

Stellar-gas kinematic misalignments in EAGLE: enhanced SMBH growth in misaligned galaxies

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ABSTRACT

Stellar-gas kinematic misalignments are a transient phenomenon observed in ~ 11 per cent of the local galaxy population. According to current models, misaligned gas is expected to lose angular momentum and relax into the galactic plane on time-scales of ~ 0.1 Gyr, driving gas toward the central regions of the galaxy. Recent observational studies have found a higher incidence of active galactic nuclei in misaligned galaxies. We use the EAGLE (Evolution and Assembly of GaLaxies and their Environments) simulation to explore the connection between stellar-gas misalignments and enhanced central black hole (BH) activity between $0 < z < 1$. We use a sample of ~ 5600 galaxies with a stellar mass of $M_* \geq 10^{9.5} M_\odot$ that feature long-lived stellar-gas alignment, counter-rotation, and unstable misalignments (non-coplanarity). Over time windows of 0.5 Gyr, we find that galaxies experiencing an unstable misalignment have systematically enhanced BH growth during relaxation. Galaxies with long-term counter-rotation show little difference in BH growth compared to aligned galaxies. We suggest that this enhanced BH growth is driven by loss of angular momentum in unstable misaligned gas discs which is able to drive gas inward toward the vicinity of the BH. At $z = 0.1$, we find a greater incidence of overmassive BHs in galaxies that have spent a greater fraction of time with unstable stellar-gas kinematic misalignments over the preceding ≈ 2 Gyr compared to control samples of aligned galaxies. In agreement with observations, we conclude that BH activity is enhanced in misaligned systems in EAGLE and suggest that the presence of overmassive BHs may be indicative of a past stellar-gas kinematic misalignment.

Key words: methods: numerical – galaxies: active – galaxies: evolution – galaxies: general – galaxies: interactions – galaxies: kinematics and dynamics.

1 INTRODUCTION

Supermassive black holes (SMBHs) are predicted to exist at the centre of most massive galaxies. Observational studies have found strong correlations between the properties of the central SMBH and the host galaxy (e.g. J. Silk & M. J. Rees 1998; L. Ferrarese & D. Merritt 2000; A. Marconi & L. K. Hunt 2003). For example, a strong linear correlation has been found between the mass of the SMBH and the mass of the galactic bulge (e.g. N. Häring & H.-W. Rix 2004; J. Kormendy & L. C. Ho 2013). Moreover, positive correlations have been found between the SMBH accretion rate and galactic star formation rate (see T. M. Heckman & P. N. Best 2014, and references therein). At high BH accretion rates, the SMBH becomes visible as an active galactic nucleus (AGN). The energy released by an AGN is thought to be able to affect their host galaxies and is a key requirement in many theoretical models to explain the quenching observed in many local red-sequence early-type galaxies (ETGs). These relations suggest that galaxies co-evolve with their SMBH (see J. Kormendy & L. C. Ho 2013, for a review) and that the supply of cold, molecular gas within galaxies is intricately linked to episodes of SF (star-forming) and AGN activity (e.g. I. Shlosman, M.

C. Begelman & J. Frank 1990; T. Storchi-Bergmann & A. Schnorr-Müller 2019).

Stellar-gas kinematic misalignments (henceforth ‘misaligned galaxies’) have become a key probe in understanding cold gas replenishment in galaxies (e.g. M. Sarzi et al. 2006; T. A. Davis et al. 2011; T. A. Davis & M. Bureau 2016; J. J. Bryant et al. 2019; C. Duckworth, R. Tojeiro & K. Kraljic 2020a; M. K. Baker et al. 2025). Kinematic misalignments are observed in ≈ 11 per cent of local galaxies (J. J. Bryant et al. 2019; S. I. Raimundo, M. Malkan & M. Vestergaard 2023), including $\sim 30 - 40$ per cent of ETGs and ~ 5 per cent of late-type galaxies (LTGs; e.g. T. A. Davis et al. 2011; J. J. Bryant et al. 2019; A. Ristea et al. 2022; S. I. Raimundo et al. 2023). Gas-rich minor mergers were frequently proposed as the dominant formation pathway for misaligned galaxies among early observational studies (e.g. T. A. Davis et al. 2011). However, recent results from hydrodynamical simulations have suggested a more diverse formation pathway with contributions from internal gas replenishment mechanisms (e.g. cooling of hot halo gas; D. Kereš et al. 2005; C. d. P. Lagos et al. 2014), external gas replenishment mechanisms (e.g. gas-rich mergers; D. J. Khim et al. 2021; M. K. Baker et al. 2025), alongside perturbations of the existing corotating gas induced from fly-bys or ram-pressure stripping (e.g. T. K. Starkenburg et al. 2019; D. J. Khim et al. 2021; C. I. Casanueva et al. 2022).

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As outlined in J. E. Tohline, G. F. Simonson & N. Caldwell (1982), any kinematically misaligned component is subject to a radially dependent torque from the mass distribution within the galactic plane (see also G. Lake & C. Norman 1983). For a misaligned gas disc, neighbouring rings of gas become non-coplanar and dissipate angular momentum through cloud–cloud collisions (F. van de Voort et al. 2015). Interactions between corotating gas from stellar mass loss (e.g. J. R. Parriott & J. N. Bregman 2008; S. N. Leitner & A. V. Kravtsov 2011) and counter-rotating gas act as a further source of dissipation (e.g. A. Negri et al. 2014; P. R. Capelo & M. Dotti 2017; P. Taylor, C. Federrath & C. Kobayashi 2018; S. Khoperskov et al. 2021; S. Peirani et al. 2025). Consequently, the misaligned gas disc is able to relax into the galactic plane, while driving gas inward (F. van de Voort et al. 2015). In the absence of ongoing smooth accretion, this relaxation process is predicted to be short-lived with time-scales of ~ 100 Myr for a typical ETG (T. A. Davis & M. Bureau 2016). Using the EAGLE (Evolution and Assembly of GaLaxies and their Environments) simulation (J. Schaye et al. 2015; R. A. Crain et al. 2015), M. K. Baker et al. (2025) found that the majority (≈ 80 per cent) of misalignments relaxes within the expected time-scales, with median relaxation time-scales of ~ 0.5 Gyr. However, longer relaxations ($\gtrsim 2$ Gyr) far exceeding theoretical predictions have also been found (F. van de Voort et al. 2015; D. J. Khim et al. 2021; M. K. Baker et al. 2025). Therefore, kinematic misalignments may provide a mechanism for transporting cold gas into the inner sub-kpc regions of the galaxy for enhanced SF and AGN activity (F. van de Voort et al. 2015; S. Khoperskov et al. 2021; C. Duckworth et al. 2020b).

Correlations between the presence of misalignments and enhanced central SF (e.g. Y.-M. Chen et al. 2016; H. Xu et al. 2022) or AGN activity (e.g. S. J. Penny et al. 2018; S. I. Raimundo 2021; A. Ristea et al. 2022; S. I. Raimundo et al. 2023; S. I. Raimundo et al. 2025; M. W. Winiarska et al. 2025) have been found in several observational studies of nearby galaxies with the use of Integral Field Surveys (IFU). For instance, kinematic mapping of radio galaxies NGC 3100 (I. Ruffa et al. 2019), MCG-06-30-15 (S. I. Raimundo et al. 2013), and NGC 5077 (S. I. Raimundo 2021) reveal the presence of misaligned gas discs, with the resulting inflows acting as a likely driver of AGN activity. Initial studies aiming to establish a connection between AGN activity and misalignments were often limited to low sample sizes (e.g. S. J. Penny et al. 2018; G. S. Ilha et al. 2019). However, a significant correlation was recently established by S. I. Raimundo et al. (2023) using a sample of ~ 1300 galaxies from the Sydney-AAO Multi-object Integral field spectrograph (SAMI) galaxy survey (S. M. Croom et al. 2012; J. J. Bryant et al. 2015). AGNs were found to preferentially reside in misaligned galaxies (≈ 17 per cent) compared to aligned galaxies (≈ 7 per cent) at the 3σ confidence level. Furthermore, the higher incidence of AGNs found among ETGs (≈ 58 per cent) compared to LTGs (≈ 42 per cent) was attributed to the higher incidence (and longevity) of misalignments in ETGs. As such, S. I. Raimundo et al. (2023) highlight the importance of stellar-gas misalignments, in addition to the presence of gas, as a mechanism to trigger AGN activity in galaxies (see also S. I. Raimundo et al. 2025).

One of the challenges with observations is the difficulty in discerning whether the observed AGN signatures in misaligned galaxies is driven by the misaligned gas disc itself or whether these are a ‘smoking gun’ from the initial gas accretion that also formed the misalignment (see S. I. Raimundo et al. 2025). This is because misaligned external accretion (e.g. minor mergers) has been shown both as an effective formation pathway for misaligned discs (e.g. F. van de Voort et al. 2015; D. J. Khim et al. 2021; M. K. Baker et al.

2025) and, more directly, as a means of driving gas inward through interactions with the *in-situ* gas and triggering central SF or AGN activity (L. V. Sales et al. 2012; S. Kaviraj 2014; P. Taylor et al. 2018; T. A. Davis & L. M. Young 2019; S. L. Ellison et al. 2025; S. Peirani et al. 2025). Additionally, outflows driven by episodes of AGN activity, especially in low-mass systems (S. J. Penny et al. 2018; T. K. Starkenburg et al. 2019), may act as a source of corotating gas depletion and disrupt the orientation of inflowing gas (A. Ristea et al. 2022; C. I. Casanueva et al. 2022; E. Cenci et al. 2024). This may increase the likelihood of misalignment formation, especially if the outflowing material is injected into a misaligned halo (e.g. C. d. P. Lagos et al. 2015; C. Duckworth et al. 2020a). As such, observed AGN signatures in misaligned galaxies are often non-trivial to interpret.

Cosmological hydrodynamical simulations such as EAGLE and ILLUSTRISTNG (D. Nelson et al. 2019) have become useful tools to understand the behaviour of misaligned galaxies in the context of gas inflows and AGN fuelling. For instance, accreting misaligned or counter-rotating gas has been established as an effective means to drive gas inward and promote SF through dissipative interactions with *in-situ*, corotating gas (e.g. A. R. Thakar & B. S. Ryden 1996; L. V. Sales et al. 2012; P. Taylor et al. 2018; T. K. Starkenburg et al. 2019; J. J. Davies, A. Pontzen & R. A. Crain 2022; S. Han et al. 2024). This mechanism has also been found to trigger associated AGN activity (e.g. P. Taylor et al. 2018; C. Duckworth et al. 2020b; S. Khoperskov et al. 2021). Likewise, misaligned galaxies have been linked with a lower specific gas angular momentum (T. K. Starkenburg et al. 2019; C. Duckworth et al. 2020a) and more compact SF regions (e.g. C. I. Casanueva et al. 2022). Using the ILLUSTRISTNG simulation, C. Duckworth et al. (2020b) find that misaligned low-mass galaxies have higher peak BH luminosities and enhanced BH growth (as a proxy for integrated BH activity) over the past 8 Gyr compared to an aligned control sample. These results suggest the formation and persistence of a stellar-gas kinematic misalignment may act as a key source of gas inflow and AGN activity in addition to any merger-driven activity.

In this paper, we use the EAGLE suite of cosmological hydrodynamical simulations to investigate the connection between AGN activity and ongoing misalignments. We use the sample of relaxing stellar-gas kinematic misalignments between $0 < z < 1$ from M. K. Baker et al. (2025). Specifically, we aim to investigate whether: (1) BH growth is enhanced in galaxies that are currently unstably misaligned (i.e. their gas discs are relaxing into the dynamically stable co- or counter-rotating regime) compared to samples of stable (aligned and counter-rotating) systems, and (2) the degree to which this imprinted on the galaxy population in the form of overmassive SMBHs at low redshifts.

