




# ALMA ACA study of the H<sub>2</sub>S/OCS ratio in low-mass protostars

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Received 29 September 2022 / Accepted 14 February 2023

## ABSTRACT

**Context.** The identification of the main sulfur reservoir on its way from the diffuse interstellar medium to the cold dense star-forming cores and, ultimately, to protostars is a long-standing problem. Despite sulfur's astrochemical relevance, the abundance of S-bearing molecules in dense cores and regions around protostars is still insufficiently constrained.

**Aims.** The goal of this investigation is to derive the gas-phase H<sub>2</sub>S/OCS ratio for several low-mass protostars, which could provide crucial information about the physical and chemical conditions in the birth cloud of Sun-like stars. This may also shed new light onto the main sulfur reservoir in low-mass star-forming systems.

**Methods.** Using Atacama Large Millimeter/submillimeter Array (ALMA) Atacama Compact Array (ACA) Band 6 observations, we searched for H<sub>2</sub>S, OCS, and their isotopologs in ten Class 0/I protostars with different source properties such as age, mass, and environmental conditions. The sample contains IRAS 16293-2422 A, IRAS 16293-2422 B, NGC 1333-IRAS 4A, RCrA IRS7B, Per-B1-c, BHR71-IRS1, Per-emb-25, NGC 1333-IRAS4B, Ser-SMM3, and TMC1. A local thermal equilibrium (LTE) model is used to fit synthetic spectra to the detected lines and to derive the column densities based solely on optically thin lines.

**Results.** The H<sub>2</sub>S and OCS column densities span four orders of magnitude across the sample. The H<sub>2</sub>S/OCS ratio is found to be in the range from 0.2 to above 9.7. IRAS 16293-2422 A and Ser-SMM3 have the lowest ratio, while BHR71-IRS1 has the highest. Only the H<sub>2</sub>S/OCS ratio of BHR71-IRS1 is in agreement with the ratio in comet 67P/Churyumov–Gerasimenko within the uncertainties.

**Conclusions.** The determined gas-phase H<sub>2</sub>S/OCS ratios can be below the upper limits on the solid-state ratios by as much as one order of magnitude. The H<sub>2</sub>S/OCS ratio depends in great measure on the environment of the birth cloud, such as UV-irradiation and heating received prior to the formation of a protostar. The highly isolated birth environment (a Bok globule) of BHR71-IRS1 is hypothesized as the reason for its high gaseous H<sub>2</sub>S/OCS ratio that is due to lower rates of photoreactions and more efficient hydrogenation reactions under such dark, cold conditions. The gaseous inventory of S-bearing molecules in BHR71-IRS1 appears to be the most similar to that of interstellar ices.

**Key words.** astrochemistry – line: identification – instrumentation: interferometers – ISM: molecules – stars: protostars

## 1. Introduction

Sulfur (S) is the tenth most abundant element in the Universe (S/H ~ 1.35 × 10<sup>-5</sup>, Yamamoto 2017). It was first detected as carbon monosulfide (CS) in the interstellar medium (Penzias et al. 1971). Since then, S-bearing species have been detected in different regions including molecular clouds (Navarro-Almáida et al. 2020; Spezzano et al. 2022), hot cores (Blake et al. 1987; Charnley 1997; Li et al. 2015; Codella et al. 2021; Drozdovskaya et al. 2018), comets (Smith et al. 1980; Bockelée-Morvan et al. 2000; Biver et al. 2021a,b), and starburst galaxies (NGC 253; Martín et al. 2005). The total abundance of an element in dust, ice, and gas is its cosmic abundance, also called its elemental abundance. The gas-phase abundance of atomic sulfur in diffuse clouds is comparable to the cosmic abundance of sulfur (~ 10<sup>-5</sup>; Savage & Sembach 1996; Howk et al. 2006). However, the observed abundance of S-bearing species in dense cores and protostellar environments is lower by a factor of ~ 1000 (Snow et al. 1986; Tieftrunk et al. 1994; Goicoechea et al. 2006; Agúndez et al. 2018) in comparison to the total S-abundance in diffuse clouds. The forms and mechanisms behind this sulfur depletion

in star-forming regions are still unknown. This is often referred to as the “missing sulfur problem.”

Different chemical models have been used to investigate this unknown form of sulfur (Woods et al. 2015; Vidal et al. 2017; Semenov et al. 2018; Vidal & Wakelam 2018; Laas & Caselli 2019). Vidal et al. (2017) proposed that a notable amount of sulfur is locked up in either HS and H<sub>2</sub>S ices or gaseous atomic sulfur in cores, depending in large part on the age of the molecular cloud. However, the only solid form of sulfur firmly detected in interstellar ices is OCS (Palumbo et al. 1995, 1997; Aikawa et al. 2012; Boogert et al. 2015) as well as potentially SO<sub>2</sub> (Boogert et al. 1997; Zasowski et al. 2009; Yang et al. 2022; McClure et al. 2023). At the same time, solid-state H<sub>2</sub>S detection remains tentative to date (Geballe et al. 1985; Smith 1991). The initial cloud abundance of S-bearing molecules has been shown to set the subsequent abundances of these molecules in protostellar regions, depending on the free-fall timescales (Vidal & Wakelam 2018). In surface layers of protoplanetary disks, the availability of gaseous S-bearing molecules appears to be strongly linked with the availability of oxygen (Semenov et al. 2018). Observational studies of gas-phase species claim either

$\text{H}_2\text{S}$  (Holdship et al. 2016) or OCS (van der Tak et al. 2003) as the main S-carrier depending on the environment being observed. Other possible reservoirs of sulfur have been proposed in the form of semi-refractory polymers up to  $\text{S}_8$  (A'Hearn et al. 1983; Druard & Wakelam 2012; Calmonte et al. 2016; Shingledecker et al. 2020), hydrated sulfuric acid (Scappini et al. 2003), atomic sulfur (Anderson et al. 2013), and mineral sulfides, FeS (Keller et al. 2002; Köhler et al. 2014; Kama et al. 2019). On the other hand, chemical models of the evolution from cloud to dense core with updated chemical networks suggest that sulfur is merely partitioned over a diverse set of simple organo-sulfur ices and no additional form is required (Laas & Caselli 2019). Matching observed and modeled cloud abundances consistently for the full inventory of gaseous S-bearing molecules to better than a factor of 10 remains challenging (Navarro-Almida et al. 2020). Laboratory experiments point to the importance of the photodissociation of  $\text{H}_2\text{S}$  ice by UV photons leading to the production of OCS ice (Ferrante et al. 2008; Garozzo et al. 2010; Jiménez-Escobar & Muñoz Caro 2011; Chen et al. 2015) and  $\text{S}_2$  in mixed ices (Grim & Greenberg 1987). Calmonte et al. (2016) claim to have recovered the full sulfur inventory in comets.

Sulfur-bearing species have been proposed to probe the physical and chemical properties of star-forming regions and to even act as chemical clocks (Charnley 1997; Hatchell et al. 1998; Viti et al. 2001; Li et al. 2015). However, it has since been shown that their abundance is sensitive to the gas-phase chemistry and the availability of atomic oxygen, which puts their reliability as chemical clocks into question (Wakelam et al. 2004, 2011). Studying S-bearing molecules in young Class 0/I protostars is crucial for two reasons. Firstly, their inner hot regions thermally desorb all the volatile ices that are otherwise hidden from gas-phase observations. Consequently, it is more likely to be able to probe the full volatile inventory of S-bearing molecules and investigate the “missing sulfur” reservoir. Secondly, these targets are a window onto the materials available for the assembly of the protoplanetary disk midplane and the cometesimals therein (Aikawa & Herbst 1999; Willacy 2007; Willacy & Woods 2009). This makes hot inner regions highly suitable targets for comparative studies with comets (Bockelée-Morvan et al. 2000; Drozdovskaya et al. 2019).

The main goal of this paper is to study the physical and chemical conditions in embedded protostars via the  $\text{H}_2\text{S}/\text{OCS}$  ratio. A sample of ten Class 0/I low-mass protostars with different physical properties (mass, age, and environment) is considered. Such protostars are in their earliest phase of formation after collapse with large envelope masses. In this work, Atacama Large Millimeter/submillimeter Array (ALMA) Atacama Compact Array (ACA) Band 6 observations towards these 10 protostars are utilized. The  $\text{H}_2\text{S}/\text{OCS}$  ratio is calculated from the column densities of  $\text{H}_2\text{S}$ , OCS, and their isotopologs. The details of the observations, model, and model parameters used for synthetic spectral fitting are introduced in Sect. 2. The detected lines of major and minor isotopologs of  $\text{H}_2\text{S}$  and OCS, their characteristics, and  $\text{H}_2\text{S}/\text{OCS}$  line ratios are presented in Sect. 3. The discussion and conclusions are presented in Sects. 4 and 5, respectively.

## 2. Methods

### 2.1. Observations

Two sets of observations are jointly analyzed in this paper. The first data set (project-id: 2017.1.00108.S; PI:

M. N. Drozdovskaya) targets IRAS 16293–2422, NGC 1333–IRAS4A, and RCrA IRS7B. The relevant observations were carried out in Band 6 (211–275 GHz) with the ALMA ACA of 7 m dishes. The data set has a spectral resolution of  $0.079\text{--}0.085\text{ km s}^{-1}$  (61 kHz), and a spatial resolution of  $(6.5\text{--}9.0)\times(4.0\text{--}6.3)''$ . The second data set (project-id: 2017.1.1350.S; PI: Ł. Tychoniec) targets Per-B1-c, BHR71, Per-emb-25, NGC 1333–IRAS4B, Ser-SMM3, and TMC1 with the ALMA ACA 7m dishes, also in Band 6. The data have a similar spatial resolution,  $(6.1\text{--}7.4)\times(4.5\text{--}6.4)''$ , but a lower spectral resolution of  $0.333\text{--}0.678\text{ km s}^{-1}$  (244–488 kHz). The observed frequency ranges of the data are given in Table 1. The data cubes were processed through the standard pipeline calibration with CASA 5.4.0–68. For each source, the noise level has been calculated by taking the standard deviation of the flux in the frequency ranges where no emission lines were detected, that is, regions with pure noise in the spectral window containing the  $\text{H}_2\text{S}$ ,  $2_{2,0}\text{--}2_{1,1}$  line. The noise level of the first data set is between 21 and 32 mJy beam $^{-1}$  channel $^{-1}$  and that of the second data set is between 7 and 13 mJy beam $^{-1}$  channel $^{-1}$  (Table 3). Both data sets have a flux uncertainty of 10%. The largest resolvable scale of the first and the second data sets are  $26.2\text{--}29.2''$  and  $24.6\text{--}29.0''$ , respectively.

### 2.2. Sources

The properties of the sources explored in this work are tabulated in Table 2. IRAS 16293–2422 (hereafter, IRAS 16293) is a triple protostellar source, consisting of protostars A and B, separated by  $5.3''$  (747 au; van der Wiel et al. 2019), and disk-like structures around the two sources, located in the Rho Ophiuchi star-forming region at a distance of 141 pc (Dzib et al. 2018). This source was studied thoroughly using ALMA under the Protostellar Interferometric Line Survey (PILS; Jørgensen et al. 2016) and many preceding observational campaigns (e.g., van Dishoeck et al. 1995; Caux et al. 2011). Both hot corinos around A and B are rich in a diverse set of complex organic molecules (Jørgensen et al. 2018; Manigand et al. 2020). The source IRAS 16293 A is itself a binary composed of sources A1 and A2 with a separation of  $0.38''$  (54 au; Maureira et al. 2020). IRAS4A is also a binary, made up of IRAS4A1 and IRAS4A2, separated by  $1.8''$  (540 au; Sahu et al. 2019) in the Perseus molecular cloud, located at a distance of 299 pc (Zucker et al. 2018) in the south-eastern edge of the complex NGC 1333 (Looney et al. 2000). IRAS4A1 has a much higher dust density in its envelope than IRAS4A2, but both contain complex organic molecules (Sahu et al. 2019; De Simone et al. 2020). IRS7B is a low-mass source, with a separation of  $14''$  (2 000 au) from IRS7A (Brown 1987), and  $\sim 8''$  (1000 au) from CXO 34 (Lindberg et al. 2014). It is situated in the Corona Australis dark cloud at a distance of 130 pc (Neuhäuser & Forbrich 2008). IRS7B has been shown to contain lower complex organic abundances as a result of being located in a strongly irradiated environment (Lindberg et al. 2015).

From the second set of sources, IRAS4B (sometimes labeled BI) has a binary component B' (or BII) that is  $11''$  (3300 au) away (Sakai et al. 2012; Anderl et al. 2016; Tobin et al. 2016). The separation between IRAS4B and IRAS4A is  $31''$  (9300 au; Coutens et al. 2013). IRAS4B displays emission from complex organic molecules (Belloche et al. 2020) and powers a high-velocity SiO jet (Podio et al. 2021). B1-c is an isolated deeply embedded protostar in the Barnard 1 clump in the western part of the Perseus molecular cloud at a distance of 301 pc (Zucker et al. 2018). B1-c contains emission from complex organic molecules and shows a high velocity outflow (Jørgensen et al. 2006;

**Table 1.** Spectral settings of the data sets.

Sky frequency (GHz)	Channel width		Number of channels
	(kHz)	(km s <sup>-1</sup> )	
Project-id: 2017.1.00108.S			
214.250-214.500	61	0.085	4096
216.646-216.771	61	0.084	2048
215.469-215.532	61	0.085	1024
216.114-216.176	61	0.084	1024
231.550-231.612	61	0.079	1024
231.308-231.370	61	0.079	1024
231.027-231.090	61	0.079	1024
230.284-230.346	61	0.080	1024
232.062-233.932 (continuum)	488	0.628	4096
Project-id: 2017.1.1350.S			
216.974-217.222	488	0.674	512
216.513-216.762	488	0.676	512
217.692-217.941	488	0.672	512
215.982-216.230	488	0.678	512
219.492-219.616	244	0.334	512
219.882-220.006	244	0.333	512
218.692-218.817	244	0.335	512
220.331-220.449	244	0.332	512
230.470-230.707	488	0.635	512
231.159-231.408	488	0.633	512
230.931-231.180	488	0.633	512
232.136-232.384	488	0.630	512
233.069-234.937 (continuum)	977	1.260	2048

**Table 2.** Observed Class 0/I protostellar systems.

Source	$d$ (pc)	$M_{\text{env}}$ ( $M_{\odot}$ )	$L_{\text{bol}}$ ( $L_{\odot}$ )	$T_{\text{bol}}$ (K)	Class	$v_{\text{LSR}}$ (km s <sup>-1</sup> )
IRAS 16293-2422 A	141 <sup>(a)</sup>	4.0 <sup>(d)</sup>	~18 <sup>(b)</sup>	–	0	+3.2 <sup>(c)</sup>
IRAS 16293-2422 B	141 <sup>(a)</sup>	4.0 <sup>(d)</sup>	~3 <sup>(b)</sup>	–	0	+2.7 <sup>(c)</sup>
NGC 1333-IRAS4A	299 <sup>(e)</sup>	5.6 <sup>(f)</sup>	9.1 <sup>(f)</sup>	29 <sup>(g)</sup>	0	+7.2 <sup>(o)</sup>
RCrA IRS7B	130 <sup>(i)</sup>	2.2 <sup>(j)</sup>	4.6 <sup>(j)</sup>	89 <sup>(j)</sup>	0/I	+5.8 <sup>(t)</sup>
Per-B1-c	301 <sup>(e)</sup>	1.8 <sup>(h)</sup>	3.84 <sup>(h)</sup>	48 <sup>(k)</sup>	0	+6.4 <sup>(k,l)</sup>
BHR71-IRS1	200 <sup>(m,n)</sup>	2.7 <sup>(o)</sup>	15 <sup>(o)</sup>	44 <sup>(o)</sup>	0	–4.4 <sup>(o)</sup>
Per-emb-25	294 <sup>(p)</sup>	0.5 <sup>(h)</sup>	1.0 <sup>(h)</sup>	68 <sup>(h)</sup>	0/I	+5.8 <sup>(k)</sup>
NGC 1333-IRAS4B	299 <sup>(e)</sup>	3.0 <sup>(o)</sup>	4.4 <sup>(o)</sup>	28 <sup>(g)</sup>	0	+7.4 <sup>(o)</sup>
Ser-SMM3	436 <sup>(q)</sup>	3.2 <sup>(o)</sup>	5.1 <sup>(o)</sup>	38 <sup>(o)</sup>	0	+7.6 <sup>(o)</sup>
TMC1	140 <sup>(r,s)</sup>	0.2 <sup>(o)</sup>	0.9 <sup>(o)</sup>	101 <sup>(o)</sup>	I	+5.2 <sup>(o)</sup>

**Notes.** The columns represent (1) source, (2)  $d$ : distance to the source in pc, (3)  $M_{\text{env}}$ : mass of the envelope in  $M_{\odot}$ , (4)  $L_{\text{bol}}$ : bolometric luminosity in  $L_{\odot}$ , (5)  $T_{\text{bol}}$ : bolometric temperature in K, (6) Class: stage of the protostar, (7)  $v_{\text{LSR}}$ : local standard of rest velocity in km s<sup>-1</sup>.

**References.** <sup>(a)</sup>Dzib et al. (2018), <sup>(b)</sup>Jacobsen et al. (2018), <sup>(c)</sup>Jørgensen et al. (2011), <sup>(d)</sup>van der Wiel et al. (2019), <sup>(e)</sup>Zucker et al. (2018), <sup>(f)</sup>Taquet et al. (2015), <sup>(g)</sup>Tobin et al. (2016), <sup>(h)</sup>Enoch et al. (2009), <sup>(i)</sup>Neuhäuser & Forbrich (2008), <sup>(j)</sup>Lindberg et al. (2014), <sup>(k)</sup>Stephens et al. (2019), <sup>(l)</sup>Hatchell et al. (2007b), <sup>(m)</sup>Seidensticker & Schmidt-Kaler (1989), <sup>(n)</sup>Straizys et al. (1994), <sup>(o)</sup>Kristensen et al. (2012), <sup>(p)</sup>Zucker et al. (2019), <sup>(q)</sup>Ortiz-León et al. (2018), <sup>(r)</sup>Elias (1978), <sup>(s)</sup>Torres et al. (2009), <sup>(t)</sup>Lindberg et al. (2015).

van Gelder et al. 2020). The next closest source, B1-a, is ~100'' (~29 500au) away (Jørgensen et al. 2006). BHR71 is a Bok globule in the Southern Coalsack dark nebulae at a distance of ~200 pc (Seidensticker & Schmidt-Kaler 1989; Straizys et al. 1994). It hosts the wide binary system of IRS1 and IRS2 with a separation of 16'' (3200 au; Bourke 2001; Parise et al. 2006; Chen et al. 2008; Tobin et al. 2019). IRS1 displays pronounced

emission from complex organic molecules (Yang et al. 2020). Emb-25 is a single source located in the Perseus molecular cloud (Enoch et al. 2009; Tobin et al. 2016). It does not show emission from complex organic molecules (Yang et al. 2021), but powers low-velocity CO outflows (Stephens et al. 2019). TMC1 is a Class I binary source, located in the Taurus molecular cloud (Chen et al. 1995; Brown & Chandler 1999) at a distance of

**Table 3.** Location of the center position (in RA and Dec) and radius (in au and arcsecond) of each circular region, from which the spectra are extracted for the studied protostars.

