (light-curve widths, line velocities) are a consequence of different viewing angles of an explosion with an off-centre ignition point.

Table 1. Observational details for previously unpublished late-time spectra observed at the VLT+XShooter.

SN	Observation date	MJD ^a	Phase ^b (d)	Exposure time (s)
SN2013aa	2014 Apr 21	56769.2	425	7480
ASASSN14jg	2015 Sept 19	57285.2	322	2900
ASASSN15be	2015 Oct 21	57317.3	266	5800
PSNJ1149	2016 Feb 02	57421.3	206	2900
ASASSN15hx	2016 June 26	57566.1	427	2980

^aMJD = Modified Julian date. ^bPhase of late-time spectrum calculated with respect to maximum light.

A tentative link between the late-time velocity shifts and SN Ia maximum-light colours has also been studied (Cartier et al. 2011; Maeda et al. 2011; Blondin et al. 2012).

In this paper, we investigate the largest sample to date of optical through near-infrared (NIR) spectral observations of SNe Ia at phases later than ~ 200 d. At these epochs, the inner Fe-rich regions of the ejecta can be used to investigate the extent of line-emitting regions, signatures of asymmetry, and the relative abundances of stable to radioactive elements. In Section 2, we present the VLT+XShooter observations, complementary maximum-light spectral, and photometric observations, as well as details of a comparison literature sample. In Section 3, the line-fitting analysis is described, combined with specific details of fitting for the prominent spectral features. The measurements of the late-time velocity shifts, the colour evolution, and calculations of the Ni to Fe abundance ratio are detailed in Sections 4, 5, and 6, respectively. A discussion of the results is presented in Section 7, and the conclusions are given in Section 8.

2 OBSERVATIONS AND DATA REDUCTION

Our analysis is performed on a sample of eight spectra of seven SNe Ia with high-signal-to-noise (S/N) optical and NIR spectral coverage at epochs greater than 200 d after maximum light. We use six spectra from the VLT+XShooter SN Ia sample presented in Maguire et al. (2016), combined with two new XShooter SN Ia spectra. Spectra of three additional SNe Ia obtained with XShooter are also presented but do not have high enough S/N in the NIR to be analysed in detail.

2.1 Late-time spectroscopy with XShooter

The spectra in our sample were obtained using the VLT+XShooter, which is an echelle spectrograph with three arms (UVB, VIS, NIR) covering the wavelength range $\sim 3000\text{--}25\,000 \text{ \AA}$ (Vernet et al. 2011). We used a setup with slit widths of 0.8, 0.9, and 0.9 arcsec in the UVB, VIS, and NIR arms, respectively. The spectra were reduced using the REFLEX pipeline to produce flux-calibrated one-dimensional spectra (Modigliani et al. 2010; Freudling et al. 2013) as described in Maguire et al. (2016).

Five late-time spectra of SNe Ia are presented for the first time here, ASASSN14jg, ASASSN15be, PSNJ11492548–0507138 (hereafter PSNJ1149), ASASSN15hx, and a second epoch of SN 2013aa (Table 1). These spectra were reduced in the same manner as in Maguire et al. (2016), and flux calibrated using photometric measurements from the acquisition images. For PSNJ1149 and SN 2013aa, the acquisition image photometry was calibrated using the XShooter acquisition camera zero-points. The spectra of ASASSN15hx were flux calibrated to coeval photometry using g'

Table 2. SN light curve, spectral, and host galaxy information.

Name	Phase ^a (d)	Host galaxy	Helio. z^b	$E(B-V)^c$ (mag)	Stretch ^d	Date of max. ^e	$B - V$ at max. (mag)	Si II vel. at max. (km s ⁻¹)	LC Ref. ^f	Spec. ref. ^f
SN 2012cg	339	NGC 4424	0.0015	0.20 ^g	1.098 ± 0.022	2012 June 03	0.14 ± 0.04	10420 ± 100	1	1
SN 2012fr	357	NGC 1365	0.00546	0.018	1.135 ± 0.010	2012 Nov 12	0.02 ± 0.03	11928 ± 100	2,3	4
SN 2012ht	433	NGC 3447a	0.00355	0.026	0.877 ± 0.022	2013 Jan 03	-0.04 ± 0.02	11000 ± 200	1	1
SN 2013aa	360,425	NGC 5643	0.003999	0.169	1.146 ± 0.019	2013 Feb 21	-0.05 ± 0.01	10600 ± 100	1	1
SN 2013cs	303	ESO 576- G 017	0.00924	0.082	0.963 ± 0.024	2013 May 26	0.03 ± 0.05	12700 ± 150	5	6
SN 2013ct	229	NGC 0428	0.00384	0.025	–	2013 May 04	–	–	–	–
PSNJ11492548	206	NGC 3915	0.0056	0.025	–	2015 July 12	–	10386 ± 100	–	4
ASASSN14jg ^h	322	PGC 128348	0.014827	0.014	1.053 ± 0.007 ⁱ	2014 Nov 02	–	10928 ± 200	7	4
ASASSN15be ^h	266	GALEX J0252 ^j	0.0219	0.017	–	2015 Jan 29	–	11900 ± 200	–	4
ASASSN15hx ^h	427	GALEX J1343 ^j	0.0084	0.043	–	2015 May 08	–	10407 ± 200	–	4

^aPhase of late-time spectrum calculated with respect to maximum light.

^bHeliocentric redshifts are from the Nasa Extragalactic Database (NED), apart from ASASSN15hx, where the redshift is measured from narrow H α emission from a WiFES (Childress et al. 2016) spectrum on 2015 July 08, and SN 2012cg from Cortés, Kenney & Hardy (2006).

^cGalactic $E(B-V)$ values from Schlafly & Finkbeiner (2011).

^dStretch is a measure of the light-curve width (Conley et al. 2008).

^eMaximum light estimated from a light curve where available or estimated from the early-time spectra if not.

^fSources for maximum-light photometry and Si II 6355 Å velocities: (1) Maguire et al. (2013) (2) Silverman et al. (2015), (3) Zhang et al. (2014) (4) Childress et al. (2013b), (5) this paper, (6) Childress et al. (2016), (7) Graham et al. (2017).

^gAdditional host galaxy extinction of $E(B-V) = 0.18$ mag was found for SN 2012cg by Silverman et al. (2012). The value quoted here is the combined Galactic and host galaxy $E(B-V)$.

^hNIR region of the spectrum is too low S/N for a detailed analysis to be performed. Used only in comparison with literature sample for line shift analysis of Maeda et al. (2010a).

ⁱConverted from a $\Delta m_{15}(B)$ value to a SiFTO stretch using the relation from Conley et al. (2008).

^jFull name of galaxies: GALEX J0252 = GALEXASC J025245.83–341850.6, GALEX J1343 = GALEXASC J134316.80–313318.2.

acquisition imaging calibrated by comparison to catalogue Pan-STARRS¹ magnitudes (Chambers et al. 2016; Magnier et al. 2016). Due to the faintness of ASASSN14jg and ASASSN15be in their acquisition images, these spectra remain without absolute flux calibration. However, since we deal only with relative line fluxes this does not affect any of our conclusions.

Telluric corrections have been applied to the spectra using the MOLECFIT package (Kausch et al. 2015; Smette et al. 2015). For the spectra of ASASSN15be and ASASSN15hx, the low S/N did not allow for accurate fitting of the telluric regions. The spectra were corrected to their rest wavelengths using their heliocentric velocities and for Galactic extinction using the Galactic $E(B - V)$ values given in Table 2.

Of the five new late-time spectra, two (PSNJ1149, SN 2013aa) are included in the full analysis, while the spectra of ASASSN14jg, ASASSN15be, and ASASSN15hx have lower S/N spectra, particularly in the NIR region of the spectrum. Therefore, they are excluded from the full analysis and only used in the comparison line-shift analysis presented in Section 4. The optical and NIR spectra for the sample used in the detailed analysis are shown in Fig. 1. The spectra are available from WISEREP, <https://wiserep.weizmann.ac.il/> (Yaron & Gal-Yam 2012).

2.2 Maximum-light spectroscopy

Maximum-light spectroscopy (within ± 5 d of max.) has been used to constrain the Si II 6355 Å velocities at maximum for our late-time spectral sample. For SNe 2012cg, 2012fr, 2012ht, 2013aa, and 2013cs, these values were obtained from the literature as detailed in Table 2. This is supplemented with unpublished data for four other SNe Ia in the sample, described below.

For PSNJ1149, a maximum-light classification spectrum was obtained with the Shane 3-m telescope of the Lick Observatory using the Visible and Near-Infrared Imaging Spectrograph (Rudy et al. 2015), from which an Si II velocity at maximum was measured. ASASSN14jg was spectroscopically classified at the 2-m Las Cumbres Observatory (LCO) telescope (Brown et al. 2013) with the FLOYDS spectrograph on 2014 Nov 02 (at maximum) and from which an Si II velocity was measured. For ASASSN15be, the Public ESO Spectroscopic Survey of Transient Objects (PESSTO; Smartt et al. 2015) obtained a spectrum on 2015 Jan 26 with EFOSC2 on ESO's New Technology Telescope. An Si II velocity was measured from this spectrum that is consistent with a phase of -3 d with respect to maximum light. ASASSN15hx was also spectroscopically observed at the 2-m LCO telescope with FLOYDS on 2015 May 05 and 2015 June 11 (-3 and $+3$ d from maximum). The Si II velocity at maximum was taken as the mean of these two values. These Si II 6355 Å velocities are provided in Table 2.

2.3 Photometry and light-curve analysis

The light-curve properties (light-curve width, optical colour at maximum) for four SNe Ia in our sample were obtained from the literature (see references in Table 2). A *gri*-band light curve of SN 2013cs obtained with the LCO 1-m telescope array is presented here for the first time (Table A1). The magnitude of the SN was measured using point spread function photometry and calibrated using reference stars in the field of the SN with magnitudes from the Pan-STARRS Data Release 1 catalogue.² The light-curve width (stretch) of SN 2013cs was calculated using the SiFTO light-curve fitting code (Conley et al. 2008) to produce a stretch value and the $B - V$ colour at maximum. No early-time photometric measurements are available for the five remaining objects in our sample.

¹ <http://archive.stsci.edu/panstarrs/search.php>

² <http://archive.stsci.edu/panstarrs/search.php>

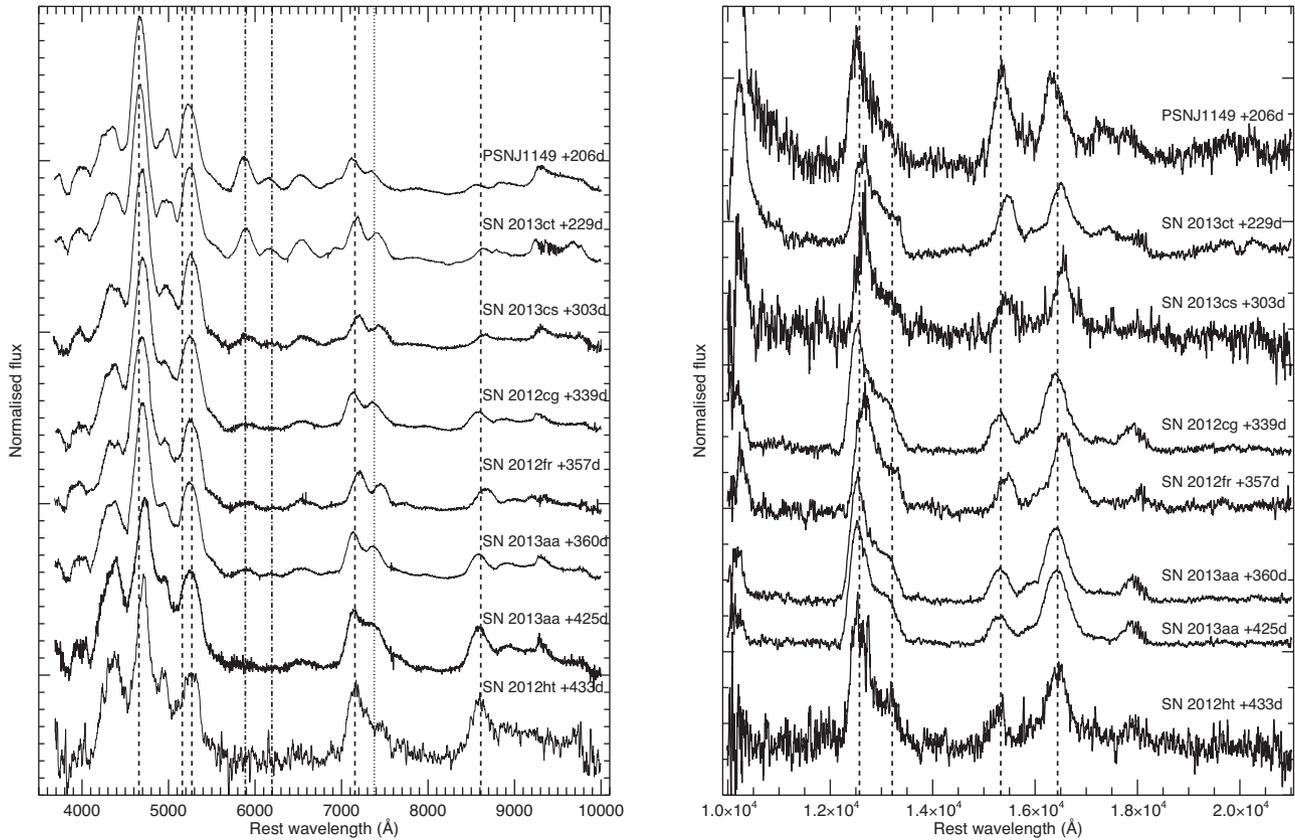


Figure 1. Optical (left) and NIR (right) spectral data of our full sample of SNe Ia at epochs greater than 200 d since maximum light. The rest wavelengths of prominent Fe-group lines are marked by the dashed lines (from blue to red wavelengths: [Fe III] 4658 Å, [Fe II] 5159 Å, [Fe III] 5270 Å, [Fe II] 7155 Å, [Fe II] 8617 Å, [Fe II] 12 567 Å, [Fe II] 13 206 Å, [Fe II] 15 335 Å and [Fe II] 16 436 Å), the [Ni II] 7378 Å line by a dashed line, and [Co III] 5890, 6197 Å by dash-dotted lines.