This paper is structured as follows. In Section 2, we give an overview of the EAGLE simulations. In Section 3, we outline the sample of galaxies we use for this work. In Section 4, we give an overview of our methods. In Section 5, we present our results. In Section 6, we discuss and compare our results to existing results. Finally, we conclude in Section 7.

2 THE EAGLE SIMULATION

EAGLE¹ is a suite of cosmological hydrodynamical simulations described in J. Schaye et al. (2015) and R. A. Crain et al. (2015). The

¹Publicly available data products available from <http://eagle.strw.leidenuniv.nl> and <https://icc.dur.ac.uk/Eagle> and described in S. McAlpine et al. (2016).

EAGLE simulations have been successful in reproducing the observed stellar mass, size, and morphology distributions (e.g. J. Schaye et al. 2015; C. A. Correa et al. 2017; J. W. Trayford & J. Schaye 2019; A. D. Hill et al. 2021). These simulations are run using a heavily modified version of GADGET-3 (V. Springel, T. Di Matteo & L. Hernquist 2005; V. Springel et al. 2008), with an SPH (smoothed particle hydrodynamics) formulation known as ANARCHY (see M. Schaller et al. 2015 for description). Galaxy evolution can be traced using the ‘subhalo’ merger trees by Y. Qu et al. (2017) that were generated using the D-TREES algorithm (L. Jiang et al. 2014). A Lambda-cold dark matter cosmology is assumed with $\Omega_m = 0.307$, $\Omega_\Lambda = 0.693$, $\Omega_b = 0.04825$, $h = 0.6777$, $\sigma_8 = 0.8288$, $n_s = 0.9611$, and $Y = 0.248$ (Planck Collaboration I 2014).

EAGLE relies on subgrid modelling in order to model physical processes below the resolution limit. These include radiative cooling and photoheating (R. P. C. Wiersma, J. Schaye & B. D. Smith 2009a), star formation (J. Schaye & C. Dalla Vecchia 2008), stellar evolution and interstellar medium (ISM) enrichment (R. P. C. Wiersma et al. 2009b), stochastic thermal stellar feedback (C. Dalla Vecchia & J. Schaye 2012), and BH accretion and feedback outlined below.

Black hole (BH) particles are seeded with a mass of $1.48 \times 10^5 M_\odot$ in the centres of dark matter haloes that have a minimum total mass of $1.48 \times 10^{10} M_\odot$ and that do not already contain a BH particle, following the method of V. Springel et al. (2005). The subgrid BH mass, M_{BH} , is allowed to grow via a modified Bondi–Hoyle accretion prescription that accounts for the velocity of surrounding gas, as described in detail by Y. M. Rosas-Guevara et al. (2015). This subgrid accretion rate, $\dot{M}_{\text{BH}}$, is then capped at the Eddington accretion rate. Upon accretion of mass, energy is stored in a reservoir that is injected stochastically into surrounding gas particles using single-mode thermal AGN feedback, as described in C. M. Booth & J. Schaye (2009). The subgrid parameters in EAGLE are tuned to reproduce observations of the galaxy stellar mass function and galaxy sizes relation at $z \approx 0$ (for details, see R. A. Crain et al. 2015).

For this work, we use the largest volume simulation of the reference model known as REF-L100N1504. This simulation has a maximum physical gravitational softening length of 0.7 pkpc (proper kiloparsec), a box length of 100 cMpc (comoving megaparsec), containing an initially equal number (1504^3) of dark matter and baryonic particles, with initial particle masses of 9.7×10^6 and $1.81 \times 10^6 M_\odot$, respectively. The outputs for these data are stored on 400 ‘snapshots’, of which we use 200 (with a mean time cadence of ≈ 120 Myr within $0 < z < 1$) for which merger trees were run by R. A. Crain et al. (2017).

3 SAMPLE

We sample galaxies within $0 < z < 1$ using the same selection criteria as M. K. Baker et al. (2025), with additional considerations made for reliable BH data. Here, we summarize the selection criteria. For full details, see M. K. Baker et al. (2025):

- (i) A minimum stellar mass of $M_* > 10^{9.5} M_\odot$ within a spherical aperture size of 30 pkpc.
- (ii) A minimum star-forming gas (gas_{SF}) particle count of 20 within the stellar half-mass radius, r_{50} , which is defined as the spherical radius enclosing half the stellar mass within 30 pkpc of the stellar centre of mass.
- (iii) A maximum spatial separation of < 2 pkpc between the stellar and gas_{SF} centres of mass within r_{50} .
- (iv) A maximum 1σ uncertainty of $< 30^\circ$ for the misalignment angle using bootstrap resampling. Misalignment angles, ψ_{3D} , are

measured in 3D space between the stellar and gas_{SF} angular momentum vectors within r_{50} .

As in M. K. Baker et al. (2025), we note that our results and conclusions are unchanged if we impose a higher particle limit of 100 gas_{SF} particles, or require a 3σ uncertainty of $< 30^\circ$.

As we intend to investigate the intrinsic behaviour of misaligned gas in the absence of bias from the viewing angle, we focus on misalignment angles in 3D space (though these are not significantly different to projected 2D angles, see e.g. M. K. Baker et al. 2025). We classify galaxies into three classifications; aligned galaxies with angles $\psi_{3D} < 30^\circ$, unstable misaligned galaxies with angles $30^\circ \leq \psi_{3D} \leq 150^\circ$, and counter-rotating galaxies with angles $\psi_{3D} > 150^\circ$, following widely used classifications (e.g. T. A. Davis et al. 2011).

For this work, we focus on the change in BH mass as a proxy for cumulative BH activity. This is because we are limited by the low temporal resolution of our snapshots ($\sim 10^8$ yr) relative to the time-scales of AGN flickering ($\sim 10^5$ yr; K. Schawinski et al. 2015). As highlighted by R. A. Crain et al. (2015), gas accretion for BH particles near the seed mass ($\sim 1.48 \times 10^5 M_\odot$) tends to be slow, because Bondi accretion rates scale with M_{BH}^2 . Likewise, mergers between low-mass BHs in this regime constitute a significant fraction of the BH’s mass, manifesting as characteristic BH masses at multiples of the BH seed mass. As such, we focus our analysis on BHs with a minimum mass of $M_{\text{BH}} > 10^6 M_\odot$ using the most massive BH particle within r_{50} . Using the BH closest to the galactic centre of potential leaves our results unchanged.

In order to compare BH growth between galaxies in varying stages of kinematic stability, we trace the BH growth over 0.50 ± 0.05 Gyr windows in galaxies that meet our selection criteria. This window size is chosen in order to probe long-term BH growth in galaxies that spend a significant amount of time (~ 0.5 Gyr) with an unstable misaligned gas disc. Furthermore, we require no BH mass decreases over this time window and no spontaneous increase in BH mass by a factor of > 5 between consecutive snapshots. These criteria ensure that we trace the same BH particle over the 0.5 Gyr window, while ensuring that we do not attribute significant BH growth if a secondary, lower mass BH particle is erroneously identified as the main BH over one or more snapshots.

Aligned and counter-rotating samples are created by selecting random 0.50 ± 0.05 Gyr windows in which $\psi_{3D} < 30^\circ$ and $> 150^\circ$ for the full 0.50 ± 0.05 Gyr period, respectively. For our sample of unstable misalignments, we utilize the parent sample of 3154 misalignment relaxations used and explained in detail in M. K. Baker et al. (2025). In short, this is a sample of galaxies that experience an unstable misalignment ($30^\circ \leq \psi_{3D} \leq 150^\circ$) and are traced while the gas disc is kinematically unstable (defined as $20^\circ \leq \psi_{3D} \leq 160^\circ$). We note that the first snapshot in the unstable regime typically coincides with the first snapshot in the unstable misaligned regime. An unstable disc is considered ‘settled’ upon returning to the kinematically stable regime (defined as $\psi_{3D} < 20^\circ$ and $\psi_{3D} > 160^\circ$), so long as the galaxy remains in the stable regime for a period of at least another 0.1 Gyr. As in M. K. Baker et al. (2025), we define the relaxation time, t_{relax} , between the first snapshot in the kinematically unstable regime, and the first snapshot in which the unstable disc has settled. In order to fit within the target window duration of 0.50 ± 0.05 Gyr, we extract a sample of relaxations with $t_{\text{relax}} \geq 0.45$ Gyr.

This leaves us with a total sample of 5570 galaxies, of which 4657 are continuously aligned, 532 are continuously counter-rotating, and 381 form our ‘unstable misaligned’ sample that are in the process of relaxing back into the stable regime (i.e. the galactic plane). Each of these samples has a median temporal window of ≈ 0.51 Gyr.

4 METHODOLOGY

As in M. K. Baker et al. (2025), we make use of the corotational energy fraction as a proxy for morphology for stellar and gas_{SF} components (see C. A. Correa et al. 2017). This describes the fraction of the total kinetic energy (K) invested in ordered corotation ($K_{\text{co}}^{\text{rot}}$) and is given by

$$\kappa_{\text{co}} = \frac{K_{\text{co}}^{\text{rot}}}{K} = \frac{1}{K} \sum_{i, L_{z,i} > 0} \frac{1}{2} m_i \left(\frac{L_{z,i}}{m_i R_i} \right)^2, \quad (1)$$

where the sum is over all particles of a given type (stars or gas_{SF}) within a spherical radius (30 pkpc for stars and r_{50} for gas_{SF}) centred on the minimum potential, m_i is the particle mass, $L_{z,i}$ is the particle angular momentum component along the direction of the total angular momentum of a given particle type of the galaxy, and R_i is the radius from the centre of potential in the plane normal to the rotation axis of the given particle type (C. A. Correa et al. 2017). The stellar and gas_{SF} corotational energy fractions are denoted as κ_{co}^* and $\kappa_{\text{co}}^{\text{SF}}$, respectively.

As shown in C. A. Correa et al. (2017), a value of $\kappa_{\text{co}}^* = 0.4$ can be used to approximately distinguish the ‘red sequence’ quiescent spheroidal galaxies (ETGs with $\kappa_{\text{co}}^* < 0.4$) from the ‘blue cloud’ discy SF galaxies (LTGs with $\kappa_{\text{co}}^* > 0.4$). Likewise, as shown in E. Jiménez et al. (2023), $\kappa_{\text{co}}^{\text{SF}} \gtrsim 0.7$ approximately corresponds to a thin gas disc which we confirmed by visual inspection. We note that during the initial stages of misalignment formation and the resulting dissipative effects, we expect $\kappa_{\text{co}}^{\text{SF}}$ to be naturally lower in unstable misaligned galaxies than in dynamically relaxed galaxies.

We define the SF gas fraction as

$$f_{\text{gas,SF}} = \frac{M_{\text{gas,SF}}}{M_{\text{gas}} + M_*}, \quad (2)$$

where M_* , M_{gas} , and $M_{\text{gas,SF}}$ are the stellar, gas, and gas_{SF} masses of the galaxy within the kinematic aperture of r_{50} akin to M. K. Baker et al. (2025).

In order to gain a rough idea of the incidence of AGN within each subsample, we can use the subgrid BH accretion rate to estimate the bolometric luminosity and Eddington ratio of the BH. A common expression (e.g. M. Habouzit et al. 2022) to convert the BH accretion rate in cosmological simulations to a bolometric luminosity is given by

$$L_{\text{bol}} = \frac{\epsilon_r}{1 - \epsilon_r} \dot{M}_{\text{BH}} c^2, \quad (3)$$

where $\epsilon_r = 0.1$ is the radiative efficiency and c is the speed of light in a vacuum.

