

# Multiparameter tests of general relativity using a principle component analysis with next-generation gravitational-wave detectors

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Principal component analysis (PCA) is an efficient tool to optimize multiparameter tests of general relativity (GR), wherein one looks for simultaneous deviations in multiple post-Newtonian phasing coefficients. This is accomplished by introducing non-GR deformation parameters in the phase evolution of the gravitational-wave templates used in the analysis. A PCA is performed to construct the “best-measured” linear combinations of the deformation parameters. This helps to set stringent limits on deviations from GR and to more readily detect possible beyond-GR physics. In this paper, we study the effectiveness of this method with the proposed next-generation gravitational-wave detectors, Cosmic Explorer (CE) and Einstein Telescope (ET). For compact binaries at a luminosity distance of 500 Mpc and the detector-frame total mass in the range 20–200 $M_{\odot}$ , CE can measure the most dominant linear combination with a 1- $\sigma$  uncertainty  $\sim 0.1\%$  and the next two subdominant linear combinations with a 1- $\sigma$  uncertainty of  $\leq 10\%$ . For a specific range of masses, constraints from ET are better by a factor of a few than CE. This improvement is because of the improved low frequency sensitivity of ET compared to CE (between 1–7 Hz). In addition, we explain the sensitivity of the PCA parameters to the different post-Newtonian deformation parameters and discuss their variation with total mass. We also discuss a criterion for quantifying the number of most dominant linear combinations that capture the information in the signal up to a threshold.

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## I. INTRODUCTION

Advanced LIGO [1] and Advanced Virgo [2] have observed  $\sim 100$  gravitational wave (GW) signals from compact binary coalescences [3–11] during their first three observing runs, providing a unique opportunity to test the validity of Einstein’s theory of general relativity (GR) in the strong field regime [12–15]. To date, GW observations have shown an excellent agreement with the predictions of GR [16–18]. However, the planned upgrades in the sensitivities of Advanced LIGO and Virgo [19], along with the prospect of KAGRA [20] and LIGO-India [21] joining the current network of GW detectors, will allow us to carry out tests of GR with greater precision.

A theory-agnostic test of GR that is routinely performed on GW signals is the parametrized tests of GR [22–28].

This is a set of null tests in which the phase of the GW signal, as predicted by the post-Newtonian (PN) approximation to GR, is deformed at every order by introducing fractional deformation parameters. These deformation parameters capture potential GR violations via the physical effects that appear at different PN orders [29] that are absent in GR. This parametrization assumes that a modified theory of gravity will result in phase evolution of the emitted signal different from that in GR [30–32]. Such deviations from GR at various PN orders may be detectable by measuring these phenomenological PN deformation parameters and testing for their consistency with the predictions of GR [22,23,33,34].

In this paper, we study tests of GR using the inspiral phase of the GW signal. During this phase, the binary components are adiabatically inspiralling towards each other under gravitational radiation reaction. Higher order PN coefficients in the formula for the phase evolution contain multiple nonlinear interactions (such as tails [35,36], tails of tails [37,38]) and physical effects (such as spins-orbit and spin-spin interactions [39–42]). Further, most of these

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effects appear at multiple PN orders [36–40]. Hence, from a theoretical perspective, it is natural to expect that multiple PN coefficients will get altered if the binary components interact differently in a modified theory of gravity [30–32]. Therefore, simultaneous estimation of *all the PN deformation parameters*, along with the GR parameters (like the component masses and spins, and other extrinsic parameters such as the distance to the source) constitute a more robust test of GR, when compared to tests that deform one PN coefficient at a time.<sup>1</sup> The resultant parametrized tests, where more than one PN deformation parameter is measured along with all the GR parameters, are often called *multiparameter tests of GR* [22,44–46].

Multiparameter tests are less sensitive because of the high correlations among the PN deformation parameters and between them and the intrinsic GR parameters [22,26,47]. Therefore, single-parameter tests are performed where one of the PN deformation parameters is measured, along with the GR parameters, fixing the other deformation parameters to their GR value, zero [18,23,26,47,48].

It has been shown that the brightest prospect for carrying out multiparameter tests is by multiband observations of binary black holes [44,45]. Multibanding combines observation of the same compact binary in the low-frequency band by the Laser Interferometer Space Antenna [49] and the high-frequency band by the next-generation terrestrial detectors like the Einstein Telescope (ET) [50] or Cosmic Explorer (CE) [51]. However, performing the multiparameter tests with only terrestrial or space-space detectors alone remains challenging because of the aforementioned parameter correlations.

One way to address this problem is to identify a new set of basis functions in terms of which the variance-covariance matrix of the problem is diagonal. We can do this by using the method of principal component analysis (PCA) on the variance-covariance matrix corresponding to the PN deformation parameters after marginalizing over the GR parameters [52,53]. The eigenvectors corresponding to the smallest eigenvalues provide the best measured (smallest measurement uncertainties) parameters, which are linear combinations of the original PN deformation parameters.

In this paper, we diagonalize the inverse of the covariance matrix, called the Fisher information matrix, to identify the principal components. In the diagonalized Fisher matrix, the most informative eigenvectors have the highest eigenvalues. Hence, ranking the eigenvalues in decreasing order and identifying the corresponding eigenvectors help us deduce the set of most informative new deformation parameters. The newly constructed parameters are linear combinations of the original ones and are zero in GR. We will call these new parameters as PCA parameters in the rest of the paper.

<sup>1</sup>See, however, Ref. [43] for a recent study on the robustness of the single-parameter tests.

The trace of the diagonalized Fisher matrix (sum of the eigenvalues associated with the PCA parameters) captures the total amount of information it carries. The ratio of the eigenvalues to the trace of the Fisher matrix helps truncate the matrix to a specified accuracy. In summary, PCA provides a way to optimize the multiparameter tests and extract maximum information from the data with fewer additional “null” deformation parameters.

In a recent paper [53], we demonstrated the use of this method to carry out multiparameter tests on GW events detected during the first (O1) and second observing (O2) runs of Advanced LIGO and Advanced Virgo, within the framework of Bayesian inference. We also discussed the effectiveness of the PCA parameters in detecting GR violations with simulated non-GR signals. The GR violations were introduced at every PN order starting from 1.5PN to 3.5PN. We found that the most dominant PCA parameter could recover the injected non-GR values by excluding the GR value (zero) with high credibility. However, due to the limited sensitivity of Advanced LIGO and Virgo during O1/O2, we could only construct the PCA parameters by simultaneously varying deformations from 1.5PN to 3.5PN (six-parameter tests). Recently, an independent work [54] also explored this method in the context of binary neutron star merger GW170817 [55] albeit with a different parametrization and procedure.