2.4 Literature sample

To perform a comparison with literature spectra, we have compiled a sample of late-time spectra with phases that match those of our XShooter data (~ 200 – 430 d). The late-time spectra come from literature spectral compilations (Maeda et al. 2010a; Blondin et al. 2012; Silverman, Ganeshalingam & Filippenko 2013; Childress et al. 2015), along with additional spectra of SN 2010ev (Gutiérrez et al. 2016) and SN 2011fe (Graham et al. 2015; Mazzali et al. 2015). Si II 6355 Å velocities within ± 5 d of maximum were also compiled from the literature. The majority of the Si II velocities are from spectra from Blondin et al. (2012), supplemented with values for individual objects: SN 2003gs (Suntzeff, Candia & Stritzinger 2003), SN 2010ev (Gutiérrez et al. 2016), SN 2011by (Silverman et al. 2013), SN 2011fe (Maguire et al. 2014), SN 2013dy (Pan et al. 2015), and SN 2014J (Ashall et al. 2014).

3 LINE-FITTING ANALYSIS OF LATE-TIME SPECTRA

The late-time optical and NIR spectra of SNe Ia are dominated by emission lines of Fe-group elements, [Fe III], [Fe II], [Co III], and [Ni II] lines (e.g. Axelrod 1980; Meyerott 1980). Our aim was to use a line-fitting code on individual Fe-group spectral regions to determine the contribution of different elements to the main emission line groups seen in the optical and NIR spectra of a sample of SNe Ia. We were interested in three regions of their optical and NIR spectra: (i) the 12 600 Å region that is dominated by [Fe II]

lines, (ii) the 7300 Å region that has contributions from [Fe II] and [Ni II], as well as a potential contribution from [Ca II], and (iii) the 6000 Å region, which is dominated by [Co III] emission. The lines identified to be important in the 7300 Å region are the same as for core-collapse SN case studied in Jerkstrand et al. (2015a), and we follow a similar analysis procedure, detailed below.

For [Fe II] and [Ni II], the line-fitting code uses the latest atomic data (e.g. rest wavelengths, transition probabilities) from the CMFGEN compilation (Hillier & Miller 1998; Hillier & Dessart 2012), which includes the forbidden line data from Quinet & Le Dourneuf (1996) and Quinet, Le Dourneuf & Zeippen (1996). For [Co III], we use the atomic data from Tyndall et al. (2016).

Using the Sobolev formula and reasonable numbers for the densities and velocity gradients, we found that an optically thin approximation for the relative line strengths is appropriate in this wavelength range. Assuming local thermodynamic equilibrium (LTE), the relative line strengths were calculated from temperatures in the range of 3000–8000 K, which is the expected range for SNe Ia at these phases (Fransson & Jerkstrand 2015). The values used in the fits are the mean values of the line strengths in this temperature range. Line strengths are calculated relative to the strongest feature in a wavelength region. Lines that have a strength of < 5 per cent of the strongest line in the region are excluded since they have a negligible influence on the fit. For the three main spectral regions discussed, the strongest two lines in each case come from the same upper energy level, and therefore their relative strengths are temperature independent.

Table 3. Best-fitting values for the line-fitting analysis.

Name	Phase (d)	Line (Å)	Velocity shift (km s ⁻¹)	FWHM (km s ⁻¹)	pEW (Å)	Name	Phase (d)	Line (Å)	Velocity shift (km s ⁻¹)	FWHM (km s ⁻¹)	pEW (Å)
PSNJ1149	+206	[Fe II] 12 567	-1272 ± 478	7112 ± 475	2383 ± 780	SN2012fr	+357	[Fe II] 12 567	1091 ± 691	7403 ± 667	9859 ± 3017
		[Fe II] 7155	-1271 ± 267	8748 ± 332	618 ± 61			[Fe II] 7155	1865 ± 205	7498 ± 80	661 ± 28
		[Ni II] 7378	-2074 ± 437	5063 ± 926	157 ± 23			[Ni II] 7378	2063 ± 222	6163 ± 46	217 ± 14
		[Co III] 5 890	-944 ± 201	9346 ± 42	1135 ± 31			[Co III] 5890	103 ± 267	11405 ± 265	360 ± 32
SN2013ct	+229	[Fe II] 12 567	834 ± 267	8020 ± 164	1941 ± 98	SN2013aa	+360	[Fe II] 12 567	-1408 ± 323	7211 ± 162	6234 ± 1601
		[Fe II] 7155	765 ± 200	7343 ± 55	417 ± 7			[Fe II] 7155	-771 ± 206	7349 ± 105	642 ± 33
		[Ni II] 7378	845 ± 206	6674 ± 25	256 ± 5			[Ni II] 7378	-1535 ± 266	7933 ± 25	366 ± 10
		[Co III] 5 890	35 ± 216	8545 ± 175	789 ± 41			[Co III] 5890	-222 ± 258	9671 ± 336	250 ± 23
SN2013cs	+303	[Fe II] 12 567	2060 ± 770	7135 ± 213	2798 ± 654	SN2013aa	+425	[Fe II] 12 567	-1189 ± 637	7781 ± 900	9763 ± 3504
		[Fe II] 7155	1656 ± 207	7122 ± 50	406 ± 5			[Fe II] 7155	-385 ± 244	8495 ± 210	738 ± 46
		[Ni II] 7378	1496 ± 207	6765 ± 25	236 ± 2			[Ni II] 7378	-1613 ± 395	7718 ± 594	329 ± 25
		[Co III] 5 890	-254 ± 216	11098 ± 328	1493 ± 271			[Co III] 5890	-	-	-
SN2012cg	+339	[Fe II] 12 567	-1203 ± 422	8001 ± 1857	6770 ± 1211	SN2012ht	+433	[Fe II] 12 567	-1179 ± 469	7497 ± 1270	3830 ± 1444
		[Fe II] 7155	-981 ± 201	7474 ± 47	474 ± 11			[Fe II] 7155	494 ± 381	8130 ± 693	893 ± 161
		[Ni II] 7378	-1381 ± 212	8801 ± 30	309 ± 3			[Ni II] 7378	424 ± 2628	6783 ± 2826	174 ± 34
		[Co III] 5 890	-65 ± 256	11906 ± 192	264 ± 16			[Co III] 5890	-	-	-

3.1 Fitting method

Gaussian profiles are used to fit the emission lines of the Fe-group elements since the emitting regions of these features are expected to be centrally concentrated. The fitting of the emission features was performed in the following way. Continuum points to the red and the blue of the spectral region under consideration were chosen interactively. These are difficult to set automatically because the velocity offsets are found to vary between objects, as well as due to changes in the background levels depending on the epoch of the spectrum. The relative strengths of lines of the same element and ionization state were constrained using the radiative transition rates and the assumption of LTE level populations, as detailed above. The widths and velocities (relative to the rest wavelength) of features of the same element and ionization state in the same spectral region were set to the same value since they are expected to be produced by the same regions. Loose priors were placed on some values as described below for the individual spectral regions.

The uncertainties in the fitting regimes were calculated using a Monte Carlo approach. The largest uncertainties arise from the placement of the continuum points to the red and the blue of the features under investigation. To produce an estimate of the uncertainties from this, the interactively chosen continuum values were varied by random amounts up to 10 Å to the blue and the red of this value and the fitting was then repeated. This was performed 3000 times to produce mean fit values and relative uncertainties. An additional uncertainty on the velocity measurements of 200 km s⁻¹ was added to account for peculiar velocity effects of the host galaxies. The best-fitting values for the spectral regions studied, in terms of the velocity offset, full width at half-maximum (FWHM), and pseudo-equivalent width (pEW), are presented in Table 3.

The ~4000–5700 Å region shows the strongest Fe emission features. However, it is also the most complex with contributions from both [Fe II] and [Fe III] lines that are difficult to disentangle. The ratio of the strongest [Fe III] lines (4658, 5270, 5011 Å) are also sensitive to temperature changes. Previous detailed spectral modelling of this region has failed to accurately model it, despite reasonably good agreement at other wavelengths (Maurer et al. 2011; Fransson & Jerkstrand 2015; Mazzali et al. 2015). We attempted to fit this region but the fits were in general not good. Therefore, we leave further analysis of this region to the detailed simulations of Shingles et al. (in preparation). For comparison with other works (e.g. Black, Fesen & Parrent 2016), we used the position of the peak as a

measure of the velocity of the 4700 Å feature (using a fit with just two lines with rest wavelengths of 4658 and 4702 Å).

We also excluded the ~15 000–18 000 Å region from detailed analysis due to the presence of [Co III] and [Co II] lines in this region, as well as because the strongest [Fe II] features in this region do come from the same upper energy levels (unlike the case in the other regions studied). This wavelength region is the focus of Flörs et al. (in preparation).

3.2 12 600 Å [Fe II]-dominated region fitting

The 12 600 Å region is dominated by the [Fe II] 12 567 Å feature, along with six other weaker [Fe II] lines (12 704, 12 788, 12 943, 12 978, 13 206, 13 278 Å) that have <30 per cent the strength of the 12 567 Å line. No lines of other elements or other ionization states of Fe are predicted to be present in this region, making it an excellent choice for constraining the properties of the [Fe II] emitting region. The second strongest line at 13 206 Å comes from the same energy level as the strongest line, so their ratio is temperature independent. Assuming LTE, temperature-induced changes in the relative line strength of the five weakest lines are of the order of ~5 per cent, resulting in variations in the strength of the strongest 12 567 Å lines of ~5 per cent, less than the uncertainty on the fit.

As we have constrained, the relative strengths from the atomic data, and the velocity shifts and widths of the [Fe II] features are tied, there are just four free parameters for the fit to the region (strength, width, velocity offset, and the height of the continuum). No priors are placed on these values. The best fits to this 12 600 Å [Fe II]-dominated region are shown in Fig. 2.

3.3 7300 Å [Fe II]/[Ni II]-dominated region

The 7300 Å line complex in late-time SN Ia spectra is made up of lines of [Fe II] (7155, 7172, 7388, 7453 Å), [Ni II] (7378, 7412 Å), and possibly [Ca II] (7291, 7324 Å). Different explosion mechanisms predict different contributions of these elements, depending on the presence of stable Fe-group material and the completeness of burning in the explosion.