Additionally, we define the Eddington ratio (λ_{Edd}) as the ratio between the current subgrid BH accretion rate and the theoretically maximum spherically symmetric Eddington accretion rate given by

$$\dot{M}_{\text{Edd}} = \frac{4\pi G M_{\text{BH}} m_p}{\epsilon_r \sigma_T c}, \quad (4)$$

where G is the gravitational constant, m_p is the proton mass, and σ_T is the Thomson scattering cross-section for an electron.

5 RESULTS

5.1 BH growth

5.1.1 Total galaxy population

In Fig. 1, we show the growth of the BH mass (ΔM_{BH}) over our ~ 0.5 Gyr windows compared to the initial BH mass at the start of

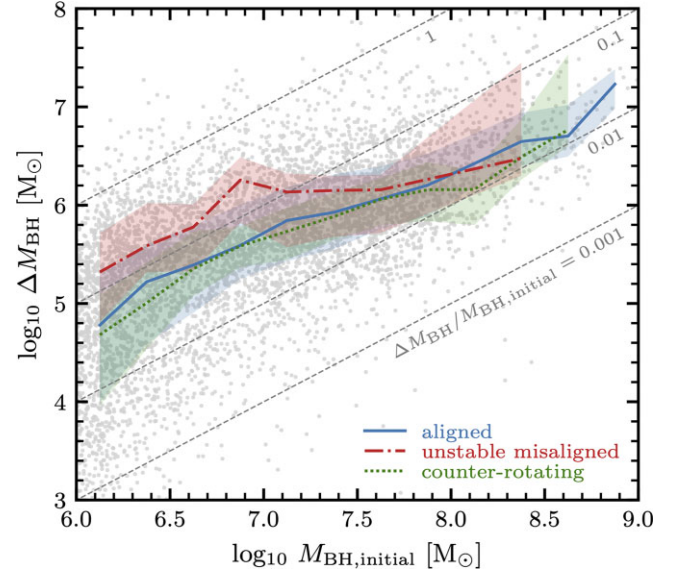


Figure 1. Median change in BH mass traced over a 0.50 ± 0.05 Gyr window as a function of the initial BH mass for samples of galaxies which are aligned (blue solid), counter-rotating (green dotted), and galaxies experiencing an unstable misalignment (red dash-dotted) for bins with a minimum bin count of 10. Shaded regions show the 25 and 75 percentiles. Grey scatter points show the total sample. Diagonal lines (dashed grey) show the relative increase in BH mass, with $\Delta M_{\text{BH}}/M_{\text{BH,initial}} = 1$ indicating a doubling of BH mass over this window. We find that counter-rotating galaxies exhibit similar BH growth to aligned galaxies, while galaxies experiencing an unstable misalignment show significantly higher (~ 0.6 dex) BH growth up to BH masses of $M_{\text{BH,initial}} \lesssim 10^{7.5} M_{\odot}$.

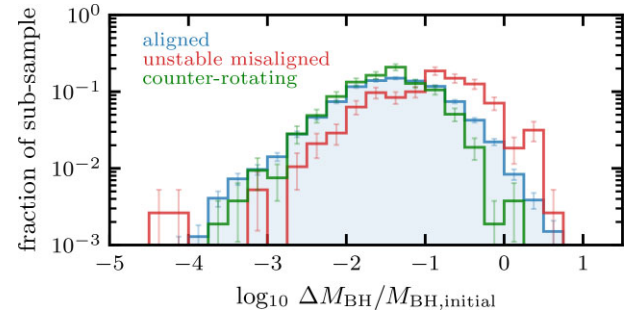


Figure 2. Fractional BH growth distributions over a 0.50 ± 0.05 Gyr window for subsamples of aligned (blue), counter-rotating (green), and galaxies experiencing an unstable misalignment (red). Errors are given as Poisson uncertainties. We find that unstable misaligned galaxies have higher fractional BH growth with a median of $0.129 \Delta M_{\text{BH}}/M_{\text{BH,initial}}$, while aligned and counter-rotating galaxies have similar growth fractions of $0.039 \Delta M_{\text{BH}}/M_{\text{BH,initial}}$ and $0.034 \Delta M_{\text{BH}}/M_{\text{BH,initial}}$, respectively.

the window ($M_{\text{BH,initial}}$). These results are also shown in Fig. 2, in which the BH growth is expressed as a fractional increase in BH mass from the initial BH mass, with distributions given for subsamples of aligned, unstable misaligned, and counter-rotating galaxies.

As seen clearly in Fig. 1, we find galaxies relaxing from an unstable misalignment experience significantly enhanced (~ 0.6 dex) BH growth compared to samples of continuously aligned and counter-rotating galaxies. This is especially true at lower masses of $M_{\text{BH,initial}} \lesssim 10^{7.4} M_{\odot}$. At higher BH masses of $M_{\text{BH,initial}} \gtrsim 10^{7.5} M_{\odot}$, unstable misaligned galaxies show comparable BH growth to their aligned and counter-rotating counterparts. However, we

caution that the sample size of unstable misaligned galaxies with $M_{\text{BH,initial}} \gtrsim 10^{7.5} M_{\odot}$ is low. Over the entire range of BH masses considered, unstable misaligned galaxies grow their BHs by ≈ 12.9 per cent on average over 0.5 Gyr, while aligned systems grew by only ≈ 3.9 per cent. In order to test the significance of this result, we perform a two-sample Kolmogorov–Smirnov (KS) test on the distributions in Fig. 2. We obtain a KS-test statistic = 0.32 and p -value $\ll 10^{-10}$, indicating a strongly significant result given the size of the sample.

Interestingly, counter-rotating systems share a similar BH growth to aligned galaxies within the full BH mass range considered (see Fig. 1). Counter-rotating galaxies grow their BHs by a median of 3.4 per cent compared to 3.9 per cent for aligned galaxies (KS-test statistic = 0.12, p -value = 2.0×10^{-6}). We discuss possible interpretations for this result in Section 5.1.3.

These results remain unchanged if we use the instantaneous subgrid BH accretion rate, $\dot{M}_{\text{BH}}$, time-averaged over the same ≈ 0.5 Gyr window (not shown). We note that the use of the subgrid BH accretion rate should be taken with caution given the significant fluctuations seen in $\dot{M}_{\text{BH}}$ that may be poorly captured within the snapshot time cadence of ≈ 120 Myr (e.g. S. McAlpine et al. 2017). Among BHs with $M_{\text{BH,initial}} \sim 10^{6.5} M_{\odot}$, both aligned and counter-rotating galaxies have BH accretion rates of $\dot{M}_{\text{BH}} \sim 1 \times 10^{-4} M_{\odot} \text{ yr}^{-1}$, averaged over ≈ 0.5 Gyr. Galaxies experiencing an unstable misalignment within this mass range have enhanced subgrid BH accretion rates of $\dot{M}_{\text{BH}} \sim 4 \times 10^{-4} M_{\odot} \text{ yr}^{-1}$. Alternatively, we can estimate the average accretion rate required to grow the BH by ΔM_{BH} between each consecutive snapshot. Taking the radiative efficiency as $\epsilon_r = 0.1$, we find an average BH accretion rate of $\sim 7 \times 10^{-4} M_{\odot} \text{ yr}^{-1}$ among unstable misaligned galaxies with $M_{\text{BH,initial}} \sim 10^{6.5} M_{\odot}$.

Evaluating the peak value of $\dot{M}_{\text{BH}}$ attained over our ≈ 0.5 Gyr window for subsamples of aligned/unstable misaligned/counter-rotating galaxies, we find $\approx 46/50/43$ per cent of BHs reaching bolometric luminosities of $L_{\text{bol}} > 10^{43} \text{ erg s}^{-1}$, and $\approx 37/54/28$ per cent of BHs with Eddington ratios of $\lambda_{\text{Edd}} > 0.01$. These results are consistent with the observation that AGN signatures are more likely to be observed in misaligned systems (S. I. Raimundo et al. 2023, 2025).

By inspection of individual BH evolutions, we find that the BH accretion rate tends to be enhanced for the duration of the unstable misalignment rather than an initial single snapshot of high accretion. We also tested using the median values of the BH accretion rate, obtained from either the subgrid BH accretion values or ΔM_{BH} , instead of averages and found our results did not change. Furthermore, misaligned galaxies in our ≈ 0.75 and ≈ 1.0 Gyr samples also clearly show this enhanced BH growth.

5.1.2 Early- and late-type morphologies

In Fig. 3, we show the growth of the BH mass for samples of ETGs (top) and LTGs (bottom) analogous to Fig. 1. Galaxies are classified as either ETGs or LTGs using κ_{co}^* at each snapshot within the ~ 0.5 Gyr window. We only include galaxies for which this morphological classification remains constant, and thus exclude galaxies that undergo a morphological transformation. For ETGs (LTGs), this returns subsamples of 2159 (2885) galaxies, of which 285 (64) are experiencing an unstable misalignment and 374 (116) are in a steady-state counter-rotating configuration.

We find the degree of BH growth enhancement varies with galaxy morphology. This enhanced growth is stronger for LTGs (~ 1.0 dex) than for ETGs (~ 0.4 dex), at least at lower masses where both can be compared. At higher BH masses, our analysis is limited to ETGs given the low number statistics of unstable misaligned LTGs beyond

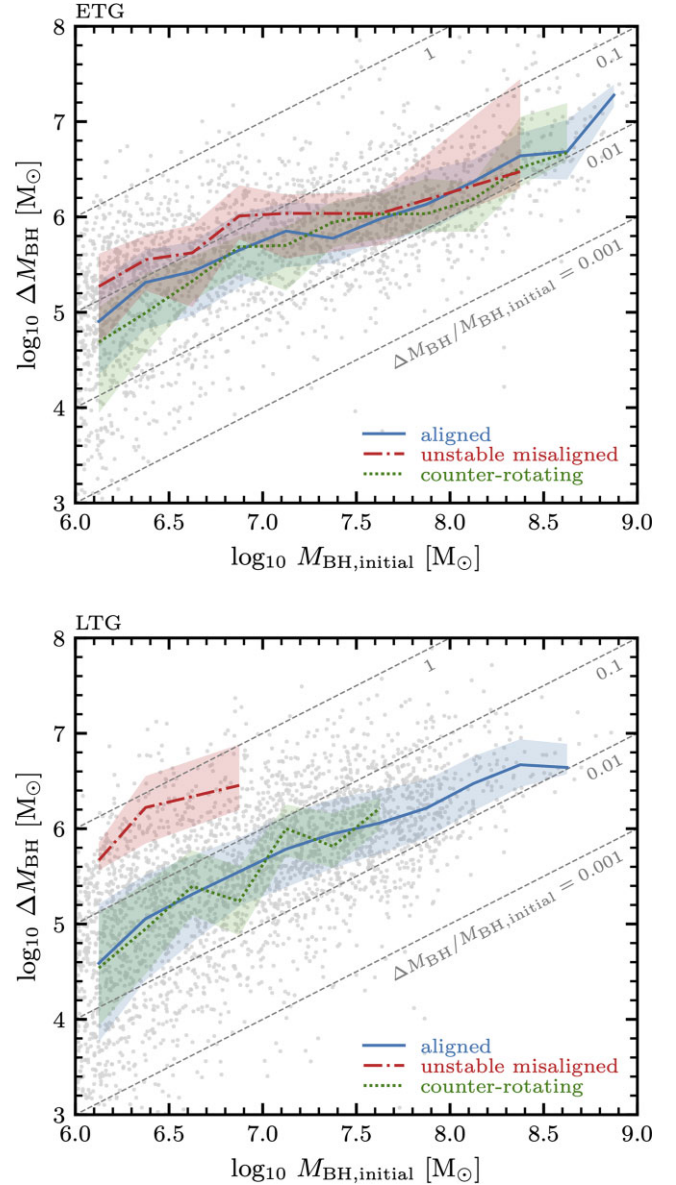


Figure 3. Median change in BH mass traced over a 0.50 ± 0.05 Gyr window as a function of the initial BH mass for samples of ETGs (top) and LTGs (bottom) which are aligned (blue solid), counter-rotating (green dotted), and galaxies experiencing an unstable misalignment (red dash-dotted) for bins with at least 10 galaxies. Shaded regions show the 25 and 75 percentiles. Grey scatter points show the total sample. Diagonal lines (dashed grey) show the relative increase in BH mass, with $\Delta M_{\text{BH}}/M_{\text{BH,initial}} = 1$ indicating a doubling of BH mass over this window. We find that BH growth is strongly enhanced (~ 1.0 dex) in unstable misaligned LTGs compared to aligned and counter-rotating morphological counterparts up to BH masses of $M_{\text{BH,initial}} \approx 10^{6.9} M_{\odot}$, beyond which we are limited by low number statistics. Compared to LTGs, the degree of enhanced BH growth among unstable misaligned ETGs is more modest at (~ 0.4 dex) up to BH masses of $M_{\text{BH,initial}} \lesssim 10^{7.3} M_{\odot}$.