In this paper, we investigate the application of PCA-based multiparameter tests with CE and ET to be built in the United States and Europe, respectively. The triangular shaped ET has a noise power spectral density similar to the CE, with a better strain sensitivity at frequencies below about 7 Hz. The improved sensitivity of these detectors should allow us to perform a powerful eight-parameter test and compute the projected bounds on the corresponding dominant PCA parameters, as opposed to the six-parameter test available to the current generation of detectors. We use the IMRPhenomD waveform model for this study, which comprises inspiral, merger, and ringdown phases of the binary evolution and considers spinning but nonprecessing configurations [56].

More specifically, we consider the following questions:

- (1) How many of the PCA parameters can be measured with a  $1\text{-}\sigma$  uncertainty of  $\leq 10\%$ , and how does this number depend on the total mass of the binary?
- (2) How does the performance of PCA-based multiparameter tests compare for the eight-parameter and six-parameter cases?
- (3) What is the relative contribution of each of the original PN deformation parameters to the PCA parameters, and how do they vary with the system’s total mass?
- (4) How can one decide the number of PCA parameters required to capture the information in the signal up to a given accuracy, and how does this number vary with the total mass of the binary system?

The rest of the paper is organized as follows. In Sec. II, we elaborate on parametrized tests of GR and the details related to the parameter space of the signal. In Sec. III, we describe our method of performing PCA on the Fisher information matrix, constructing the linear combinations of the original PN deformation parameters to build new ones, and computation of bounds on the new PCA parameters. In Sec. IV we study the bounds obtained on the PCA parameters as a function of total mass and investigate the sensitivity of these new linear combinations on PN deformation parameters. Finally, we investigate the hierarchy in the significance of these new PCA parameters in Sec. V, followed by a summary of results in Sec. VI.

## II. PARAMETRIZED TESTS OF GR WITH INSPIRAL-MERGER-RINGDOWN WAVEFORMS

Compact binary systems emit gravitational waves as they inspiral towards each other and finally coalesce. In the frequency domain, the waveform can be schematically written as

$$\tilde{h}(f) = \mathcal{A}(f)e^{i\Phi(f)}, \quad (2.1)$$

where  $\mathcal{A}(f)$  and  $\Phi(f)$  are the amplitude and phase of the gravitational waves. The amplitude  $\mathcal{A}(f)$  depends on the source's intrinsic parameters such as the component masses  $m_1$  and  $m_2$ , dimensionless spins  $\chi_1$  and  $\chi_2$  and the extrinsic parameters like the luminosity distance  $D_L$ , location and orientation of the source binary with respect to the interferometer. At the leading order, the Fourier domain amplitude of the waveform in the inspiral regime is proportional to  $\mathcal{M}^{5/6}D_L^{-1}f^{-7/6}$ , where  $\mathcal{M} = M\eta^{3/5}$  is called the chirp mass,  $\eta = m_1m_2/(M)^2$  is the symmetric mass ratio and  $M = m_1 + m_2$  is the detector-frame total mass, or the redshifted mass of the binary related to the source-frame total mass via  $M_{\text{source}}$  by  $M_z = M_{\text{source}}(1+z)$ .

When the compact binary components slowly inspiral towards each other, the rate of change of the orbital period is much smaller than the orbital period itself ( $\dot{\omega}/\omega^2 \ll 1$ ). In this regime, the phase and amplitude in Eq. (2.1) can be expressed in terms of the PN expansion parameter,  $v \equiv (\pi Mf)^{1/3}$ . The frequency domain phase evolution for binaries in quasicircular orbits, constructed using inputs from the PN theory, reads as [57]

$$\Phi(f) = 2\pi f t_c - \phi_c + \frac{3}{128\eta v^5} \sum_{k=0}^N [(\phi_k + \phi_{kL} \ln v)v^k], \quad (2.2)$$

where,  $t_c$  and  $\phi_c$  are the coalescence time and phase of the signal, respectively. The PN coefficients,  $\phi_k$  ( $k = 0, 2, 3, 4, 6, 7$ ) and  $\phi_{kL}$  ( $k = 5, 6$ ), are unique functions of the source properties, such as the component masses and dimensionless spin magnitudes and directions [58].

An alternative theory of gravity can differ from GR in several ways. These include the presence of extra fields, higher dimensions, and violation of diffeomorphism invariance [15,28]. Modifications due to these may show up in the conservative and the dissipative dynamics of the binary, leading to modifications to the GR phase evolution Eq. (2.2), resulting in a modified set of PN coefficients,  $\phi_k$  and  $\phi_{kL}$ . One of the most generic ways to test for these deviations is using a parametrized waveform model, which has deformation parameters at every PN order. The PN coefficients under this ‘‘null parametrization’’ can be written as

$$\phi_k \rightarrow \phi_k^{\text{GR}}(1 + \delta\hat{\phi}_k), \quad (2.3a)$$

$$\phi_{kL} \rightarrow \phi_{kL}^{\text{GR}}(1 + \delta\hat{\phi}_{kL}). \quad (2.3b)$$

Here,  $\delta\hat{\phi}_k$  and  $\delta\hat{\phi}_{kL}$  are the fractional non-GR deformation parameters. We do not consider deformation to the non-logarithmic part of the 2.5PN coefficient as this term is independent of frequency and hence can be absorbed into the redefinition of  $\phi_c$  [59,60]. Hence, we have eight non-GR PN deformation parameters to constrain using GW data. These PN deformation parameters encapsulate the deviations from GR and are zero in GR [26,61]. If these parameters measured by matching the above-mentioned parametrized waveforms with the GW data are found to be consistent with zero, then we may say the dynamics of the corresponding compact binary merger is consistent with the predictions of GR.

We employ the IMRPHEMOMD [56] waveform family in this study. These semianalytical waveforms account for the inspiral, merger, and ringdown phases of the evolution of the binary in GR. The amplitude of this family of waveforms has only the leading quadrupolar mode and assumes that the binary is composed of BHs whose spins are aligned or antialigned with the orbital angular momentum and hence the companions are nonprecessing. The inspiral part of the IMRPHEMOMD waveform follows from Eq. (2.2) correct up to 3.5PN order. Since our method involves testing only the structure of the inspiral phase as predicted by the PN approximation to GR, we only need to change the ansatz to the inspiral part of the IMRPhenomD model by introducing PN deformation parameters as shown in Eq. (2.3), while leaving the postmerger phase unchanged. In a self-consistent modified theory of gravity, the inspiral and postinspiral dynamics of the binary may both differ from their GR predictions. However, our better understanding of the inspiral dynamics motivates the current choice. In the spirit of null tests, this inspiral-only parametrization should help us detect a deviation, if present in the observed signal, although deviations are not introduced in the postmerger part of the waveform.