The relative strengths of the two strongest lines of the [Fe II] lines are temperature independent. There is a small temperature dependence of the two weaker lines relative to the strongest line, [Fe II] 7155 Å, but changes in temperature in the range 3000–8000 K

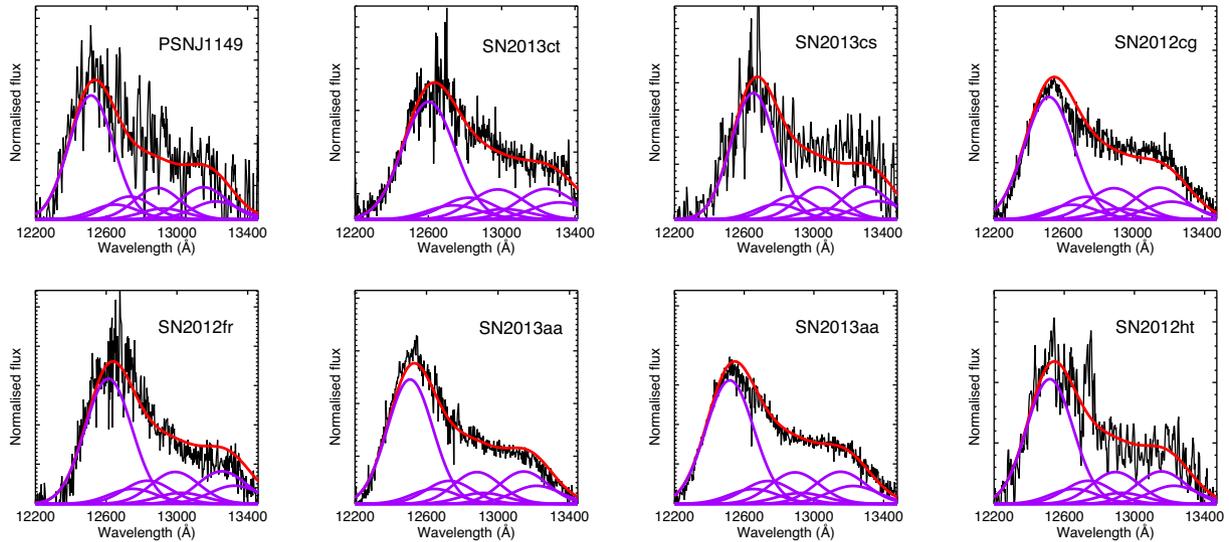


Figure 2. Best fits to the [Fe II] 12600 Å region for the XShooter sample. The observed spectra are shown in black, the overall fits are shown in red, and the individual [Fe II] features are shown in purple.

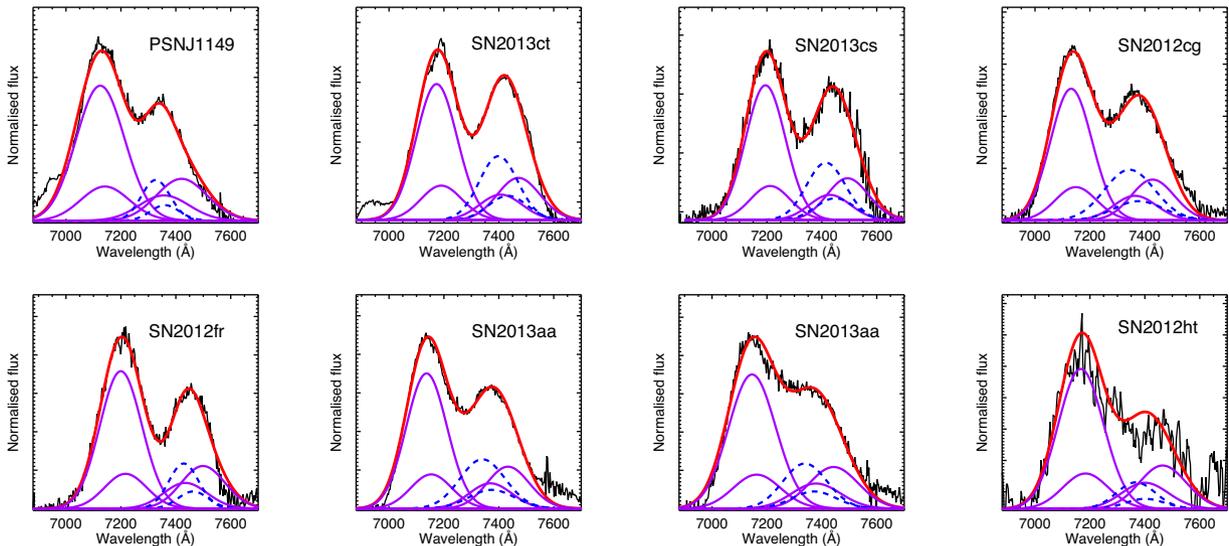


Figure 3. Best fits to the 7300 Å region containing [Fe II] and [Ni II] features. The observed spectra are shown in black, the overall fits are shown in red, the [Fe II] features are shown as solid purple lines, and the [Ni II] features are shown as blue dashed lines.

results in a change in the relative strength of the strongest line of <5 per cent. These [Fe II] lines are expected to be dominated by emission of the decay of radioactive ^{56}Co . A smaller contribution from stable primordial Fe (^{54}Fe , ^{56}Fe) may be present but it is impossible to disentangle these contributions. The initial guess for the velocity and width of [Fe II] lines in the 7300 Å complex were set as the output of the [Fe II] 12600 Å region but were allowed to vary within ± 5000 km s $^{-1}$ of the NIR [Fe II] velocity and width values (much greater than the velocity shifts or the range of widths observed).

Since ^{56}Ni decays with a half-life of 6.1 d, the contribution from radioactive Ni at late times is negligible. Therefore, any Ni emission lines seen in late-time spectra must come from stable ^{58}Ni (and to a much smaller extent from stable ^{60}Ni). For example, the mass ratio of ^{60}Ni to ^{58}Ni is $\lesssim 10$ per cent for the M_{ch} delayed-detonation (DDT) models of Seitzzahl et al. (2013a), so we refer to ^{58}Ni to

mean the total stable Ni contribution. Two [Ni II] features (7378, 7412 Å) are expected to emit in this region if sufficient stable Ni is present. No priors were placed on the velocity offset of the [Ni II] lines. Different explosion models predict different extents of the stable Fe-group element emitting regions. W7 predicts that the width of the stable Fe-group element emitting region should be <6000 km s $^{-1}$ (Iwamoto et al. 1999), while sub- M_{ch} models prefer values of <7000 km s $^{-1}$ (Sim et al. 2010). To be agnostic as to the explosion mechanism, the width of the [Ni II] was left to vary between 2000 km s $^{-1}$ and the width of the radioactive-chain-dominated [Fe II] 12567 Å line (+10 per cent), which is typically of the order of 7000–9000 km s $^{-1}$.

Fig. 3 shows the best fit to the 7300 Å [Fe II]/[Ni II] region under the assumptions and constraints detailed above. Since constraining the contribution of [Ni II] to the region was one of our primary aims, we repeated the fitting removing the two [Ni II] features

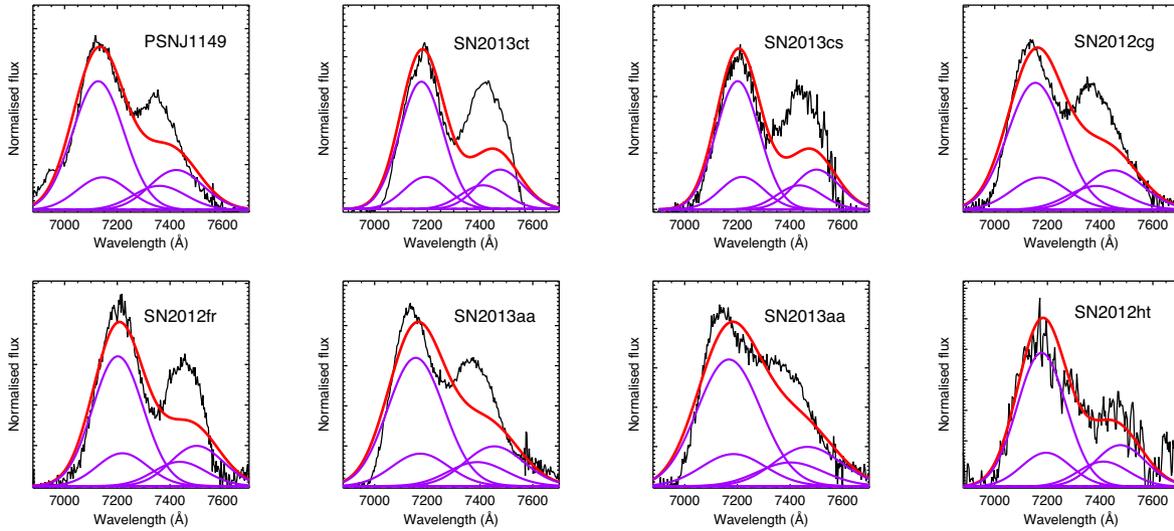


Figure 4. Best fits to the 7300 Å region with no [Ni II] included in the fit (just [Fe II]). The observed spectra are shown in black, the overall fits are shown in red, and the [Fe II] features are shown in purple.

from the fit and obtained the results shown in Fig. 4. It can be seen that the fits are significantly worse for all SNe apart from SN 2012ht (this is because it contained only a small contribution of [Ni II] in the best fit when the presence of [Ni II] was allowed and is also the lowest S/N of our spectral sample). Therefore, for six out of seven SNe in our sample, a non-negligible contribution from [Ni II] emission is required to produce a reasonable fit to the region.

[Ca II] features (7291, 7324 Å) could also be present in this wavelength region. The Ca-emitting regions are expected to be further out in the ejecta than the Fe-group elements with a shell-like structure for most explosion scenarios for normal luminosity SNe Ia that are studied in this work. We approximated this as a shell with a width in the range 5000–15 000 km s⁻¹ covering the range predicted by Chandrasekhar and sub-Chandrasekhar mass explosions. We used a boxy profile with tapered edges with the flux for a given wavelength of the form, $F_\lambda = 1 - \Delta v^2$, where Δv is the velocity shift from the rest velocity (Jerkstrand et al. 2017). We tested the contribution of the [Ca II] lines to the fit by performing the fits again with a contribution from [Ca II]. To test if the fit is strengthened by including [Ca II], we used the Bayesian Information Criterion (BIC; Schwarz 1978) to select between models with and without a contribution from [Ca II]. The BIC penalizes models with a greater number of parameters. The model with the lower information criterion score is considered to be the better model for explaining a given data set. We measured the BIC for models with and without a contribution from Ca II and find ΔBIC values of $\gg 10$ in favour of the no Ca model, providing strong evidence against the need to include [Ca II] in the fit to this region.

If we assume instead that [Ca II] emission was responsible for the ‘[Ni II]-feature’, the redder peak in this region, and that it was allowed to have a Gaussian instead of boxy profile (although this is not expected for the main explosion models), then the velocity offsets to fit the ‘[Ni II]-feature’ would be shifted to the red by a further ~ 3500 km s⁻¹. For the most extreme redshifted case of SN 2012fr, this feature would have a velocity offset from the rest wavelength of [Ca II] of ~ 5500 km s⁻¹. When the red peak is assumed to be dominated by [Ni II], then the distribution of velocity offsets is centred on zero with three of the XShooter sample offsets lying to the blue and five to the red. If the origin of

this feature was [Ca II] emission, then all the velocity shifts of the sample would be shifted strongly towards the red, again arguing against a centrally peaked Ca distribution being the origin of this feature.

3.4 6000 Å [Co III]-dominated region

The ~ 5700 – 6300 Å region displays two peaks dominated by [Co III] emission that radioactively decays with time from ~ 60 d after explosion (Childress et al. 2015), and are undetectable after ~ 400 d (Fig. 5). Since there are no non-negligible stable Co isotopes, the [Co III] lines trace regions rich in radioactive material. Along with composition changes, the regions emitting [Co III] compared to [Fe II] and [Ni II] lines are also likely to differ in their thermodynamic conditions since they result in the production of different ionization states. The strengths of the [Co III] lines (5890, 5908, 6129, 6197 Å) are constrained relative to the strongest feature. The second strongest line at 6197 Å has a fixed ratio to the strongest feature because it comes from the same upper energy level. The weakest two lines have a small temperature dependence in the 3000–8000 K regime, causing a ~ 6 per cent uncertainty in the relative strength of the strongest line. No constraints were placed on the strength of the strongest line, its velocity offset, or width. In agreement with Dessart et al. (2014), we find that this spectral region is well fitted without a contribution from Na I D.