$M_{\text{BH,initial}} \gtrsim 10^{6.9} M_{\odot}$. Consistent with results from Section 5.1.1, BH growth among unstable misaligned ETGs is comparable to their aligned and counter-rotating counterparts at higher BH masses.

Over the entire range of BH masses considered, unstable misaligned LTGs show by far the largest median BH growth (33.7 per cent), followed by unstable misaligned ETGs (10.6 per cent), and this result is significant (KS-test statistic = 0.41,

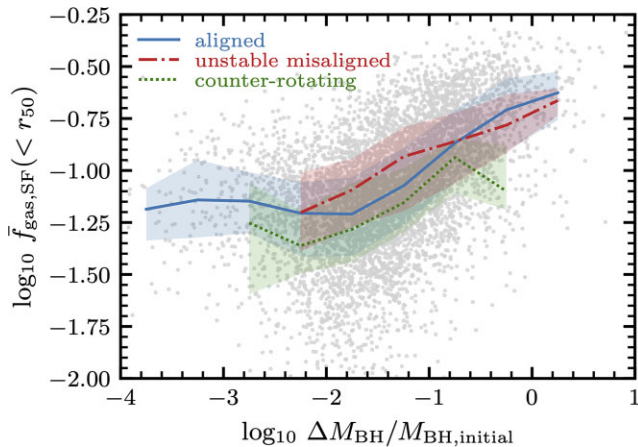


Figure 4. Median SF gas fraction within r_{50} as a function of fractional BH growth averaged over a 0.50 ± 0.05 Gyr window for subsamples of aligned (blue solid), counter-rotating (green dotted), and galaxies relaxing from an unstable misalignment (red dash-dotted). Shaded regions show the 25 and 75 percentiles. Grey scatter points show the total sample. Galaxies experiencing an unstable misalignment tend to be more gas rich overall, but have similar $f_{\text{gas,SF}}$ compared to aligned galaxies among the fastest growing BHs.

p -value = 1.5×10^{-8}). Aligned and counter-rotating galaxies show little dependence on morphology with similar BH growths of ~ 3.5 per cent. These results suggest misaligned gas is driven more effectively toward the centre of an LTG. This is expected given that a more disc-like stellar mass distribution will exert stronger torques on misaligned components (J. E. Tohline et al. 1982; G. Lake & C. Norman 1983). This results in stronger dissipative forces acting on the gas and shorter relaxation time-scales in LTGs compared to ETGs (for details, see M. K. Baker et al. 2025).

Given that unstable misalignments are found to be most effective in growing BHs in LTGs, it is worth considering the degree to which unstable misalignments may dominate the AGN population among discy galaxies at low redshifts. In Appendix A, we briefly explore the degree to which AGN are more likely to be found in misaligned systems using a sample of galaxies at $z = 0.1$. In summary, we find that the presence of an unstable misalignment increases the likelihood of fuelling AGN activity. This trend is significantly weaker for counter-rotation. Misalignments as a whole are only found in a minority (≈ 15) per cent of LTGs with AGN, whereas they are found over half of ETGs with AGN. Therefore, while unstable misalignments can drive efficient BH growth in LTGs in EAGLE, they are not the dominant driver of AGN in LTGs at $z = 0.1$.

5.1.3 Gas reservoirs and dynamics

In Fig. 4, we show the correlation between average gas_{SF} fraction, $f_{\text{gas,SF}}$, and fractional BH growth over ~ 0.5 Gyr windows. We find a broadly positive correlation between enhanced BH growth and galaxies that are more gas-rich.

On average, galaxies experiencing an unstable misalignment (shown in red) tend to be more gas-rich within r_{50} (median $\bar{f}_{\text{gas,SF}} = 0.123$) compared to aligned galaxies (median $\bar{f}_{\text{gas,SF}} = 0.084$). We note that this may be a selection effect as our sample does not include misalignments with $t_{\text{relax}} < 0.45$ Gyr which are common among more gas-poor galaxies and which dominate the overall relaxation time distributions (M. K. Baker et al. 2025). None the less, given that misalignments can be associated with gas replenishment (e.g. D. J.

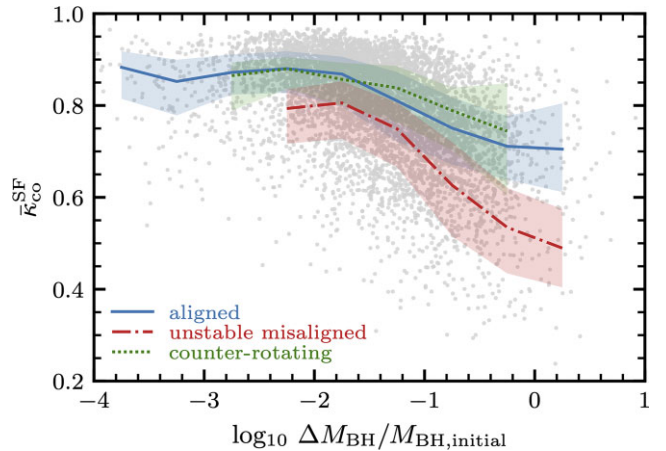


Figure 5. Median corotational energy fraction for SF gas within r_{50} as a function of fractional BH growth averaged over a 0.50 ± 0.05 Gyr window for subsamples of aligned (blue solid), counter-rotating (green dotted), and galaxies relaxing from an unstable misalignment (red dash-dotted). Shaded regions show the 25 and 75 percentiles. Grey scatter points show the total sample. Galaxies experiencing an unstable misalignment tend to have gas_{SF} discs with less rotational support compared to aligned and counter-rotating galaxies.

Khim et al. 2021; C. I. Casanueva et al. 2022; M. K. Baker et al. 2025), this result is not entirely surprising.

However, among galaxies with the largest BH growth, unstable misaligned galaxies have comparable gas_{SF} fractions. Likewise, among gas rich-galaxies ($\bar{f}_{\text{gas,SF}} > 0.1$), unstable misaligned galaxies continue to show boosted BH growth compared to aligned galaxies (not shown). This hints at an underlying difference in the behaviour of the gas in aligned versus relaxing galaxies in addition to the degree of gas richness of galaxies.

AGN fuelling requires both the presence of gas in the galaxy and a means to transport it to the inner regions of the galaxy (e.g. loss of angular momentum) in the vicinity of the BH (T. Storchi-Bergmann & A. Schnorr-Müller 2019). In Fig. 5, we show the correlation between the corotational energy fraction for gas_{SF} ($\kappa_{\text{co}}^{\text{SF}}$) and fractional BH growth, averaged over ~ 0.5 Gyr windows.

We find that BH growth tends to be enhanced in systems with a lower $\kappa_{\text{co}}^{\text{SF}}$, regardless of kinematic (mis-)alignment. This may be indicative of greater contributions from non-circular motions (i.e. radial inflows), or the increased presence of gas_{SF} particles in circular orbits that are non-coplanar. Aligned and counter-rotating galaxies show a noticeable decrease ($\Delta \kappa_{\text{co}}^{\text{SF}} \sim -0.2$) in the average gas_{SF} corotational energy fraction over the range of BH growth fractions considered. Among the fastest growing BHs, median $\bar{\kappa}_{\text{co}}^{\text{SF}}$ values of ≈ 0.70 indicate that, despite the loss of rotational-support, gas discs in aligned galaxies remain thin and discy.

As seen in Fig. 5, we find that the gas discs of galaxies experiencing an unstable misalignment show weaker (co)rotational support compared to aligned and counter-rotating galaxies with similar BH growth. Consequently, the median gas_{SF} corotational energy fraction for unstable misaligned galaxies is lower ($\bar{\kappa}_{\text{co}}^{\text{SF}} \approx 0.66$) compared to that of aligned ($\bar{\kappa}_{\text{co}}^{\text{SF}} \approx 0.84$) and counter-rotating galaxies ($\bar{\kappa}_{\text{co}}^{\text{SF}} \approx 0.82$). Unstable misaligned galaxies also show a larger decrease ($\Delta \kappa_{\text{co}}^{\text{SF}} \sim -0.3$) in the average corotational energy fraction over the range of BH growth within the sample compared to aligned galaxies. Morphologically, this approximately indicates a transition from thin gas discs at low BH growth fractions to thicker gas discs with more

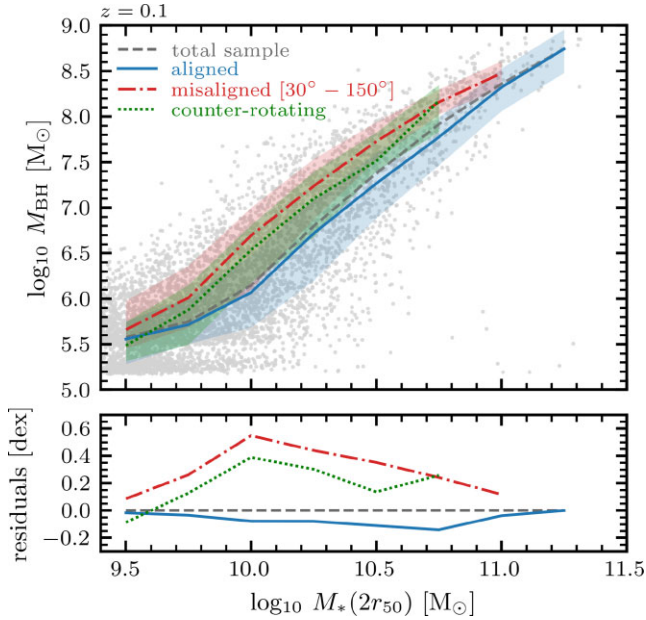


Figure 6. $M_* - M_{\text{BH}}$ relation at $z = 0.1$ for galaxies in various stellar-gas kinematic states. Top: median lines for the total sample (black dashed) and subsamples of aligned (blue solid), unstable misaligned (red dash-dotted), and counter-rotating (green dotted) galaxies for bins with a minimum bin count of 10. Shaded regions show the 25 and 75 percentiles for aligned, unstable misaligned, and counter-rotating galaxies. Grey scatter points show the total sample. Bottom: residuals of aligned, unstable misaligned, and counter-rotating medians, with respect to the total sample. We find overmassive BHs preferentially reside in galaxies that are currently in the unstable misaligned or counter-rotating regime, with a maximum BH mass difference of ~ 0.6 dex at a turnover mass of $\sim 10^{10} M_{\odot}$.

dispersion at the highest BH growth fractions (see also A. D. Hill et al. 2021).