Having introduced the null parametrization, our parameter space for this analysis consists of seven GR and eight non-GR deformation parameters:



$$\theta^i = \left\{ \ln D_L, t_c, \phi_c, \ln \mathcal{M}_c, \eta, \chi_1, \chi_2, \{\delta\hat{\phi}_k\}, \{\delta\hat{\phi}_{kL}\} \right\}, \quad (2.4)$$

where  $\chi_1$  and  $\chi_2$  are magnitudes of the companion spins in the direction of the orbital angular momentum. The seven GR parameters include the source's intrinsic and extrinsic parameters relevant for the signal observed in a single detector. The curly brackets enclosing  $\delta\hat{\phi}_k$  and  $\delta\hat{\phi}_{kL}$  denote the set of deformation parameters corresponding to the nonlog and log types.

### III. FISHER INFORMATION MATRIX AND THE PRINCIPAL COMPONENT ANALYSIS

Fisher information matrix formalism [62,63] has been widely used in GW research to forecast measurement uncertainties in the parameter estimation problems. The formalism assumes a waveform model for the astrophysical signal in question and the detector noise's power spectral density (PSD) [57,60,64]. Fisher matrix is the noise-weighted inner product of the derivatives of the gravitational waveforms with respect to the physical parameters of the waveform. More explicitly, the components of the Fisher matrix can be written as [57,60,64,65],

$$\Gamma_{mn} = \left\langle \frac{\partial \tilde{h}(f)}{\partial \theta^m}, \frac{\partial \tilde{h}(f)}{\partial \theta^n} \right\rangle, \quad (3.1)$$

where the noise weighted inner product is defined as

$$\langle a|b \rangle = 2 \int_{f_{\text{low}}}^{f_{\text{high}}} \frac{a(f)b(f)^* + a(f)^*b(f)}{S_n(f)} df. \quad (3.2)$$

In the above definition,  $f_{\text{low}}$  and  $f_{\text{high}}$  are suitably chosen lower and upper frequency cutoffs such that the contribution to the integral from outside this range is negligibly small,  $a(f)^*$  denotes complex conjugate of the function  $a(f)$  and  $S_n(f)$  is the one-sided noise PSD of the detector in question. We consider the noise PSD of CE given in Ref. [66] and choose the lower frequency cutoff to be 5 Hz. The noise PSD for ET is taken from Ref. [67], and the lower frequency cutoff is set at 1 Hz. The upper frequency cutoffs for both the detectors are chosen such that the ratio between the characteristic amplitude  $2\sqrt{f}|\tilde{h}(f)|$  and the CE/ET noise amplitude spectral density is at most 10% [44,45].

If the detector noise is stationary and Gaussian and in the limit of high signal-to-noise ratio (SNR), then the measurement uncertainties in the parameters of the signal are given by the square-root of the diagonal elements of the variance-covariance matrix, which is the inverse of the Fisher matrix. This is expected to be a reasonable approximation to the uncertainties one would get from a numerical sampling of the likelihood function [68]. (See Ref. [65] for a detailed discussion of possible caveats.)

In our problem, the dimensionality of the parameter space is 15: seven GR parameters ( $\vec{\theta}_{\text{GR}}$ ) and eight deformation

parameters  $\vec{\theta}_{\text{NGR}}$ . The parameters of our primary interest here are the set of eight PN deformation parameters,  $\vec{\theta}_{\text{NGR}} = \{\delta\hat{\phi}_0, \delta\hat{\phi}_2, \delta\hat{\phi}_3, \delta\hat{\phi}_4, \delta\hat{\phi}_{5L}, \delta\hat{\phi}_6, \delta\hat{\phi}_{6L}, \delta\hat{\phi}_7\}$ . The Fisher matrix corresponding to this eight-dimensional space can be obtained by marginalizing the full 15-dimensional Fisher matrix in Eq. (3.1) over the seven GR parameters  $\vec{\theta}_{\text{GR}}$ . We use the Schur complement method [69] to marginalize out the GR parameters [44,45]. The marginalized  $p$ -dimensional Fisher matrix,  $\tilde{\Gamma}$  can be calculated as

$$\tilde{\Gamma}_{p \times p} = \Gamma_{p \times p} - \Gamma_{p \times q} \Gamma_{q \times q}^{-1} (\Gamma^T)_{q \times p}, \quad (3.3)$$

where  $\Gamma_{p \times p}$  and  $\Gamma_{q \times q}$  are the non-GR and GR parameter submatrices of the full Fisher matrix, respectively, and  $\Gamma_{p \times q}$  denotes cross terms between the two blocks. The variance-covariance matrix for the non-GR parameters is the inverse of this eight-dimensional marginalized Fisher matrix,  $\tilde{\Gamma}$ . It represents an eight-dimensional error ellipsoid in a space spanned by the deformation parameters  $\vec{\theta}_{\text{NGR}}$ .

The GR parameters, such as the chirp mass, mass ratio, and component spins, appear at different PN orders, leading to direct correlations with the corresponding PN deformation parameters. For instance, there are high correlations between  $\delta\hat{\phi}_0$  and chirp mass,  $\delta\hat{\phi}_2$  and symmetric mass ratio,  $\delta\hat{\phi}_3$  and the component spins, and so on. Therefore, the resulting Fisher matrix is ill conditioned and non-invertible when one tries to simultaneously estimate all the PN deformation parameters and the GR parameters.

Principal component analysis offers a way of addressing the problem of correlations. We compute the eigenvalues of the Fisher information matrix and keep only the dominant few that carry most of the information. The corresponding eigenvectors provide the new deformation parameters, which are the best-measured linear combination of the original ones. The number of new parameters retained will depend on the accuracy with which we want to represent the given data. These new parameters are uncorrelated by construction and lead to a diagonal Fisher matrix in the new representation.

We rewrite the  $8 \times 8$  matrix  $\tilde{\Gamma}$  in terms of basis vectors that diagonalize  $\tilde{\Gamma}$ , which reads

$$\tilde{\Gamma} = \text{USU}^T, \quad (3.4)$$

where  $\text{U}$  is an orthogonal transformation matrix whose columns are the eigenvectors of  $\tilde{\Gamma}$  and  $\text{S}$  is the diagonal singular matrix whose elements are the eigenvalues. Without loss of generality, we will assume that  $\text{S}$  is arranged in decreasing order of the eigenvalues, i.e.,  $S_{11} \geq S_{22} \geq \dots \geq S_{pp}$ . The relative magnitude of the eigenvalues helps us determine the dominant linear combinations of the deformation parameters, which we will call the PCA parameters in this paper. Equation (3.4) represents the transformation of the highly correlated Fisher matrix,  $\tilde{\Gamma}$ ,

to a space where it is diagonal (represented by  $S$ ) with the basis that are linear combinations of the original deformation parameters:

$$\delta\hat{\phi}_{\text{PCA}}^{(i)} = \sum_k U^{ik} \delta\hat{\phi}_k, \quad (3.5)$$

where  $U^{ik}$ ,  $k = 1, \dots, p$ , are the elements of the  $i$ th column of the transformation matrix  $U$ , corresponding to the  $i$ th dominant PCA parameter. They determine the relative weights of the original PN deformation parameters in the PCA parameters. The values of the coefficients  $U^{ik}$  will, in general, depend on the properties of the source binary and is studied in detail in Sec. IV.