4 LATE-TIME VELOCITY SHIFTS AND LINE WIDTHS

An asymmetric distribution of Fe in SNe Ia was first suggested by Spyromilio et al. (2004) by studying the shapes of NIR [Fe II] lines in late-time spectra. Velocity shifts of the central wavelength of the [Fe II] 16 440 Å feature with respect to its rest wavelength were identified by Motohara et al. (2006) for three SNe Ia, 2003du, 2003hv, and 2005W. In SN 2003hv, a shift was observed for [Co III] at mid-infrared wavelengths (Gerardy et al. 2007) and [Fe II] at optical wavelengths (Leloudas et al. 2009). Optical line shifts have also been identified for the [Fe II] 7155 Å and [Ni II] 7378 Å features with peak velocity shifts from rest of ± 3000 km s⁻¹ that correlate with Si II velocities measured at maximum light (Maeda et al. 2010a,b;

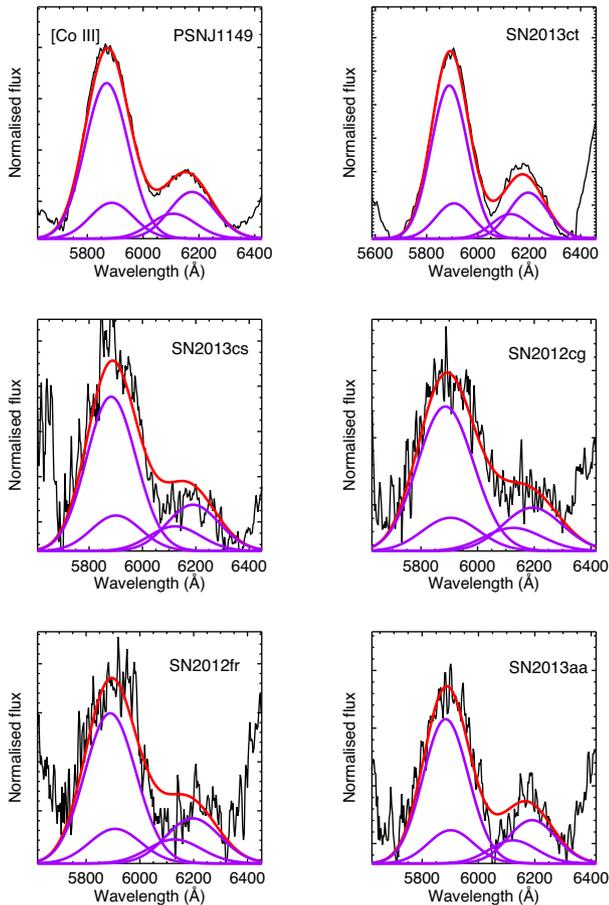


Figure 5. Best fits to the 6000 Å region for the four [Co III] lines. The observed spectra are shown in black, the overall fits are shown in red, and the four individual [Co III] profiles are shown in purple.

Blondin et al. 2012; Silverman et al. 2013). Here, we discuss the velocity shifts, as well as measurements of the line widths, for our sample of late-time spectra, and also perform a comparison with literature objects.

4.1 Velocity shift evolution with phase

In the top panel of Fig. 6, the velocity shifts of the 7155, 12567 Å [Fe II] lines, as well as those of [Ni II] 7378 Å and [Co III] 5890 Å, are shown as a function of phase for the eight spectra of the seven SNe Ia in our sample. No evolution with time is identified for our sample and the [Fe II] and [Ni II] features show large positive and negative offsets that do not appear phase dependent. Black et al. (2016) investigated the velocity offset with time for the [Co III] 5890 Å region for a sample of SNe Ia with multiple spectra per object and found a flat trend (within the uncertainties) of their velocities in the range 200–350 d post-maximum. In the bottom panel of the same plot, the peak of the [Fe III] 4700 Å is shown as a function of phase. This peak value was not measured from a full fit to the region but from a simple Gaussian fit to the strongest 4700 Å peak. A clear trend of increasing redshift with time is identified for the [Fe III] 4700 Å feature. This is in agreement with previous studies, where the wavelength of the peak of the feature was found to shift towards the red with time from ~50 to 450 d past maximum light (Maeda et al. 2010b; Black et al. 2016; Graham et al. 2017). This could be due to contributions from other features in this region, such

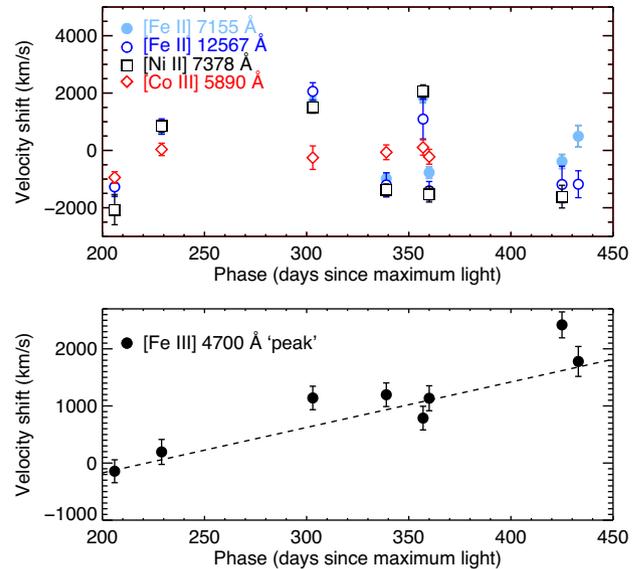


Figure 6. Velocity shift with respect to their rest wavelength versus phase for the [Fe II] 7155 Å (pale blue solid circles), [Fe II] 12567 Å (blue open circles), [Ni II] 7378 Å (black squares), [Co III] 5890 Å (red squares) in the top panel, and [Fe III] 4700 Å features in the bottom panel. The [Fe II] 7155 Å, [Fe II] 12567 Å, and [Co III] 5890 Å velocities are from full fitting to the region, while the [Fe III] 4700 Å values are from a ‘peak’ fit to the strongest emission feature in this region. The best-fitting line to the [Fe III] is shown, demonstrating a shift redward with time of this feature, as has been previously identified.

as permitted Fe and [Co II] lines. A redshifting with time of line peaks has previously been observed in a number of core-collapse SNe (e.g. Taubenberger et al. 2009), and has been suggested to be caused by a decrease in opacity with time, which mainly affects bluer wavelengths (Jerkstrand et al. 2015b).

4.2 Velocity offsets and widths of Fe-group elements

We wished to investigate in our sample the relative velocity shifts of different Fe-group elements (and ionization states) to determine if their emitting regions are the same. Previous studies have suggested that [Ni II], [Fe II], and [Fe III] features are produced in different regions of the ejecta that depend on the density (Maeda et al. 2010b). As shown in Section 4.1, the [Fe III] 4700 Å feature is not suitable for this study because its central position evolves with time.

In Fig. 7, we show the results of our comparative study of the velocity and widths of singly and doubly ionized elements. We have compared the velocities of the [Co III] 5890 Å line to those of the [Ni II] 7378 Å, [Fe II] 7155, 12567 Å features. We find that the mean absolute offset from zero velocity ($\langle |v| \rangle$) for [Co III] is much smaller than that for the lower ionization lines of [Fe II] and [Ni II]. The [Fe II] 7155 and 12567 Å features have mean absolute velocity offsets from zero of 1024 ± 533 and 1279 ± 355 km s⁻¹, respectively. The [Ni II] 7378 Å feature has a mean absolute velocity offset of 1429 ± 563 km s⁻¹ that is slightly higher than the [Fe II] features but consistent within the uncertainties. However, the [Co III] 5890 Å feature has a significantly lower mean absolute velocity offset of 271 ± 341 km s⁻¹, which is consistent with zero offset from its rest wavelength.

The ~15 000–8 000 Å region is also dominated by [Fe II] features. As discussed in Section 3, we have excluded this region from our detailed analysis because of potential contributions from [Co II]

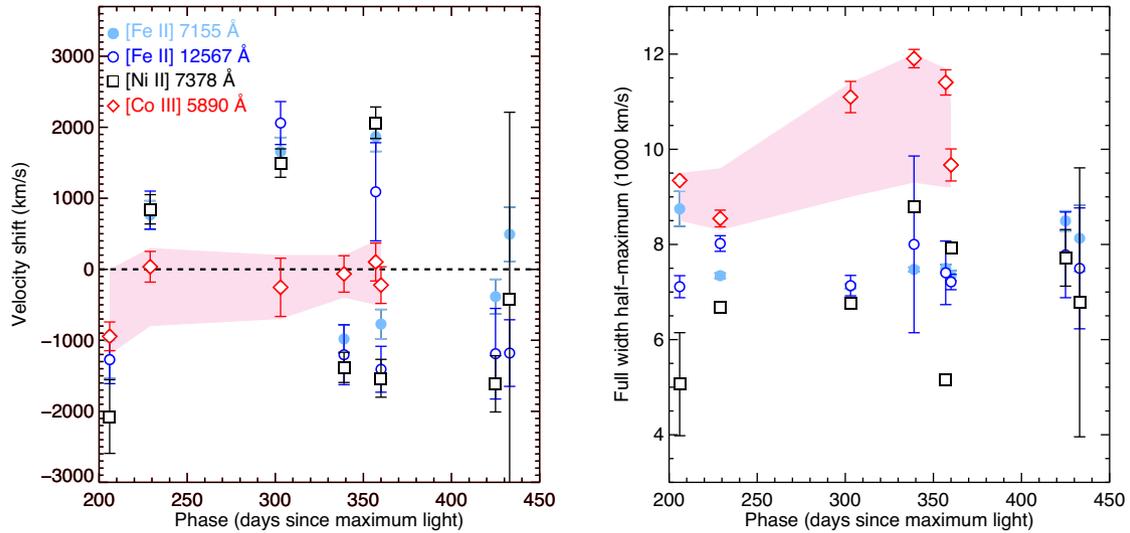


Figure 7. Left: velocity shifts of the [Fe II] 7155 Å, [Fe II] 12567 Å [Co III] 5890 Å, and [Ni II] 7378 Å features as a function of time. The pink shaded region highlights the smaller relative velocities for the [Co III] lines. Right: FWHM of the [Fe II] 7155 Å, [Fe II] 12567 Å [Co III] 5890 Å, and [Ni II] 7378 Å features as a function of time. The pink shaded region highlights the larger FWHM values for the [Co III] lines.

and [Co III] that makes the fitting of the [Fe II] features less clean. However, if we assume that the main peak in this region is dominated by emission from the [Fe II] 16440 Å line, then it can be seen from Fig. 11 that its peak also shifted in a similar way to the peak of the 12600 Å region. The peaks of the 12600 and 16400 Å regions are shifted consistently to the red for SNe 2013ct, 2013cs, and 2012fr. This confirms our previous results that the distribution of [Fe II] emission lines is shifted in SNe Ia both positively and negatively with respect to the rest wavelengths of the lines.

We have also investigated the extent of the emitting regions of these Fe-group features using their FWHM (Fig. 7, right-hand panel). We found that the [Co III] lines (mean FWHM of $10\,300 \pm 1300$ km s⁻¹) are significantly broader than the [Fe II] 7155 Å, [Fe II] 12567 Å, and [Ni II] 7378 Å lines, with mean FWHM of 7800 ± 600 , 7500 ± 400 , and 6900 ± 1300 km s⁻¹, respectively.

4.3 Late-time velocity shifts of [Fe II] and the connection to maximum-light Si II velocities

A connection between the average peak line shift of the two main components of the 7300 Å region (suggested to be due to [Fe II] 7155 Å and [Ni II] 7378 Å for the blue and red peaks, respectively) in the late-time spectra (150–400 d past max.) and the Si II velocity gradient at maximum light was first identified by Maeda et al. (2010a). SNe Ia displaying redshifted late-time 7300 Å region spectral features have higher velocity gradients at early times, suggested to originate from asymmetric and off-centre explosion mechanisms. Silverman et al. (2013) extended this work to show that the relation also holds if Si II velocities at maximum are used instead of the harder to measure velocity gradient.

We wished to measure velocity offsets for our new late-time spectral sample (>200 d after max.) and compare with Si II velocities at maximum (presented in Table 2). We included the three additional XShooter spectra (described in Section 2), with lower S/N spectra in the NIR region, to increase the sample. Unlike Maeda et al. (2010a), we chose to focus on the blueward peak that is predominantly [Fe II] because the redward peak of this feature is not due to solely [Fe II] or [Ni II] but is a blend of both (Fig. 3). Therefore, it is not as clean

a measurement of Fe-group velocity offsets as the [Fe II] 7155 Å peak.

We measured the [Fe II] 7155 Å velocity shift using two methods: the ‘peak’ method of Maeda et al. (2010a) and the more detailed multiline fitting method we presented in Section 3. We also remeasured values of the late-time velocity offset of the [Fe II] 7155 Å feature for the literature sample described in Section 2.4, which includes SNe Ia falling in the same phase range of our late-time sample of ~200–430 d post-max. This is a stricter lower phase cut than that used by Maeda et al. (2010a).

A comparison between the two methods (‘peak’ and full fit) shows that there is a systematic offset such that the values measured from the peak of the blueward feature ([Fe II] 7155 Å) are ~150–300 km s⁻¹ offset to the red compared to the full fitting method. This is evidence for line blending skewing the measured peak value of the emission profile. However, this is a small offset and generally within the uncertainties of the line fitting so we conclude that the choice of method does not significantly impact the final results. We used the full fitting method in our further analysis.