The weaker (co)rotational support in gas discs among galaxies experiencing unstable misalignments is unsurprising. During the initial stages of misalignment formation, interactions between any *in-situ* corotating gas and stochastic accretion of misaligned gas naturally increase the dispersion within the gas disc while it settles into a more coherent disc. These interactions have been shown to be an efficient way for gas to dissipate angular momentum (e.g. L. V. Sales et al. 2012; P. Taylor et al. 2018; T. K. Starkenburg et al. 2019). This effect is strongest in the initial snippets of the ≈ 0.5 Gyr window and likely accounts for many of the systems with relatively low values of $\bar{\kappa}_{\text{co}}^{\text{SF}}$. However, once a misaligned gas disc forms and for the remainder of the unstable misalignment, decreased values of $\bar{\kappa}_{\text{co}}^{\text{SF}}$, in excess of aligned galaxy counterparts, are likely indicative of interactions between adjacent rings of non-coplanar gas (see e.g. F. van de Voort et al. 2015). Beyond the initial phase of misalignment formation, this is expected to become the dominant method of angular momentum dissipation which allows the gas disc to relax into the galactic plane. The lower values of $\bar{\kappa}_{\text{co}}^{\text{SF}}$ in galaxies experiencing an unstable misalignment suggests that misaligned gas is able to efficiently dissipate angular momentum. This causes the central gas density to increase (as evidenced by the relative absence of gas-poor unstable misaligned galaxies in our sample), resulting in higher subgrid BH accretion rates.

Although this is one interpretation, an alternative explanation for the weaker values of $\bar{\kappa}_{\text{co}}^{\text{SF}}$ among BHs with larger mass increases may come, at least in part, from the AGN feedback implementation in

EAGLE (for details see J. Schaye et al. 2015). The stochastic thermal heating of gas particles in the immediate vicinity of a fast-growing BH may increase the velocity dispersion of gas and subsequently decrease the gas_{SF} corotational energy fraction. However, this effect is likely limited to the innermost regions of the galaxy. Furthermore, we tend to see $\bar{\kappa}_{\text{co}}^{\text{SF}}$ increase alongside the increased BH growth following misalignment formation as the gas reforms a disc. This suggests that this effect does not dominate the low values of $\bar{\kappa}_{\text{co}}^{\text{SF}}$ seen in galaxies experiencing an unstable misalignment. We discuss this further in Section 6.

In contrast to the unstable misaligned sample, continuously counter-rotating galaxies tend to contain gas discs that are dynamically relaxed into thin discs ($\bar{\kappa}_{\text{co}}^{\text{SF}} \gtrsim 0.7$). In the absence of external factors, gas particles in retrograde orbits must therefore rely on subgrid stellar wind modelling to induce change in the angular momentum. Stellar-gas counter-rotation in EAGLE tends to be long-lived, forming a significant fraction of the galaxy population by $z = 0.1$ (for details, see M. K. Baker et al. 2025). While the fraction of counter-rotating galaxies is comparable to observations, the stellar mass-loss model in EAGLE is unable to accurately capture the long-term angular momentum dissipation of a counter-rotating gas disc imbedded in a corotating stellar disc. This is likely a result of the limited resolution of large-scale cosmological simulations which results in similar growth of BHs in aligned and counter-rotating galaxies (see Fig. 1).

5.2 BH growth history

5.2.1 Overmassive BHs at $z = 0.1$

It is worth investigating whether the enhanced BH growth found in galaxies experiencing unstable misalignments is imprinted on the galaxy population at $z = 0.1$. This redshift was chosen arbitrarily and lies within the redshift limit probed by recent IFU surveys such as SAMI and Mapping Nearby Galaxies at Apache Point Observatory (MaNGA; K. Bundy et al. 2015). Given the relatively low numbers of misaligned galaxies that coincide with this redshift in our pre-existing sample (due to the relaxation requirement, see M. K. Baker et al. 2025), we extract a new sample of 5372 galaxies at $z = 0.1$ matching the criteria in Section 3 with no minimum BH mass requirement. These galaxies are classified according to their current kinematic state at $z = 0.1$, of which 4166 are aligned ($\psi_{3\text{D}} < 30^\circ$), 308 are counter-rotating ($\psi_{3\text{D}} > 150^\circ$), and 898 are in the unstable misaligned regime ($30^\circ < \psi_{3\text{D}} < 150^\circ$).

In Fig. 6 (top), we show the $M_* - M_{\text{BH}}$ relationship of this sample alongside medians for each subsample of kinematic state. Stellar masses are measured within $2r_{50}$. Residuals of these medians are shown with respect to the total sample (bottom). We find overmassive BHs preferentially residing in galaxies that are currently misaligned and, to a lesser degree, counter-rotating. This trend is seen most strongly for intermediate-mass galaxies with $10^{9.7} \lesssim M_*/M_{\odot} \lesssim 10^{10.5}$. For instance, misaligned and aligned galaxies show a maximum difference of ~ 0.6 dex between BH masses at a turnover mass of $M_* \sim 10^{10} M_{\odot}$. Comparing the distributions of M_{BH}/M_* for aligned and misaligned galaxies, we find the differences to be significant (KS-test statistic = 0.21, p -value $\ll 10^{-10}$). Similarly, for aligned and counter-rotating galaxies (KS-test statistic = 0.27, p -value $\ll 10^{-10}$). As before, this enhancement is seen more strongly among LTGs than in ETGs. These results remain unchanged if we crudely mimic observational measurements by using projected angles in 2D space instead of 3D angles.

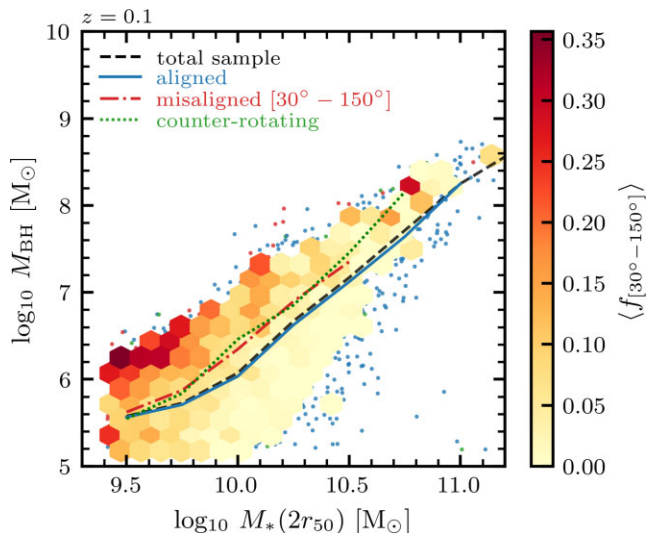


Figure 7. M_* – M_{BH} relation at $z = 0.1$ for a representative sample of galaxies with reliable kinematic measurements over $0.10 \lesssim z \lesssim 0.28$. Bins coloured according to the fractional time a galaxy spends in the unstable misaligned regime ($f_{[30^\circ-150^\circ]}$) within this redshift range, averaged for all galaxies within a given bin (minimum count of 5). Medians are shown for the total sample (black dashed) and subsamples of aligned (blue solid), unstable misaligned (red dash-dotted), and counter-rotating (green dotted) galaxies for bins with a minimum bin count of 10. We find a strong correlation between galaxies hosting overmassive BHs and galaxies with a history of spending more time in the unstable misaligned regime.

Overmassive BHs are also found in counter-rotating systems at $z = 0.1$, but with a weaker deviation from the total population. Counter-rotating galaxies are commonly formed from the relaxation of past unstable misalignments (e.g. M. K. Baker et al. 2025), assuming the counter-rotating gas disc did not form directly from retrograde gas accretion onto a gas-poor galaxy. As such, many of these present-day counter-rotating systems may have experienced a past unstable misalignment and corresponding phase of enhanced BH accretion. While their current BH growth may be indistinguishable from the aligned galaxy population, counter-rotating galaxies in EAGLE encapsulate their past misalignment history in the form of overmassive BHs.

5.2.2 Galaxy kinematic history at $z = 0.1$

We can also investigate the degree to which the mass of the BH traces the history of past unstable misalignments. Beginning with our sample of 5372 galaxies at $z = 0.1$, we select a subset of 3931 galaxies that follow the selection criteria of M. K. Baker et al. (2025) for reliable kinematic measurements over the preceding 2 Gyr ($0.10 \lesssim z \lesssim 0.28$). Of this subset, 3367 are aligned, 200 are counter-rotating, and 364 are in the unstable misaligned regime. With the exception of low-mass galaxies with $M_* \lesssim 10^{9.7} M_\odot$, this creates a sample of galaxies with stellar masses representative of those at $z = 0.1$.

In Fig. 7, we show the relationship between galaxies with overmassive BHs and their stellar-gas kinematic histories over the past 2 Gyr. We denote the fraction of time a galaxy has spent in the unstable misaligned regime over the window $0.10 \lesssim z \lesssim 0.28$ as $f_{[30^\circ-150^\circ]}$. We find a clear trend between galaxies hosting overmassive BHs and the average time spent in the unstable misaligned regime. This fraction tends to increase with distance offset from the median

$M_* - M_{\text{BH}}$ relationship. Correspondingly, galaxies with undermassive BHs rarely show any sign of past unstable misalignments within $0.10 \lesssim z \lesssim 0.28$. These trends are seen more strongly in galaxies with $M_* \lesssim 10^{10.3} M_\odot$.

We also find a trend between overmassive BHs and a rich history of past counter-rotation (not shown), although this trend is weaker compared to that of unstable misalignments. Similarly for the misaligned result in Fig. 7, this is seen most strongly for intermediate-mass galaxies with $M_* \sim 10^{10} M_\odot$. As explained above, this result is unsurprising given stellar-gas counter-rotation can be the result of a past relaxation. We find that both trends are also seen for galaxy histories extending over the past 4 Gyr ($0.10 \lesssim z \lesssim 0.51$) and 6 Gyr ($0.10 \lesssim z \lesssim 0.85$), although these are not shown. These samples tend to exclude an increasing number of low-mass galaxies as lookback time increases, and thus we caution that these longer time-scale results are less certain due to lower number statistics.

Furthermore, by only selecting the subset of galaxies that are aligned at $z = 0.1$, we continue to find that overmassive BHs trace galaxies with a history of more unstable misalignments (not shown). Indeed, we find the past kinematic (mis)alignment history of a galaxy is a significantly stronger indicator of the relative mass of present-day BHs than compared to gas-richness ($f_{\text{gas,SF}} > 0.1$) history. Therefore, at least in EAGLE, the relative deviation of a BH from the $M_* - M_{\text{BH}}$ relationship at $z = 0.1$ can provide clues about the past kinematic history of gas within the galaxy, regardless of the current state of stellar-gas (mis)alignment.