As the Fisher matrix is now in the diagonal form, the corresponding covariance matrix can be easily computed. The  $1-\sigma$  uncertainties on the PCA parameters are given by

$$\Delta[\delta\hat{\phi}_{\text{PCA}}^{(i)}] = \sqrt{\frac{1}{S_{ii}}}. \quad (3.6)$$

As mentioned earlier, the diagonal matrix  $S$  may be truncated depending on the relative dominance of the eigenvalues and the accuracy requirement of the problem. The ratio of the eigenvalues is also expected to depend on the source properties, such as masses and spins, and are discussed in detail in Sec. V. We propose retaining as many of the PCA parameters as is required to capture at least 99% of the information in the marginalized non-GR Fisher matrix,  $\tilde{\Gamma}$ . This procedure ensures that the eigenvectors which carry the maximum information are kept and discard those that make the Fisher matrix ill conditioned.

In the following section, we will employ the method described above to compute the bounds on the new PCA parameters and investigate how our ability to measure them with CE and ET observatories varies with the total mass of the binary.

#### IV. PROJECTED BOUNDS ON THE NEW PCA PARAMETERS

In this section, we use the formalism explained in Sec. III to investigate the variation in the bounds of the PCA parameters as a function of the total mass of the binary. We also study the composition of the PCA parameters in terms of the original PN deformation parameters.

##### A. Variation of the PCA parameters with total mass

We assume all binaries to have a mass ratio  $q$  equal to 2 and individual dimensionless spins,  $\chi_1$  and  $\chi_2$  set to 0.2 and 0.1, respectively. We consider sources at a luminosity distance of 500 Mpc, with their redshifted total masses  $M_z$  between  $20 - 200M_\odot$ . Binary systems in this mass range have the highest SNR  $\mathcal{O}(10^2) - \mathcal{O}(10^3)$  accumulated

during the inspiral phase and hence are ideal for this inspiral-based tests of GR.

We consider two kinds of multiparameter tests: (1) an eight-parameter test, where we measure all the eight PN deformation parameters simultaneously, and (2) a six-parameter test, where only the last six PN deformation parameters (1.5PN to 3.5PN) are measured together. In the six-parameter test, the PN deformation parameters at 0PN and 1PN are assumed to be consistent with their GR value of zero. This is a scenario that allows the leading two PN orders in the phase evolution to remain as in GR, but allows simultaneous deviation in PN coefficients beyond 1PN, broadly motivated by effective field theoretic extensions of GR [70,71]. This also allows us to directly compare the bounds obtained with CE and ET observatories with that of Ref. [21], which also used this parametrization.

Figure 1 shows the projected  $1-\sigma$  bounds on the PCA parameters as a function of the total mass of the binary. The solid (dashed) lines represent the bounds on the first three (two) dominant PCA parameters for the eight-parameter (six-parameter) test computed with the CE observatory. The fainter lines denote the same quantities computed with the ET observatory.

We will first discuss the bounds obtained with the CE observatory. We find that for the eight-parameter test, the three most dominant PCA parameters can be measured with  $1-\sigma$  uncertainties less than 10%. For the six-parameter test, we can constrain only two of the dominant PCA parameters with similar uncertainties. The most dominant PCA parameter,  $\delta\hat{\phi}_{\text{PCA}}^{(1)}$  for both the eight-parameter and six-parameter cases have the best bounds of  $\mathcal{O}(10^{-3})$ . Hence, we will call

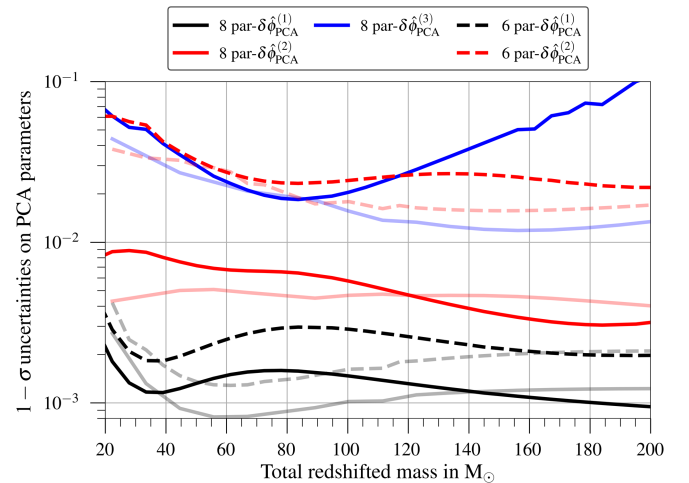


FIG. 1. The solid (fainter solid) lines show uncertainties in the three most-dominant PCA parameters for the eight-parameter case and the dashed (fainter dashed) lines show the uncertainties in the two most-dominant PCA parameters for the six-parameter case, computed with the CE (ET) noise PSD, as a function of total redshifted mass  $M_z$ . The sources are all placed at a luminosity distance of 500 Mpc.

$\delta\hat{\phi}_{\text{PCA}}^{(1)}$  the best measured linear combination for the corresponding multiparameter test. We do not show the bounds on the rest of the subdominant PCA parameters, as they have negligible information on the binary dynamics and are dominated by numerical noise, as discussed in more detail in Sec. V.

The uncertainties in the dominant PCA parameter for both the eight- (solid black line) and six-parameter (dashed black line) cases have similar features, as shown in Fig. 1. The bounds slowly fall off with increasing total mass, reaching a minimum around  $\sim 35M_{\odot}$ . The gradual improvement in the constraints occurs for the following reasons: (1) the signal amplitude and the SNR increase with the total mass of the system and (2) the late time dynamics of larger mass systems is observed with greater sensitivity, making the higher PN orders carry significant information. These effects together help in resolving the degeneracies between the lower order PN deformation parameters with the chirp mass and mass ratio—predominantly  $M_c$  with  $\delta\hat{\phi}_0$ , and  $\eta$  with  $\delta\hat{\phi}_0$ ,  $\delta\hat{\phi}_2$ , and  $\delta\hat{\phi}_3$ . In the next subsection, we explore the varying degree to which this happens across the mass range and influences the bounds on the PCA parameters. The minimum around  $\sim 35M_{\odot}$  results from a complex interplay between the two factors mentioned above. The uncertainties worsen above  $\sim 35M_{\odot}$  due to the decrease in the in-band number of inspiral cycles as the total mass increases.