Source	RA	Dec	Pixel size ( $''$ )	Radius		FWHM ( $\text{km s}^{-1}$ )	Noise level	
	(J2000)	(J2000)		(au)	( $''$ )		( $\text{mJy beam}^{-1} \text{ channel}^{-1}$ )	( $\text{mJy beam}^{-1} \text{ km s}^{-1}$ )
IRAS 16293-2422 A	16 h 32 m 22.854 s	$-24^{\circ}28' 36.465''$	0.8	522	3.7	4.5	28	17.0
IRAS 16293-2422 B	16 h 32 m 22.671 s	$-24^{\circ}28' 33.145''$	0.8	522	3.7	1.0	32	9.2
NGC 1333-IRAS4A	03 h 29 m 10.509 s	$+31^{\circ}13' 30.918''$	1.1	822	3.5	1.8	29	11.2
RCrA IRS7B	19 h 01 m 56.402 s	$-36^{\circ}57' 28.276''$	0.8	730	4.3	1.0	21	6.0
Per-B1-c	03 h 33 m 17.880 s	$+31^{\circ}09' 31.795''$	1.2	879	3.0	2.2	11	13.4
BHR71-IRS1	12 h 01 m 36.516 s	$-65^{\circ}08' 49.298''$	1.0	700	3.5	2.5	7	9.2
Per-emb-25	03 h 26 m 37.514 s	$+30^{\circ}15' 27.792''$	1.1	600	3.0	1.0	11	9.0
NGC 1333-IRAS4B	03 h 29 m 12.019 s	$+31^{\circ}13' 08.010''$	1.2	879	3.0	2.0	11	12.9
Ser-SMM3	18 h 29 m 59.311 s	$+01^{\circ}14' 00.365''$	0.9	1526	3.5	2.5	10	13.5
TMC1	04 h 41 m 12.700 s	$+25^{\circ}46' 34.800''$	1.1	303	3.0	1.0	13	10.6

**Notes.** The pixel size (in arcsecond) for each source is also given. The noise levels in  $\text{mJy beam}^{-1} \text{ channel}^{-1}$  and  $\text{mJy beam}^{-1} \text{ km s}^{-1}$  of the protostars are deduced according to  $\sqrt{\frac{\sum (\text{flux in line-free channel } i)^2}{\text{number of line-free channels}}}$  and noise ( $\text{mJy}^{-1} \text{ beam}^{-1} \text{ channel}^{-1}$ )  $\times \sqrt{n} \times \text{spectral resolution}$  ( $\text{km s}^{-1}$ ), where  $n$  is the number of channels computed with  $\frac{\text{FWHM} (\text{km s}^{-1})}{\text{spectral resolution} (\text{km s}^{-1})}$ , using the FWHM of the  $\text{H}_2\text{S}$  line at 216.710 GHz, respectively.

140 pc (Elias 1978; Torres et al. 2009). The separation between the two components, TMC1E and TMC1W, is  $\sim 0.6''$  ( $\sim 85$  au); and neither of the two display complex organic emission (van't Hoff et al. 2020). SMM3 is a single, embedded protostar located in the SE region in Serpens region; 436 pc away (Ortiz-León et al. 2018). The next closest-lying source is SMM6 at a separation of  $20''$  ( $\sim 8700$  au; Davis et al. 1999; Kristensen et al. 2010; Mirocha et al. 2021). SMM3 launches a powerful jet (Tychoniec et al. 2021), but does not display complex organic molecule emission, which may be obscured by the enveloping dust (van Gelder et al. 2020).

### 2.3. Synthetic spectral fitting

For the spectral analysis, on-source spectra were extracted from the data cubes of the sources of the two data sets. Circular regions centered on source positions (in RA and Dec) with the radius of spectrum extraction corresponding to one-beam on-source are given in Table 3. The number of pixels in radius  $r$  of the circular region to be used was computed by dividing the radius of the circular region with the size of one pixel in arcsecond. The spectroscopy used for the targeted molecules and their isotopologs stems from the Cologne Database of Molecular Spectroscopy (CDMS; Müller et al. 2001, 2005; Endres et al. 2016)<sup>1</sup> and the Jet Propulsion Laboratory (JPL) catalog (Pickett et al. 1998)<sup>2</sup>. Line blending in the detected lines was checked with the online database Splatalogue<sup>3</sup>.

Synthetic spectral fitting was performed with custom-made Python scripts based on the assumption of local thermal equilibrium (LTE). The input parameters include the full width half-maximum (FWHM) of the line, column density ( $N$ ), excitation temperature ( $T_{\text{ex}}$ ), source size, beam size, and spectral resolution of the observations. Line profiles are assumed to be Gaussian. Further details are provided in Sect. 2.3 of Drozdovskaya et al. (2022). For some sources, the number of free parameters (e.g., source size or  $T_{\text{ex}}$ ) can be reduced based on information from other observing programs. These are detailed

on a source-by-source basis in the corresponding Appendices. Typically, just two free parameters were fitted at a time by means of a visual inspection and an exploration of a grid of possible values. The considered range for  $N$  was  $10^{13}$ – $10^{19} \text{ cm}^{-2}$  in steps of  $0.1 \times N$ . Simultaneously, the FWHM of the synthetic fit was adjusted to match the FWHM of the detected line. For some sources (e.g., NGC 1333-IRAS4A and Ser-SMM3), the excitation temperature could not be constrained. Hence, a grid of excitation temperatures was considered with a range between 50 and 300 K.

The line optical depth ( $\tau$ ) was calculated for the best-fitting combination of parameters to check for optical thickness. If a transition of a certain molecule was found to be optically thick, its column density was computed as the average of the column density of the main species derived from its minor isotopologs, given by:

$$\overline{N(X)} = \frac{1}{n} \sum_{i=0}^n N(X)_i, \quad (1)$$

where  $X$  is  $\text{H}_2\text{S}$  or  $\text{OCS}$  and  $N(X)_i$  is the column density of  $\text{H}_2\text{S}$  or  $\text{OCS}$  derived from its minor isotopologs. Adopted isotopic ratios for the derivation of main isotopologs from minor isotopologs are given alongside Table 4.

## 3. Results

The spectral setup of the first data set allows the targeted sources to be probed for the emission of the main isotopologs of  $\text{H}_2\text{S}$  and  $\text{OCS}$ ,  $v=0$ , their minor isotopologs ( $\text{HDS}$ ,  $\text{HD}^{34}\text{S}$ ,  $\text{H}_2^{33}\text{S}$ ,  $\text{H}_2^{34}\text{S}$ ,  $^{18}\text{OCS}$ ,  $\text{O}^{13}\text{CS}$ ,  $\text{OC}^{33}\text{S}$ ,  $^{18}\text{OC}^{34}\text{S}$ ), and also the vibrationally excited states of  $\text{OCS}$  ( $v_2=1^{\pm}$ ). Consequently, sources IRAS 16293 A, IRAS 16293 B, IRAS4A, and IRS7B were probed for all these species.

The spectral setup of the second data set allows the other targeted sources (B1-c, BHR71-IRS1, Per-emb-25, IRAS4B, SMM3, and TMC1) to be probed for the main isotopologs of  $\text{H}_2\text{S}$  and  $\text{OCS}$ ,  $v=0$ , their minor isotopologs ( $\text{HDS}$ ,  $\text{HD}^{34}\text{S}$ ,  $^{18}\text{OCS}$ ,  $\text{OC}^{33}\text{S}$ ,  $^{18}\text{OC}^{34}\text{S}$ ,  $^{18}\text{O}^{13}\text{CS}$ ), as well as the vibrationally excited state of  $\text{OCS}$  ( $v_2=1^-$ ). All the transitions of the detected molecules have  $E_{\text{up}}$  in the range of 84–123 K and  $A_{ij}$  values

<sup>1</sup> <https://cdms.astro.uni-koeln.de/>

<sup>2</sup> <https://spec.jpl.nasa.gov/>

<sup>3</sup> <https://splatalogue.online/>



**Table 4.** Results from the modeling of synthetic spectra of the detected S-bearing species toward IRAS 16293 B for source sizes of 1'' and 2'', an excitation temperature of 125 K, and a line width of 1 km s<sup>-1</sup>.

Species	Transition	Freq. (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ (s <sup>-1</sup> )	Beam size ('')	$N$ (cm <sup>-2</sup> )		Derived $N$ of isotopologs (cm <sup>-2</sup> )		$\tau$	
						(7)	(8)	(9)	(10)		
(1)	(2)	(3)	(4)	(5)	(6)	1''	2''	1''	2''	1''	2''
H <sub>2</sub> S	2 <sub>2,0</sub> -2 <sub>1,1</sub>	216.710	84	4.9×10 <sup>-5</sup>	6.0	op. thick	op. thick	$\overline{N(\text{H}_2\text{S})} = (3.6 \pm 0.6) \times 10^{17}$	$\overline{N(\text{H}_2\text{S})} = (9.2 \pm 1.7) \times 10^{16}$	30.0	8.00
H <sub>2</sub> <sup>33</sup> S	2 <sub>2,0,1</sub> -2 <sub>1,1,2</sub>	215.494	84	2.4×10 <sup>-5</sup>	6.0	2.7 <sup>+0.7</sup> <sub>-0.4</sub> ×10 <sup>15</sup>	7.0 <sup>+1.7</sup> <sub>-1.1</sub> ×10 <sup>14</sup>	$N(\text{H}_2\text{S}) = 3.4^{+0.8}_{-0.5} \times 10^{17}$ (a)	$N(\text{H}_2\text{S}) = 8.8^{+2.1}_{-1.4} \times 10^{16}$ (a)	0.02	0.004
	2 <sub>2,0,1</sub> -2 <sub>1,1,1</sub>	215.497	84	2.4×10 <sup>-5</sup>						0.02	0.004
	2 <sub>2,0,4</sub> -2 <sub>1,1,3</sub>	215.501	84	6.9×10 <sup>-6</sup>						0.02	0.005
	2 <sub>2,0,4</sub> -2 <sub>1,1,4</sub>	215.503	84	4.2×10 <sup>-5</sup>						0.10	0.030
	2 <sub>2,0,2</sub> -2 <sub>1,1,3</sub>	215.504	84	1.7×10 <sup>-5</sup>						0.02	0.010
	2 <sub>2,0,2</sub> -2 <sub>1,1,2</sub>	215.505	84	1.9×10 <sup>-5</sup>						0.03	0.010
	2 <sub>2,0,2</sub> -2 <sub>1,1,1</sub>	215.508	84	1.2×10 <sup>-5</sup>						0.02	0.004
	2 <sub>2,0,3</sub> -2 <sub>1,1,3</sub>	215.512	84	2.8×10 <sup>-5</sup>						0.06	0.020
	2 <sub>2,0,3</sub> -2 <sub>1,1,2</sub>	215.513	84	1.1×10 <sup>-5</sup>						0.02	0.010
2 <sub>2,0,3</sub> -2 <sub>1,1,4</sub>	215.513	84	9.1×10 <sup>-6</sup>						0.02	0.005	
H <sub>2</sub> <sup>34</sup> S	2 <sub>2,0</sub> -2 <sub>1,1</sub>	214.377	84	4.7×10 <sup>-5</sup>	6.0	op. thick	>1.5×10 <sup>15</sup>			1.00	0.20
OCS, $v=0$	19-18	231.061	111	3.6×10 <sup>-5</sup>	5.6	op. thick	op. thick	$\overline{N(\text{OCS})} = (2.7 \pm 0.3) \times 10^{17}$	$\overline{N(\text{OCS})} = (7.0 \pm 0.8) \times 10^{16}$	46.0	11.0
OC <sup>33</sup> S	18-17	216.147	99	2.9×10 <sup>-5</sup>	6.0	>2.4×10 <sup>15</sup>	>5.6×10 <sup>14</sup>			0.40	0.10
O <sup>13</sup> CS	19-18	230.318	110	3.5×10 <sup>-5</sup>	5.7	>3.8×10 <sup>15</sup>	>8.2×10 <sup>14</sup>			0.60	0.10
<sup>18</sup> OCS	19-18	216.753	104	3.0×10 <sup>-5</sup>	6.0	4.7 <sup>+0.6</sup> <sub>-0.3</sub> ×10 <sup>14</sup>	1.2 <sup>+0.2</sup> <sub>-0.1</sub> ×10 <sup>14</sup>	$N(\text{OCS}) = 2.6^{+0.4}_{-0.2} \times 10^{17}$ (b)	$N(\text{OCS}) = 6.7^{+1.8}_{-0.6} \times 10^{16}$ (b)	0.07	0.02
OCS, $v_2 = 1^-$	19-18	231.342	860	3.5×10 <sup>-5</sup>	5.6	8.5 <sup>+14.5</sup> <sub>-4.3</sub> ×10 <sup>16</sup>	2.5 <sup>+3.3</sup> <sub>-1.4</sub> ×10 <sup>16</sup>			0.03	0.01
OCS, $v_2 = 1^+$	19-18	231.584	860	3.5×10 <sup>-5</sup>	5.6					0.03	0.01

**Notes.** Columns in the table of line parameters denote the following: (1) detected S-bearing species, (2) transition of the emission line, (3) frequency of the emission line, (4) upper energy level, (5) Einstein A coefficient, (6) beam size, (7) column density, (8) derived column density of the main isotopologs, and (9) optical depth of the emission line. Directly across from a specific minor isotopolog under “Derived  $N$  of isotopologs” comes the column density of the main isotopolog upon the assumption of the standard isotopic ratio. In bold, in the same column, we have the average column density of the main isotopolog based on all the available minor isotopologs (only if the minor isotopolog is optically thin and including the uncertainties). The isotopic ratios assumed to derive column densities of the main isotopologs from their minor isotopologs: (<sup>a</sup>)<sup>32</sup>S/<sup>33</sup>S = 125 (Asplund et al. 2009), (<sup>b</sup>)<sup>16</sup>O/<sup>18</sup>O = 557 (Wilson 1999). Other isotopic ratios considered in the analysis of the data are <sup>12</sup>C/<sup>13</sup>C = 69 (Wilson 1999) and <sup>32</sup>S/<sup>34</sup>S = 22 (Wilson 1999).

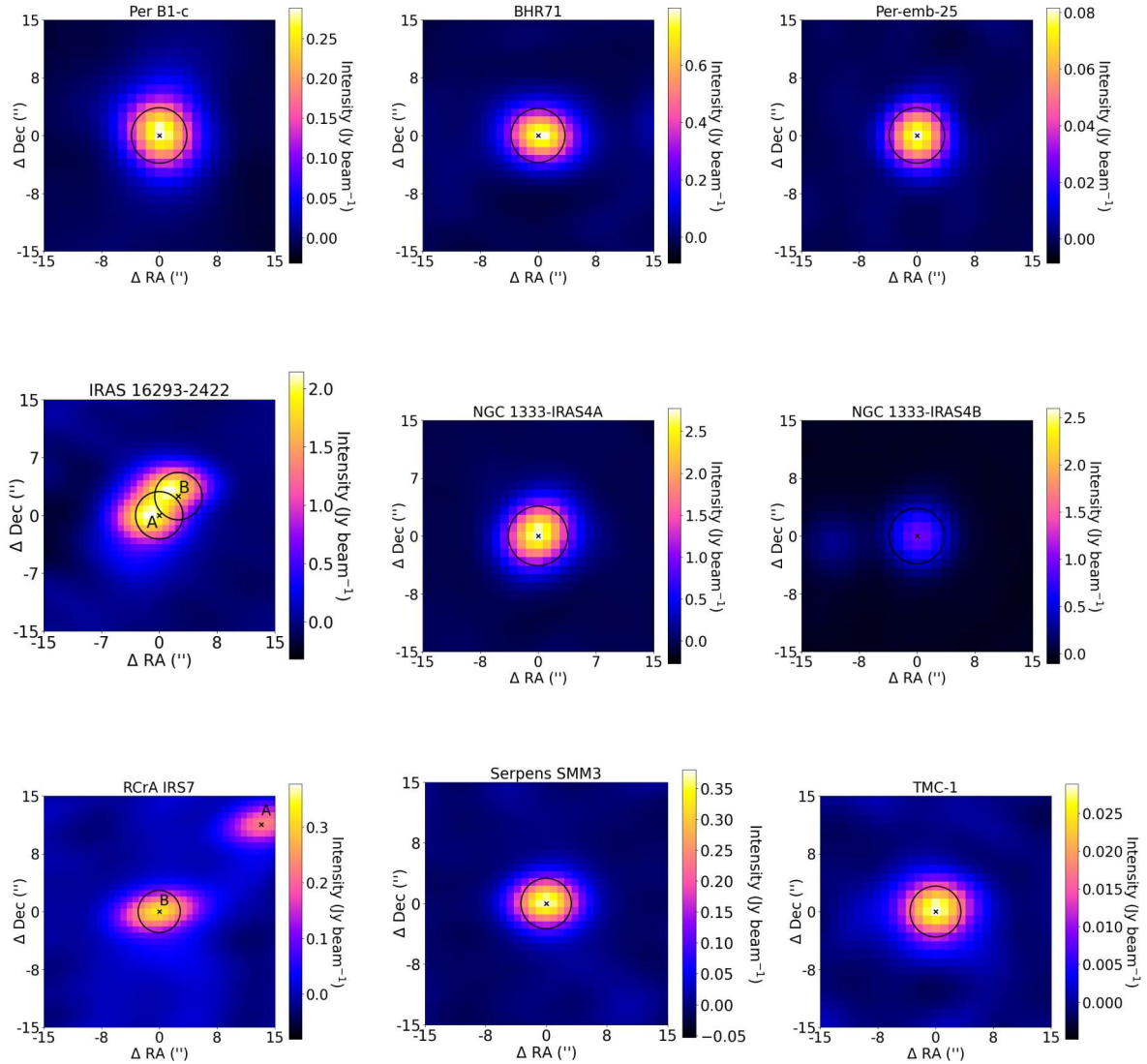
of 0.69–4.9 × 10<sup>-5</sup> s<sup>-1</sup>. The details of the targeted molecular lines are presented in Appendix A. We note that the HDS lines probed in the two data sets are not the same – the first data set was probed for HDS, 14<sub>2,12</sub>–13<sub>4,9</sub> transition at rest frequency 214.325 GHz, while the second data set was probed for HDS, 7<sub>3,4</sub>–7<sub>3,5</sub> and 12<sub>5,7</sub>–12<sub>5,8</sub> transitions at 234.046 and 234.528 GHz, respectively. Nevertheless, the  $E_{\text{up}}$  is high (>400 K) for all three transitions of HDS; and it was not detected in any of these lines in any of the sources. The HD<sup>34</sup>S and OCS,  $v_2 = 1^\pm$  transitions also have high  $E_{\text{up}}$  (>400 K). Also, HD<sup>34</sup>S was not detected in any of these lines in any of the sources, but OCS,  $v_2 = 1^\pm$  was detected in IRAS 16293 A, IRAS 16293 B, and IRAS4A (owing to high OCS column densities and higher sensitivity of the first data set).

All the main and minor isotopologs were detected in IRAS 16293 A, IRAS 16293 B, and IRAS4A except HDS, HD<sup>34</sup>S, <sup>18</sup>OC<sup>34</sup>S, and <sup>18</sup>O<sup>13</sup>CS. In IRS7B, only H<sub>2</sub>S was detected, the rest of the molecular lines were undetected including OCS,  $v=0$ . However, only the main S-bearing species, H<sub>2</sub>S and OCS,  $v=0$ , were detected in B1-c, BHR71-IRS1, and SMM3. IRAS4B showed the rotational transition of OC<sup>33</sup>S ( $J = 18$ – $17$ ), in addition to the H<sub>2</sub>S and OCS,  $v=0$  lines. Then, Emb-25 and TMC1 showed no detections of main S-bearing species and their minor isotopologs. Thus, 1- $\sigma$  upper limits on the column densities of H<sub>2</sub>S and OCS,  $v=0$  were derived for Emb-25 and TMC1; and an upper limit on the column density of OCS,  $v=0$  was derived for IRS7B. Table A.1 provides the CDMS entry, transition

quantum numbers, rest frequency, upper energy level, Einstein A coefficient, and the detection or non-detection of each line of all the targeted S-bearing molecules toward all of the sources in the sample.