Observational studies of local galaxies have shown that overmassive BHs tend to preferentially reside in ETGs (e.g. A. W. Graham & N. Sahu 2023). This has been attributed to the cumulative effects of mergers acting as a source of both cold gas to fuel BH growth and disruption of pre-existing cold gas reservoirs during the assembly of these galaxies. Our findings suggest that the presence and persistence of an unstable kinematic misalignment is a complementary channel for galaxies to acquire an overmassive BH without the prerequisite of a galaxy merger. Furthermore, the presence of an undermassive BH is a strong indicator of past stellar-gas alignment. Given both mergers and stellar-gas kinematic misalignments preferentially occur among ETGs compared to LTGs in EAGLE (e.g. C. I. Casanueva et al. 2022; M. K. Baker et al. 2025), this further complements the observed relationship between galaxy morphology and location on the $M_* - M_{\text{BH}}$ plane. We note that observational measurements of single-epoch BH masses using the broad H β line typically have mass uncertainties of 0.3 – 0.4 dex (e.g. E. Dalla Bontà et al. 2020). As such, the difference of ~ 0.6 dex we find between intermediate-mass misaligned and aligned galaxies is likely to be detected in observations.

6 DISCUSSION

Overall, our results are consistent with the picture that the persistence of an unstable stellar-gas kinematic misalignment is an effective mechanism to drive gas inwards for increased BH activity. This is in good agreement with results from other cosmological simulations (e.g. T. K. Starkenburg et al. 2019; C. Duckworth et al. 2020a, b; S. Khoperskov et al. 2021) and recent observational findings (e.g. I. Ruffa et al. 2019; S. I. Raimundo et al. 2017, 2023; S. I. Raimundo 2021).

In previous work, C. Duckworth et al. (2020b) used the ILLUSTRISTNG-100 simulation and a sample of ~ 2500 galaxies (with masses of $10^{8.5} < M_*/M_\odot < 10^{11.5}$) to investigate AGN activity in misaligned galaxies. This was done through the construction of mock observational images to mimic results from MaNGA, with 2D

position angles (PAs) used to classify galaxies as misaligned ($\Delta\text{PA} \geq 30^\circ$). In agreement with our results, low-mass ($M_* < 10^{10.2} M_\odot$) misaligned galaxies show boosted AGN luminosity and BH growth at $z \sim 0.1$ with results becoming less conclusive in higher-mass galaxies. Similarly, S. Khoperskov et al. (2021) find that non-aligned gas is efficiently funnelled inward in ILLUSTRISTNG through interactions with aligned gas. We note that ILLUSTRISTNG uses a two-mode AGN feedback model, with $M_{\text{BH}} \sim 10^8 M_\odot$ corresponding to the typical transition between thermal and a more violent kinetic-mode feedback. The implementation of such a feedback mode can have a strong effect on the gas content of galaxies, especially at the more massive end (e.g. R. Davé et al. 2020). As such, while our results may be sensitive to the BH feedback model used, it is reassuring that we find good qualitative agreement between the two simulations within the range of galaxy masses considered for this paper. A detailed comparison between massive galaxies (with correspondingly massive BHs), in which the feedback regime may have a stronger impact on our results, is strongly limited by our small sample of $M_{\text{BH}} \sim 10^8 M_\odot$ BHs in misaligned galaxies.

Counter-rotating systems in ILLUSTRISTNG have also been found with boosted AGN luminosities and BH growth (S. Khoperskov et al. 2021; C. Duckworth et al. 2020b). While we do find overmassive BHs in counter-rotating galaxies at $z = 0.1$ in EAGLE, we do not find significantly enhanced BH growth or luminosity in these systems. Instead, these overmassive BHs trace prior unstable misalignments. This discrepancy may be due to differences in the subgrid models and underlying hydrodynamic scheme of these simulations. However, our results are supported by recent observational results that found no evidence of enhanced AGN luminosities among AGN host galaxies with larger misalignment angles (M. W. Winiarska et al. 2025).

Misaligned gas has also been associated with lower angular momentum in simulations (e.g. F. van de Voort et al. 2015; T. K. Starkenburg et al. 2019; C. Duckworth et al. 2020a) and recent observational studies (e.g. H. Xu et al. 2022). Specifically for EAGLE, C. I. Casanueva et al. (2022) find misaligned galaxies have SF discs that are more compact than aligned counterparts. This hints at the efficiency of misaligned gas to lose angular momentum as it relaxes into a dynamically stable configuration, driving gas inward. Again, our results are in good qualitative agreement with this interpretation.

Galaxy mergers have long been established as an effective means for gas to dissipate angular momentum and potentially fuel AGN activity (see T. Storchi-Bergmann & A. Schnorr-Müller 2019, and references therein). For instance, mergers have been found as a key driver of overmassive BHs in EAGLE through the disruption of pre-existing corotating gas (e.g. J. J. Davies et al. 2022; J. J. Davies, A. Pontzen & R. A. Crain 2024). Likewise, mergers have been found to increase the incidence of luminous AGN in EAGLE (S. McAlpine et al. 2020). As such, it is important to consider any bias associated with our sample of relaxing unstable misalignments as these may be associated with significantly more mergers.

For this, we compare the incidence of mergers with stellar mass ratios of $> 1/10$ within our ~ 0.5 Gyr windows between our samples of aligned, unstable misaligned, and counter-rotating galaxies. We find a higher incidence of mergers in our unstable misaligned sample (≈ 9 per cent) compared to our aligned (≈ 4 per cent) and counter-rotating samples (≈ 3 per cent). This is consistent with previous results in cosmological simulations that found mergers to be subdominant in the formation of misalignments (e.g. D. J. Khim et al. 2021; M. K. Baker et al. 2025). Our results remain unchanged if we remove all unstable misaligned galaxies that experience a merger within our ~ 0.5 Gyr windows from our sample. As BH growth has been shown to increase significantly within ± 0.12 Gyr

of coalescence (see S. McAlpine et al. 2020), we reason that the lack of change in our results emphasizes the dominance of misalignment-driven growth over that of in-plane gas inflow driven by galaxy mergers in our sample. We note that this does not preclude a higher incidence of galaxy interactions within the unstable misaligned sample. While a higher incidence of interactions have been shown to increase BH luminosities, the increase is significantly weaker than for mergers (S. McAlpine et al. 2020). Therefore, we reason that a ~ 5 per cent increase in the incidence of mergers between aligned and relaxing unstable misaligned galaxies is unable to account for the systematically higher BH growth found in Fig. 1. We conclude that the boosted BH growth seen in our results is likely not due to a higher incidence of mergers in misaligned galaxies (see also S. I. Raimundo et al. 2025).

As explored in S. McAlpine et al. (2020), despite galaxy mergers increasing the incidence of luminous AGN, merger-driven BH growth was found to contribute only ~ 15 per cent toward present-day BH mass in EAGLE. Similar results have also been found in the HORIZON-AGN (Y. Dubois et al. 2014) simulation (see G. Martin et al. 2018). Therefore, it is worth considering how misalignment-driven growth may compare. Compared to a control sample, S. McAlpine et al. (2020) report a 2 – 3 times increase in BH accretion rates in merging systems. For comparison, we report an increase of ~ 4 times (~ 0.6 dex) among unstable misaligned galaxies compared to aligned counterparts. In both cases, BH growth was found to be increased predominantly in less massive BHs ($M_{\text{BH}} \lesssim 10^7 M_\odot$). Unstable misalignments in EAGLE were found with median lifetimes of ~ 0.5 Gyr, but with a large standard deviation of ~ 0.6 Gyr (see M. K. Baker et al. 2025). Furthermore, misalignments have been found with diverse formation pathways in which mergers are subdominant (e.g. D. J. Khim et al. 2021; C. I. Casanueva et al. 2022; M. K. Baker et al. 2025). As such, misalignments not only provide more BH growth than mergers with the ability to spend considerable ($\gtrsim 0.5$ Gyr) time with these features, but can also occur more frequently and independently of galaxy mergers. While quantifying in detail the contribution of misalignments toward present-day BH mass is beyond the scope of this paper, we suggest that misalignments may contribute more toward present-day BH mass in EAGLE than mergers alone.

AGN feedback is likely to play a role in reducing the degree of BH growth in more massive misaligned systems. As summarized in S. McAlpine et al. (2018), lower mass BHs ($M_{\text{BH}} \lesssim 10^7 M_\odot$) tend to grow rapidly as AGN feedback is unable to self-regulate gas inflows, leading to a rise in central gas density (see also R. G. Bower et al. 2017). In more massive BHs ($M_{\text{BH}} \gtrsim 10^7 M_\odot$), AGN feedback is able to self-regulate inflows and growth. As such, we speculate that the increased efficiency of AGN feedback for higher mass BHs is likely reducing the efficiency of unstable misalignments to enhance BH growth over the time periods considered in this work. Consequently, this is imprinted on the $M_* - M_{\text{BH}}$ relationship at $z = 0.1$ which results in a weaker BH mass offset between misaligned and aligned galaxies for higher mass galaxies (see Fig. 6).

Finally, while the increased AGN feedback in misaligned systems may be contributing to the lower gas corotational energy fraction, we do not believe it is the key driver of the unstable misalignments over ~ 0.5 Gyr periods. This is because the angular momentum of the outflows would need to dominate over the angular momentum of the remaining gas disc. This would imply that a large amount of gas_{SF} would be carried away within r_{50} , starving the BH of long-term fuel. Yet, we continue to clearly identify enhanced BH growth in misaligned galaxies for the largest time window considered (≈ 1.0 Gyr). Secondly, the thermal BH feedback implementation in

EAGLE means outflows naturally follow the path of least resistance perpendicular to the gas disc (T. Hartwig, M. Volonteri & G. Dashyan 2018; P. D. Mitchell et al. 2020a). Thus, it may be difficult for AGN-driven outflows to misalign the gas disc. This is supported by the finding that wind recycling rates in EAGLE tend to be low and that outflows at galactic scales tend to be ejective out to high radii (P. D. Mitchell et al. 2020a; P. D. Mitchell, J. Schaye & R. G. Bower 2020b). In contrast, outflows in ILLUSTRISTNG galaxies are more violent and tend to follow ballistic orbits (P. D. Mitchell et al. 2020a) which manifest as galactic fountains (P. R. Shapiro & Field 1976). While this makes AGN-driven misalignments more feasible in ILLUSTRISTNG (see also C. Duckworth et al. 2020b), we argue that this process is likely rare in EAGLE. However, we do not rule out the possibility that significant AGN feedback preceded the misalignment, decreasing the gas fraction and easing the formation of a misalignment (as proposed by e.g. C. I. Casanueva et al. 2022; E. Cenci et al. 2024). Or, alternatively, that ongoing AGN feedback is disrupting the axis of gas accretion onto the misaligned gas disc as suggested by previous authors (e.g. T. K. Starkenburg et al. 2019; D. J. Khim et al. 2021; C. I. Casanueva et al. 2022).