Fig. 8 shows the results of our analysis of the relation between Si II velocity at maximum and late-time velocity offset of the [Fe II] 7155 Å lines for our SN sample and remeasured literature sample. We find results in agreement with that of previous studies but strengthened (with no clear outliers) by the inclusion of additional objects, the stricter phase cut, and the consistent measurements of solely the [Fe II] 7155 Å feature using our more detailed fitting technique. We also tested for relations between the maximum-light light-curve width parameter, ‘stretch’, the $B - V$ colour at maximum, and the late-time velocities but find no significant correlations.

A similar relation is seen between Si II velocity at maximum and the late-time velocity offset of the NIR [Fe II] 12567 Å line – not surprisingly given the clear correlation between the optical and NIR [Fe II] lines observed in Fig. 7. However, the sample size is small since the literature spectral sample generally did not cover combined optical and NIR wavelengths. We have also investigated the relation between the Si II velocity at maximum and the late-time velocity shift of the [Co III] 5890 Å lines for the combined XShooter and literature sample. We identify a tentative trend in the sense that

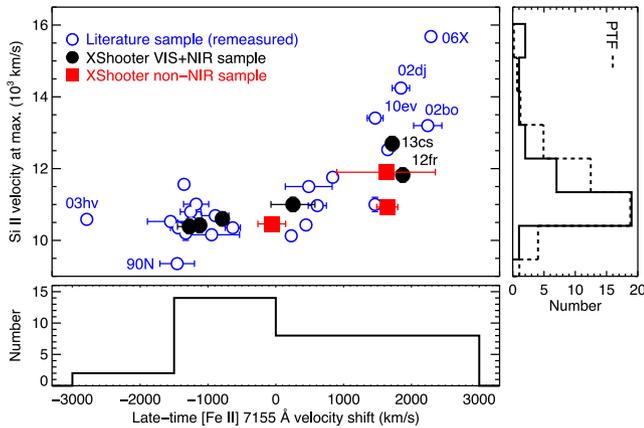


Figure 8. Main: Si II velocity at maximum light versus the velocity shift of the [Fe II] 7155 Å feature from spectra at >200 d after maximum (measured from a multi-Gaussian fit to the feature). The new XShooter sample with high S/N NIR spectra are shown in black solid circles, the lower XShooter S/N sample without good NIR data is shown as red squares, and the remeasured literature sample is shown as blue open circles. Bottom: histogram of the late-time [Fe II] 7155 Å velocity shifts for the combined XShooter and literature samples (black solid line). Side: histogram for the Si II velocity at maximum light for the combined sample (solid black line), with the PTF SN Ia sample of 247 SNe Ia within ± 5 d of maximum light from Maguire et al. (2014) normalized to the peak of the combined sample (black dashed line).

SNe Ia with positive (redward shifted) [Co III] late-time velocities are more likely to have higher Si II velocities at maximum (in the same sense as the [Fe II] lines). However, when we measure the mean Si II velocity for samples with redward or blueward [Co III] shifts, the means agree within the uncertainties. The [Co III] lines also have significantly smaller offsets from zero velocity (Fig. 7).

To determine if the combined XShooter and literature sample is unbiased with respect to selection effects, we have compared it to the untargeted Palomar Transient Factory (PTF) sample of 247 SNe Ia within 5 d of maximum light (Maguire et al. 2014). We have found that the percentage of SNe Ia with Si II velocities above $12\,000\text{ km s}^{-1}$ is consistent between the samples at ~ 18 and ~ 16 per cent for the late-time spectral sample and the maximum-light PTF sample, respectively. This suggests that there is no significant bias in the sample selection for the late-time sample.

5 LATE-TIME COLOURS OF SNe Ia

The percentage flux in the NIR bands was compared to the total flux in the optical and NIR by measuring the flux through a broad optical filter (3550–9500 Å), as used in Dimitriadis et al. (2017), and the sum of the flux through the *JHK* bands. The percentage flux in the NIR compared to the combined optical+NIR is shown as a function of phase in the range of 200–550 d post-maximum in Fig. 9. This figure includes measurements from the new spectral sample presented here, as well as the late-time optical and NIR measurements of SNe 2000cx, 2001el, 2003hv, 2011fe (Sollerman et al. 2004; Stritzinger & Sollerman 2007; Leloudas et al. 2009) that were presented in Dimitriadis et al. (2017).

An increase in the NIR flux contribution with time is seen over the phase range studied. The NIR contribution to the combined optical and NIR flux is relatively constant in the range of 200–350 d with a mean of 7 ± 2 per cent. In the range 350–450 d, there is a larger spread with a mean of 13 ± 10 per cent, while in the range 450–

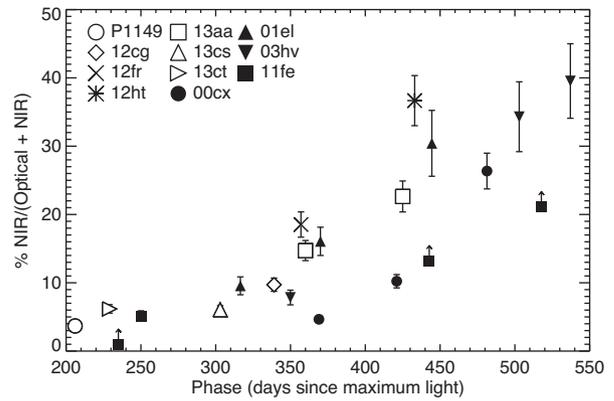


Figure 9. Percentage flux through the NIR *JHK* filters compared to the combined flux from a broad optical filter (3550–9500 Å) and the *JHK* filters. Three epochs of data for SN 2011fe are marked as lower limits because only *J*-band data are available as the NIR contribution.

550 d, the NIR flux contribution has increased significantly with a mean of 33 ± 7 per cent.

6 CALCULATION OF Ni TO Fe ABUNDANCE RATIO

We have produced fits to the 7300 Å spectral region, which output relative line strengths, velocity shifts, and line widths of [Fe II] and [Ni II]. The pseudo-equivalent widths of the line profiles can be used to estimate the Ni/Fe abundance for SNe Ia in the sample, under some basic assumptions. Assuming optically thin emission, we used equation 2 of Jerkstrand et al. (2015a), with an additional term to account for the departure coefficients, to give the equation

$$\frac{L_{7378}}{L_{7155}} = 4.9 \frac{n_{\text{Ni II}}}{n_{\text{Fe II}}} \exp\left(\frac{0.28}{kT}\right) \frac{dc_{\text{Ni II}}}{dc_{\text{Fe II}}} \quad (1)$$

where L_{7378}/L_{7155} is the measured flux ratio of the [Ni II] 7378 Å to [Fe II] 7155 Å line, $n_{\text{Ni II}}$ and $n_{\text{Fe II}}$ are the number densities of Ni II and Fe II, k is the Boltzmann constant in units of eV per K, T is the temperature in K, and $dc_{\text{Ni II}}/dc_{\text{Fe II}}$ is the ratio of departure coefficients from LTE for the appropriate states of Ni II and Fe II, respectively. The numerical constants are the same as were used in Jerkstrand et al. (2015a).

The ratio of departure coefficients for the Fe and Ni-emitting zones is estimated from the W7 model at 330 d using the modelling of Fransson & Jerkstrand (2015) to be in the range of 1.2–2.4 over the phase range of interest with smaller coefficients at earlier times when the conditions are slightly closer to LTE (private communication). We use a mean value of the departure coefficient of 1.8. Since Ni II and Fe II are both singly ionized and have similar ionization potentials, we assumed that the $n_{\text{Ni II}}/n_{\text{Fe II}}$ ratio should closely follow the total Ni/Fe ratio. This is supported by the nebular phase modelling of Fransson & Jerkstrand (2015) at 330 d past maximum light, which gives values in the range of 0.8–1.2, with a mean of 1.0 (private communication). Similar ionization rates are also found in the NLTE modelling of Shingles et al. (in preparation).

The uncertainties on the Ni/Fe ratio were calculated by varying the temperature in the range 3000–8000 K and the departure coefficient in the range 1.2–2.4 (the range of departure coefficients in different zones from the model of Fransson & Jerkstrand 2015), and the ionization balance in the range of 0.8–1.2, also from the

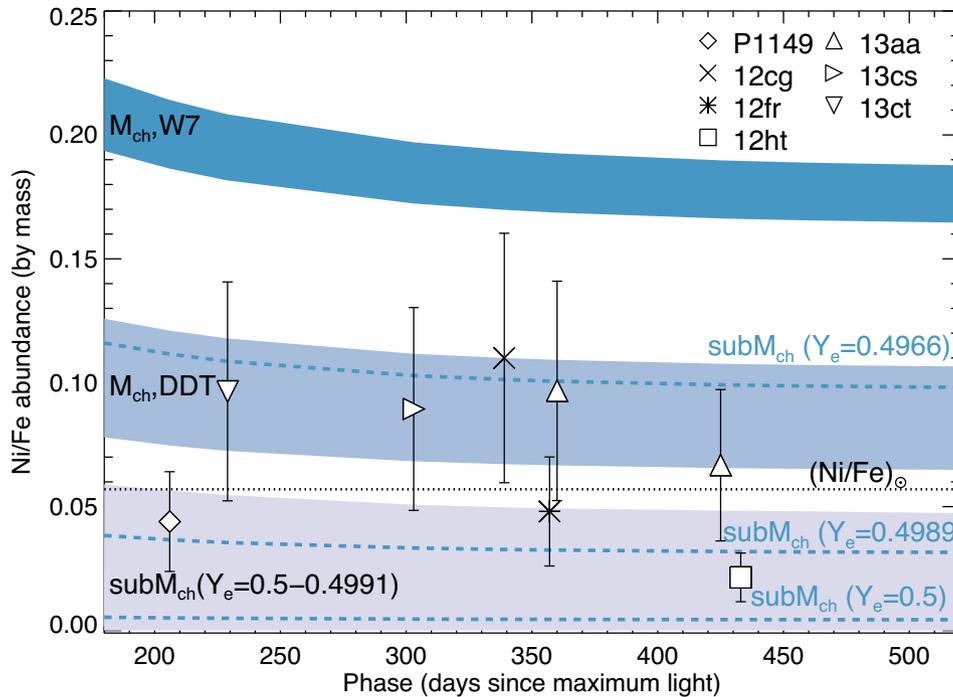


Figure 10. Ratio of Ni to Fe abundance (by mass) compared to bands of the M_{ch} ‘DDT’ models (Seitenzahl et al. 2013a), and ‘W7’ models (Nomoto et al. 1997; Iwamoto et al. 1999). The sub- M_{ch} detonation models of Sim et al. (2010) and Ruiter et al. (2013) are marked as blue dashed lines (sub- M_{ch}) at $Y_e = 0.5, 0.4989, \text{ and } 0.4966$ (equivalent to $\sim 0, 2, \text{ and } 6$ times solar metallicity, respectively). The sub- M_{ch} detonation models of Shen et al. (2018) in the range of $Y_e = 0.5\text{--}0.4991$ (equivalent to $\sim 0\text{--}2$ times solar metallicity) are marked as a solid band (sub- M_{ch}). For reference, the solar metallicity from Asplund et al. (2009) is equivalent to $Y_e=0.499$ and the solar Ni/Fe abundance ratio is 0.057 (black dotted line).

Fransson & Jerkstrand (2015) model. The uncertainties on the line fitting are negligible in comparison to these uncertainties.

Fig. 10 shows the measured ratio of Ni to Fe abundance compared to four SN Ia explosion models (a M_{ch} deflagration model, a M_{ch} DDT model, and two different versions of sub- M_{ch} detonation models). The M_{ch} deflagration explosion model, ‘W7’, is shown with a range of values from Nomoto et al. (1997) and Iwamoto et al. (1999), along with the three-dimensional M_{ch} DDT models, ‘N100’, ‘N100l’, ‘N100h’, ‘N150’, of Seitenzahl et al. (2013a). In the DDT models, the neutron-rich isotopes, such as ^{58}Ni , are synthesized mainly in the explosion itself in the highest density regions, and these models set the range of the models labelled as ‘DDT’ in Fig. 10.

Neutronization can also occur due to a increased progenitor star metallicity that results in a higher ^{22}Ne content, which in turn introduces an excess of neutrons and different electron fractions, Y_e . This results in the production of neutron-rich isotopes but in smaller amounts than those produced during the explosion itself. For the M_{ch} DDT models of Seitenzahl et al. (2013a), variations in initial progenitor star metallicity are dwarfed by the variations produced by varying the central density at explosion.