In Fig. 1, the pipeline-produced integrated intensity maps (with the line channels excluded) are shown for all the sources. These are dominated by the dust emission, but with some degree of contamination by the line emission, especially for some of the line-rich sources. The circular regions used to extract the spectra of each individual source are also shown. The pixel size of the integrated maps of the sources varies from 0.8 to 1.2''. To match the beam size of the observations, the radius of the circular regions was also varied from 3.0 to 4.3''. The spatial resolution of the presented ACA observations allowed the binary IRAS 16293 A and B (separated by 5.3'') to be resolved as single sources; however, the resolution was not high enough to disentangle the binary components A1 and A2 of IRAS 16293 A. Similarly, the binary components of IRAS4A (with a separation of 1.8'') and of TMC1 (with a separation of 0.6'') could not be disentangled due to the insufficient level of spatial resolution. All other sources are either single sources or binaries separated by large distances; hence, they are spatially resolved as individual sources.

The lower and upper uncertainties on the fitted column densities are derived assuming an error of  $\pm 20$  K on the assumed excitation temperature and a 1- $\sigma$  noise level. The analysis of the spectra extracted toward IRAS 16293-2422 B is presented



**Fig. 1.** ALMA pipeline-produced integrated intensity maps (color scale) with the line channels excluded, which are dominated by dust emission, for the studied sample of sources. On-source spectra are extracted by averaging the flux from the pixels within the circular area centered on “X.”

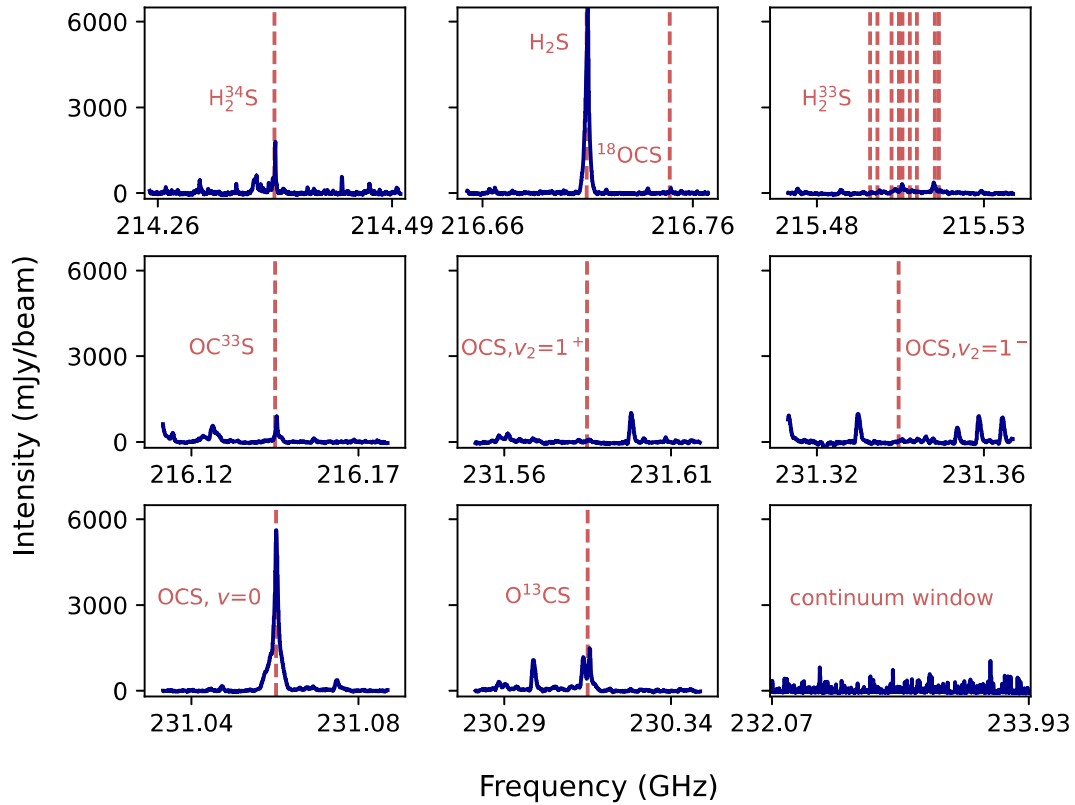
in Sect. 3.1 and Appendix E. The full observed spectral windows toward IRAS 16293 B are shown in Fig. 2. For the other sources, the analysis is presented in Appendices C–K.

### 3.1. IRAS 16293-2422 B

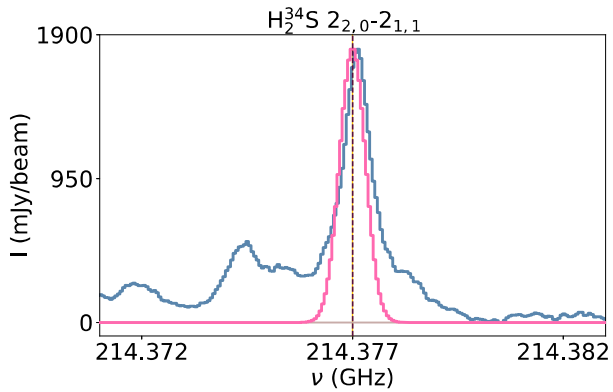
Towards IRAS 16293 B, the main S-bearing species ( $\text{H}_2\text{S}$  and  $\text{OCS}$ ,  $v=0$ ) and all the targeted minor isotopologs were securely detected, except for HDS due to a very high  $E_{\text{up}}$  value (1277 K) and the double isotopologs of  $\text{HD}^{34}\text{S}$ ,  $^{18}\text{OC}^{34}\text{S}$ ,  $^{18}\text{O}^{13}\text{CS}$ , due to their low abundances (and high  $E_{\text{up}}$  for the case of  $\text{HD}^{34}\text{S}$ ). The detected transitions of  $\text{H}_2\text{S}$ ,  $2_{2,0}-2_{1,1}$  and  $\text{OCS}$ ,  $v=0$ ,  $J=19-18$  are bright and optically thick ( $\tau \gg 1$ ). The  $\text{H}_2^{34}\text{S}$  line is marginally optically thick ( $\tau = 0.2$ ), as shown in Fig. 3. The vibrationally excited  $\text{OCS}$ ,  $v_2 = 1^\pm$  lines were detected. The lines of the detected molecules do not suffer from blending, except the  $\text{H}_2^{33}\text{S}$ ,  $2_{2,0,3}-2_{1,1,3}$  transition at 215.512 GHz, which is contaminated by the  $\text{CH}_3\text{CHO}$ ,  $11_{2,9}-10_{2,8}$  transition.  $\text{HD}^{34}\text{S}$ ,  $7_{3,4}-7_{3,5}$  transition at 232.964 GHz is heavily blended with the  $\text{CH}_3\text{CN}$ ,  $v_8 = 1$ ,  $J = 15-15$ ,  $K = 7-5$  transition. It is most likely that all the emission seen around the rest frequency of  $\text{HD}^{34}\text{S}$  comes from  $\text{CH}_3\text{CN}$ , because  $\text{HD}^{34}\text{S}$  is a minor species ( $^{32}\text{S}/^{34}\text{S} = 22$ ,

Wilson 1999, and  $D/H \sim 0.04$  incl. the statistical correction by a factor of 2 to account for the two indistinguishable D atom positions, Drozdovskaya et al. 2018) and the  $E_{\text{up}}$  of this transition is high (416 K). The spectra of detected and undetected lines are in Figs. B.1 and B.2, respectively.

For the analysis of the targeted S-bearing molecules toward IRAS 16293 B, a  $T_{\text{ex}}$  value of 125 K is assumed. This value has been deduced to be the best-fitting on the basis of ALMA-PILS observations at higher spatial resolution obtained with the 12m array and a full inventory of S-bearing molecules (Drozdovskaya et al. 2018). A FWHM of  $1 \text{ km s}^{-1}$  was adopted, as it has been shown that this value consistently fits nearly all the molecules investigated toward the hot inner regions of IRAS 16293 B (e.g., Jørgensen et al. 2018). For the larger scales probed by the present ALMA ACA observations, a deviation by  $2 \text{ km s}^{-1}$  from this FWHM can be seen for optically thick lines. This broadening in FWHM is likely due to the opacity broadening effects, which are dominant in optically thick lines, but this can be neglected in optically thin lines (Hacar et al. 2016). The synthetic spectral fitting has been carried out for two potential source sizes,  $1''$  and  $2''$  (Table 4). Column densities depend on the assumed source size and are lower for the larger source size. However, the



**Fig. 2.** Observed spectral windows of IRAS 16293-2422 B (Table 3) obtained with ALMA ACA at Band 6 frequencies (Table 1). A Doppler shift by  $v_{\text{LSR}} = 2.7 \text{ km s}^{-1}$  has been applied (Table 2).



**Fig. 3.**  $\text{H}_2^{34}\text{S}$  line detected in IRAS 16293 B. The observed spectrum (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), and fitted synthetic spectrum (in pink) for source size:  $2''$ , excitation temperature: 125 K, and FWHM:  $1 \text{ km s}^{-1}$ .

$N(\text{H}_2\text{S})/N(\text{OCS})$  ratio is  $1.3 \pm 0.27$  and  $1.3 \pm 0.28$  for source sizes of  $1''$  and  $2''$ , respectively. Thus, the ratio is independent of the assumed source size and is robustly determined with the ALMA ACA data.

For a source size of  $2''$ , the column density of the vibrationally excited state of OCS,  $v_2 = 1^\pm$  derived for IRAS 16293 B ( $2.5 \times 10^{16} \text{ cm}^{-2}$ ) is an order of magnitude lower than the OCS,  $v_2 = 1$  column density ( $2.0 \times 10^{17} \text{ cm}^{-2}$ ) derived in Drozdovskaya et al. (2018). For a source size of  $1''$ , the value obtained here ( $8.5 \times 10^{16} \text{ cm}^{-2}$ ) is in closer agreement with Drozdovskaya et al. (2018). Likewise, the OCS,  $v=0$  column density determined from the minor isotopologs

of OCS for a source size of  $1''$  ( $2.7 \times 10^{17} \text{ cm}^{-2}$ ) is in a closer agreement with the column density of OCS,  $v = 0$  ( $2.8 \times 10^{17} \text{ cm}^{-2}$ ) derived in Drozdovskaya et al. (2018), also based on minor isotopologs than for source size of  $2''$  ( $7.0 \times 10^{16} \text{ cm}^{-2}$ ). Drozdovskaya et al. (2018) used a smaller source size ( $0.5''$ ) to constrain the column densities of OCS and  $\text{H}_2\text{S}$ . These comparisons suggest that the ALMA ACA observations in this work are subject to beam dilution, hence, the column densities are likely to be somewhat underestimated. The column density of  $\text{H}_2\text{S}$  could not be constrained to better than a factor of 10 in Drozdovskaya et al. (2018), namely  $1.6 \times 10^{17} - 2.2 \times 10^{18} \text{ cm}^{-2}$ . This was due to the fact that only deuterated isotopologs of  $\text{H}_2\text{S}$  were covered by the PILS observations and the D/H ratio of  $\text{H}_2\text{S}$  is only constrained to within a factor of 10. Based on values of the  $\text{H}_2\text{S}$  column densities for  $1''$  and  $2''$  source sizes obtained in this work, the lower estimate for the  $\text{H}_2\text{S}$  column density in Drozdovskaya et al. (2018) seems to be more accurate. In turn, the  $\text{H}_2\text{S}/\text{OCS}$  ratio obtained in this work (1.3) is closer to the lower end of the 0.7–7 range computed in Drozdovskaya et al. (2018).

### 3.2. Line profiles

For the synthetic spectral modeling, Gaussian line profiles were assumed (Sect. 2.3). However, even for optically thin lines, a deviation from Gaussian line profiles is seen in some cases. Two prominent examples are  $\text{H}_2^{34}\text{S}$  and  $\text{O}^{13}\text{CS}$  in IRAS 16293 A (Fig. C.1), where the high spectral resolution of the data set clearly allows multiple peaks to be spectrally resolved in these lines. Likely, the reason for this is that this source is a compact binary (Maureira et al. 2020) with multiple components within the ACA beam of these observations. Another prominent

**Table 5.** H<sub>2</sub>S/OCS ratio for the studied sources, including their evolutionary class, binarity, environment, and the derived column densities of H<sub>2</sub>S and OCS for the stated excitation temperatures.

Source	Class	Binarity	Environment	$T_{\text{ex}}$ (K)	$N(\text{H}_2\text{S})$ (cm <sup>-2</sup> )	$N(\text{OCS})$ (cm <sup>-2</sup> )	$N(\text{H}_2\text{S})/N(\text{OCS})$
IRAS 16293-2422 A	0	CB	Clustered	125±20	(2.4±0.4)×10 <sup>17</sup>	(3.6±1.4)×10 <sup>17</sup>	0.7±0.3
IRAS 16293-2422 B	0	WB	Clustered	125±20	(9.2±1.7)×10 <sup>16</sup>	(7.0±0.8)×10 <sup>16</sup>	1.3±0.3
NGC 1333-IRAS4A	0	CB	Clustered	150–300	(3.4±0.8)×10 <sup>16</sup>	(1.8±0.2)×10 <sup>16</sup>	1.9±0.5
RCrA IRS7B	0/I	WB	Clustered	100±20	(5.6±0.8)×10 <sup>13</sup>	≤3.6×10 <sup>13</sup>	≥1.5
Per-B1-c	0	S	Clustered	60	>9.7×10 <sup>15</sup>	>5.0×10 <sup>15</sup>	(1.9)
				200	>1.2×10 <sup>16</sup>	>3.6×10 <sup>15</sup>	(3.3)
BHR71-IRS1	0	WB	Isolated	100	>2.4×10 <sup>16</sup>	>2.7×10 <sup>15</sup>	(8.9)
				250	>3.3×10 <sup>16</sup>	(3.4±0.3)×10 <sup>15</sup>	≥9.7
Per-emb-25	0/I	S	Clustered	50–300	≤8.3×10 <sup>13</sup>	≤3.2×10 <sup>14</sup>	–
NGC 1333-IRAS4B	0	WB	Clustered	100±20	>5.8×10 <sup>15</sup>	(2.8±0.6)×10 <sup>16</sup>	≥0.21
Ser-SMM3	0	S	Clustered	100–250	(5.8±3.2)×10 <sup>14</sup>	(8.7±4.9)×10 <sup>14</sup>	0.7±0.5
TMC1	I	CB	Clustered	40	≤1.5×10 <sup>13</sup>	≤2.6×10 <sup>13</sup>	–
Comet (67P/C-G)							26.8 <sup>+47.5</sup> <sub>-21.6</sub> <sup>(a)</sup>
					$N(\text{H}_2\text{S})/N(\text{H}_2\text{O})$	$N(\text{OCS})/N(\text{H}_2\text{O})$	$N(\text{H}_2\text{S})/N(\text{OCS})$
ISM ices (W33A)					<3.0×10 <sup>-2</sup> , <sup>(b)</sup>	4.0×10 <sup>-4</sup> , <sup>(c)</sup>	<75
ISM ices (Mon R2 IRS2)					<4.7×10 <sup>-3</sup> , <sup>(d)</sup>	5.5×10 <sup>-4</sup> , <sup>(c)</sup>	<8.5

**Notes.** The H<sub>2</sub>S/OCS ratios for the cold and warm components of B1-c, and the cold component of BHR71-IRS1 are the best-available estimates pending opacity issues. CB: close binary (<500 au) that is not spatially resolved in the data used for this work; WB: wide binary (500–5000 au) that is spatially resolved in the data used for this work; S: single source within 5 000 au. ISM stands for interstellar medium.

**References.** <sup>(a)</sup>Rubin et al. (2019), <sup>(b)</sup>van der Tak et al. (2003), <sup>(c)</sup>Palumbo et al. (1995), <sup>(d)</sup>Smith (1991).

example is the OCS,  $v = 0$  line in Ser-SMM3 (Fig. J.1), which has a double-peaked profile centered around the source velocity. Such a line profile is typical for a rotating structure around its protostar (which could resemble an envelope or disk in nature). Detailed modeling of line profiles is out of scope of this paper, as additional observations would be necessary in order to achieve meaningful results. For the purpose of studying the H<sub>2</sub>S/OCS ratio, these effects are secondary and likely do not significantly affect the calculated ratio and the conclusions of this paper. For IRAS 16293 A, the column density of H<sub>2</sub><sup>34</sup>S is not used to get the column density of H<sub>2</sub>S, because it is computed to be partially optically thick. Meanwhile, the column density of OCS as obtained from O<sup>13</sup>CS is within a factor of 2 of what is obtained from OC<sup>33</sup>S and <sup>18</sup>OCS. For Ser-SMM3, the lack of constraints on the excitation temperature dominates the uncertainty in the H<sub>2</sub>S/OCS ratio.

### 3.3. H<sub>2</sub>S/OCS ratio determination

The column densities of H<sub>2</sub>S and OCS derived on the basis of ALMA ACA observations have been used to constrain the ratio of H<sub>2</sub>S to OCS (Table 5). It was possible to compute this ratio for five out of ten sources in the considered sample. Neither H<sub>2</sub>S nor OCS were detected in Emb-25 and TMC1, consequently, the H<sub>2</sub>S/OCS ratio could not be constrained. The non-detection of OCS in IRS7B allowed us to derive only a lower limit on the H<sub>2</sub>S/OCS ratio. Table 5 also contains the best-available estimates of the H<sub>2</sub>S/OCS ratio for the warm and cold components of B1-c, and for the cold component of BHR71-IRS1, although these numbers carry a higher level of uncertainty due to line opacity that could not be resolved on the basis of these observations. For the warm component of BHR71-IRS1, a lower limit on the H<sub>2</sub>S/OCS ratio could be computed. For further analysis, the sample has been divided into three sub-samples: compact

binary, wide binary, and single, based on the separations between components of multiple sources or closest neighbors.

## 4. Discussion

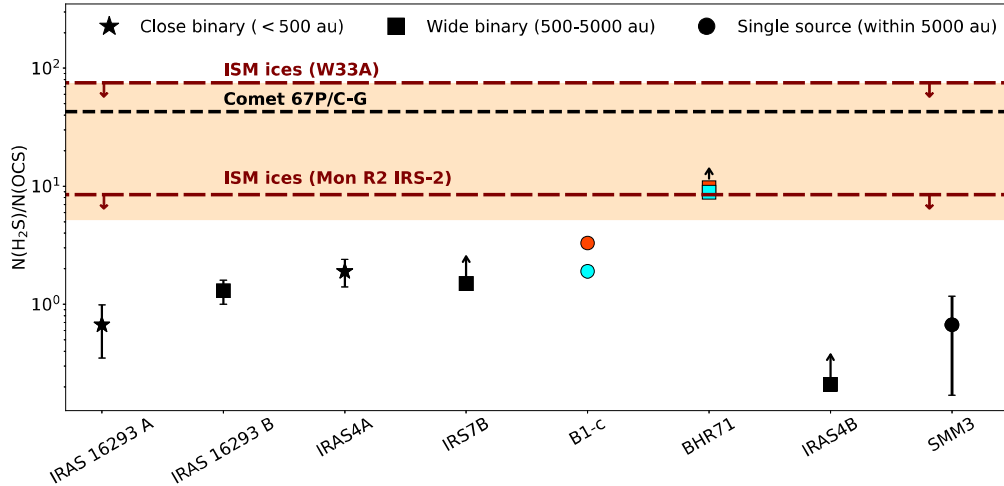
Figure 4 displays the derived protostellar H<sub>2</sub>S/OCS ratios, as well as the cometary (67P/Churyumov–Gerasimenko, hereafter 67P/C-G) and interstellar ice (W33A and Mon R2 IRS2) H<sub>2</sub>S/OCS ratios. The derived protostellar H<sub>2</sub>S/OCS ratios span a range from 0.2 to above 9.7. The ratios show a variation of approximately one order of magnitude, being the lowest in IRAS 16293 A and SMM3, and the highest in BHR71-IRS1.

In Sects. 4.1 and 4.2, we make a comparison of the protostellar H<sub>2</sub>S/OCS ratios with this ratio in interstellar and cometary ices, respectively. Comets are thought to preserve the chemical composition of the Sun’s birth cloud (Mumma & Charnley 2011). By comparing the H<sub>2</sub>S/OCS ratio of comet 67P/C-G with the ratios in nascent solar-like protostellar systems, an assessment can be made whether such an inheritance is true in the case of S-bearing molecules.