7 CONCLUSIONS

In this work, we have used the EAGLE cosmological simulations to investigate the growth of the central BH in galaxies with stellar-gas kinematic misalignments between $0 < z < 1$. We used a sample of 5570 galaxies with masses of $> 10^{9.5} M_{\odot}$ and reliable kinematic and BH mass measurements. We traced the change in BH mass over ≈ 0.5 Gyr windows in galaxies experiencing an unstable kinematic misalignment ($30^{\circ} \leq \psi_{3D} \leq 150^{\circ}$) and compared these to a sample of steady-state aligned ($\psi_{3D} < 30^{\circ}$) and counter-rotating ($\psi_{3D} > 150^{\circ}$) galaxies. Using this sample, we established correlations between the amount and kinematics of SF gas and the degree of BH growth. Finally, we used a sample of 3931 galaxies at $z = 0.1$ with reliable kinematic measurements over the previous 2 Gyr to investigate the relationship between overmassive BHs and the stellar-gas kinematic history of galaxies. Our results can be summarized as follows:

- (i) BHs residing in galaxies experiencing an unstable misalignment experience significantly enhanced (~ 0.6 dex) growth and accretion rates, growing their BH masses by ≈ 13 per cent over ≈ 0.5 Gyr windows. In comparison, aligned ($< 30^{\circ}$) and counter-rotating ($> 150^{\circ}$) galaxies grow their BHs by ≈ 3.9 per cent and ≈ 3.4 per cent over the same time period, respectively. This boosted BH growth (~ 0.6 dex) is seen most strongly for BHs of mass $\lesssim 10^{7.2} M_{\odot}$, weakening at higher BH masses (see Figs 1 and 2).
- (ii) The degree of enhanced BH growth depends on morphology only in galaxies experiencing an unstable misalignment. Unstable misaligned LTGs provide a greater enhancement (~ 1.0 dex) to their BH growth than unstable misaligned ETGs (~ 0.4 dex) over ≈ 0.5 Gyr windows. Consequently, unstable misaligned LTGs show the greatest BH mass increase (≈ 34 per cent), followed by unstable misaligned ETGs (≈ 11 per cent), with the remaining aligned/counter-rotating LTGs and ETGs experiencing typical BH mass increases of ~ 3.5 per cent over ≈ 0.5 Gyr windows. We attribute this to the greater torques experienced by the misaligned gas from more discy stellar distributions, resulting in stronger angular momentum dissipation and inflow of gas (see Fig. 3).
- (iii) While our sample of unstable misaligned galaxies tend to be more gas-rich than aligned or counter-rotating galaxies, we find

only a weak correlation between BH growth and SF gas fraction irrespective of kinematic (mis)alignment (see Fig. 4).

(iv) SF gas in galaxies experiencing an unstable misalignment has less rotational support. This is likely indicative of the angular momentum dissipation experienced by misaligned gas, driving gas inward to be accreted by the central BH as the disc relaxes into the galactic plane (see Fig. 5).

(v) At $z = 0.1$, BH masses are larger (~ 0.6 dex) among intermediate-mass galaxies with unstable misalignments compared to aligned counterparts (see Fig. 6). Likewise, the population of overmassive BHs at $z = 0.1$ is dominated by galaxies that have spent more time in the unstable misaligned regime over the past 2 Gyr ($0.10 \lesssim z \lesssim 0.28$, see Fig. 7).

(vi) At $z = 0.1$, counter-rotating galaxies tend to host overmassive BHs. This is despite finding similar BH growth between aligned and counter-rotating galaxies. We suggest that the BHs of counter-rotating galaxies were affected by their past unstable misalignment and subsequent enhanced growth, followed by a relaxation into the counter-rotating regime (see Fig. 6).

(vii) At $z = 0.1$, galaxies with AGN show a significantly higher fraction (~ 38 per cent) of misaligned ($\psi_{3D} > 30^{\circ}$) galaxies than the total galaxy population (~ 23 per cent). Among galaxies with AGN, unstable misalignments ($30^{\circ} \leq \psi_{3D} \leq 150^{\circ}$) are found in about half of ETGs but in only ~ 9 per cent of LTGs. Despite the efficient BH growth experienced by LTGs with unstable misalignments, we suggest that the relative rarity of misalignments in LTGs results in other processes dominating AGN activity in these systems (see Appendix A).

We conclude that, at least in EAGLE, the presence of an unstable stellar-gas kinematic misalignment is an effective means to drive gas inward to boost the growth of the central BH and trigger AGN activity. Thus, at least statistically, the relative deviation from the $M_{*} - M_{BH}$ relation could be used to infer a galaxy's past (mis)alignment history.

It is clear that more work is needed to fully understand the connection between AGN activity and the presence of stellar-gas kinematic misalignments. For instance, the degree to which misalignments contribute toward non-merger BH growth over a galaxy's lifetime is an exciting question for future studies. Future observational studies with, e.g. the Atacama Large Millimeter/submillimeter Array (ALMA), Multi Unit Spectroscopic Explorer (MUSE), and *James Webb Space Telescope* alongside simulations that resolve the multiphase ISM will allow for more thorough investigations of this subject.

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DATA AVAILABILITY

The EAGLE simulations and database are publicly available at: <http://icc.dur.ac.uk/Eagle/database.php>. Detailed guides to simulation parameters, models, and how to access and query data can be found in The EAGLE team (2017) and S. McAlpine et al. (2016). EAGLE ‘snapshots’ are available on request from members of the EAGLE team.

REFERENCES

- Baker M. K., Davis T. A., van de Voort F., Ruffa I., 2025, *MNRAS*, 541, 494
- Booth C. M., Schaye J., 2009, *MNRAS*, 398, 53
- Bower R. G., Schaye J., Frenk C. S., Theuns T., Schaller M., Crain R. A., McAlpine S., 2017, *MNRAS*, 465, 32
- Bryant J. J. et al., 2015, *MNRAS*, 447, 2857
- Bryant J. J. et al., 2019, *MNRAS*, 483, 458
- Bundy K. et al., 2015, *ApJ*, 798, 7
- Capelo P. R., Dotti M., 2017, *MNRAS*, 465, 2643
- Casanueva C. I., Lagos C. d. P., Padilla N. D., Davison T. A., 2022, *MNRAS*, 514, 2031
- Cenci E., Feldmann R., Gensior J., Bullock J. S., Moreno J., Bassini L., Bernardini M., 2024, *ApJ*, 961, L40
- Chen Y.-M. et al., 2016, *Nat. Commun.*, 7, 13269
- Correa C. A., Schaye J., Clauwens B., Bower R. G., Crain R. A., Schaller M., Theuns T., Thob A. C. R., 2017, *MNRAS*, 472, L45
- Crain R. A. et al., 2015, *MNRAS*, 450, 1937
- Crain R. A. et al., 2017, *MNRAS*, 464, 4204
- Croom S. M. et al., 2012, *MNRAS*, 421, 872
- Dalla Bontà E. et al., 2020, *ApJ*, 903, 112
- Dalla Vecchia C., Schaye J., 2012, *MNRAS*, 426, 140
- Davé R., Crain R. A., Stevens A. R. H., Narayanan D., Saintonge A., Catinella B., Cortese L., 2020, *MNRAS*, 497, 146
- Davies J. J., Pontzen A., Crain R. A., 2022, *MNRAS*, 515, 1430
- Davies J. J., Pontzen A., Crain R. A., 2024, *MNRAS*, 527, 4705
- Davis T. A. et al., 2011, *MNRAS*, 417, 882
- Davis T. A., Bureau M., 2016, *MNRAS*, 457, 272
- Davis T. A., Young L. M., 2019, *MNRAS*, 489, L108
- Dubois Y. et al., 2014, *MNRAS*, 444, 1453
- Duckworth C., Starkenburg T. K., Genel S., Davis T. A., Habouzit M., Kraljic K., Tojeiro R., 2020b, *MNRAS*, 495, 4542
- Duckworth C., Tojeiro R., Kraljic K., 2020a, *MNRAS*, 492, 1869
- Ellison S. L. et al., 2025, *Open J. Astrophys.*, 8, 12
- Ferrarese L., Merritt D., 2000, *ApJ*, 539, L9
- Graham A. W., Sahu N., 2023, *MNRAS*, 518, 2177
- Habouzit M. et al., 2022, *MNRAS*, 509, 3015
- Han S. et al., 2024, *ApJ*, 977, 116
- Häring N., Rix H.-W., 2004, *ApJ*, 604, L89
- Hartwig T., Volonteri M., Dashyan G., 2018, *MNRAS*, 476, 2288
- Heckman T. M., Best P. N., 2014, *ARA&A*, 52, 589
- Hill A. D., Crain R. A., Kwan J., McCarthy I. G., 2021, *MNRAS*, 505, 65
- Ilha G. S. et al., 2019, *MNRAS*, 484, 252
- Jiang L., Helly J. C., Cole S., Frenk C. S., 2014, *MNRAS*, 440, 2115
- Jiménez E., Lagos C. d. P., Ludlow A. D., Wisnioski E., 2023, *MNRAS*, 524, 4346
- Kaviraj S., 2014, *MNRAS*, 437, L41
- Kereš D., Katz N., Weinberg D. H., Davé R., 2005, *MNRAS*, 363, 2
- Khim D. J., Yi S. K., Pichon C., Dubois Y., Devriendt J., Choi H., Bryant J. J., Croom S. M., 2021, *ApJS*, 254, 27
- Khoperskov S. et al., 2021, *MNRAS*, 500, 3870
- Kormendy J., Ho L. C., 2013, *ARA&A*, 51, 511
- Lagos C. d. P., Davis T. A., Lacey C. G., Zwaan M. A., Baugh C. M., Gonzalez-Perez V., Padilla N. D., 2014, *MNRAS*, 443, 1002
- Lagos C. d. P., Padilla N. D., Davis T. A., Lacey C. G., Baugh C. M., Gonzalez-Perez V., Zwaan M. A., Contreras S., 2015, *MNRAS*, 448, 1271
- Lake G., Norman C., 1983, *ApJ*, 270, 51
- Leitner S. N., Kravtsov A. V., 2011, *ApJ*, 734, 48
- Marconi A., Hunt L. K., 2003, *ApJ*, 589, L21
- Martin G. et al., 2018, *MNRAS*, 476, 2801
- McAlpine S. et al., 2016, *Astron. Comput.*, 15, 72
- McAlpine S., Bower R. G., Harrison C. M., Crain R. A., Schaller M., Schaye J., Theuns T., 2017, *MNRAS*, 468, 3395
- McAlpine S., Bower R. G., Rosario D. J., Crain R. A., Schaye J., Theuns T., 2018, *MNRAS*, 481, 3118
- McAlpine S., Harrison C. M., Rosario D. J., Alexander D. M., Ellison S. L., Johansson P. H., Patton D. R., 2020, *MNRAS*, 494, 5713
- Mitchell P. D., Schaye J., Bower R. G., 2020b, *MNRAS*, 497, 4495
- Mitchell P. D., Schaye J., Bower R. G., Crain R. A., 2020a, *MNRAS*, 494, 3971
- Negri A., Posacki S., Pellegrini S., Ciotti L., 2014, *MNRAS*, 445, 1351
- Nelson D. et al., 2019, *Comput. Astrophys. Cosmol.*, 6, 2
- Parriott J. R., Bregman J. N., 2008, *ApJ*, 681, 1215
- Peirani S., Suto Y., Han S., Yi S. K., Dubois Y., Kraljic K., Park M., Pichon C., 2025, *A&A*, 696, A45
- Penny S. J. et al., 2018, *MNRAS*, 476, 979
- Planck Collaboration I., 2014, *A&A*, 571, A1
- Qu Y. et al., 2017, *MNRAS*, 464, 1659
- Raimundo S. I. et al., 2025, *MNRAS*, 542, L103
- Raimundo S. I., 2021, *A&A*, 650, A34
- Raimundo S. I., Davies R. I., Canning R. E. A., Celotti A., Fabian A. C., Gandhi P., 2017, *MNRAS*, 464, 4227
- Raimundo S. I., Davies R. I., Gandhi P., Fabian A. C., Canning R. E. A., Ivanov V. D., 2013, *MNRAS*, 431, 2294
- Raimundo S. I., Malkan M., Vestergaard M., 2023, *Nat. Astron.*, 7, 463
- Riste A. et al., 2022, *MNRAS*, 517, 2677
- Rosas-Guevara Y. M. et al., 2015, *MNRAS*, 454, 1038
- Ruffa I. et al., 2019, *MNRAS*, 489, 3739
- Sales L. V., Navarro J. F., Theuns T., Schaye J., White S. D. M., Frenk C. S., Crain R. A., Dalla Vecchia C., 2012, *MNRAS*, 423, 1544
- Sarzi M. et al., 2006, *MNRAS*, 366, 1151
- Schaller M., Dalla Vecchia C., Schaye J., Bower R. G., Theuns T., Crain R. A., Furlong M., McCarthy I. G., 2015, *MNRAS*, 454, 2277
- Schawinski K., Koss M., Berney S., Sartori L. F., 2015, *MNRAS*, 451, 2517
- Schaye J. et al., 2015, *MNRAS*, 446, 521
- Schaye J., Dalla Vecchia C., 2008, *MNRAS*, 383, 1210
- Shapiro P. R., Field G. B., 1976, *ApJ*, 205, 762
- Shlosman I., Begelman M. C., Frank J., 1990, *Nature*, 345, 679
- Silk J., Rees M. J., 1998, *A&A*, 331, L1
- Springel V. et al., 2008, *MNRAS*, 391, 1685
- Springel V., Di Matteo T., Hernquist L., 2005, *MNRAS*, 361, 776
- Starkenburg T. K., Sales L. V., Genel S., Manzano-King C., Canalizo G., Hernquist L., 2019, *ApJ*, 878, 143
- Storchi-Bergmann T., Schnorr-Müller A., 2019, *Nat. Astron.*, 3, 48
- Taylor P., Federrath C., Kobayashi C., 2018, *MNRAS*, 479, 141
- Thakar A. R., Ryden B. S., 1996, *ApJ*, 461, 55
- The EAGLE team, 2017, preprint (arXiv:1706.09899)
- Tohline J. E., Simonson G. F., Caldwell N., 1982, *ApJ*, 252, 92
- Trayford J. W., Schaye J., 2019, *MNRAS*, 485, 5715
- van de Voort F., Davis T. A., Kereš D., Quataert E., Faucher-Giguère C.-A., Hopkins P. F., 2015, *MNRAS*, 451, 3269
- Wiersma R. P. C., Schaye J., Smith B. D., 2009a, *MNRAS*, 393, 99
- Wiersma R. P. C., Schaye J., Theuns T., Dalla Vecchia C., Tornatore L., 2009b, *MNRAS*, 399, 574
- Winiarska M. W., Raimundo S. I., Davis T. A., Riffel R., Shankar F., Wiseman P., 2025, *MNRAS*, 538, 1191
- Xu H. et al., 2022, *MNRAS*, 511, 4685