Two sets of sub- M_{ch} detonation models are shown from Sim et al. (2010) and Shen et al. (2018). In these models, the neutron-rich ^{58}Ni production depends on the progenitor star metallicity. The models of Sim et al. (2010) are of a $1.06 M_{\odot}$ CO white dwarf calculated for the case of zero metallicity ($Y_e = 0.5$), and two neutron-rich environments of $Y_e = 0.4989$ and 0.4966 , roughly equivalent to twice and six times solar metallicity, respectively. The models of Shen et al. (2018) are of 0.9 and $1.0 M_{\odot}$ CO white dwarfs, calculated in the range, $Y_e = 0.5\text{--}0.4991$, roughly equivalent to zero to twice solar metallicity.

All the model ratios of Ni over Fe abundance decline with phase because while the stable Ni mass stays constant, the amount of Fe is increasing with time due to decay from Co. ‘W7’ overpredicts the Ni/Fe abundance for all the SNe Ia in the sample. The data are found to be broadly consistent with the ‘DDT’ model range, as well as some of the sub- M_{ch} models. However, to explain the data, the sub- M_{ch} detonation models would need neutron excesses equivalent to a few times solar metallicity.

7 DISCUSSION

We have presented an analysis of the largest sample to date of late-time optical and NIR SN Ia spectra, where we used a line-fitting code to estimate the widths, velocity offsets, and strengths of prominent spectral features. In this section, we review the evidence for asymmetries in the ejecta and the constraints on the presence of stable material. We also discuss the rate of cooling of the ejecta, as well as the absence of flat-topped profiles. These results are linked to a number of explosion models to constrain which models are the most viable to explain the bulk of normal SNe Ia.

7.1 Stratification of stable and radioactive material

In the W7 model (often used as a reference model), the stable material is more centrally concentrated than the radioactive decay products. The velocities of the stable material are predicted to be $\lesssim 4000 \text{ km s}^{-1}$ (Nomoto et al. 1997; Iwamoto et al. 1999). The stable Ni and Fe material in SN 2011fe was found using a modified density W7 model to have velocities below $\sim 4000 \text{ km s}^{-1}$ (Mazzali et al. 2015). The DDT models predict distributions of stable ^{58}Ni that when spherically averaged are approximately co-spatial with

the radioactive material (see fig. 1 of Seitenzahl et al. 2014). The sub- M_{ch} detonation models of Sim et al. (2010) and Shen et al. (2018) predict distributions of stable Fe-group material that cover a similar velocity range to that of the radioactive isotopes.

For our SN Ia spectral sample, we have measured the width of Fe-group element features to constrain the extent of their emitting region. The width of the [Co III] line is found to be significantly broader than those of the [Fe II] and [Ni II] features. The mean FWHM of the [Fe II] 7155 Å and [Ni II] 7378 Å lines are 7600 ± 700 and 6600 ± 1500 km s⁻¹, respectively, agreeing within the uncertainties. The mean FWHM of the [Co III] 5890 line is 10300 ± 1300 km s⁻¹. This means the region producing the [Fe II] and [Ni II] features is smaller than the region producing the [Co III] lines.

The [Ni II] emission is dominated by stable ⁵⁸Ni since the radioactive ⁵⁶Ni is decayed by these epochs but the [Co III] lines are produced from ⁵⁶Co, the product of radioactive decay. To first order, since the [Ni II] emitting region is narrower than those of the [Co III] region, we can assume that the underlying composition between the two emitting regions is different. Since the difference in mean FWHM for the [Ni II] 7378 Å and [Co III] 5890 Å lines is of the order of a thousand to a few thousand km s⁻¹ (taking into account the uncertainties), this suggests that the stable material is (at least slightly) more concentrated and coming from a narrower region than the radioactive material. We have not used [Fe III] in our study because of the time-dependent evolution of the most prominent [Fe III] feature (see Section 4.1).

This suspected stratification of the stable and radioactive material is in quantitative agreement with most of the models (W7, DDT, sub- M_{ch} detonation). However, W7 predicts that the velocities of the stable material are $\lesssim 4000$ km s⁻¹, which we do not see in our sample. The lowest FWHM of [Ni II] 7378 Å for an SN Ia in our sample is SN 2012fr, with an FWHM of the [Ni II] 7378 Å line of ~ 4650 km s⁻¹. The gravitationally confined detonation models of Seitenzahl et al. (2016) are ruled out for the SNe Ia in our sample; these models have no stable Fe or Ni at low velocities but significant high-velocity stable material, which is at odds with our results.

However, there are a number of caveats with connecting the underlying distribution of stable and radioactive material to line emissions of [Ni II] and [Co III]. Since their ionization states are different (singly ionized [Ni II] compared to doubly ionized [Co III]), the density of their emitting regions is expected to be different (Maeda et al. 2010b). An isolated blob of stable material will not emit without an energy source – there must be some radioactive material mixed in with the stable material to produce line emission. Therefore, although different line emitting regions are likely dominated by differences in the composition, the role of thermodynamic conditions can of course not be neglected to focus only on differences in composition. Full non-LTE radiative-transfer modelling of different explosion models is needed to confirm these results.

The [Fe II] emitting region is more difficult to link to stable or radioactive material, since most of the models predict that there is a non-negligible amount of stable ⁵⁴Fe and ⁵⁶Fe produced in the explosion. These contributions cannot be easily disentangled. We have found that the [Fe II] and [Ni II] lines have significant positive and negative velocity shifts with respect to the rest wavelengths. Interestingly, the shifts of the two features correlate, i.e. when the [Fe II] emission lines are shifted to the blue, the [Ni II] lines usually do as well. The [Fe II] and [Ni II] velocity shifts are also significantly higher (up to ± 2000 km s⁻¹) than those of the [Co III] lines, which are consistent with zero shift.

A toy model was set up by Maeda et al. (2010b) with (i) an offset (from zero velocity) high-density region dominated by electron-capture elements (90 per cent by mass as stable ⁵⁸Ni, 10 per cent of ⁵⁶Fe), (ii) an offset relatively high-density region with ⁵⁶Fe, and (iii) an extended, non-offset low-density region with ⁵⁶Fe. They assumed in this model that the NIR lines of [Fe II] (12 567 Å, 16 440 Å) and the [Fe II] 7155 Å line should come from the offset high-density ⁵⁶Fe-dominated region. Our result that the [Co III] lines come from a non-offset (zero velocity shift), extended (broader width) ⁵⁶Fe-dominated region agrees with the toy model of Maeda et al. (2010b). The shifts to the red and blue for the [Fe II] and [Ni II] lines can be interpreted as viewing angle effects of an asymmetric explosion (see further discussion below in Section 7.2).

7.2 Asymmetry predictions from explosions models

Bulla et al. (2016) investigated the line velocities and polarization signals of the DDT model (N100) of Seitenzahl et al. (2013a) and the sub- M_{ch} double-detonation model of Fink et al. (2010). They found that both models have continuum polarization levels that match well with the observations but differ significantly in the level of asymmetries, with the sub- M_{ch} double-detonation model having significantly higher levels of global asymmetry. These asymmetries are introduced by the explosion mechanism, where the detonation is ignited at a single point on the surface and the detonation wave then travels on this surface to converge at the opposite side. This results in intermediate-mass elements that have a wider range of velocities and are more abundant on one side of the ejecta than on the other (Fink et al. 2010). The N100 DDT model of Seitenzahl et al. (2013a) does not predict significant variations in the offset of stable material (the stable ⁵⁸Ni offset averaged over the ejecta is of the order of just tens of km s⁻¹) or variation in the Si II velocity with viewing angle. For one of the most asymmetric models (N0003) of Seitenzahl et al. (2013a), the average velocity offset is larger but still not very large: ~ 200 km s⁻¹.

We have confirmed and strengthened the trend first identified by Maeda et al. (2010b) that the SNe Ia with more redshifted velocities of the late-time [Fe II] lines are found to have higher maximum-light Si II velocities (see Fig. 8). This suggests that the Si II velocity is greatest when the bulk of the Fe-group elements are moving away from the observer. This is in good qualitative agreement with the sub- M_{ch} double-detonation model predictions, where an observer would see a high Si II velocity when viewing the explosion from the opposite side of the ejecta to where the core detonation has occurred. The Fe material would in this case to be moving away from the observer (see fig. 2 of Bulla et al. 2016).

Dong et al. (2015) investigated the shape of the [Co III] 5890 Å feature in a sample of 18 SNe Ia from the literature and found that some SNe Ia with $\Delta m_{15}(B) > 1.3$ (stretch < 0.85) have double-peaked profiles. They suggested this could be due to different viewing angles of a bimodal explosion of the direct collisions of white dwarfs. They suggested that [Fe II] lines should also show double-peaked features. We have not identified double-peaked features of [Co III] or [Fe II] in the SNe Ia in our sample. Recent triple system population synthesis suggests that direct collisions are unlikely to explain more than 0.1 per cent of normal SNe Ia (Toonen, Perets & Hamers 2018).

7.3 Constraints on Ni to Fe abundance

Different explosion models predict different production rates of stable material due to different central densities at the time of

explosion, with only M_{ch} mass white dwarfs able to produce neutron-rich isotopes, such as ^{58}Ni , during the explosion itself. However, sub- M_{ch} models can have non-zero amounts of neutron-rich isotopes, but this is dependent on the white dwarf metallicity that results in higher neutronization for higher metallicity progenitors. The ratio of Ni to Fe is much lower for sub- M_{ch} models than for M_{ch} models, except for when very high metallicity (many times solar) sub- M_{ch} progenitors are invoked (Sim et al. 2010; Ruiter et al. 2013; Shen et al. 2018).

Stable Fe is difficult to measure at late times because radioactively produced ^{56}Fe from the decay of ^{56}Ni via ^{56}Co is also present. However, at late times, all the radioactively produced ^{56}Ni has decayed and so any features of Ni seen in SN Ia spectra should be dominated by stable ^{58}Ni . The most prominent lines to study are the doublet of $[\text{Ni II}]$ 7378, 7412 Å. We find that a contribution from $[\text{Ni II}]$ is required to fit this region.

We estimate the Ni/Fe abundance ratio to be in the range 0.02–0.10 for the seven SNe Ia studied here, within uncertainties on the values of ~ 40 per cent, excluding systematic uncertainties, such as those due to choosing a certain set of atomic data. In Fig. 10, we show that Ni/Fe abundance for our sample is broadly consistent with the M_{ch} DDT models of Seitenzahl et al. (2013a). However, as discussed in Section 7.2, the measured velocity offsets seen for the $[\text{Fe II}]$ and $[\text{Ni II}]$ lines appear inconsistent with this model. While there is overlap between the data and the sub- M_{ch} explosion models of Sim et al. (2010) and Shen et al. (2018), they generally under-produce the Ni/Fe abundance compared to the observations and the sub- M_{ch} models that match best are those with progenitor star metallicities significantly above solar. This increase in metallicity boosts the neutron excess available to produce stable Fe-group elements. The ‘W7’ model significantly overpredicts the Ni/Fe abundance for our sample, because of its large stable Ni core. Previous studies of SNe Ia have also found that the high nickel content of W7 leads to too strong emission lines of $[\text{Ni II}]$ at late times (Ruiz-Lapuente & Lucy 1992; Liu et al. 1997; Leloudas et al. 2009; Fransson & Jerkstrand 2015).

To make a crude estimate of the number of SNe Ia occurring in supersolar metallicity environments, we used the results of a number of spectroscopic metallicity studies of the global host galaxy properties of SNe Ia. We find that the percentage of SNe Ia in high-metallicity galaxies is approximately half (46–64 per cent; Childress et al. 2013a; Pan et al. 2014; Campbell, Fraser & Gilmore 2016). However, to explain the majority of Ni over Fe abundances in our sample, a metallicity of >2 times solar would be required, which is reasonably rare for SNe Ia host galaxy locations – estimates of the number of SNe Ia exploding in galaxies at greater than twice solar metallicity are of the order of ~ 1 –25 per cent (Childress et al. 2013a; Pan et al. 2014; Campbell et al. 2016).

Studies of the local host galaxy properties of SNe Ia find similar numbers of events at greater than solar metallicity (~ 50 per cent), and at greater than twice solar metallicity ($\lesssim 5$ per cent; Galbany et al. 2016, Galbany et al. in preparation). We also found that our SN samples all occur in emission-line, morphologically spiral galaxies, suggesting that they are not biased towards very high metallicity environments. Therefore, with the current sub- M_{ch} detonation models it appears difficult to explain the bulk of the SN Ia Ni over Fe abundances observed in our late-time spectral sample.