### 4.1. Interstellar ices

Observations toward the cold, outer protostellar envelopes of high-mass protostars W33A and Mon R2 IRS2 are used to acquire the H<sub>2</sub>S/OCS ratio in interstellar ices. The ratio is computed based on the ice abundances of OCS detected as an absorption feature at 4.9 μm (Palumbo et al. 1995) using the Infrared Telescope Facility (IRTF) and the upper limits on the H<sub>2</sub>S abundance derived based on the non-detection of the 3.98 μm band. The column density of solid OCS with respect to solid H<sub>2</sub>O is  $N_{\text{solid}}(\text{OCS})/N_{\text{solid}}(\text{H}_2\text{O}) = 4 \times 10^{-4}$  (Palumbo et al. 1995) and  $N_{\text{solid}}(\text{OCS})/N_{\text{solid}}(\text{H}_2\text{O}) = 5.5 \times 10^{-4}$  (Palumbo et al. 1997) for W33A and Mon R2 IRS2, respectively. Based on the





**Fig. 4.**  $N(\text{H}_2\text{S})/N(\text{OCS})$  of the studied sources. Different symbols represent different types of sources: “star” for close binary (<500 au), “square” for wide binary (500–5000 au), and “circle” for single sources (within 5000 au). The upper limits on the interstellar ice (W33A and Mon R2 IRS2) ratios are shown by downward arrows. The uncertainty on the  $\text{H}_2\text{S}/\text{OCS}$  ratio in comet 67P/C-G is shown by the coral shaded region. The lower limit on the ratio in IRS7B is shown by an upwards arrow. The  $\text{H}_2\text{S}/\text{OCS}$  ratios for the cold (cyan) and warm (orange) components of B1-c, and cold (cyan) component of BHR71-IRS1 are the best-available estimates pending opacity issues. These latter three data points do not have error bars associated to them in the figure (so as to indicate that they are merely estimates).

non-detection of solid  $\text{H}_2\text{S}$  toward W33A in the Infrared Space Observatory (ISO) Spectra from the Short Wavelength Spectrometer (SWS),  $N_{\text{solid}}(\text{H}_2\text{S})/N_{\text{solid}}(\text{H}_2\text{O})_{\text{solid}} < 0.03$  (van der Tak et al. 2003). The upper limits on the solid  $\text{H}_2\text{S}$  and solid  $\text{H}_2\text{O}$  in Mon R2 IRS2 are  $< 0.2 \times 10^{17}$  and  $42.7 \times 10^{17} \text{ cm}^{-2}$  (Smith 1991), respectively, yielding  $N_{\text{solid}}(\text{H}_2\text{S})/N_{\text{solid}}(\text{H}_2\text{O}) < 4.7 \times 10^{-3}$ . The  $\text{H}_2\text{S}/\text{OCS}$  ratio in interstellar ices is poorly constrained due to the non-detection of solid  $\text{H}_2\text{S}$  to date. The upper limits on the interstellar ices ratio are within the uncertainties of the cometary ices ratio. The derived protostellar ratios for all the sources are lower than the upper limit on the  $\text{H}_2\text{S}/\text{OCS}$  ratio determined for interstellar ices – except BHR71-IRS1, with the  $\text{H}_2\text{S}/\text{OCS}$  ratio exceeding the upper limit on the ratio for Mon R2 IRS2.

#### 4.2. Comet 67P/Churyumov-Gerasimenko

Comets are thought to be the most unprocessed objects in the Solar System (Mumma & Charnley 2011). Cometary chemical composition has been shown to be similar to a degree to that of star-forming regions (Bockelée-Morvan et al. 2000; Drozdovskaya et al. 2019). Consequently, the cometary  $\text{H}_2\text{S}/\text{OCS}$  ratio is thought to provide an independent measurement of this ratio in interstellar ices. The  $\text{H}_2\text{S}$  and OCS abundances from the ESA Rosetta mission were used to compute the  $\text{H}_2\text{S}/\text{OCS}$  ratio for the Jupiter-family comet 67P/C-G. The  $\text{H}_2\text{S}$  and OCS abundances relative to  $\text{H}_2\text{O}$  are  $1.10 \pm 0.46\%$  and  $0.041^{+0.082}_{-0.020}\%$ , respectively (Rubin et al. 2019). These molecules are typical constituents of comets (Lis et al. 1997; Bockelée-Morvan et al. 2000; Boissier et al. 2007; Mumma & Charnley 2011). BHR71-IRS1 is the only protostar in the sample with a  $\text{H}_2\text{S}/\text{OCS}$  ratio within the uncertainties of the cometary ices ratio. The  $\text{H}_2\text{S}/\text{OCS}$  ratio for the other sources is at least an order of magnitude lower than for 67P/C-G, while even considering the high uncertainties on the cometary value. The availability of  $\text{H}_2\text{S}$  relative to  $\text{H}_2\text{O}$  in cometary ice (0.0064–0.0156) appears to be higher than in the interstellar ices toward Mon R2 IRS2 ( $< 0.0047$ ). For W33A, the currently available upper limit ( $< 0.03$ ) is less constraining and, hence, no conclusion can be drawn about how its ices compare to those of comet 67P/C-G. The relative ratio of  $\text{H}_2\text{S}$  to OCS is only one window

onto the inventory of S-bearing molecules in gas and ice at different stages of star and planet formation, meanwhile the overall availability relative to, for example,  $\text{H}_2\text{O}$  is another window that requires dedicated exploration.

#### 4.3. $\text{H}_2\text{S}/\text{OCS}$ ratio as an environmental (clustered or isolated) tracer

The measured gas-phase  $\text{H}_2\text{S}/\text{OCS}$  ratios in the sample of young, low-mass protostars explored in this paper are predominantly lower (by as much as an order of magnitude) than the solid-state ratio measured through direct infrared observations of interstellar ices and indirectly via comets (Fig. 4). There appears to be no correlation with binarity nor a specific host cloud (Table 5). The dependence with evolutionary stage could not be properly explored, as the sample contains only one Class I source (TMC1), which did not result in detections of neither  $\text{H}_2\text{S}$  nor OCS.

The highest ratio of  $\geq 9.7$  is found for the warm component (250 K) of BHR71-IRS1, which is a wide binary ( $\sim 3$  200 au; Bourke 2001; Parise et al. 2006; Chen et al. 2008; Tobin et al. 2019) Class 0 protostar. The ratio in BHR71-IRS1 resides within the uncertainty of the cometary ratio, but is between the two upper limits derived for the interstellar ices. The overall envelope mass and bolometric luminosity of BHR71-IRS1 is comparable to those of other compact binary and wide binary systems. The similarity of its gas-phase  $\text{H}_2\text{S}/\text{OCS}$  ratio to the ratio in ices may suggest that it is displaying the most recently thermally desorbed volatiles that have not been subjected to gas-phase processing for long. However, what makes BHR71-IRS1 stand out is that it is located in an isolated cloud, namely, it is not associated with processes typical for clustered environments such as dynamical interactions, mechanical and chemical feedback from outflows, or enhanced irradiation.

It is possible that isolation resulted in lower irradiation of the ice grains during the prestellar phase in BHR71-IRS1, thus converting less  $\text{H}_2\text{S}$  ices to OCS ices by photodissociation in the presence of CO ice. Consequently, leaving a higher  $\text{H}_2\text{S}/\text{OCS}$  ratio in ices, which after evaporation resulted in a higher  $\text{H}_2\text{S}/\text{OCS}$  ratio in the gas phase. Another reason could be

more efficient hydrogenation chemistry in such a colder environment. On dust grains, hydrogenation is expected to be the most effective process leading to the formation of H<sub>2</sub>S (Wakelam et al. 2011; Esplugues et al. 2014). Hence, BHR71-IRS1 may have a higher H<sub>2</sub>S content and a lower OCS content, which results in a higher H<sub>2</sub>S/OCS ratio. Water deuteration is also higher by a factor of 2–4 in isolated protostars such as BHR71-IRS1 in comparison to those in clustered environments, such as IRAS 16293 and IRAS4 (Jensen et al. 2019).

One alternative cause of lower H<sub>2</sub>S/OCS ratios toward clustered low-mass protostars could be local temperature differences in their birth clouds, for instance, due to enhanced irradiation from the neighbouring protostars. Laboratory experiments have proven that OCS forms readily in ices when interstellar ice-analogs are irradiated by high-energy photons (Ferrante et al. 2008; Garozzo et al. 2010; Jiménez-Escobar & Muñoz Caro 2011; Chen et al. 2015). This would lead to a lower H<sub>2</sub>S/OCS ratio.

Additionally, cosmic rays and other forms of radiation (UV and X-ray photons) are a ubiquitous source of ionization in the interstellar gas. It is a pivotal factor in the dynamical and chemical evolution of molecular clouds (Padovani et al. 2018, 2020). Cosmic rays are not attenuated in the molecular clouds as strongly as UV photons (Ivlev et al. 2018; Padovani et al. 2018; Silsbee et al. 2018). Thus, dust grains in the interstellar medium can be heated by impinging cosmic rays, thereby heating up the icy grain mantles and resulting in calamitous explosions (Leger et al. 1985; Ivlev et al. 2015) and activating chemistry in solids (Shingledecker et al. 2017). Magnetohydrodynamic simulations have shown a higher cosmic ray production in protostars in a clustered environment (Kuffmeier et al. 2020), which would be consistent with the lower H<sub>2</sub>S/OCS ratios for such protostars found in this work. The results suggest that the H<sub>2</sub>S/OCS ratio traces the environment (isolated or clustered) of the protostellar systems. However, a follow-up study is needed as the sample consisted of only one isolated source.

## 5. Conclusions

In this work, we probe a sample of ten low-mass protostars for the presence of H<sub>2</sub>S, OCS, and their isotopologs using ALMA ACA Band 6 observations. For 5 out of 10 protostars, the H<sub>2</sub>S/OCS ratio was firmly constrained, and for an additional 3, best-possible estimates were obtained. This ratio is thought to be a potential chemical and physical clock of star-forming regions, which sheds light on the sulfur depletion that transpires from the diffuse medium to the dense core stage. Our main conclusions are:

- The main S-bearing species, H<sub>2</sub>S and OCS, are detected in IRAS 16293-2422 A, IRAS 16293-2422 B, NGC 1333-IRAS4A, NGC 1333-IRAS4B, Per-B1-c, BHR71-IRS1, and Ser-SMM3. There are 1- $\sigma$  upper limits on the column densities of OCS derived for RCrA IRS7B, TMC1, and Per-emb-25; and 1- $\sigma$  upper limits on the column densities of H<sub>2</sub>S are derived for TMC1 and Per-emb-25;
- The gas-phase H<sub>2</sub>S/OCS ratio ranges from 0.2 to above 9.7, and is typically at least one order of magnitude lower than that of ices. The lowest ratio is obtained for IRAS 16293 A and Ser-SMM3, while the highest is for BHR71-IRS1. The environment of the natal cloud, prior to the onset of star formation, may have played a major role in the distribution of sulfur across various S-bearing molecules, which have resulted in a spread in the H<sub>2</sub>S/OCS ratio by one order of magnitude;
- The upper limits derived for the interstellar ices (Mon R2 IRS2 and W33A) lie within the uncertainties of the cometary ices ratio, specifically, that of comet 67P/Churyumov-Gerasimenko. The protostellar ratios are lower than the upper limits on the interstellar ices ratio and the cometary ices ratio by at least one order of magnitude for all sources, except BHR71-IRS1;
- The lower ratio in clustered protostellar regions could be due to elevated birth cloud temperatures or due to additional radiation from nearby protostars, thereby enhancing the photodissociation pathways from H<sub>2</sub>S to OCS;
- The high H<sub>2</sub>S/OCS ratio in BHR71-IRS1 could be the result of less efficient photodissociation of H<sub>2</sub>S to OCS in the presence of CO ice in its isolated birth cloud or more efficient hydrogenation chemistry leading to more efficient H<sub>2</sub>S formation.

Follow-up high spatial resolution observations are required toward several sources to better constrain the spatial distribution and excitation temperatures associated with the H<sub>2</sub>S and OCS detections. Furthermore, the difference of more than tenfold in the H<sub>2</sub>S/OCS ratio toward Class 0 protostars in clustered and isolated environments is a strong motivation for performing more spectroscopic observations toward such sources, thereby understanding the physical and chemical differences in the two types of environments. Observations from *James Webb* Space Telescope (JWST) could play an important role in constraining the H<sub>2</sub>S/OCS ice ratio in low- and intermediate-mass stars. More studies of the H<sub>2</sub>S/OCS ratio in a larger sample of Class 0 and Class I protostars in clustered and isolated environments should also be performed to further understand the sulfur chemistry in star-forming regions.

*Acknowledgements.* The research was started as part of the Leiden/ESA Astrophysics Program for Summer Students (LEAPS) 2021. M.N.D. acknowledges the support by the Swiss National Science Foundation (SNSF) Ambizione grant no. 180079, the Center for Space and Habitability (CSH) Fellowship, and the IAU Gruber Foundation Fellowship. This paper makes use of the following ALMA data: ADS/JAO.ALMA#2017.1.00108.S and ADS/JAO.ALMA#2017.1.01350.S. ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada), MOST and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO and NAOJ. This research made use of Astropy, a community-developed core Python package for Astronomy (Astropy Collaboration 2013, 2018; <http://www.astropy.org>). The authors would like to thank Prof. Dr. Ewine van Dishoeck for useful discussions about the H<sub>2</sub>S/OCS ratio, and the anonymous referee for constructive feedback.

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## Appendix A: Inventory of targeted sources and detected and undetected emission lines

Species	CDMS entry	Transition	Freq. (GHz)	$E_{up}$ (K)	$A_{ij}$ ( $s^{-1}$ )	IRAS 16293 A	IRAS 16293 B	IRAS4A	IRS7B	B1-c	BHR71-IRS1	Emb25	IRAS4B	SMM3	TMC1	
H <sub>2</sub> S	46519	2 <sub>2,0</sub> -2 <sub>1,1</sub>	216.710	84	$4.9 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
H <sub>2</sub> <sup>33</sup> S	35503	2 <sub>2,0,1</sub> -2 <sub>1,1,2</sub>	215.494	84	$2.4 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
		2 <sub>2,0,1</sub> -2 <sub>1,1,1</sub>	215.497	84	$2.4 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
		2 <sub>2,0,4</sub> -2 <sub>1,1,3</sub>	215.501	84	$6.9 \times 10^{-6}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
		2 <sub>2,0,4</sub> -2 <sub>1,1,4</sub>	215.503	84	$4.1 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
		2 <sub>2,0,2</sub> -2 <sub>1,1,3</sub>	215.504	84	$1.7 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
		2 <sub>2,0,2</sub> -2 <sub>1,1,2</sub>	215.505	84	$1.9 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
		2 <sub>2,0,2</sub> -2 <sub>1,1,1</sub>	215.508	84	$1.2 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
		2 <sub>2,0,3</sub> -2 <sub>1,1,3</sub>	215.512	84	$2.8 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
		2 <sub>2,0,3</sub> -2 <sub>1,1,2</sub>	215.513	84	$1.1 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
		2 <sub>2,0,3</sub> -2 <sub>1,1,4</sub>	215.513	84	$9.1 \times 10^{-6}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓
H <sub>2</sub> <sup>34</sup> S	36504	2 <sub>2,0</sub> -2 <sub>1,1</sub>	214.377	84	$4.7 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
HDS	35502	7 <sub>3,4</sub> -7 <sub>3,5</sub>	234.046	417	$6.6 \times 10^{-6}$	-	-	-	-	✓	✓	✓	✓	✓	✓	
		12 <sub>5,7</sub> -12 <sub>5,8</sub>	234.528	1148	$6.3 \times 10^{-6}$	-	-	-	-	✓	✓	✓	✓	✓	✓	
		14 <sub>2,12</sub> -13 <sub>4,9</sub>	214.325	1277	$5.0 \times 10^{-8}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
HD <sup>34</sup> S	37503	7 <sub>3,4</sub> -7 <sub>3,5</sub>	232.964	416	$6.5 \times 10^{-6}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
		12 <sub>5,7</sub> -12 <sub>5,8</sub>	233.088	1145	$6.2 \times 10^{-6}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
OC <sup>33</sup> S	61503	18-17	216.147	99	$2.9 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
O <sup>13</sup> CS	61502	19-18	230.318	110	$3.5 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
<sup>18</sup> OCS	62506	19-18	216.753	104	$2.9 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
OC <sub>s</sub> , $\nu=0$	60503	19-18	231.061	111	$3.6 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
OC <sub>s</sub> , $\nu_2=1^-$ OC <sub>s</sub> , $\nu_2=1^+$	60504	19-18	231.342	860	$3.5 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
		19-18	231.584	860	$3.5 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
<sup>18</sup> OC <sup>34</sup> S	64511	21-20	233.479	123	$3.7 \times 10^{-5}$	✓	✓	✓	✓	✓	✓	✓	✓	✓	✓	
<sup>18</sup> O <sup>13</sup> CS	63503	19-18	216.229	104	$2.9 \times 10^{-5}$	-	-	-	-	✓	✓	✓	✓	✓	✓	

**Table A.1:** Lines of H<sub>2</sub>S, OCS, and their isotopologs towards the targeted sources. The columns show 1) the CDMS entry, 2) the quantum numbers associated to the transition, 3) the rest frequency of the line, 4) the upper energy level of the transition, 5) the Einstein A coefficient of the transition, and the line evaluation done for different protostars given in columns 6-14. Note: ✓ represents **detected lines**, ✗ represents **undetected lines** and - represents **unobserved lines** (i.e., not covered in the spectral windows observed for that source).



## Appendix B: IRAS 16293-2422 B

The synthetic spectral fitting of the detected species toward IRAS 16293 B is presented thoroughly in Section 3.1. An extensive study of the effect of source size on column density and  $\text{H}_2\text{S}/\text{OCS}$  is also presented in Section 3.1. The excitation temperature and FWHM assumed for fitting synthetic spectra to the detected lines toward IRAS 16293 B are 125 K and  $1 \text{ km s}^{-1}$ , respectively, based on the ALMA-PILS observations of IRAS 16293 (Drozdovskaya et al. 2018).