The above two reasons also apply to the bounds on the second-dominant PCA parameter  $\delta\hat{\phi}_{\text{PCA}}^{(2)}$  for both the eight- (solid pink line) and six-parameter (dashed pink line) cases, which also decrease with increasing total mass. However, the improvement in the bounds is steeper compared to the dominant PCA parameter. This feature can be attributed to the second dominant PCA parameter being more sensitive to the higher PN order deformation parameters (subdominant dynamics of the system), which only become more pronounced for systems with larger masses. We will study this aspect in more detail in the following subsection.

The bounds on the dominant PCA parameter for the eight-parameter case are better than the bounds from the six-parameter case for every system. This is because in the eight-parameter case,  $\delta\hat{\phi}_{\text{PCA}}^{(1)}$  is a linear combination of all the eight PN deformation parameters, which includes the lowest two PN deformation parameters  $\delta\hat{\phi}_0$  and  $\delta\hat{\phi}_2$ . As these two lower order PN deformation parameters are well measured and mostly contribute to the most dominant PCA parameter, they yield better bounds than the six parameter case. The same argument applies for the second-dominant PCA parameter,  $\delta\hat{\phi}_{\text{PCA}}^{(2)}$ , which is also a function of these two PN parameters and provides a stronger bound in the eight parameter case.

The order of magnitude estimates in measuring the PCA parameters with the ET observatory do not differ much from that of the CE observatory. For the eight-parameter

(six-parameter) case, the bounds on the leading PCA parameter with the ET, improve by a factor of 2–3 (1.5) for masses between  $40 - 100M_{\odot}$ , as compared to the CE. The bounds on the third-dominant PCA parameter for the eight-parameter case improve the most—by a factor of 2–9, for higher mass systems ( $\gtrsim 100M_{\odot}$ ), as shown in Fig. 1. This improvement for the ET compared to the CE is due to the improved low-frequency sensitivity of the former, leading to an increase in the number of cycles between 1 and 5 Hz.

## B. Sensitivity of the PCA parameters to the different PN deformation parameters

Next, we explore in depth the composition of the PCA parameters. The PCA parameters are optimal linear combinations of the PN deformation parameters, as shown in Eq. (3.5). In this section, we study the relative contribution of the original PN deformation parameters to the bounds on the PCA parameters and the variation with the total mass of the system. We can study the composition of the PCA parameters by comparing the magnitude of the elements of the transformation matrix  $U^{ik}$  in Eq. (3.5).

The transformation matrix is unique for every system and strongly depends on the total mass. The structure of the PCA parameters will also tell us about the dynamics they probe. For example, a PCA parameter which is composed of coefficients with large magnitudes corresponding to lower-order PN deformation parameters, e.g.,  $\delta\hat{\phi}_0$ ,  $\delta\hat{\phi}_2$ , and  $\delta\hat{\phi}_3$ , and negligible magnitudes corresponding to higher-order deformation parameters, can only probe the early inspiral phase.

We next study the composition of the different PCA parameters in terms of the original PN ones in the case of the CE observatory. The plots on the left and right columns of the top row of Fig. 2 show the variation in the magnitude of the coefficients of the dominant PCA parameter for the eight- and six-parameter cases, respectively. The top-left panel of Fig. 2 clearly shows that for the eight-parameter case, only 0PN ( $\delta\hat{\phi}_0$ ), 1PN ( $\delta\hat{\phi}_2$ ), and 1.5PN ( $\delta\hat{\phi}_3$ ) deformation parameters contribute significantly to the bounds on the dominant PCA parameter,  $\delta\hat{\phi}_{\text{PCA}}^{(1)}$ . Similarly, from the top-right panel of Fig. 2, we see that the most dominant PCA parameter for the six-parameter case is mostly sensitive to the 1.5PN ( $\delta\hat{\phi}_3$ ) and 2.5PN-log ( $\delta\hat{\phi}_{5\text{L}}$ ) deformation parameters.

From the values of  $|U^{ik}|$  in the top-left panel of Fig. 2, 1PN and 1.5PN deformation parameters dominate over 0PN for lower masses (up to  $\sim 50M_{\odot}$ ). With increasing total mass, 0PN rises in dominance and dictates the overall characteristics of the bounds on  $\delta\hat{\phi}_{\text{PCA}}^{(1)}$ . This happens because, for lower masses the contribution from the higher PN orders is not enough to break the degeneracy between the chirp mass and 0PN deformation parameter  $\delta\hat{\phi}_0$ . For higher mass systems, the higher order PN coefficients