Previous modelling efforts have suggested that $[\text{Ca II}]$ emission could play a role in this region (Liu et al. 1997; Fransson & Jerkstrand 2015; Mazzali et al. 2015; Botyánszki & Kasen 2017). The explosion models generally predict that intermediate-mass elements, such as Ca, have a shell-like ejecta structure. The presence

of any broad $[\text{Ca II}]$ emission would have a pedestal effect, increasing the level of the underlying continuum. We have found that a contribution is not needed in our analysis to fit the 7300 Å region (Fig. 3) and is rejected as a component in our fits using a model-selection technique. If we assumed that the redward peak in the 7300 Å region had a contribution from Gaussian-shaped $[\text{Ca II}]$ lines instead of $[\text{Ni II}]$ lines, then it would produce much larger velocity offsets, with the maximum redshifted SN 2012fr being shifted by $\sim 5500 \text{ km s}^{-1}$ – a very large shift for any explosion model. Therefore, our conclusion is that $[\text{Ca II}]$ emission does not significantly contribute to the 7300 Å region of late-time SN Ia spectra.

The measured Ni/Fe abundance ratios for our sample of SNe Ia at late times are broadly consistent (mean of 0.06 ± 0.03) with the solar Ni/Fe value of 0.057 (Asplund et al. 2009). A previous study of Ni/Fe abundances in a sample of six core-collapse SNe found significantly supersolar ratios in two of these objects, and around solar for the other four (Jerkstrand et al. 2015a). It was speculated that SNe Ia may have, on average, lower than solar Ni/Fe ratios that would bring the combined yield of SNe Ia and core-collapse SNe to the solar value (Jerkstrand et al. 2015c). However, we have shown that the Ni/Fe ratio for SNe Ia is broadly consistent with the solar value. Future studies of the Ni/Fe abundances in both core-collapse SNe and SNe Ia are needed so that we can put stronger constraints on their explosions, and their relative contributions to Galactic chemical evolution.

7.4 Absence of flat-topped NIR $[\text{Fe II}]$ features

‘Flat-topped’ profiles have been suggested to arise in SNe Ia originating from M_{ch} explosions, where the high central density at the time of explosion results in significant electron capture resulting in a core that is dominated by stable Fe-group elements (Höflich et al. 2004). At >150 –200 d, local deposition by positrons should dominate over gamma-ray deposition (e.g. Penney & Höflich 2014). Therefore, if there is no mixing of the ejecta, this highest density region could produce nebular-phase emission lines of radioactivity-dominated species that have flat-topped profiles (due to the lack of radioactive elements in this region). Previous studies have identified flat-topped profiles at NIR wavelengths for SNe 2003du and 2003hv (Höflich et al. 2004; Motohara et al. 2006). They were not seen at optical wavelengths, which was suggested to be due to more line blending at optical wavelengths. Flat-topped profiles were not seen in the NIR for SN 2005df (Diamond, Höflich & Gerardy 2015) or SN 2011fe (Mazzali et al. 2015).

For our XShooter sample, we have found that the 12 600 Å region is well fitted by multiple Gaussians of $[\text{Fe II}]$ lines (see Fig. 2). The main peak of this feature comes nearly exclusively from the $[\text{Fe II}]$ 12 567 Å line and we identified no evidence for a flat-topped profile of this or the other weaker $[\text{Fe II}]$ features in this region. In Fig. 11, we show the NIR region from $\sim 11\,000$ – $18\,000$ Å for our sample of eight late-time SN Ia spectra, as well as the +394 d spectrum of SN 2003hv (Motohara et al. 2006; Leloudas et al. 2009) and +217 d spectrum of SN 2005df (Diamond et al. 2015). As previously discussed, we did not include the 15 000–18 000 Å region in our detailed fitting analysis due to the more complicated temperature dependence of $[\text{Fe II}]$ lines in this region and potential contributions from other elements. However, from a qualitative analysis of Fig. 11, we do not identify flat-topped profiles for the SNe in our sample in any of the main $[\text{Fe II}]$ emission features in the NIR. Although it cannot be ruled out that the previously identified flat-topped profiles are real, it is possible that they were due to low S/N in these regions.

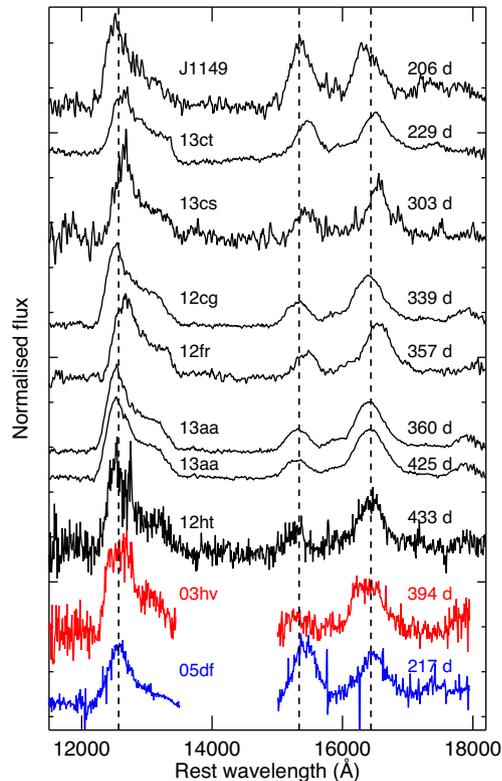


Figure 11. Comparison of SN Ia NIR spectra, focusing on the $\sim 11\,000$ – $18\,000$ Å wavelength region. The eight new XShooter spectra of seven SNe Ia are shown in black compared to two from the literature (SN 2003hv in red and SN 2005df in blue). The dashed vertical lines marked the position of the most prominent [Fe II] emission features. No flat-topped profiles are identified for the XShooter spectral sample. Shifts of the peaks of the main [Fe II] features with respect to their rest wavelength can also be observed. In particular, shifts to the red for the peaks of the $12\,567$ Å and $16\,440$ Å lines can be seen for SN 2013ct, SN 2013cs, and SN 2012fr.

We also do not require flat-topped profiles for the [Fe II] at optical wavelengths, which are predicted by models that only have stable material in their highest density regions (Höfllich et al. 2004). A central region solely comprised of stable material is not required for our spectra. The widths of the [Ni II] lines of $\sim 5\,000$ – $9\,000$ km s $^{-1}$ also overlap with those of the [Fe II] features of $\sim 7\,000$ – $9\,000$ km s $^{-1}$, again suggesting that the stable material that we have identified in our spectra is not confined to a high-density region without any radioactive material present.

7.5 Cooling at late times

As has been seen in other studies of late-time data of SNe Ia, there is a marked increase in flux emerging at NIR compared to optical wavelengths at >350 d past explosion (Fig. 9). This is seen for SN 2013aa in Fig. 12, which has spectra at 360 and 425 d post-maximum. The flux in the NIR [Fe II] lines remains constant between the two epochs but the flux at bluer wavelengths has significantly decreased. This levelling off in the flux of the NIR bands of SNe Ia at late times was previously seen in SN 1998bu (Spyromilio et al. 2004) and SN 2000cx (Sollerman et al. 2004) in the phase range of ~ 250 – 480 d. Since most of the emission at these phases is still from thermal collisions, this relative increase in the flux in the NIR bands is expected to be due to the continued cooling of the ejecta, resulting in the transition to emission at longer wavelengths. This

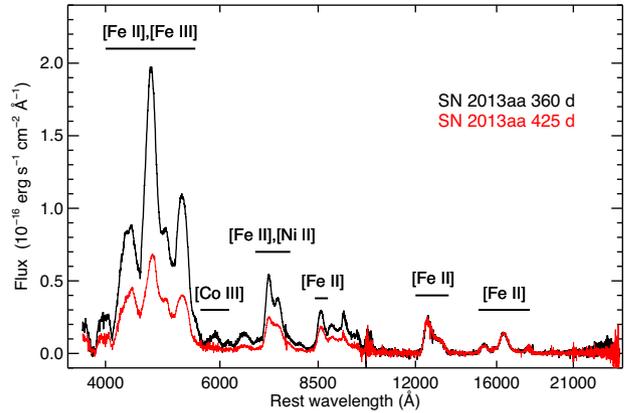


Figure 12. Comparison of the optical and NIR spectra of SN 2013aa at 360 and 425 d post-maximum. The main contributors to individual spectral regions are marked.

temporarily cancels out the decreasing radioactive heating at NIR wavelengths resulting in little evolution in the NIR bands at these epochs.

We have also identified a hint of the transition from [Fe III]/[Fe II] to [Fe II]/[Fe I]-dominated spectra that was previously observed in SN 2011fe between 365 and 593 d post-explosion (Graham et al. 2015). The main peak in the blue at $\sim 4\,700$ Å is dominated by [Fe III] emission while the peak to the red of this at $\sim 5\,300$ Å is mixture of [Fe III] and [Fe II]. If we normalize to this peak at $5\,300$ Å we find that the peak at $4\,700$ Å has decreased with respect to it, meaning that [Fe III]/[Fe II] ratio has decreased between these epochs. We have not identified any [Fe I] emission features in our spectra but speculate based on the results of Graham et al. (2015) that we may be seeing the beginning of a transition from an [Fe III]/[Fe II] to [Fe II]/[Fe I]-dominated spectrum in the range of 360 and 425 d post-maximum for SN 2013aa.

8 CONCLUSIONS

In this paper, we have presented a detailed analysis of the late-time (>200 d post-explosion) optical and NIR spectra of seven SNe Ia. The late-time XShooter spectra were supplemented by maximum-light photometric and spectroscopic observations. Using a multi-Gaussian fitting code, we have performed fits to three prominent spectral regions: the $12\,600$ Å [Fe II]-dominated region, the $7\,300$ Å [Fe II]/[Ni II]-dominated region, and the $6\,000$ Å [Co III] region. This has allowed us to constrain the stable Ni over Fe abundance ratio, the velocity shifts and widths of the measured features, and their connection to maximum-light observations. We have investigated the NIR compared to optical flux contribution, along with a study of the presence of flat-topped NIR profiles and the double-peaked [Co III] features.

Our main results are as follows:

(i) Multi-Gaussian fits to the $7\,300$ Å [Fe II]/[Ni II]-dominated region show that a contribution from stable [Ni II] is required to fit the line profiles (Figs. 3 and 4). [Ca II] is found to make a negligible contribution to this region.

(ii) The abundance ratio of Ni to Fe is found to be in the range of 0.02 – 0.10 for the seven SNe Ia in our sample, with uncertainties of ~ 40 per cent, making them broadly consistent with DDT and some supersolar sub- M_{ch} models. The mean value of our sample is consistent with solar Ni/Fe abundance measurements.

(iii) Using a simple fit to the 4700 Å [Fe III] main peak, we confirm the results of previous studies that the peak of this feature shifts to the red with time.

(iv) We have constrained the velocity offsets (from rest wavelength) of the [Fe II], [Ni II], and [Co III] features. The [Fe II] 7155 Å, 12 567 Å, and the [Ni II] 7378 Å features are found to have significant offsets ($\sim 500\text{--}2000\text{ km s}^{-1}$) both positive and negative from their rest wavelengths (Fig. 7). The [Co III] 5890 Å feature does not show such significant offsets with a mean absolute velocity offset of $271 \pm 341\text{ km s}^{-1}$, suggesting a different emitting region for the [Co III] compared to the [Fe II] and [Ni II] features.

(v) The [Co III] 5890 Å feature is found to have significantly broader FWHM of $\sim 10\,000\text{ km s}^{-1}$ than the [Fe II] and [Ni II] features of $6600\text{--}7600\text{ km s}^{-1}$, suggesting a broader emitting region for [Co III] compared to [Fe II] and [Ni II] features.

(vi) By connecting the late-time [Fe II] line shift velocities with the velocity of the Si II absorption feature at maximum light, we confirm (and strengthen) the trend of Maeda et al. (2010a), where SNe Ia with blueshifted [Fe II] features have lower Si II velocities.

(vii) A trend of increasing NIR flux with time at epochs $>350\text{ d}$ post explosion is identified (from <10 at $\sim 350\text{ d}$ up to 40 per cent at $\sim 550\text{ d}$). The onset of the transformation from [Fe III]/[Fe II]- to [Fe II]/[Fe I]-dominated spectra is potentially identified for SN 2013aa in the range of 360–425 d post-maximum light.

(viii) In contradiction with some previous studies, we do not see flat-topped profiles for the [Fe II] features at NIR wavelengths, which may be due to the higher S/N spectra in our sample.

(ix) No double-peaked [Co III] or [Fe II] features are identified that have been suggested to be present in fainter SNe Ia in direct collision models.