### Appendix B.1: Detected lines in IRAS 16293-2422 B

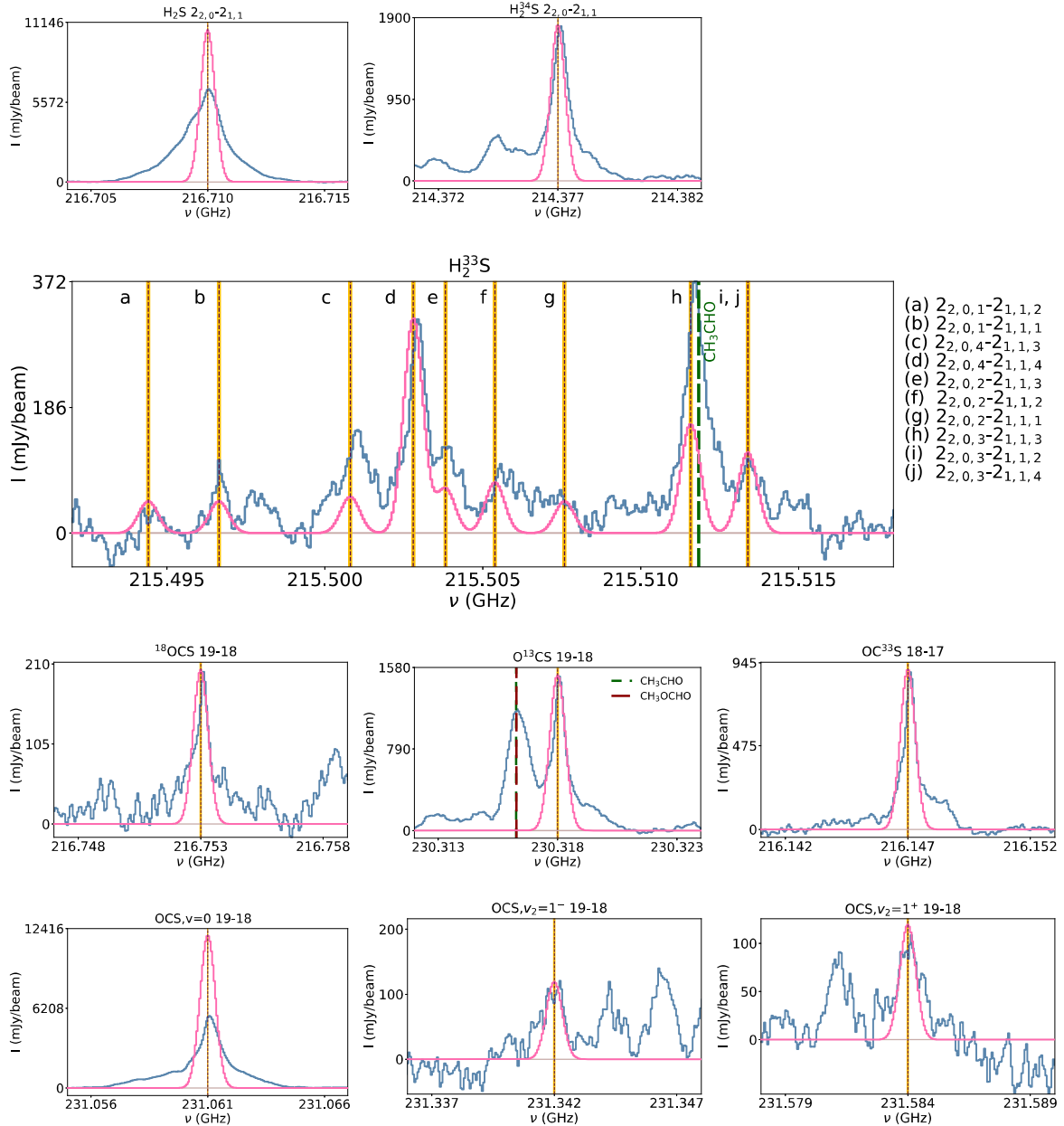


Figure B.1: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), blending species (green dash-dotted line and red dashed line), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward IRAS 16293-2422 B.  $\bar{N}(\text{H}_2\text{S})$  and  $\bar{N}(\text{OCS})$  are used for the synthetic spectra of the optically thick lines. For the displayed fits, a source size of  $2''$  is assumed.

## Appendix B.2: Undetected lines in IRAS 16293-2422 B

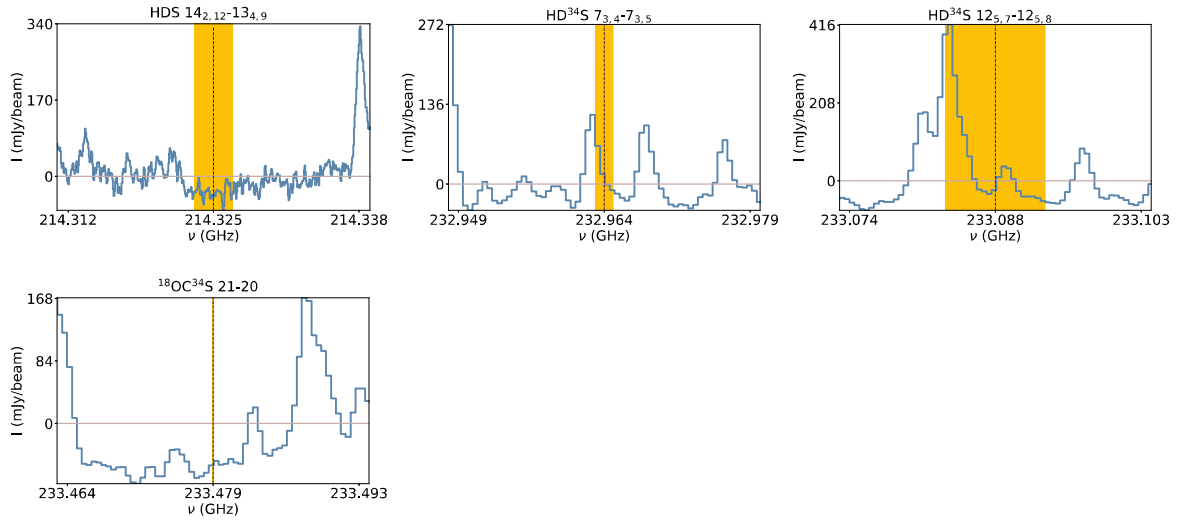


Figure B.2: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward IRAS 16293-2422 B.

## Appendix C: IRAS 16293-2422 A

The strengths of the lines detected toward IRAS 16293 A are comparable to those toward IRAS 16293 B. The lines of  $\text{H}_2\text{S}$ ,  $2_{2,0}-2_{1,1}$  and  $\text{OCS}$ ,  $v=0$ ,  $J=19-18$  are bright and optically thick ( $\tau \gg 1$ ). The line of  $\text{H}_2^{34}\text{S}$  is marginally optically thick ( $\tau=0.1$ ). For the analysis of IRAS 16293 A, we have assumed  $T_{\text{ex}} = 125$  K, which is typical for the hot inner regions of this source (Calcutt et al. 2018; Manigand et al. 2020), and a source size of  $2''$  in diameter (based on the investigation of IRAS 16293 B; Section 3.1). The fitted FWHM and  $N$ , and the computed  $\tau$  are given in Table C.1 for the detected S-bearing molecules. Due to the high spectral resolution of the data, all detected lines are spectrally resolved to be double- or multi-peaked. This is possibly caused by the multiple components of the compact binary IRAS 16293 A (Maureira et al. 2020). One of the detected  $\text{H}_2^{33}\text{S}$  lines (at 215.512 GHz) is contaminated by the  $\text{CH}_3\text{CHO}$  molecule. The line of  $\text{HD}^{34}\text{S}$  is heavily blended with that of  $\text{CH}_3\text{CN}$ ,  $v_8 = 1$  at 232.965 GHz. Having  $^{32}\text{S}/^{34}\text{S}=22$  (Wilson 1999) and  $\text{D}/\text{H} \sim 0.04$  (incl. statistical correction; Drozdovskaya et al. 2018), it is most likely that all the emission seen around the rest frequency of  $\text{HD}^{34}\text{S}$  stems from  $\text{CH}_3\text{CN}$ . The spectra of detected and undetected molecules can be found in Figure C.1 and Figure C.2, respectively.

**Table C.1:** Synthetic fitting of the detected S-bearing species toward IRAS 16293-2422 A for an excitation temperature ( $T_{\text{ex}}$ ) of 125 K and a source size of  $2''$ . Directly across from a specific minor isotopolog under “Derived  $N$  of isotopologs” follows the column density of the main isotopolog upon the assumption of the standard isotopic ratio. In bold in the same column, we show the average column density of the main isotopolog based on all the available minor isotopologs (only if the minor isotopolog is optically thin and including the uncertainties).

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ ( $\text{s}^{-1}$ )	Beam size ( $''$ )	FWHM ( $\text{km s}^{-1}$ )	$N$ ( $\text{cm}^{-2}$ )	Derived $N$ of isotopologs ( $\text{cm}^{-2}$ )	$\tau$
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.0	4.5	op. thick	$\overline{N(\text{H}_2\text{S})} = (2.4 \pm 0.4) \times 10^{17}$	6.0
$\text{H}_2 \text{ }^{33}\text{S}$	$2_{2,0,1}-2_{1,1,2}$	215.494	84	$2.4 \times 10^{-5}$	6.0	6.0	$2.0_{-0.5}^{+0.2} \times 10^{15}$	$N(\text{H}_2\text{S}) = 2.5_{-0.6}^{+0.3} \times 10^{17}$	0.002
	$2_{2,0,1}-2_{1,1,1}$	215.497	84	$2.4 \times 10^{-5}$					0.002
	$2_{2,0,4}-2_{1,1,3}$	215.501	84	$6.9 \times 10^{-6}$					0.002
	$2_{2,0,4}-2_{1,1,4}$	215.503	84	$4.1 \times 10^{-5}$					0.010
	$2_{2,0,2}-2_{1,1,3}$	215.504	84	$1.7 \times 10^{-5}$					0.003
	$2_{2,0,2}-2_{1,1,2}$	215.505	84	$1.9 \times 10^{-5}$					0.003
	$2_{2,0,2}-2_{1,1,1}$	215.508	84	$1.2 \times 10^{-5}$					0.002
	$2_{2,0,3}-2_{1,1,3}$	215.512	84	$2.8 \times 10^{-5}$					0.007
	$2_{2,0,3}-2_{1,1,2}$	215.513	84	$1.1 \times 10^{-5}$					0.003
	$2_{2,0,3}-2_{1,1,4}$	215.513	84	$9.1 \times 10^{-6}$					0.002
$\text{H}_2 \text{ }^{34}\text{S}$	$2_{2,0}-2_{1,1}$	214.377	84	$4.7 \times 10^{-5}$	6.0	7.0	$> 6.6 \times 10^{15}$		0.1
$\text{OCS}, v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	5.6	4.5	op. thick	$\overline{N(\text{OCS})} = (3.6 \pm 1.4) \times 10^{17}$	9.00
$\text{OC}^{33}\text{S}$	18-17	216.147	99	$2.9 \times 10^{-5}$	6.0	6.0	$1.9_{-0.3}^{+0.1} \times 10^{15}$	$N(\text{OCS}) = 2.4_{-0.2}^{+0.1} \times 10^{17}$	0.05
$\text{O}^{13}\text{CS}$	19-18	230.318	110	$3.5 \times 10^{-5}$	5.7	8.0	$3.8_{-0.2}^{+0.3} \times 10^{15}$	$N(\text{OCS}) = 2.6_{-0.1}^{+0.2} \times 10^{17}$	0.08
$^{18}\text{OCS}$	19-18	216.753	104	$3.0 \times 10^{-5}$	6.0	9.0	$8.5_{-2.0}^{+0.5} \times 10^{14}$	$N(\text{OCS}) = 4.7_{-1.1}^{+0.3} \times 10^{17}$	0.01
$\text{OCS}, v_2 = 1^-$	19-18	231.342	860	$3.5 \times 10^{-5}$	5.6	6.0	$2.5_{-1.8}^{+1.5} \times 10^{17}$		0.02
$\text{OCS}, v_2 = 1^+$	19-18	231.584	860	$3.5 \times 10^{-5}$	5.6				0.02

Appendix C.1: Detected lines in IRAS 16293-2422 A

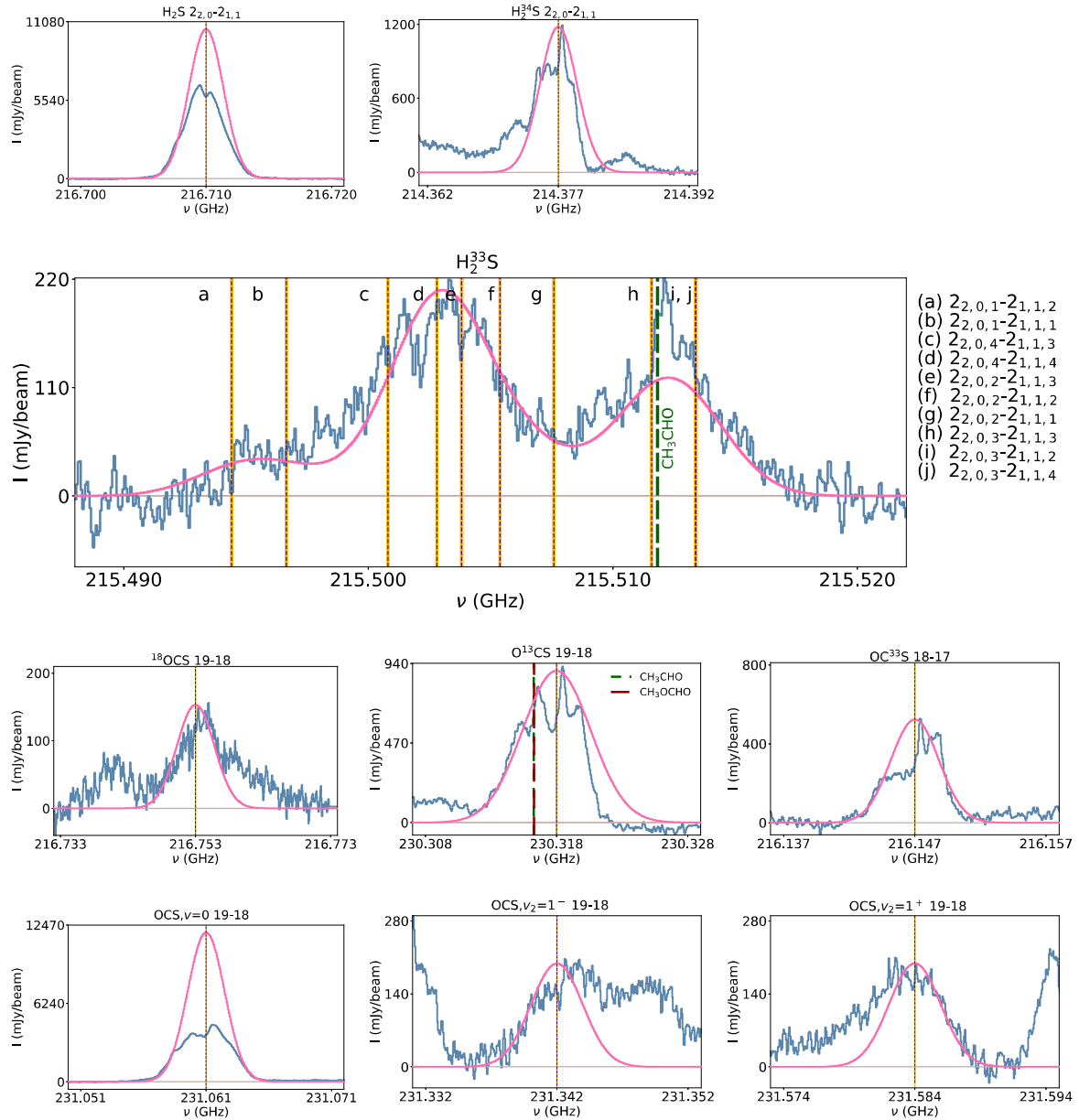


Figure C.1: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), blending species (green dash-dotted line and red dashed line), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward IRAS 16293-2422 A.  $N(\text{H}_2\text{S})$  and  $N(\text{OCS})$  are for the synthetic spectra of the optically thick lines.



## Appendix C.2: Undetected lines in IRAS 16293-2422 A

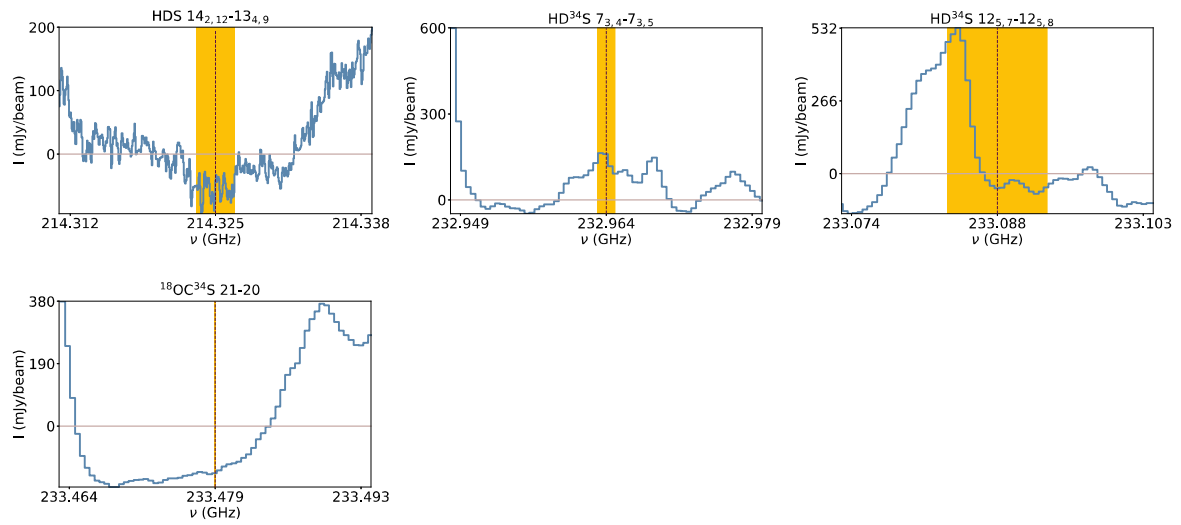


Figure C.2: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward IRAS 16293-2422 A.

## Appendix D: NGC 1333-IRAS4A

Line detections in IRAS4A are nearly the same as in IRAS 16293 A and IRAS 16293 B. The main sulfur species,  $\text{H}_2\text{S}$ ,  $2_{2,0}-2_{1,1}$  and  $\text{OCS}$ ,  $v=0$ ,  $J=19-18$ , are bright and optically thick,  $\tau > 1$  for  $50 \leq T_{\text{ex}} < 200$  K or partially optically thick,  $0.1 \leq \tau < 1$  for  $T_{\text{ex}} \geq 200$  K. The two protostars in IRAS4A, 4A1 and 4A2, could not be disentangled due to insufficient angular resolution of the ALMA ACA observations toward this source in comparison to the separation between the two components ( $1.8''$ , i.e., 540 au; Sahu et al. 2019). Due to a lack of small-scale constraints on the distribution of S-bearing molecules in IRAS4A, a range of excitation temperature between 150 and 300 K is considered (López-Sepulcre et al. 2017). Temperatures of 50 and 100 K have been ruled out on the grounds that the observed signal at the frequency corresponding to the  $\text{OCS}$ ,  $v_2=1^\pm$  is incompatible with these lowest temperatures for column densities in the physical range of values. Looking at high spatial resolution maps from Sahu et al. (2019), a  $2''$  source size is assumed to cover the circumbinary envelope around 4A1 and 4A2. The spectra fitted to the line profiles have a FWHM of  $1.8 \text{ km s}^{-1}$ . The computed column densities of the detected species are reported in Table D.1 for excitation temperatures in the 150 – 300 K range. The lines of  $\text{H}_2^{33}\text{S}$  and  $\text{HD}^{34}\text{S}$  face the same blending effects as in IRAS 16293 A and IRAS 16293 B.

### Appendix D.1: Results of the synthetic fitting of the line profiles toward NGC 1333-IRAS4A for $T_{\text{ex}} = 150 - 300$ K

**Table D.1:** Synthetic fitting of the detected S-bearing species toward NGC 1333-IRAS4A for a range of excitation temperatures between 150 and 300 K, a FWHM of  $1.8''$ , and a source size of  $2''$ . Directly across from a specific minor isotopolog under “Derived  $N$  of isotopologs” follows the column density of the main isotopolog upon the assumption of the standard isotopic ratio. In bold in the same column is the average column density of the main isotopolog based on all the available minor isotopologs (only if the minor isotopolog is optically thin and including the uncertainties).