APPENDIX A: AGN AMONG MISALIGNED GALAXIES AT $z = 0.1$

As shown in Sections 5.1.1 and 5.1.2, galaxies with unstable stellar-gas kinematic misalignments ($30^\circ < \psi_{3D} < 150^\circ$) show enhanced BH growth. It is therefore interesting to consider whether we see more AGN signatures at $z = 0.1$ in galaxies that are misaligned.

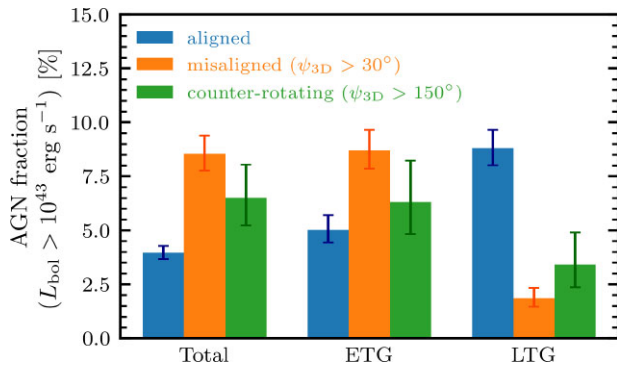


Figure A1. Fraction of AGN (defined as BH luminosities of $L_{\text{bol}} > 10^{43} \text{ erg s}^{-1}$) among the total galaxy sample at $z = 0.1$ among populations of aligned (blue, first bar of each group), misaligned (orange, second bar of each group), and counter-rotating galaxies (green, third bar of each group). Additionally, we show the fraction of AGN among subsamples of aligned/misaligned/counter-rotating ETGs and LTGs. The error bars indicate the 1σ (68 per cent) binomial confidence intervals estimated using the Jeffreys prior. We find higher AGN fractions among misaligned galaxies for the total galaxy population and for ETGs. Conversely, AGN fractions are lowest among misaligned LTGs.

We consider two approaches; first, the degree to which AGN fractions are increased among aligned ($\psi_{3D} < 30^\circ$), misaligned ($\psi_{3D} > 30^\circ$), and counter-rotating galaxies ($\psi_{3D} > 150^\circ$); and secondly, the degree to which misalignment and counter-rotating fractions vary among the AGN and total galaxy population. In both cases, we use the pre-existing sample of 5372 galaxies at $z = 0.1$ as outlined in Section 5.2.1. We define AGN host galaxies using two separate criteria as commonly used in other studies (e.g. S. McAlpine et al. 2020). First, using a threshold of $L_{\text{bol}} > 10^{43} \text{ erg s}^{-1}$, and secondly, using a threshold of $\lambda_{\text{Edd}} > 0.01$. In total, we find 268 L_{bol} -defined AGN (of which 145 are ETGs and 123 are LTGs), and 299 λ_{Edd} -defined AGN (of which 152 are ETGs and 147 are LTGs).

In Fig. A1, we show the L_{bol} -defined AGN fraction among galaxies in three states of (mis)alignment for the total sample, and split among ETGs and LTGs. Among all galaxies studied, we see a clear increase in the AGN fraction (8.5 ± 0.8 per cent) among misaligned galaxies, compared to aligned galaxies (4.0 ± 0.3 per cent). This is also seen for ETGs, with AGN fractions of 8.7 ± 1.0 per cent and $5.0^{+0.7}_{-0.6}$ per cent for misaligned and aligned ETGs, respectively. This trend is reversed for LTGs, with higher AGN fractions among aligned ($8.8^{+0.9}_{-0.8}$ per cent) LTGs than in misaligned LTGs ($1.8^{+0.5}_{-0.4}$ per cent). These results are all significant at the 2σ confidence level. Counter-rotating galaxies show AGN fractions intermediate between aligned and misaligned galaxies, but this is only significant to the 1σ confidence level. Among counter-rotating ETGs and LTGs, we are limited by low number statistics and the increased AGN fraction is not significant. These results are largely unchanged for λ_{Edd} -defined AGN (not shown). We note that these states are classified using the galaxy's inherent (i.e. 3D) misalignment angles but that similar results are obtained using projected (i.e. 2D) misalignment angles.

In Fig. A2, we show the misaligned (top) and counter-rotating (bottom) galaxy fraction among the control sample (consisting of AGN and non-AGN galaxies) and both AGN samples at $z = 0.1$. In total, misalignment fractions are significantly increased among both the L_{bol} -defined AGN (38.4 ± 3.0 per cent) and the λ_{Edd} -defined AGN ($33.8^{+2.8}_{-2.7}$ per cent) compared to the control sample (22.4 ± 0.6 per cent). This increase is also seen individually among ETG and LTG populations. For ETGs, misalignment fractions of

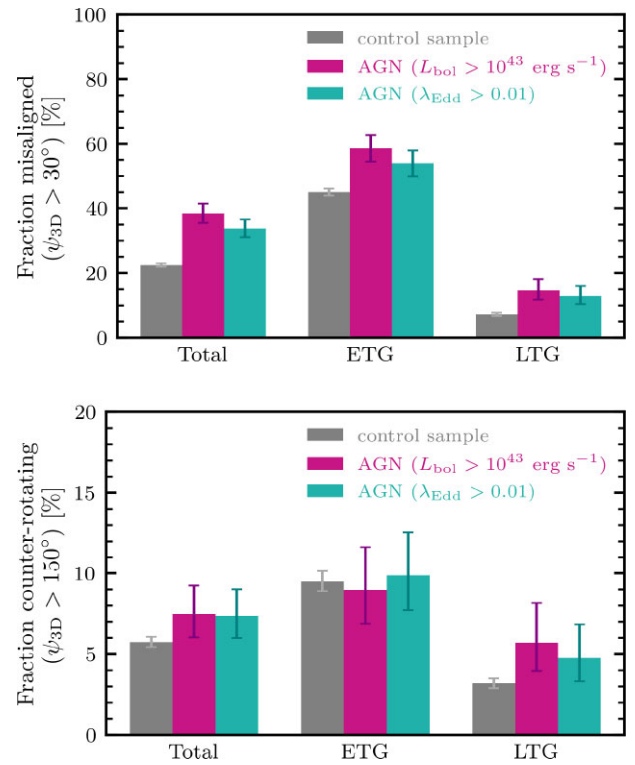


Figure A2. Fraction of misaligned (top) and counter-rotating (bottom) galaxies among the control sample (non-AGN plus AGN) at $z = 0.1$ (grey, first bar of each group), and among two samples of AGN defined using $L_{\text{bol}} > 10^{43} \text{ erg s}^{-1}$ (purple, second bar of each group), and $\lambda_{\text{Edd}} > 0.01$ (cyan, third bar of each group). Additionally, we show misalignment fractions among subsamples of ETGs and LTGs. The error bars indicate the 1σ (68 per cent) binomial confidence intervals estimated using the Jeffreys prior. We find a significantly higher fraction of misaligned galaxies in both AGN samples when compared to the control samples. We find only a significantly increased fraction of counter-rotating galaxies among the total and LTG samples for L_{bol} -defined AGN, compared to their respective control samples.

$58.6^{+4.0}_{-4.1}$ per cent are found among L_{bol} -defined AGN, compared to 45.0 ± 1.1 per cent among the total ETG population (see also M. K. Baker et al. 2025). For LTGs, misalignment fractions among L_{bol} -defined AGN ($14.6^{+3.5}_{-2.9}$ per cent) are almost double that of the typical LTG population 7.2 ± 0.5 per cent. These results are all significant at the 2σ confidence level.

These results suggest that while unstable misalignments in LTGs can increase the likelihood of AGN activity, overall, other processes dominate AGN activity in LTGs. The most likely explanation for this is the relative rarity and short duration (~ 0.4 Gyr) of unstable misalignments in LTGs (see M. K. Baker et al. 2025). Consequently, many of these unstable misalignments in LTGs are not included as part of our analysis as they do not meet our window selection criteria. Furthermore, we note that our strict morphological classification excludes some systems that are morphologically transforming during these windows and which may be classified differently using instantaneous values of κ_{co}^* .

We note that in all cases, the fraction of unstable misaligned galaxies dominates the misaligned galaxy fraction. This is evident by the significantly lower counter-rotating fractions found among all galaxies when compared to their respective misalignment fractions. Indeed, the slightly increased counter-rotating fractions are only significant for L_{bol} -defined AGN among the total and LTG populations. This is consistent with our results in Sections 5.1.1 and 5.1.2

that showed BH growth is boosted only in systems with unstable misalignments. Furthermore, the weaker increase in counter-rotating fractions seen among our AGN samples is consistent with our results from Section 5.2.

We conclude that the presence of an unstable misalignment increases the likelihood of AGN signatures at low redshifts in EAGLE. While a more detailed comparison of these results to observations is beyond the scope of this paper, these results are consistent with the observation that AGN signatures are more likely to be

observed in misaligned systems (e.g. S. I. Raimundo et al. 2023; S. I. Raimundo et al. 2025) and that AGN luminosities are not boosted in counter-rotating systems (M. W. Winiarska et al. 2025). This is despite different morphological thresholds, AGN classifications, and misalignment uncertainties compared to observations.

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