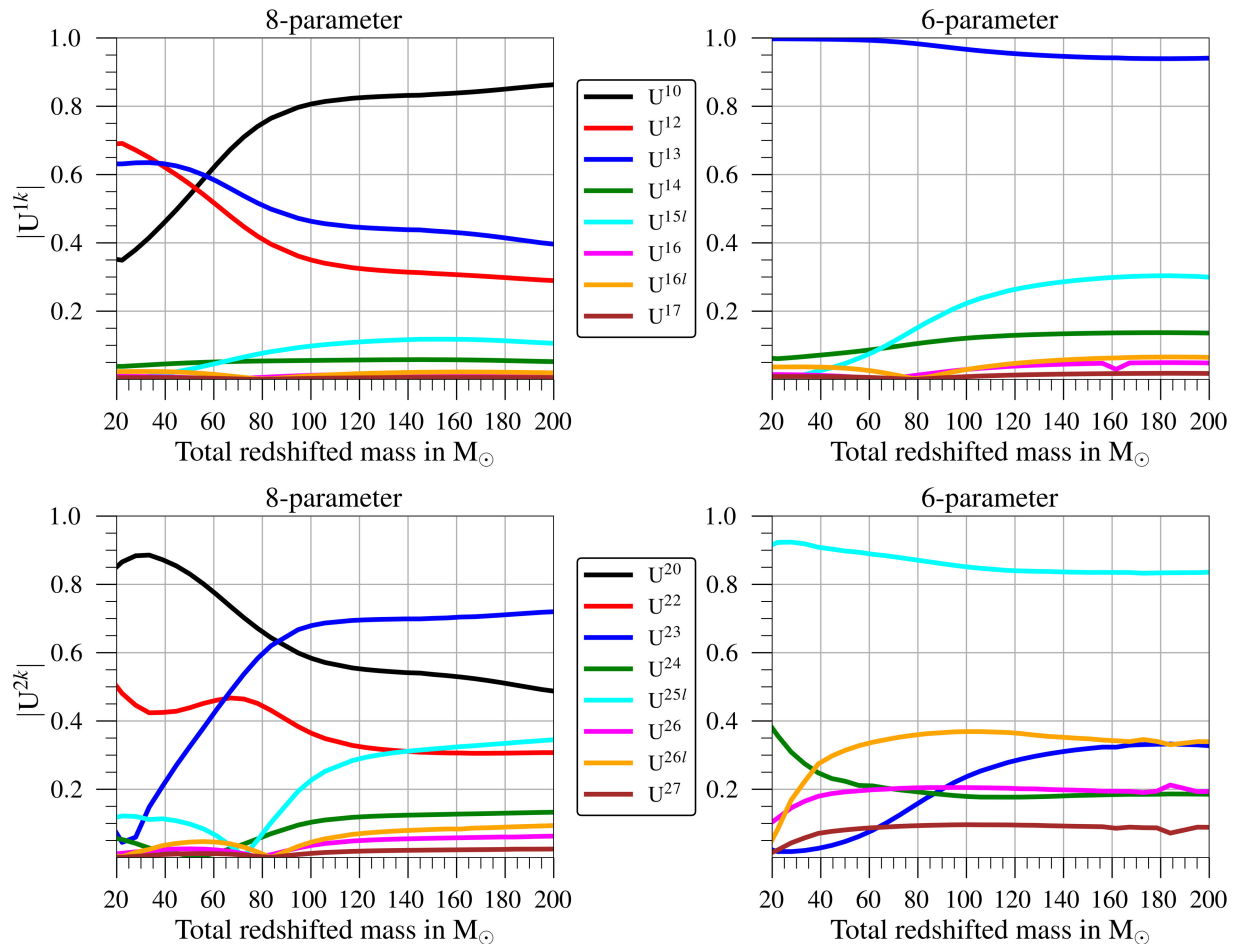


FIG. 2. Absolute values of the coefficients of the most dominant (top row) and second-dominant (bottom row) PCA parameter from eight-parameter (left column) and six-parameter (right column) cases, as a function of redshifted total mass of the systems,  $M_z$ .

become relevant to break the  $M_c$ - $\delta\hat{\phi}_0$  degeneracy along with an increase in the SNR. This causes the OPN deformation parameter to influence the bounds on the eight-parameter- $\delta\hat{\phi}_{\text{PCA}}^{(1)}$  the most (denoted by solid black line in Fig. 1).

The relative values of  $|U^{1k}|$  in the top-right panel of Fig. 2 convey that, for the six-parameter case, it is the 1.5PN deformation parameter that has the strongest influence on the measurability of the most dominant PCA parameter. However, as the total mass of the system increases, the 2.5PN-log deformation parameter rises in significance, although its relative contribution is almost three times lower compared to the 1.5PN deformation parameter, even for the highest masses.

The variation in the magnitude of the coefficients of the second dominant PCA parameter  $\delta\hat{\phi}_{\text{PCA}}^{(2)}$  for the eight- and six-parameter cases are shown in the bottom-left and bottom-right panels of Fig. 2, respectively. First important observation from the bottom-left panel of Fig. 2 is that 0PN, 1PN, 1.5PN, and 2.5PN-log deformation parameters are the dominant contributors to the eight-parameter  $\delta\hat{\phi}_{\text{PCA}}^{(2)}$ .

Second, the contribution of  $\delta\hat{\phi}^{(0)}$  decreases with increasing mass in contrast to what we saw for the dominant PCA parameter,  $\delta\hat{\phi}_{\text{PCA}}^{(1)}$ . The values of  $|U^{23}|$  and  $|U^{25L}|$  sharply increase as a function of mass, indicating increased sensitivity of  $\delta\hat{\phi}_{\text{PCA}}^{(2)}$  to higher PN effects.

The bottom-right panel of Fig. 2 tells that for the six-parameter case, all the deformation parameters starting from 1.5PN contribute to the bounds on  $\delta\hat{\phi}_{\text{PCA}}^{(2)}$ , with the exception of  $|U^{27}|$  which contributes the least. The bounds on the second dominant PCA parameter for the six-parameter case have the dominant contribution from the 2.5PN-log deformation parameter, unlike the dominant PCA parameter, where the 1.5PN deformation parameter contributes the most (shown in the top-right panel of Fig. 2).

The rest of the subdominant PCA parameters should be sensitive to higher PN deformation parameters. However, they are uninformative because of the significant uncertainties associated with their measurement. This reflects the limitation of what one can do with the sensitivities of next-generation detectors. Lastly, given the similarity in the



overall features between the ET and CE curves, we do not repeat the analysis for the ET.

## V. DETERMINING THE NUMBER OF SIGNIFICANT PCA PARAMETERS

We have seen in the previous section that the principal components are unique to each binary and are determined by the specific parameters characterizing the system. Hence, it is only natural to expect that the relative dominance of the PCA parameters will also depend on the source properties. Evaluation of the relative dominance of the PCA parameters will help us determine the number of significant PCA parameters needed for an accurate reconstruction of the information in the signal. In this section we will describe our proposal to determine the number of principal components that carry most of the information relevant to parametrized tests of GR.

The eigenvalues of the Fisher information matrix encode the relative importance of the eigenvectors in terms of the amount of information they carry. The trace of the diagonal Fisher matrix captures the total information contained in the signal and it is invariant under the transformation in Eq. (3.4). For any given system, let  $\{\lambda_k\}$ ,  $k = 1, \dots, n$ , denote the eigenvalues obtained after finding the PCA parameters from the  $n$ -dimensional marginalized Fisher matrix calculated from Eq. (3.3). We shall further assume that the eigenvalues have been ordered such that  $\lambda_1 \geq \lambda_2 \geq \dots \geq \lambda_n$ . The relative information carried by the  $k$ th eigenvector is calculated by dividing the corresponding eigenvalue by the trace of the matrix:

$$I_k = \frac{\lambda_k}{\sum_{k=1}^n \lambda_k}. \quad (5.1)$$

The cumulative information carried by the  $m$  dominant PCA parameters is calculated as

$$\sum_{k=1}^m I_k = \frac{\sum_{k=1}^m \lambda_k}{\sum_{k=1}^n \lambda_k} \quad (m \leq n). \quad (5.2)$$

We calculate the cumulative information content of the PCA parameters according to Eq. (5.2) for all the systems and plot them as a function of mass in Fig. 3 (these plots are called Scree plots [72]). We use these plots to determine the number of significant PCA parameters for a given system depending on the amount of information we want to capture. We set a threshold of 99% for the total information to be retained by the PCA parameters denoted by the gray shaded line.