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ ( $\text{s}^{-1}$ )	Beam size ( $''$ )	$N$ ( $\text{cm}^{-2}$ )	Derived $N$ of isotopologs ( $\text{cm}^{-2}$ )	$\tau$
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.1	op. thick	$\overline{N(\text{H}_2\text{S})} = \mathbf{(2.6-4.2)} \times 10^{16}$	0.4-1.2
$\text{H}_2^{33}\text{S}$	$2_{2,0,1}-2_{1,1,2}$	215.494	84	$2.4 \times 10^{-5}$	6.8	$3 \times 10^{14}$	$N(\text{H}_2\text{S}) = (3.8-8.0) \times 10^{16}$	0.0003-0.0007
	$2_{2,0,1}-2_{1,1,1}$	215.497	84	$2.4 \times 10^{-5}$				0.0003-0.0007
	$2_{2,0,4}-2_{1,1,3}$	215.501	84	$6.9 \times 10^{-6}$				0.0004-0.0008
	$2_{2,0,4}-2_{1,1,4}$	215.503	84	$4.1 \times 10^{-5}$				0.0020-0.0080
	$2_{2,0,2}-2_{1,1,3}$	215.504	84	$1.7 \times 10^{-5}$				0.0005-0.0010
	$2_{2,0,2}-2_{1,1,2}$	215.505	84	$1.9 \times 10^{-5}$				0.0005-0.0010
	$2_{2,0,2}-2_{1,1,1}$	215.508	84	$1.2 \times 10^{-5}$				0.0003-0.0007
	$2_{2,0,3}-2_{1,1,3}$	215.512	84	$2.8 \times 10^{-5}$				0.0010-0.0020
	$2_{2,0,3}-2_{1,1,2}$	215.513	84	$1.1 \times 10^{-5}$				0.0005-0.0010
	$2_{2,0,3}-2_{1,1,4}$	215.513	84	$9.1 \times 10^{-6}$				0.0004-0.0008
$\text{H}_2^{34}\text{S}$	$2_{2,0}-2_{1,1}$	214.377	84	$4.7 \times 10^{-5}$	6.8	$6.3 \times 10^{14}$	$N(\text{H}_2\text{S}) = (1.4-2.9) \times 10^{16}$	0.01-0.03
$\text{OCS}$ , $v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	6.2	op. thick	$\overline{N(\text{OCS})} = \mathbf{(1.6-2.0)} \times 10^{16}$	0.5-1.1
$\text{OC}^{33}\text{S}$	18-17	216.147	99	$2.9 \times 10^{-5}$	6.8	$1.0 \times 10^{14}$	$N(\text{OCS}) = (1.3-1.9) \times 10^{16}$	0.003-0.007
$\text{O}^{13}\text{CS}$	19-18	230.318	110	$3.5 \times 10^{-5}$	6.3	$1.4 \times 10^{14}$	$N(\text{OCS}) = (0.1-1.3) \times 10^{16}$	0.005-0.010
$^{18}\text{OCS}$	19-18	216.753	104	$3.0 \times 10^{-5}$	6.1	$3.8 \times 10^{13}$	$N(\text{OCS}) = (2.3-3.0) \times 10^{16}$	0.001-0.002
$\text{OCS}$ , $v_2=1^-$	19-18	231.342	860	$3.5 \times 10^{-5}$	5.6	$(0.4-1.2) \times 10^{15}$		0.001-0.002
$\text{OCS}$ , $v_2=1^+$	19-18	231.584	860	$3.5 \times 10^{-5}$	5.6			0.001-0.004

## Appendix D.2: Detected lines in NGC 1333-IRAS4A

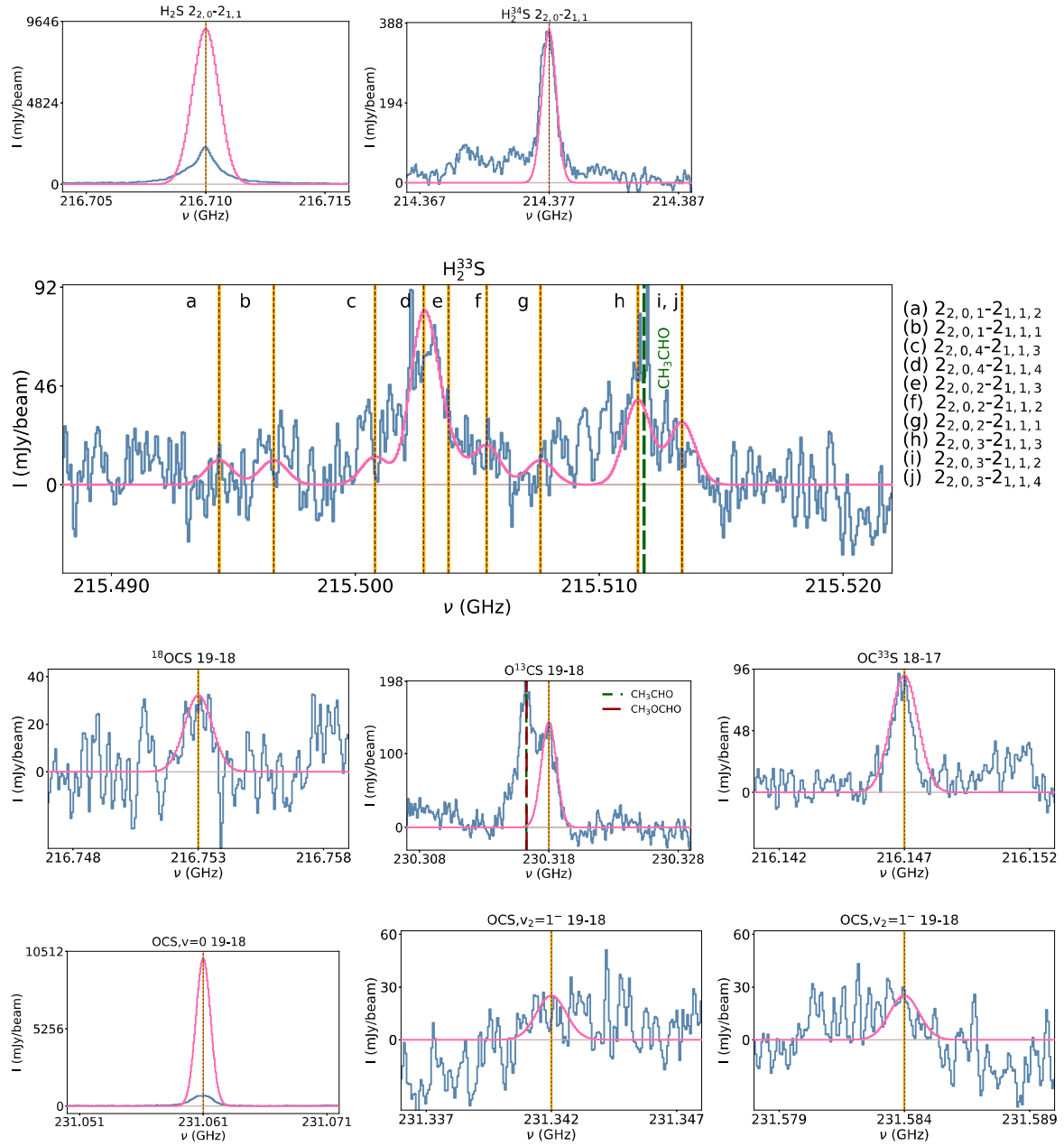


Figure D.1: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), blending species (green dash-dotted line and red dashed line), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward NGC 1333-IRAS4A for  $T_{\text{ex}} = 150$  K.  $\overline{N}(\text{H}_2\text{S})$  and  $\overline{N}(\text{OCS})$  are used to fit the synthetic spectra of the optically thick lines.

## Appendix D.3: Undetected lines in NGC 1333-IRAS4A

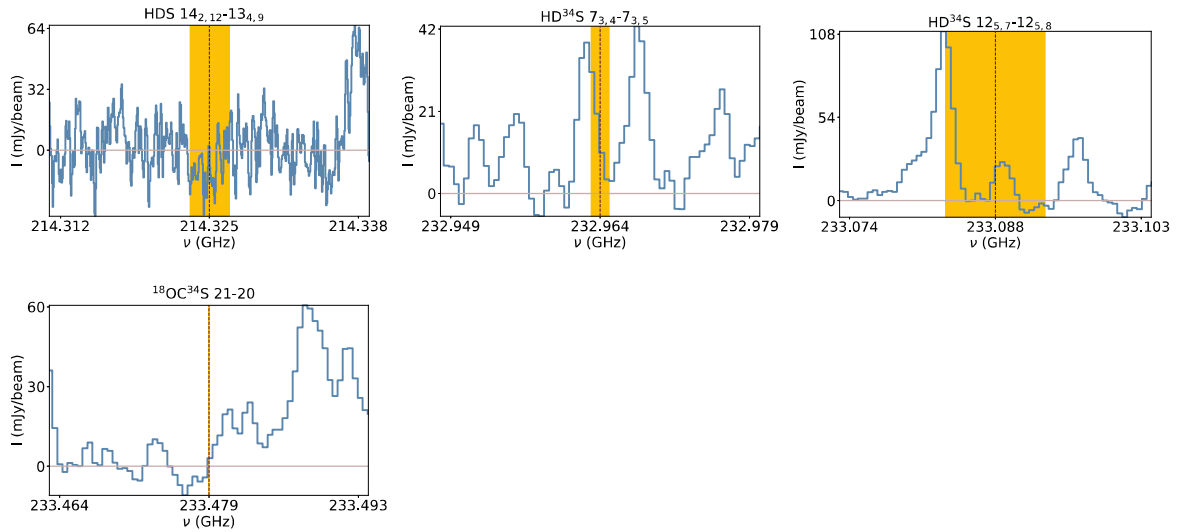


Figure D.2: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward NGC 1333-IRAS4A.

## Appendix E: RCrA IRS7B

$\text{H}_2\text{S}$ ,  $2_{2,0}-2_{1,1}$  is the only emission line detected toward IRS7B (Figure E.1). The synthetic spectrum for this line is best-fitted for a FWHM of  $1 \text{ km s}^{-1}$ , an excitation temperature of 100 K, and a source size of  $1.5''$ , as determined from ALMA 12m observations (project id: 2017.1.00108.S; PI: M. N. Drozdovskaya) toward IRS7B. The upper limit on the column density of OCS,  $v=0$  is stated in Table 5 and E.1. IRS7B exhibits weak hot corino activity, resulting in few detections of complex organic molecules and their low abundances (Lindberg et al. 2015). The chemical composition in the envelope of IRS7B could be affected by the enhanced UV radiation (photodissociation) from the nearby Herbig Ae star RCrA located at a distance of  $39''$  ( $6\,630 \text{ au}$ ) NW of IRS7B (Watanabe et al. 2012).

**Table E.1:** Synthetic fitting of the detected S-bearing species toward RCrA IRS7B for an excitation temperature of 100 K, FWHM of  $1 \text{ km s}^{-1}$ , and source size of  $1.5''$ .

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ ( $\text{s}^{-1}$ )	Beam size ( $''$ )	$N$ ( $\text{cm}^{-2}$ )	$\tau$
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.5	$5.5^{+0.9}_{-0.6} \times 10^{13}$	0.01
OCS, $v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	6.2	$\leq 3.6 \times 10^{13}$	–

## Appendix E.1: Detected lines in RCrA IRS7B

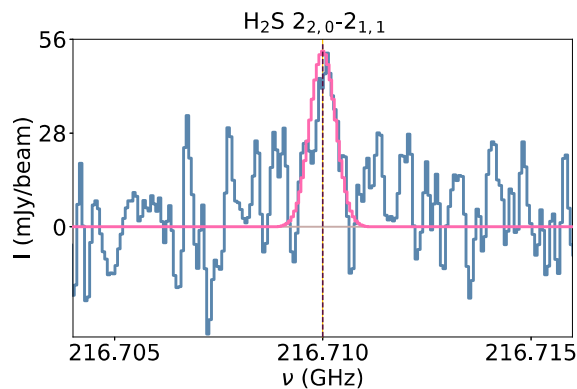


Figure E.1: Observed spectrum (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), and fitted synthetic spectrum (in pink) plotted for the sulfur-bearing species detected toward RCrA IRS7B.



## Appendix E.2: Undetected lines in RCrA IRS7B

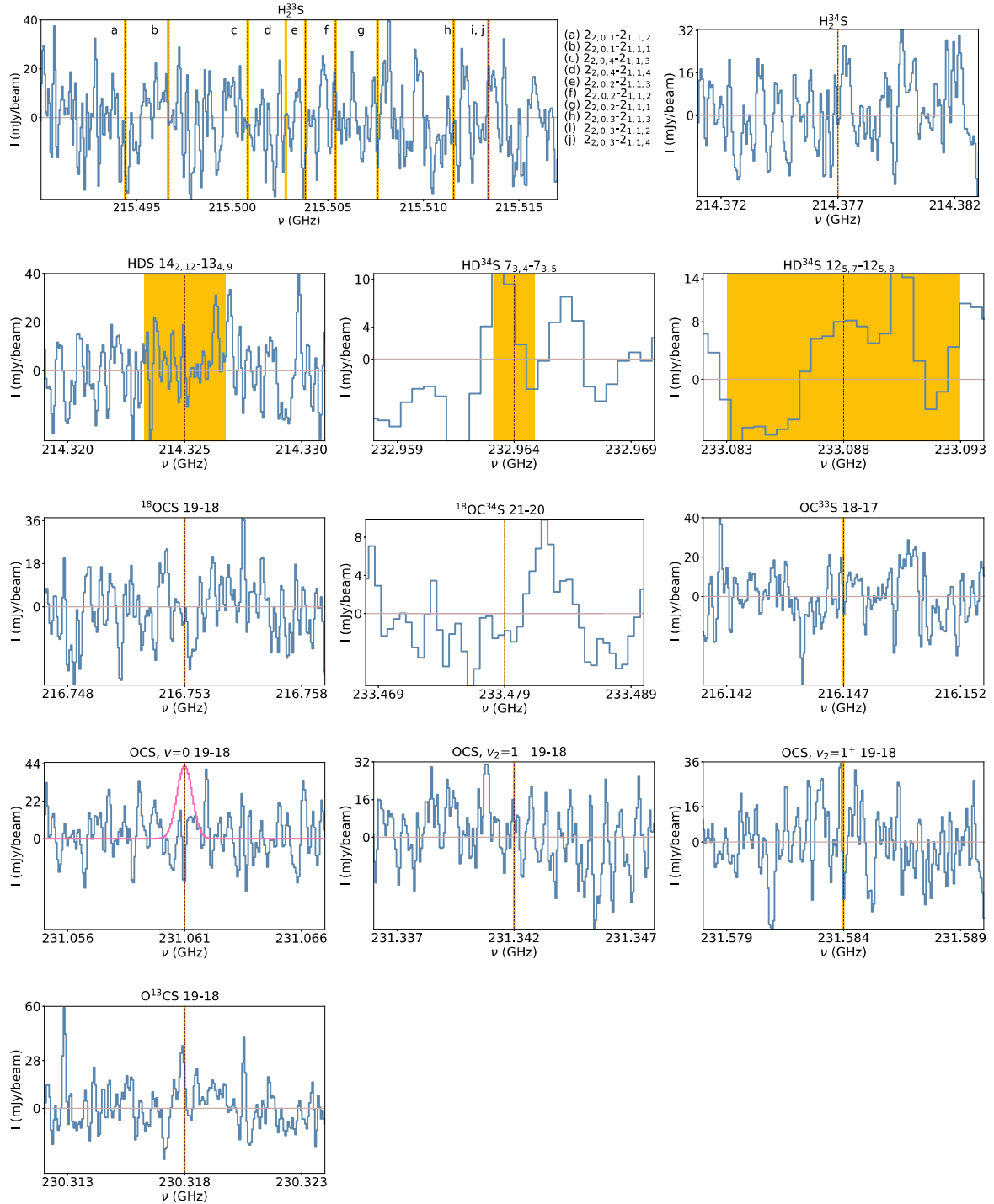


Figure E.2: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward RCrA IRS7B. Synthetic spectrum (in pink) fitted to the OCS,  $v=0$  line with the  $1-\sigma$  upper limit on its column density.

## Appendix F: B1-c

$\text{H}_2\text{S}$ ,  $2_{2,0}-2_{1,1}$  and OCS,  $v=0$ ,  $J = 19 - 18$  lines are detected toward B1-c. Unfortunately, no optically thin isotopologs of  $\text{H}_2\text{S}$  and OCS were observed toward B1-c in this observational data set. Synthetic spectra for the detected  $\text{H}_2\text{S}$  and OCS lines are fitted for two components of B1-c, a cold component ( $T_{\text{ex}} = 60$  K) and a warm central component ( $T_{\text{ex}} = 200$  K) for a source size of  $0.45''$  (as assumed in van Gelder et al. 2020 for the same observations). The best-fitting FWHM is  $2.2 \text{ km s}^{-1}$ . The fitting parameters, computed  $N$  and  $\tau$  values are tabulated in Table F.1 and F.2. For the cold component, the  $\text{H}_2\text{S}$  line is optically thick ( $\tau=2.0$ ); whereas

the OCS,  $v=0$ ,  $J = 19 - 18$  line is partially optically thick ( $\tau = 0.7$ ). For the warm component, both H<sub>2</sub>S and OCS,  $v=0$  lines are partially optically thick ( $\tau = 0.26$  and  $\tau = 0.10$ , respectively). The undetected lines are shown in Figure F.3.

**Table F.1:** Synthetic fitting of the detected S-bearing species toward B1-c for the cold component ( $T_{\text{ex}} = 60$  K), a FWHM of  $2.2 \text{ km s}^{-1}$ , and a source size of  $0.45''$ .

Species	Transition	Frequency	$E_{\text{up}}$	$A_{ij}$	Beam size	$N$	$\tau$
		(GHz)	(K)	( $\text{s}^{-1}$ )	( $''$ )	( $\text{cm}^{-2}$ )	
H <sub>2</sub> S	2 <sub>2,0</sub> -2 <sub>1,1</sub>	216.710	84	$4.9 \times 10^{-5}$	6.5	$>9.7 \times 10^{15,*}$	2.0
OCS, $v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	6.1	$>5.0 \times 10^{15,**}$	0.7

**Notes.** \*optically thick, \*\*partially optically thick

**Table F.2:** Synthetic fitting of the detected S-bearing species toward B1-c for the warm component ( $T_{\text{ex}} = 200$  K), a FWHM of  $2.2 \text{ km s}^{-1}$ , and a source size of  $0.45''$ .

Species	Transition	Frequency	$E_{\text{up}}$	$A_{ij}$	Beam size	$N$	$\tau$
		(GHz)	(K)	( $\text{s}^{-1}$ )	( $''$ )	( $\text{cm}^{-2}$ )	
H <sub>2</sub> S	2 <sub>2,0</sub> -2 <sub>1,1</sub>	216.710	84	$4.9 \times 10^{-5}$	6.5	$>1.2 \times 10^{16,*}$	0.26
OCS, $v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	6.1	$>3.6 \times 10^{15,*}$	0.10

**Notes.** \*partially optically thick

#### Appendix F.1: Detected lines in B1-c

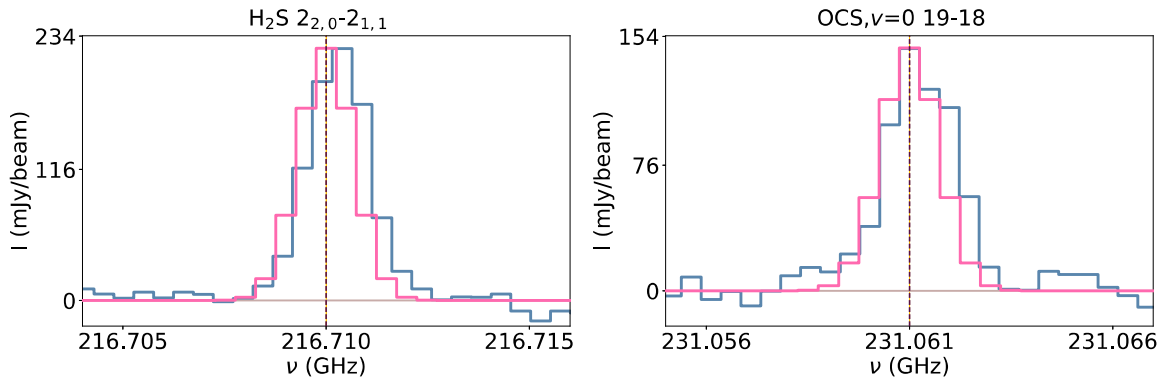


Figure F.1: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward the cold component ( $T_{\text{ex}} = 60$  K) of B1-c.