The largest sample to date of late-time combined optical and NIR spectra of SNe Ia has allowed us to test a number of theories related to density structures, asymmetries, and elemental abundances. We have identified signatures of potential asymmetries in their explosions through measurements of the velocities and widths of key Fe-group elements. Our results suggest that the explosions of normal SNe Ia, as studied here, may be intrinsically non-spherically symmetric with significant offsets from zero velocity for some elements and ionization states. Future work will involve larger samples of SNe Ia at late times, combined with state-of-the-art spectral modelling (Shingles et al. in preparation), to allow the geometry and explosion products to be quantified to higher precision, and place tighter constraints on viable SN Ia explosion models.

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REFERENCES

- Ashall C., Mazzali P., Bersier D., Hachinger S., Phillips M., Percival S., James P., Maguire K., 2014, *MNRAS*, 445, 4427
- Asplund M., Grevesse N., Sauval A. J., Scott P., 2009, *ARA&A*, 47, 481
- Axelrod T. S., 1980, PhD thesis, California Univ., Santa Cruz
- Betoule M. et al., 2014, *A&A*, 568, A22
- Bildsten L., Shen K. J., Weinberg N. N., Nelemans G., 2007, *ApJ*, 662, L95
- Black C. S., Fesen R. A., Parrent J. T., 2016, *MNRAS*, 462, 649
- Blondin S. et al., 2012, *AJ*, 143, 126
- Botyánszki J., Kasen D., 2017, *ApJ*, 845, 176
- Brown T. M. et al., 2013, *PASP*, 125, 1031
- Bulla M. et al., 2016, *MNRAS*, 462, 1039
- Campbell H., Fraser M., Gilmore G., 2016, *MNRAS*, 457, 3470
- Cartier R., Förster F., Coppi P., Hamuy M., Maeda K., Pignata G., Folatelli G., 2011, *A&A*, 534, L15
- Chambers K. C. et al., 2016, preprint ([arXiv: e-prints](https://arxiv.org/abs/1608.07948))
- Childress M. et al., 2013a, *ApJ*, 770, 107
- Childress M. et al., 2013b, *ApJ*, 770, 108
- Childress M. J. et al., 2015, *MNRAS*, 454, 3816
- Childress M. J. et al., 2016, *PASA*, 33, e055
- Conley A. et al., 2008, *ApJ*, 681, 482
- Cortés J. R., Kenney J. D. P., Hardy E., 2006, *AJ*, 131, 747
- Dan M., Rosswog S., Guillochon J., Ramirez-Ruiz E., 2011, *ApJ*, 737, 89
- Dessart L., Hillier D. J., Blondin S., Khokhlov A., 2014, *MNRAS*, 439, 3114
- Diamond T. R., Höflich P., Gerardy C. L., 2015, *ApJ*, 806, 107
- Dimitriadis G. et al., 2017, *MNRAS*, 468, 3798
- Dong S., Katz B., Kushnir D., Prieto J. L., 2015, *MNRAS*, 454, L61
- Fink M., Röpke F. K., Hillebrandt W., Seitenzahl I. R., Sim S. A., Kromer M., 2010, *A&A*, 514, 53
- Fransson C., Jerkstrand A., 2015, *ApJ*, 814, L2
- Freudling W., Romaniello M., Bramich D. M., Ballester P., Forchi V., García-Dabó C. E., Moehler S., Neesser M. J., 2013, *A&A*, 559, A96
- Galbany L. et al., 2016, *A&A*, 591, A48
- Gerardy C. L. et al., 2007, *ApJ*, 661, 995
- Graham M. L., Nugent P. E., Sullivan M., Filippenko A. V., Cenko S. B., Silverman J. M., Clubb K. I., Zheng W., 2015, *MNRAS*, 454, 1948
- Graham M. L. et al., 2017, *MNRAS*, 472, 3437
- Gutiérrez C. P. et al., 2016, *A&A*, 590, A5
- Hillier D. J., Dessart L., 2012, *MNRAS*, 424, 252
- Hillier D. J., Miller D. L., 1998, *ApJ*, 496, 407

- Höfllich P., Gerardy C. L., Nomoto K., Motohara K., Fesen R. A., Maeda K., Ohkubo T., Tominaga N., 2004, *ApJ*, 617, 1258
- Iben I., Tutukov A. V., 1984, *ApJS*, 54, 335
- Iwamoto K., Brachwitz F., Nomoto K., Kishimoto N., Umeda H., Hix W. R., Thielemann F.-K., 1999, *ApJS*, 125, 439
- Jerkstrand A. et al., 2015a, *MNRAS*, 448, 2482
- Jerkstrand A., Ergon M., Smartt S. J., Fransson C., Sollerman J., Taubenberger S., Bersten M., Spyromilio J., 2015b, *A&A*, 573, A12
- Jerkstrand A. et al., 2015c, *ApJ*, 807, 110
- Jerkstrand A., Alsabti A. W., Murdin P. eds, 2017, *Spectra of Supernovae in the Nebular Phase*. Springer International Publishing, Switzerland, p. 1
- Kausch W. et al., 2015, *A&A*, 576, A78
- Leloudas G. et al., 2009, *A&A*, 505, 265
- Leonard D. C., 2007, *ApJ*, 670, 1275
- Liu W., Jeffery D. J., Schultz D. R., 1997, *ApJ*, 483, L107
- Livne E., Glasner A. S., 1990, *ApJ*, 361, 244
- Maeda K. et al., 2010a, *Nature*, 466, 82
- Maeda K., Taubenberger S., Sollerman J., Mazzali P. A., Leloudas G., Nomoto K., Motohara K., 2010b, *ApJ*, 708, 1703
- Maeda K. et al., 2011, *MNRAS*, 413, 3075
- Magnier E. A. et al., 2016, preprint ([arXiv: e-prints](https://arxiv.org/abs/1608.07447))
- Maguire K. et al., 2013, *MNRAS*, 436, 222
- Maguire K. et al., 2014, *MNRAS*, 444, 3258
- Maguire K., Taubenberger S., Sullivan M., Mazzali P. A., 2016, *MNRAS*, 457, 3254
- Marion G. H. et al., 2016, *ApJ*, 820, 92
- Martínez-Rodríguez H. et al., 2017, *ApJ*, 843, 35
- Mattila S., Lundqvist P., Sollerman J., Kozma C., Baron E., Fransson C., Leibundgut B., Nomoto K., 2005, *A&A*, 443, 649
- Maurer I., Jerkstrand A., Mazzali P. A., Taubenberger S., Hachinger S., Kromer M., Sim S., Hillebrandt W., 2011, *MNRAS*, 418, 1517
- Mazzali P. A., Röpke F. K., Benetti S., Hillebrandt W., 2007, *Science*, 315, 825
- Mazzali P. A. et al., 2015, *MNRAS*, 450, 2631
- Meyerott R. E., 1980, *ApJ*, 239, 257
- Modigliani A. et al., 2010, in Silva D. R., Peck A. B., Soifer B. T., Proc. SPIE Conf. Ser. Vol. 7737, *Observatory Operations: Strategies, Processes, and Systems III*. SPIE, Bellingham
- Motohara K. et al., 2006, *ApJ*, 652, L101
- Nomoto K., 1982, *ApJ*, 253, 798
- Nomoto K., Iwamoto K., Nakasato N., Thielemann F.-K., Brachwitz F., Tsujimoto T., Kubo Y., Kishimoto N., 1997, *Nucl. Phys. A*, 621, 467
- Olling R. P. et al., 2015, *Nature*, 521, 332
- Pakmor R., Kromer M., Taubenberger S., Sim S. A., Röpke F. K., Hillebrandt W., 2012, *ApJ*, 747, L10
- Pan Y.-C. et al., 2014, *MNRAS*, 438, 1391
- Pan Y.-C. et al., 2015, *MNRAS*, 452, 4307
- Penney R., Höfllich P., 2014, *ApJ*, 795, 84
- Quinet P., Le Dourneuf M., 1996, *A&AS*, 119, 99
- Quinet P., Le Dourneuf M., Zeppen C. J., 1996, *A&AS*, 120, 361
- Rosswog S., Kasen D., Guillochon J., Ramirez-Ruiz E., 2009, *ApJ*, 705, L128
- Rudy R. J., Ardila D. R., Crawford K. B., Moody M. S. L., Safrit T. K., Puetter R. C., 2015, *Astron. Telegram*, 7825
- Ruiter A. J. et al., 2013, *MNRAS*, 429, 1425
- Ruiz-Lapuente P., Lucy L. B., 1992, *ApJ*, 400, 127
- Scalzo R. et al., 2014, *MNRAS*, 440, 1498
- Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Schwarz G., 1978, *Ann. Stat.*, 6, 461
- Seitenzahl I. R. et al., 2013a, *MNRAS*, 429, 1156
- Seitenzahl I. R., Cescutti G., Röpke F. K., Ruiter A. J., Pakmor R., 2013b, *A&A*, 559, L5
- Seitenzahl I. R. et al., 2014, *MNRAS*, 444, 350
- Seitenzahl I. R. et al., 2016, *A&A*, 592, A57
- Shen K. J., Bildsten L., 2014, *ApJ*, 785, 61
- Shen K. J., Kasen D., Miles B. J., Townsley D. M., 2018, *ApJ*, 854, 52
- Silverman J. M. et al., 2012, *ApJ*, 756, L7
- Silverman J. M., Ganeshalingam M., Filippenko A. V., 2013, *MNRAS*, 430, 1030
- Silverman J. M., Vinkó J., Marion G. H., Wheeler J. C., Barna B., Szalai T., Mulligan B. W., Filippenko A. V., 2015, *MNRAS*, 451, 1973
- Sim S. A., Röpke F. K., Hillebrandt W., Kromer M., Pakmor R., Fink M., Ruiter A. J., Seitenzahl I. R., 2010, *ApJ*, 714, L52
- Smartt S. J. et al., 2015, *A&A*, 579, A40
- Smette A. et al., 2015, *A&A*, 576, A77
- Sollerman J. et al., 2004, *A&A*, 428, 555
- Spyromilio J., Meikle W. P. S., Allen D. A., Graham J. R., 1992, *MNRAS*, 258, 53P
- Spyromilio J., Gilmozzi R., Sollerman J., Leibundgut B., Fransson C., Cuby J.-G., 2004, *A&A*, 426, 547
- Sternberg A. et al., 2011, *Science*, 333, 856
- Stritzinger M., Sollerman J., 2007, *A&A*, 470, L1
- Suntzeff N. B., Candia P., Stritzinger M., 2003, *IAU Circ.*, 8171, 1
- Taubenberger S. et al., 2009, *MNRAS*, 397, 677
- Toonen S., Perets H. B., Hamers A. S., 2018, *A&A*, 610, A22
- Tyndall N. B., Ramsbottom C. A., Ballance C. P., Hibbert A., 2016, *MNRAS*, 462, 3350
- Yaron O., Gal-Yam A., 2012, *PASP*, 124, 668
- Vernet J. et al., 2011, *A&A*, 536, A105
- Webbink R. F., 1984, *ApJ*, 277, 355
- Whelan J., Iben I. J., 1973, *ApJ*, 186, 1007
- Wilk K. D., Hillier D. J., Dessart L., 2018, *MNRAS*, 474, 3187
- Yoon S.-C., Podsiadlowski P., Rosswog S., 2007, *MNRAS*, 380, 933
- Zhang J. et al., 2014, *ApJ*, 797, 5

APPENDIX: PHOTOMETRY OF SN 2013CS

Table A1. Photometry for SN 2013cs obtained with the LCO 1-m array.

MJD ^a	<i>g</i> (mag)	<i>r</i> (mag)	<i>i</i> (mag)
56427.1	14.82 ± 0.06	14.74 ± 0.08	15.17 ± 0.11
56427.7	14.71 ± 0.05	14.63 ± 0.06	15.04 ± 0.17
56428.8	14.50 ± 0.05	14.47 ± 0.07	14.80 ± 0.12
56431.1	14.15 ± 0.04	14.14 ± 0.04	–
56448.0	14.45 ± 0.07	14.48 ± 0.09	15.27 ± 0.06
56453.7	14.82 ± 0.04	14.59 ± 0.06	15.43 ± 0.20
56454.0	14.91 ± 0.05	14.64 ± 0.06	15.31 ± 0.12
56459.0	15.29 ± 0.06	–	15.18 ± 0.06
56461.0	15.43 ± 0.06	14.92 ± 0.06	–
56463.0	15.63 ± 0.05	14.58 ± 0.08	15.13 ± 0.09
56465.0	15.70 ± 0.05	14.82 ± 0.06	15.00 ± 0.10
56467.0	15.99 ± 0.04	14.99 ± 0.06	15.17 ± 0.13
56469.0	16.22 ± 0.05	15.16 ± 0.08	15.31 ± 0.09

^aMJD = Modified Julian date.

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