The cumulative information monotonically increases as we include more subdominant PCA parameters. The sharpest increase in the information is between the most-dominant and the second-most-dominant PCA parameters. The increment beyond this point is relatively slower and

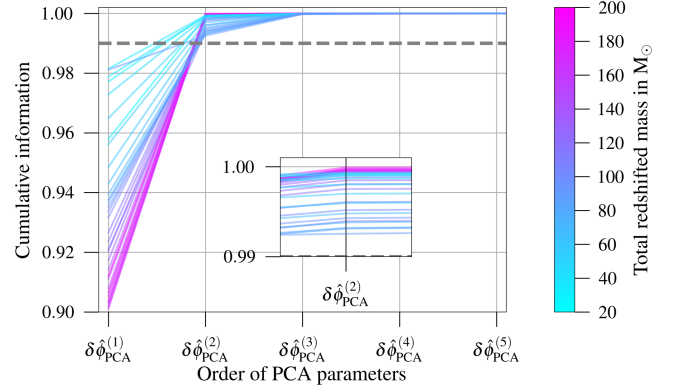


FIG. 3. Cumulative information carried by the PCA parameters as defined in Eq. (5.2) for different masses. Dashed horizontal lines indicate 99% information.

quickly saturates to values very close to one as more of the subdominant PCA parameters are included. This is why we have shown only the five most dominant PCA parameters in the plot.

The amount of information captured by  $\delta\hat{\phi}_{\text{PCA}}^{(1)}$  gradually increases from high to low mass systems. This is because the dominant PCA parameter is mostly influenced by the first three PN deformation parameters ( $\delta\hat{\phi}_0, \delta\hat{\phi}_2, \delta\hat{\phi}_3$ ) as shown in the top-left plot in Fig. 2. The GW signals become increasingly inspiral dominated with decreasing mass, allowing the lower PN orders to carry most of the information. On the other hand,  $\delta\hat{\phi}_{\text{PCA}}^{(2)}$  carries more information for higher mass systems as compared to lower ones (see inset of Fig. 3). This is because high-mass systems probe more relativistic dynamics encoded in the higher order PN deformation parameters, to which  $\delta\hat{\phi}_{\text{PCA}}^{(2)}$  is more sensitive as shown in the bottom-left panel of Fig. 2.

We find that the first PCA parameter captures over 90% of the information for all the systems analyzed, but none cross the required threshold (all the points lie below the gray dashed line marking 99% of the information). This suggests we must consider at least one of the subdominant PCA parameters. The most-dominant and the second-most-dominant PCA parameters together capture more than 99% of the information as shown in the Fig. 3 and hence meet our criteria.

## VI. CONCLUSION

We discussed the optimization of multiparameter tests of GR, where one simultaneously measures parametrized departures from GR in the GW phase evolution predicted by the PN theory, using the principal component analysis. We studied this in the context of the Cosmic Explorer and Einstein Telescope observatories. Principal component analysis helps identify the best-measured deformation parameters that are linear combinations of the original PN deformation parameters. These new PCA parameters



not only have the least statistical uncertainties but are also sensitive to multiple PN deformation parameters, making the test of GR more robust.

Below we list the most important findings of this paper which answer the questions we posed in the introduction:

- (1) Three (two) of the most dominant PCA parameters can be constrained within a  $1\text{-}\sigma$  uncertainty of less than 10% for the case where all the eight (six) PN deformation parameters were varied simultaneously. The variation of the measurement uncertainties with the total mass of the binary were explored in detail in Sec. IV A and Fig. 1.
- (2) The PCA parameters for the six-parameter test, where the six PN deformation parameters (1.5PN to 3.5PN) were varied, have weaker bounds compared to the case where all the eight PN deformation parameters are simultaneously measured. This difference arises because the 0PN and 1PN deformation parameters that are very well measured from the GW signal, are assumed to follow GR (fixed to their GR value of zero) in the six-parameter test. In the eight-parameter test the contribution from 0PN and 1PN deformation parameters help in constraining the PCA parameters better. This is explained in more detail in Sec. IV A.
- (3) The most dominant PCA parameter is sensitive to lower order PN deformation parameters, whereas the second dominant PCA parameter, captures higher order PN effects better with increasing total mass. Figure 2 shows these features and a detailed discussion can be found in Sec. IV B.
- (4) Lastly, we proposed to use a truncation criterion based on the information carried by each of the PCA parameters. We retain only those PCA parameters that together account for at least 99% of the information contained in the full Fisher matrix. This criterion may be changed depending on the exact science one is interested in (see Fig. 3 and Sec. V for details).

Besides these salient points, we find that the relative weights of PN deformation parameters beyond the 2.5PN log-term are weak. This is true even in the case of six-parameter tests, which implies that the next-generation ground based detectors have limited ability to bound deviations at the higher PN end.

Our waveform model does not account for spin-induced precession and nonquadrupolar modes. Our method should work even with a waveform that captures more complex physics, though the constraints on the PCA parameters will likely differ. Our goal here is to demonstrate the method's applicability in the era of next generation observatories with a simple waveform model without having to deal with technical issues that arise when dealing with a more complex waveform.

In a future publication, we will discuss methods to carry out PCA-assisted multiparameter tests of GR from a BBH population. In a previous paper [21], we demonstrated a

method to combine the bounds on the PCA parameters under the assumption that the deviation from GR at a particular PN order is the same across all events, independent of the source properties. We take the individual joint posterior distributions of the PN deformation parameters measured simultaneously for every event and combine them by multiplying the individual likelihoods. The bounds on the PCA parameters are then computed by diagonalizing this combined joint posterior distribution of the PN deformation parameters. In an upcoming publication, we will demonstrate the use of hierarchical inference [73,74] to combine a population of events to bound the PCA parameters. In this method, we need not conform to the assumption of a shared value of deviation from GR across different events. This assumption is a natural fit for the PCA-based tests where the dependence on the source properties is explicit through  $U^{ik}$  elements.

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## APPENDIX: VARIATION IN THE BOUNDS ON THE PCA PARAMETERS WITH SYMMETRIC MASS RATIO

In Fig. 1, we studied the variation in the uncertainties on the PCA parameters with total mass for sources with mass ratio,  $q = 2$ , and individual spin magnitudes,  $\chi_1 = 0.2$ , and  $\chi_1 = 0.1$ . In this section, we study how the bounds vary with changing mass ratio and spin magnitudes. We explore the parameter space with only CE, as CE and ET yielded similar constraints on the leading PCA parameters.