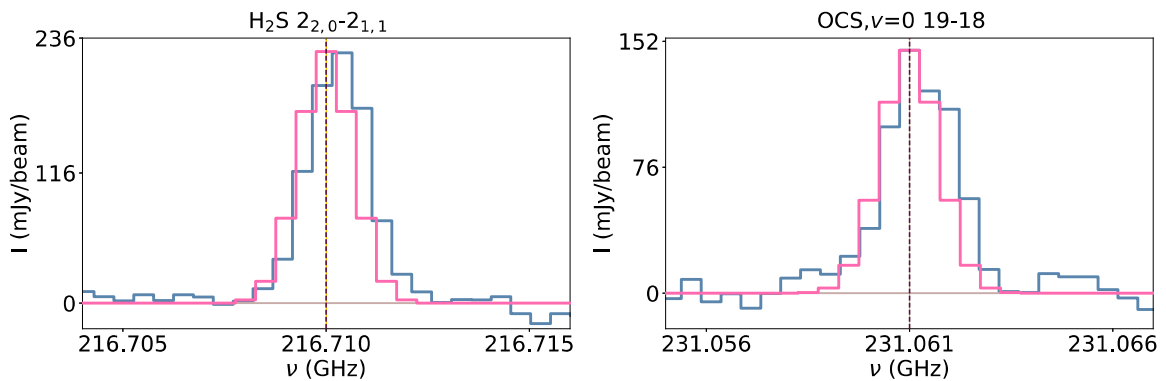


Figure F.2: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward the warm component ( $T_{\text{ex}} = 200$  K) of B1-c.

## Appendix F.2: Undetected lines in B1-c

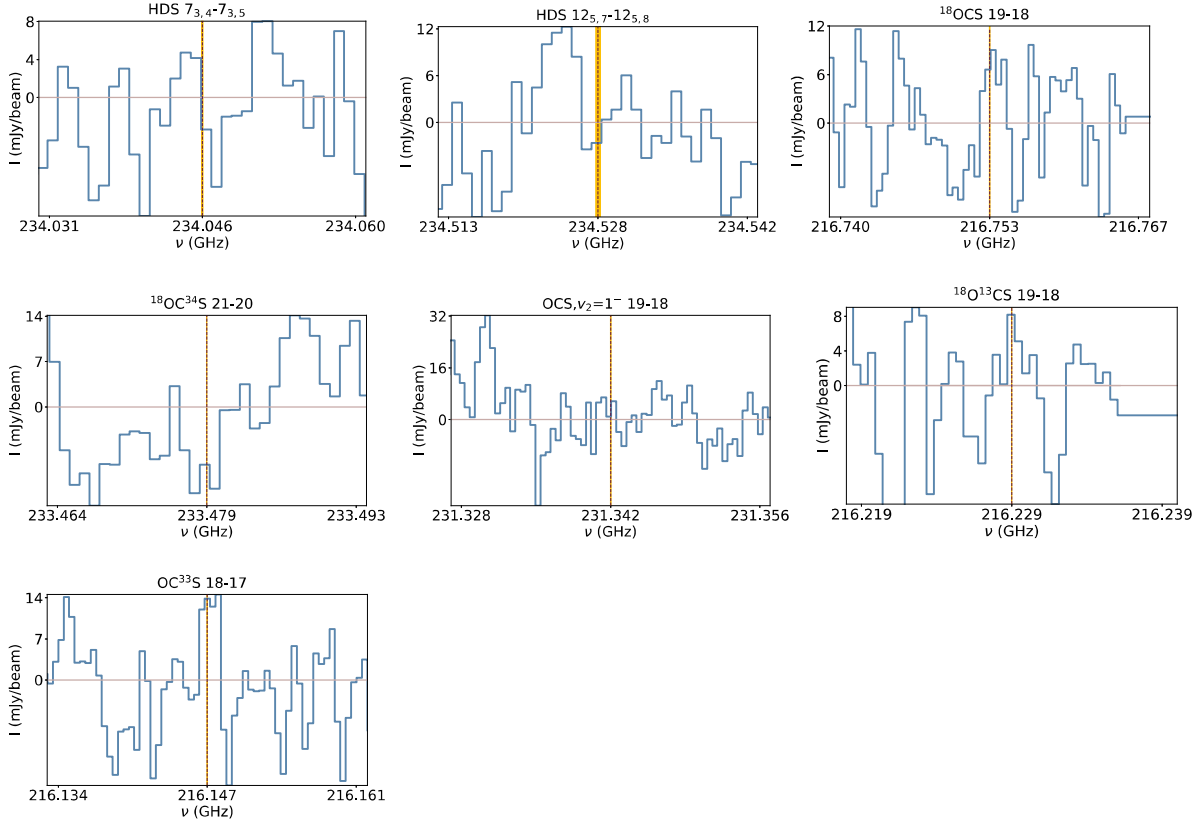


Figure F.3: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward B1-c.

## Appendix G: BHR71-IRS1

$\text{H}_2\text{S}$ ,  $2_{2,0}-2_{1,1}$  and  $\text{OCS}$ ,  $v = 0$ ,  $J = 19 - 18$  emission lines are detected toward BHR71-IRS1 (Figure G.1, G.2). As for B1-c, two excitation temperatures (100 and 250 K) probing two different regions are inputted into the modeling of synthetic spectra for BHR71-IRS1. The excitation temperatures are taken from the detection of methanol at  $T_{\text{ex}} = 100$  K and gauche-ethanol at  $T_{\text{ex}} = 250$  K (Yang et al. 2020). Yang et al. (2020) fitted CS for a source size of  $0.32''$ ; however, in order to cover a larger area of the envelope a source size of  $0.6''$ , and a FWHM of  $2.5 \text{ km s}^{-1}$  is assumed here. The  $\text{H}_2\text{S}$  line of the colder component seems to be optically thick ( $\tau = 1.8$ ), while the  $\text{OCS}$ ,  $v=0$  line is partially optically thick for this component ( $\tau = 0.23$ ).  $\text{H}_2\text{S}$  is partially optically thin for the warmer component ( $\tau = 0.39$ ). The spectra of the undetected lines are shown in Figure G.3.

**Table G.1:** Synthetic fitting of the detected S-bearing species toward BHR71-IRS1 for an excitation temperature of 100 K, FWHM of  $2.5 \text{ km s}^{-1}$ , and source size of  $0.6''$ .

Species	Transition	Frequency	$E_{\text{up}}$	$A_{ij}$	Beam size	$N$	$\tau$
		(GHz)	(K)	( $\text{s}^{-1}$ )	( $''$ )	( $\text{cm}^{-2}$ )	
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.2	$>2.4 \times 10^{16}$ , *	1.80
$\text{OCS}$ , $v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	5.8	$>2.7 \times 10^{15}$ , *	0.23

**Notes.** \*optically thick, \*\*partially optically thick

**Table G.2:** Synthetic fitting of the detected S-bearing species toward BHR71-IRS1 for an excitation temperature of 250 K, FWHM of 2.5 km s<sup>-1</sup>, and source size of 0.6".

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ (s <sup>-1</sup> )	Beam size (")	$N$ (cm <sup>-2</sup> )	$\tau$
H <sub>2</sub> S	2 <sub>2,0</sub> -2 <sub>1,1</sub>	216.710	84	$4.9 \times 10^{-5}$	6.2	$>3.3 \times 10^{16}$ , *	0.39
OCS, $\nu=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	5.8	$3.4^{+0.3}_{-0.3} \times 10^{15}$	0.08

Notes. \*partially optically thick

#### Appendix G.1: Detected lines in BHR71-IRS1

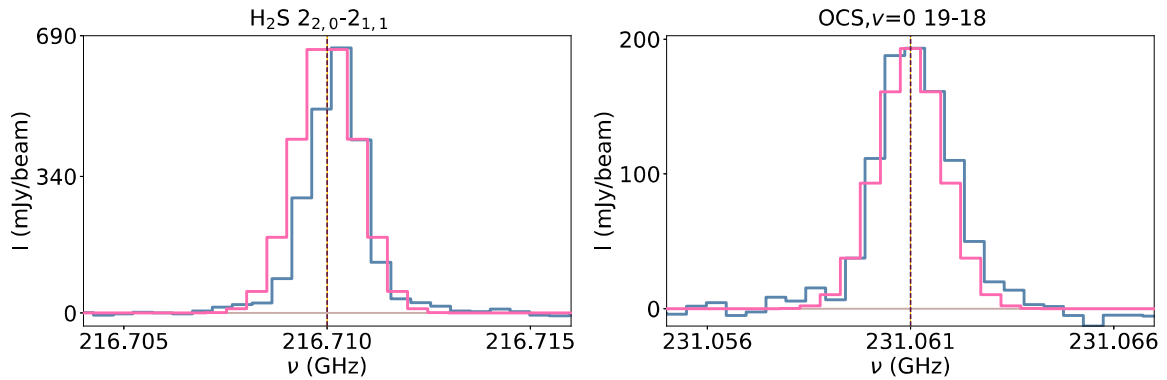


Figure G.1: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward the cold component ( $T_{\text{ex}} = 100$  K) of BHR71-IRS1.

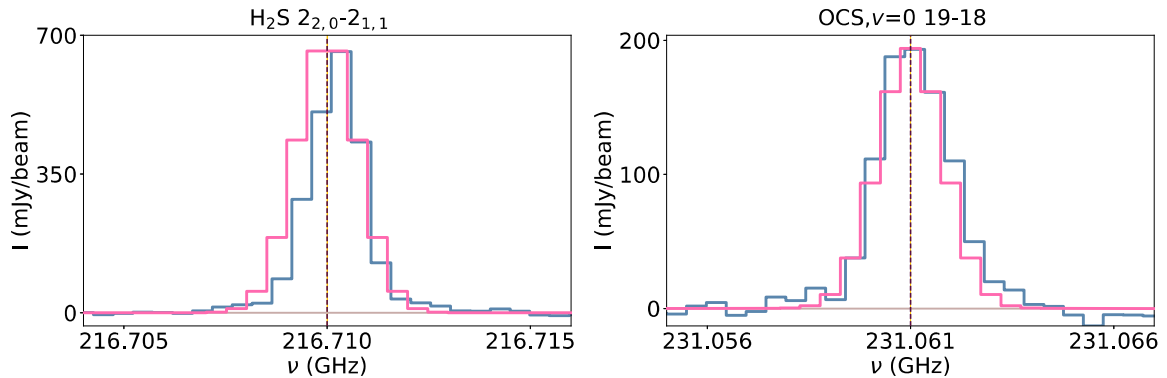


Figure G.2: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward the warm component ( $T_{\text{ex}} = 250$  K) of BHR71-IRS1.

## Appendix G.2: Undetected lines in BHR71-IRS1

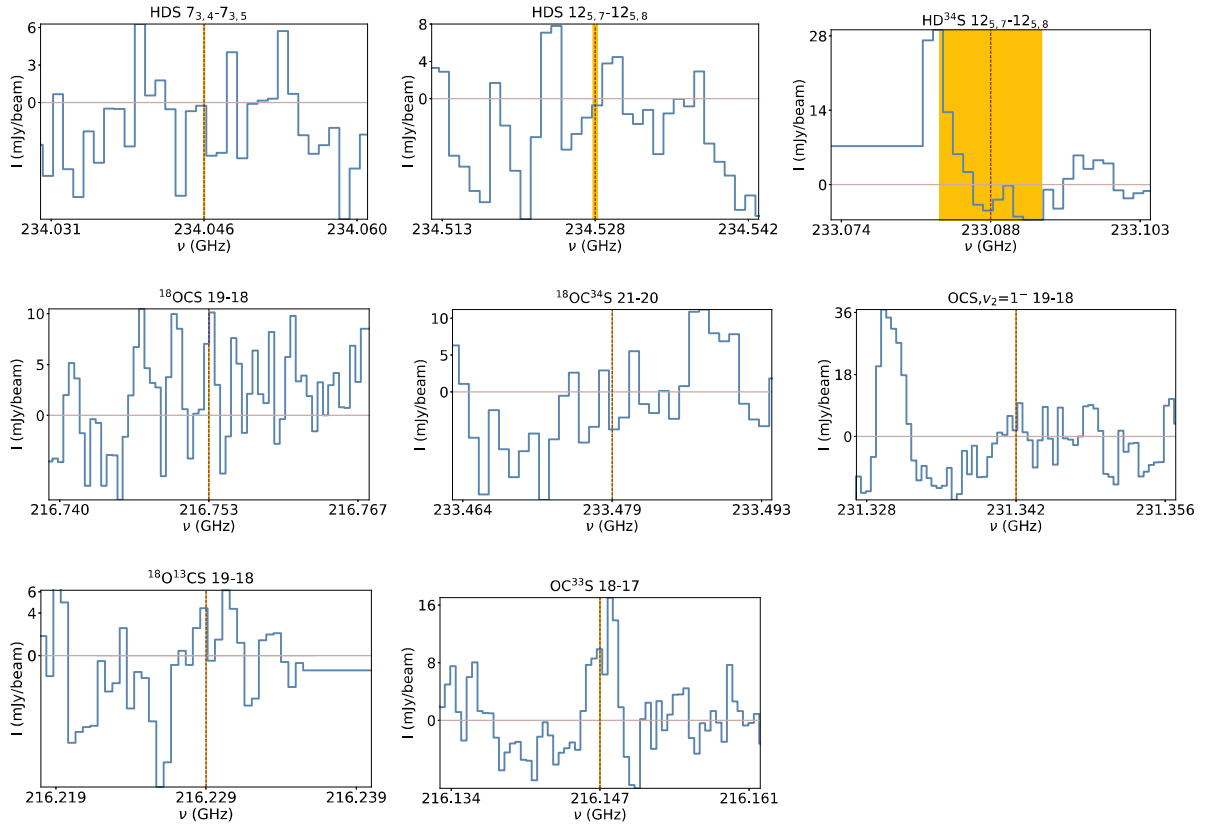


Figure G.3: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward BHR71-IRS1.

## Appendix H: Per-emb-25

No S-bearing species are detected toward Emb-25. Consequently,  $1-\sigma$  upper limits on the column densities of  $\text{H}_2\text{S}$  and  $\text{OCS}$ ,  $v=0$  are computed assuming a source size of  $1''$  (Yang et al. 2021), a FWHM of  $1 \text{ km s}^{-1}$ , and excitation temperatures in the range of  $50 - 300 \text{ K}$ . The upper limits on  $\text{H}_2\text{S}$  and  $\text{OCS}$ ,  $v=0$  column densities are  $\leq 8.3 \times 10^{13}$  and  $\leq 3.2 \times 10^{14} \text{ cm}^{-2}$ , respectively (Table 5 and H.1). The spectra of all the undetected species are presented in Figure H.1.

Appendix H.1:  $1-\sigma$  upper limits on the column densities of  $\text{H}_2\text{S}$  and  $\text{OCS}$ ,  $v=0$  of Per-emb-25

**Table H.1:** Synthetic fitting of the non-detected main S-bearing species toward Per-emb-25 for a range of excitation temperatures between 50 and 300 K, a FWHM of  $1 \text{ km s}^{-1}$ , and a source size of  $1''$ .

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ ( $\text{s}^{-1}$ )	Beam size ( $''$ )	source size ( $''$ )	$N$ ( $\text{cm}^{-2}$ )
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.5	2.0	$\leq 8.3 \times 10^{13}$
$\text{OCS}$ , $v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	6.0	0.5	$\leq 3.2 \times 10^{13}$



## Appendix H.2: Undetected lines in Per-emb-25

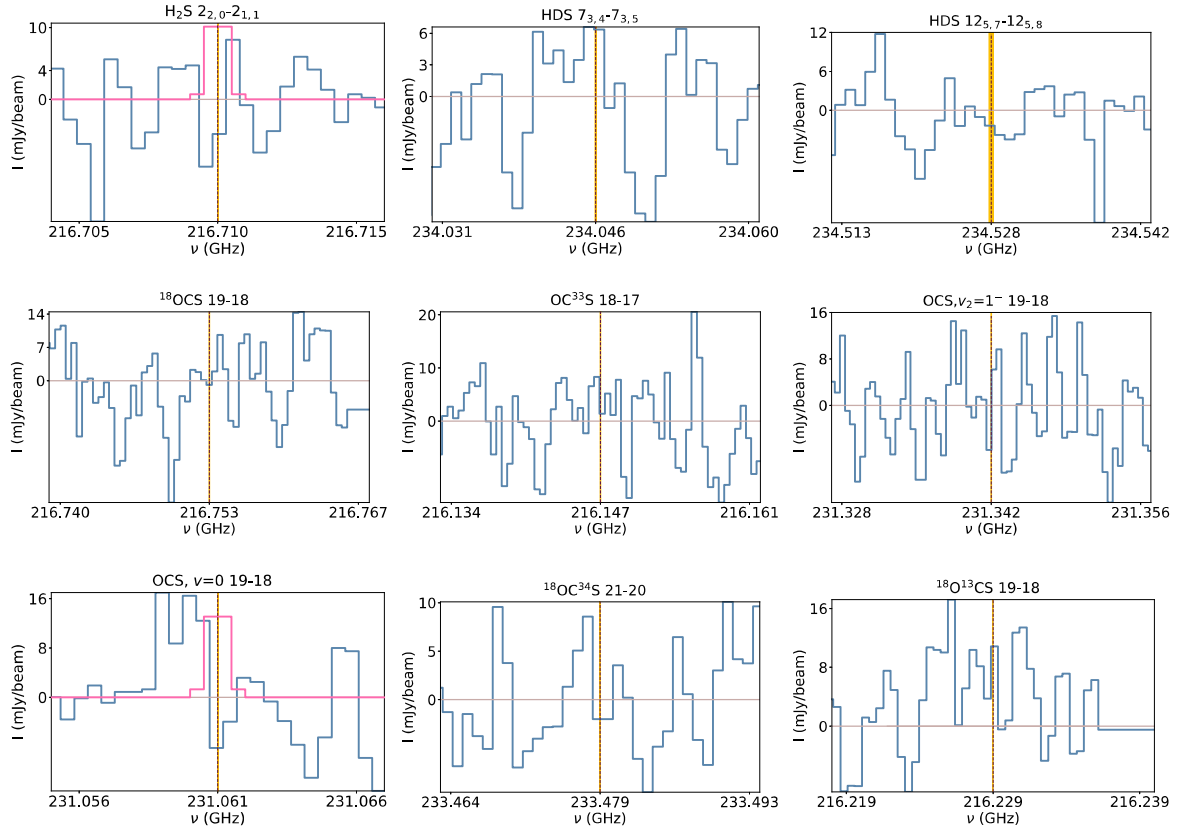


Figure H.1: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward Per-emb-25. Synthetic spectra (in pink) fitted to  $\text{H}_2\text{S}$  and  $\text{OCS}, v=0$  lines with the  $1\text{-}\sigma$  upper limit on their column densities.

## Appendix I: NGC 1333-IRAS4B

$\text{H}_2\text{S}$ ,  $\text{OCS}, v=0, J = 19 - 18$ , and  $\text{OC}^{33}\text{S}, J = 18 - 17$  emission lines are detected in IRAS4B. Initially, an excitation temperature and a source size of 210 K and  $2''$ , respectively, were assumed for the modeling based on the rotation temperature of the detected  $\text{CH}_3\text{OH}$  lines toward this source (Yang et al. 2021). In this case, the computed  $\tau$  values indicate that all the lines are optically thin (Table I.1). However, there is a mismatch by an order of magnitude in the column density of  $\text{OCS}$ , as determined from its  $v=0$  line and the  $\text{OC}^{33}\text{S}$  line. It is unlikely that the isotopic ratio of  $^{32}\text{S}/^{33}\text{S}$  deviates from the canonical value for this one source in the well-studied NGC 1333. It is more likely that the assumed source size and excitation temperature are not appropriate for  $\text{OCS}$  (and  $\text{H}_2\text{S}$  by extension). An exploration of the parameter space demonstrated that an excitation temperature of 100 K and a source size of  $1''$  are more appropriate for these S-bearing molecules (Table I.2). In this case, the  $\text{OCS}, v=0$  line is partially optically thick and thus the column density determined from its optically thin minor  $\text{OC}^{33}\text{S}$  isotopolog is more reliable. The fitted synthetic spectra and spectra of undetected species are shown in Figure I.1 and Figure I.2, respectively. Figure I.1 shows that the lines of S-bearing molecules deviate from the literature  $v_{\text{LSR}} = 7.4$  km/s value that was determined based on CO observations Kristensen et al. (2012). A single  $v_{\text{LSR}}$  value does not fit all the emission lines detected in this work. This discrepancy can only be understood with a more thorough characterisation of the structure of this source on large and small spatial scales.