We first keep the spin magnitudes fixed at  $\chi_1 = 0.2$  and  $\chi_1 = 0.1$ , as before, and vary the mass ratio ( $q = m_1/m_2$ , where  $m_1 > m_2$ ) from 1.01 to 10. The mass ratio,  $q = 1.1$ ,

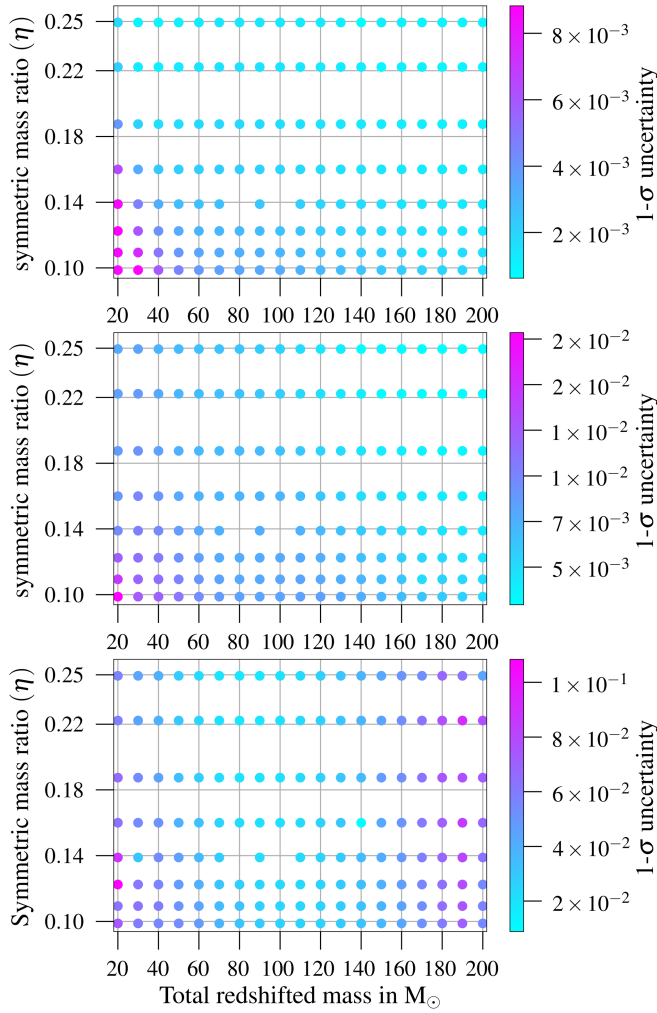


FIG. 4. The top, middle, and bottom panels show the variation in the bounds of the first, second, and third most dominant PCA parameters, respectively, with varying mass and symmetric mass ratio. The individual spin magnitudes are fixed at  $\chi_1 = 0.2$ , and  $\chi_2 = 0.1$ . All the sources are placed at a luminosity distance of 500 Mpc.

denotes an almost equal mass source. We avoid exactly equal mass cases as they are prone to numerical inaccuracies arising from numerical differentiation methods. For every value of the mass ratio, we vary the detector frame masses from  $20$ – $200M_\odot$ . The luminosity distance is fixed at 500 Mpc. With these source parameter configurations,

TABLE I. The table shows the bounds on the three most dominant PCA parameters for sources with detector frame masses of  $20$  and  $200M_\odot$  and for combinations of two mass ratios ( $q = 1.1$  and  $q = 10$ ) and two component spin magnitude configurations ( $\chi_1 = 0.9, \chi_2 = 0.8$  and  $\chi_1 = 0.2, \chi_2 = 0.1$ ). The sources are placed at a luminosity distance of 500 Mpc.

	$q = 1.1$		$q = 10$	
	$\chi_1 = 0.9,$ $\chi_2 = 0.8$	$\chi_1 = 0.2,$ $\chi_2 = 0.1$	$\chi_1 = 0.9,$ $\chi_2 = 0.8$	$\chi_1 = 0.2,$ $\chi_2 = 0.1$
Total mass				
$\delta\hat{\phi}_{\text{PCA}}^{(1)}$				
$20M_\odot$	$1.38 \times 10^{-3}$	$1.87 \times 10^{-3}$	$7.8 \times 10^{-4}$	$8.1 \times 10^{-3}$
$200M_\odot$	$4.9 \times 10^{-4}$	$9.6 \times 10^{-4}$	$7.2 \times 10^{-4}$	$2.2 \times 10^{-3}$
$\delta\hat{\phi}_{\text{PCA}}^{(2)}$				
$20M_\odot$	$7.2 \times 10^{-3}$	$7.9 \times 10^{-3}$	$7.3 \times 10^{-4}$	$2.3 \times 10^{-2}$
$200M_\odot$	$2.9 \times 10^{-3}$	$3.1 \times 10^{-3}$	$2.3 \times 10^{-3}$	$6.1 \times 10^{-3}$
$\delta\hat{\phi}_{\text{PCA}}^{(3)}$				
$20M_\odot$	$2.6 \times 10^{-2}$	$5.7 \times 10^{-2}$	$1.7 \times 10^{-2}$	0.13
$200M_\odot$	$5.7 \times 10^{-2}$	$4.4 \times 10^{-2}$	$3.7 \times 10^{-2}$	$5 \times 10^{-2}$

we calculate the bounds on the first three dominant PCA parameters for CE, as shown in Fig. 1.

The top, middle, and bottom panels in Fig. 4 show the bounds on the first, second, and third dominant PCA parameters for CE. The y axis of the plots denotes the symmetric mass ratio ( $\eta$ ), computed from the corresponding mass ratio using  $\eta = q/(1+q)^2$ . Symmetric mass ratio,  $\eta = 0.25$ , represents sources with equal mass components. The plots show that for a given mass, the bounds improve for the more symmetric systems compared to the asymmetric ones. Although, asymmetric systems have larger number of inspiral cycles, the higher SNRs for symmetric systems contribute more to improving the bounds.

In Table I, we also show the variation in the bounds of the first three most dominant PCA parameters with changing magnitudes of the individual spin components and mass ratio. For both unequal mass ( $q = 10$ ) and almost equal mass cases ( $q = 1.1$ ), high spin magnitude configurations ( $\chi_1 = 0.9, \chi_2 = 0.8$ ) give better bounds than low spin magnitude configurations. This is because high spin magnitudes lead to higher SNR and help break the correlations between the component spins and higher PN order deformation parameters starting from 1.5PN. Antialigned configurations ( $\chi_1 = 0.9, \chi_2 = -0.8$  and  $\chi_1 = 0.2, \chi_2 = -0.1$ ) of the systems considered in Table I give worse bounds as the SNR decreases.

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