**Table I.1:** Synthetic fitting of the detected S-bearing species toward NGC 1333-IRAS4B for an excitation temperature of 210 K and a source size of  $2''$ . Directly across from a specific minor isotopolog under ‘‘Derived  $N$  of isotopologs’’ follows the column density of the main isotopolog upon the assumption of the standard isotopic ratio.

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ ( $\text{s}^{-1}$ )	Beam size ( $''$ )	FWHM $\text{km s}^{-1}$	$N$ ( $\text{cm}^{-2}$ )	Derived $N$ of isotopologs ( $\text{cm}^{-2}$ )	$\tau$
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.5	2.0	$2.5^{+0.3}_{-0.3} \times 10^{15}$	–	0.050
$\text{OC}^{33}\text{S}$	18-17	216.147	99	$2.9 \times 10^{-5}$	6.5	1.5	$7.6^{+3.4}_{-1.3} \times 10^{13}$	$N(\text{OCS})=9.5^{+4.3}_{-1.6} \times 10^{15}$	0.004
$\text{OCS}, v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	6.1	2.0	$5.7^{+0.3}_{-0.3} \times 10^{14}$	–	0.020

**Table I.2:** Synthetic fitting of the detected S-bearing species toward NGC 1333-IRAS4B for an excitation temperature of 100 K and a source size of  $1''$ . Directly across from a specific minor isotopolog under ‘‘Derived N of isotopologs’’ follows the column density of the main isotopolog upon the assumption of the standard isotopic ratio.

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ ( $\text{s}^{-1}$ )	Beam size ( $''$ )	FWHM $\text{km s}^{-1}$	$N$ ( $\text{cm}^{-2}$ )	Derived $N$ of isotopologs ( $\text{cm}^{-2}$ )	$\tau$
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.5	2.0	$>5.8 \times 10^{15}$	–	0.56
$\text{OC}^{33}\text{S}$	18-17	216.147	99	$2.9 \times 10^{-5}$	6.5	1.5	$2.2_{-0.4}^{+0.5} \times 10^{14}$	$N(\text{OCS})=2.75_{-0.55}^{+0.65} \times 10^{16}$	0.03
$\text{OCS}, v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	6.1	2.0	$>1.9 \times 10^{15}$	–	0.20

### Appendix I.1: Detected lines in NGC 1333-IRAS4B

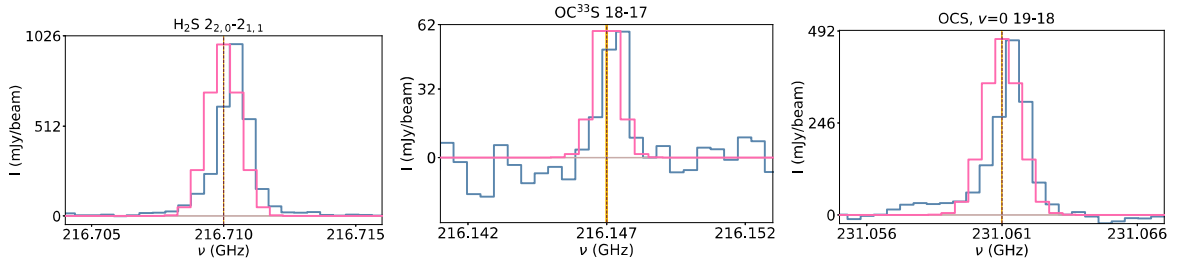


Figure I.1: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward NGC 1333-IRAS4B for  $T_{\text{ex}} = 100$  K and a source size of  $1''$ .

### Appendix I.2: Undetected lines in NGC 1333-IRAS4B

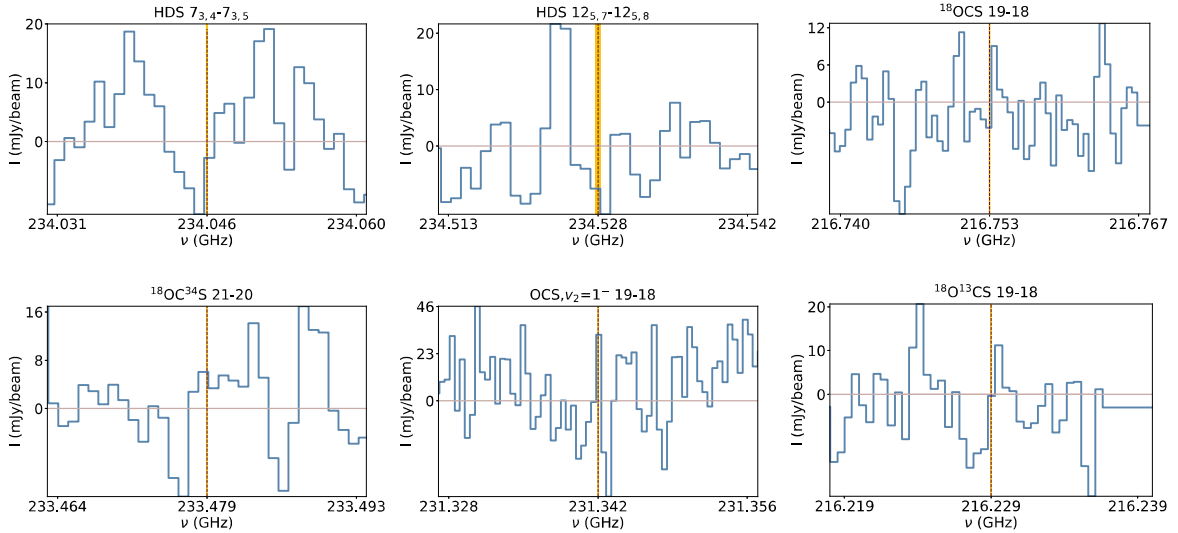


Figure I.2: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward NGC 1333-IRAS4B.

## Appendix J: Ser-SMM3

The excitation temperature was difficult to estimate for SMM3, as no molecular tracer of the inner envelope has been detected. As a result, the column density of the detected  $\text{H}_2\text{S}$ ,  $2_{2,0}-2_{1,1}$  and  $\text{OCS}, v=0, J = 19 - 18$  lines is evaluated for a range of excitation temperatures between 100 and 250 K. The  $\text{OCS}, v=0$  line shows a double peaked feature and the synthetic spectrum is fitted covering both peaks. The best-fitting models to  $\text{H}_2\text{S}$  and  $\text{OCS}, v=0$  have line widths of 2.5 and 3.0  $\text{km s}^{-1}$ , and source sizes of 2.0'' and 0.5'', respectively. Different source sizes for  $\text{H}_2\text{S}$  and  $\text{OCS}$  were assumed in this source, because 12m ALMA data at a spatial resolution of 0.5'' show very different spatial distributions for these molecules. The  $\text{H}_2\text{S}$  and  $\text{OCS}$  integrated intensity maps in fig. 9 of [Tychońiec et al. \(2021\)](#) show extended emission for  $\text{H}_2\text{S}$  covering  $\sim 2''$  and much more compact emission for  $\text{OCS} \sim 0.5''$ . The lines are optically thin ( $\tau < 0.1$ ) with line opacities tabulated in Table J.1.

**Table J.1:** Synthetic fitting of the detected S-bearing species toward Ser-SMM3 for a range of excitation temperatures between 100 and 250 K.

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ ( $\text{s}^{-1}$ )	Beam size ( $''$ )	source size ( $''$ )	FWHM ( $\text{km s}^{-1}$ )	$N$ ( $\text{cm}^{-2}$ )	$\tau$
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.0	2.0	2.5	$5.8^{+3.1}_{-3.2} \times 10^{14}$	0.01-0.02
$\text{OCS}, v=0$	19-18	231.060	111	$3.6 \times 10^{-5}$	5.6	0.5	3.0	$8.7^{+4.9}_{-4.9} \times 10^{14}$	0.02-0.05

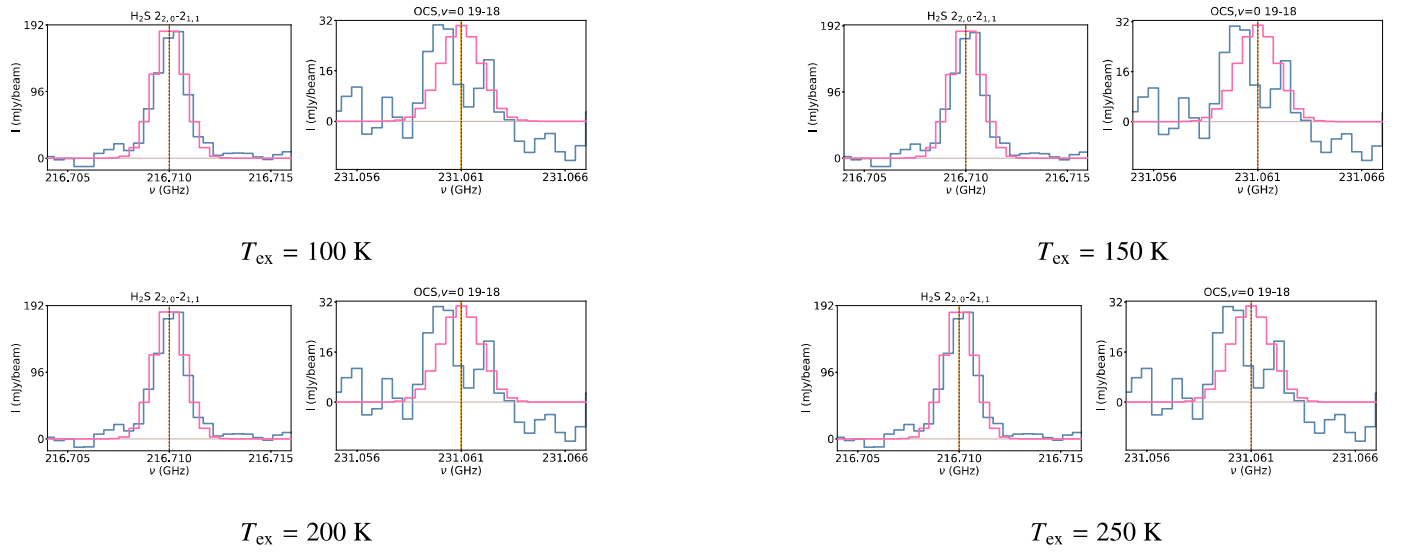
*Appendix J.1: Detected lines in Ser-SMM3*

Figure J.1: Observed spectra (in blue), rest frequency of the detected line (brown dashed line), spectroscopic uncertainty on the rest frequency of the detected line (yellow shaded region), and fitted synthetic spectra (in pink) plotted for the sulfur-bearing species detected toward Ser-SMM3 for a range of excitation temperature between 100 and 250 K.

Appendix J.2: Undetected lines in Ser-SMM3

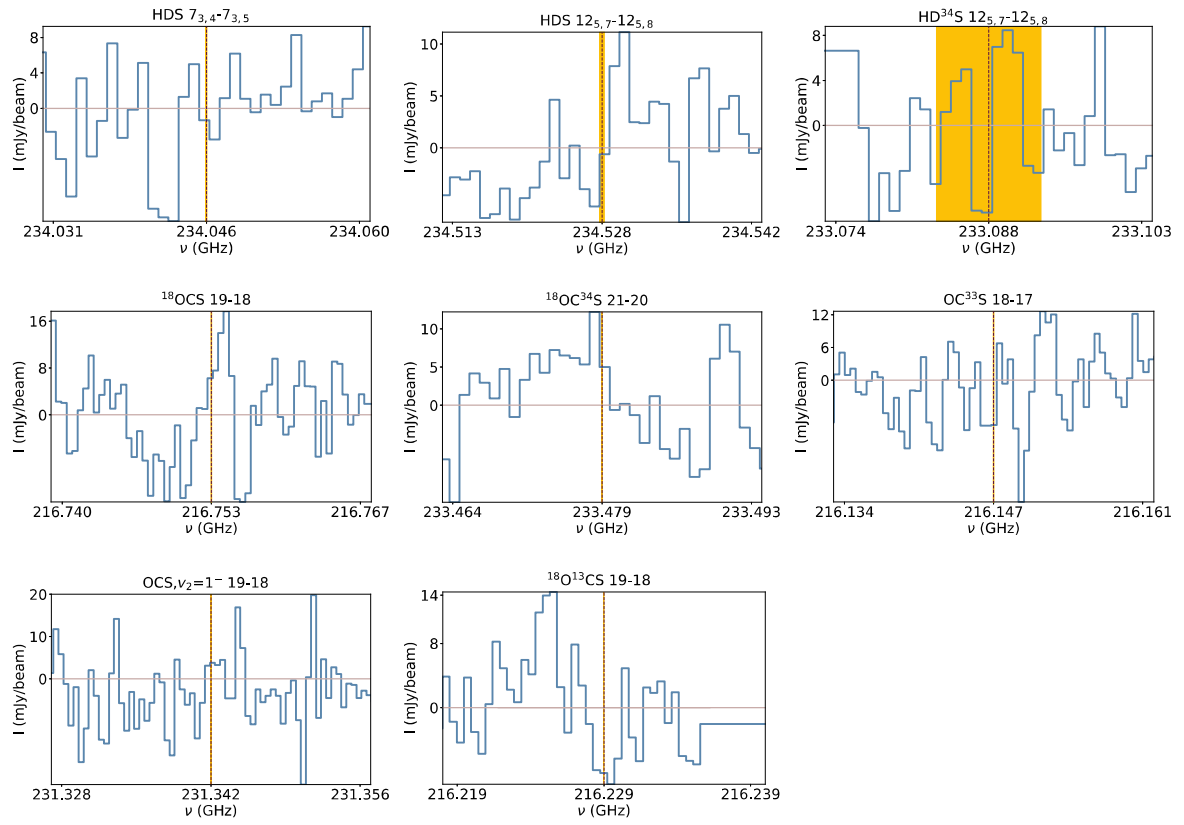


Figure J.2: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward Ser-SMM3.

## Appendix K: TMC1

No S-bearing species are detected toward TMC1. Thus,  $1-\sigma$  upper limit on the column density of  $\text{H}_2\text{S}$ ,  $2_{2,0}-1_{1,1}$  and  $\text{OCS}$ ,  $v=0$ ,  $J=19-18$  are computed with a FWHM of  $1\text{ km s}^{-1}$ , for an excitation temperature of 40 K, and a source size of  $1.5''$  to include both the components of TMC1 (Harsono et al. 2014). The upper limits on the column density of  $\text{H}_2\text{S}$  and  $\text{OCS}$  are given in the Table K.1. The undetected lines are shown in Figure K.1.

**Table K.1:** Synthetic fitting of the undetected main S-bearing species toward TMC1 for an excitation temperature of 40 K, FWHM of  $1\text{ km s}^{-1}$ , and source size of  $1.5''$ .

Species	Transition	Frequency (GHz)	$E_{\text{up}}$ (K)	$A_{ij}$ ( $\text{s}^{-1}$ )	Beam size ( $''$ )	$N$ ( $\text{cm}^{-2}$ )
$\text{H}_2\text{S}$	$2_{2,0}-2_{1,1}$	216.710	84	$4.9 \times 10^{-5}$	6.4	$\leq 1.5 \times 10^{13}$
$\text{OCS}$ , $v=0$	19-18	231.061	111	$3.6 \times 10^{-5}$	5.9	$\leq 2.6 \times 10^{13}$

### Appendix K.1: Undetected lines in TMC1

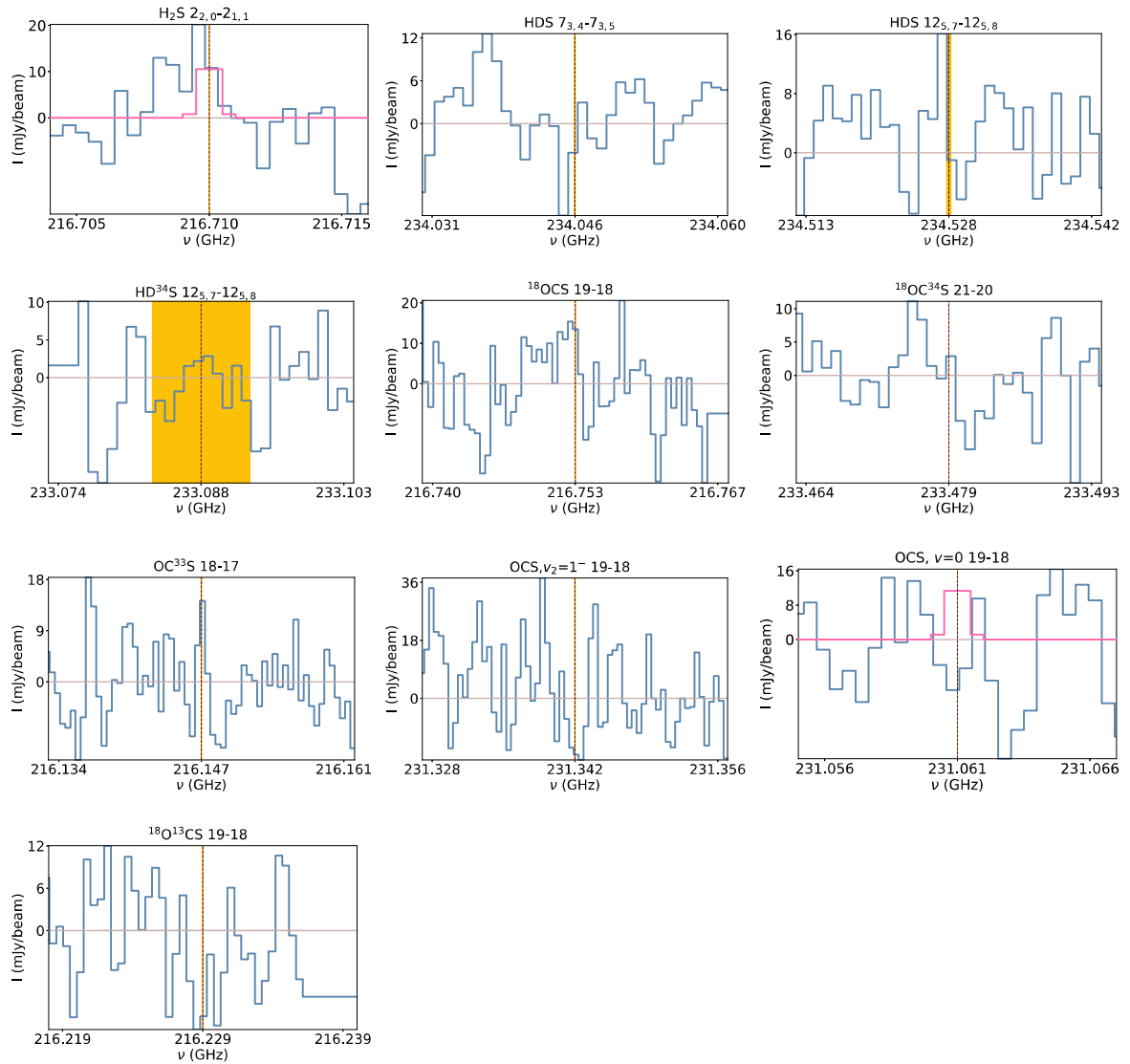


Figure K.1: Observed spectra (in blue), rest frequency of the undetected line (brown dashed line), and spectroscopic uncertainty on the rest frequency of the undetected line (yellow shaded region) plotted for the sulfur-bearing species undetected toward TMC1. Synthetic spectra (in pink) fitted to the  $\text{H}_2\text{S}$ , and the  $\text{OCS}$ ,  $v=0$  line with the  $1-\sigma$  upper limit on its column density.