

# WEAVE imaging spectroscopy of NGC 6720: an iron bar in the Ring

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## ABSTRACT

We present spatially resolved spectroscopic observations of the planetary nebula NGC 6720, the Ring Nebula, taken during the science verification phase of WEAVE, a new instrument mounted on the William Herschel Telescope on La Palma. We use the instrument's Large Integral Field Unit (LIFU) to obtain spectra of the Ring Nebula, covering its entire optically bright inner regions as well as parts of its much fainter outer molecular halo. We report the discovery of emission from [Fe V] and [Fe VI] confined to a narrow ‘bar’ extending across the central regions of the nebula. No lines of other elements share this morphology or, at the spectral resolving power used ( $R \sim 2500$ ), the same radial velocity. The extent to which iron in this bar is depleted is presently unclear; comparison with *JWST*-detected dust continuum emission suggests that some dust grain destruction may be occurring in the region, but there is currently no observational evidence for the  $>50 \text{ km s}^{-1}$  shock waves or  $T > 10^6 \text{ K}$  X-ray emitting gas needed to enable this. Where the bar is located along the line of sight through the nebula, and how it was created, are new puzzles to be solved for this iconic planetary nebula.

**Key words:** techniques: imaging spectroscopy – stars: evolution – planetary nebulae: general – planetary nebulae: individual: NGC 6720.

## 1 INTRODUCTION

Planetary nebulae (PNe) represent an important point in the evolution of low- and intermediate-mass stars, just as their ejecta begin to enrich the interstellar medium. Spectroscopic observations of PNe

are a critical step towards a full understanding of their structure and composition. So far, few have been observed spectroscopically using integral field units capable of capturing their full extent. Here, we present the first such optical panchromatic observations of NGC 6720, the Ring Nebula, obtained using the William Herschel Telescope (WHT) during science verification of the WEAVE Large Integral Field Unit (LIFU). WEAVE is the recently installed WHT Enhanced Area Radial Velocity Explorer (G. Dalton et al. 2016;

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S. Jin et al. 2024), and the LIFU is its first-commissioned mode of operation (M. I. Arnaudova et al. 2024). This follows on from another recent instance of observations of the Ring Nebula with new instrumentation—those on board the *JWST* (R. Wesson et al. 2024; N. Clark et al. 2025; R. Sahai et al. 2025). Broad-band *JWST* infrared images are used here to complement the new spatially resolved optical spectroscopy of this object.

NGC 6720 has long served as a benchmark object for the study of planetary nebulae, owing to its proximity<sup>1</sup>, brightness, and nearly face-on orientation. Its detailed morphology – including the bright ring, outer halo, and inner cavity – has been extensively analysed across the electromagnetic spectrum (e.g. M. A. Guerrero, A. Manchado & Y. H. Chu 1997; C. R. O’Dell et al. 2013a, 2013b, 2013c; X. Fang et al. 2018; R. Wesson et al. 2024). Despite this, several key questions remain about its three-dimensional structure, ionization stratification, and internal kinematics.

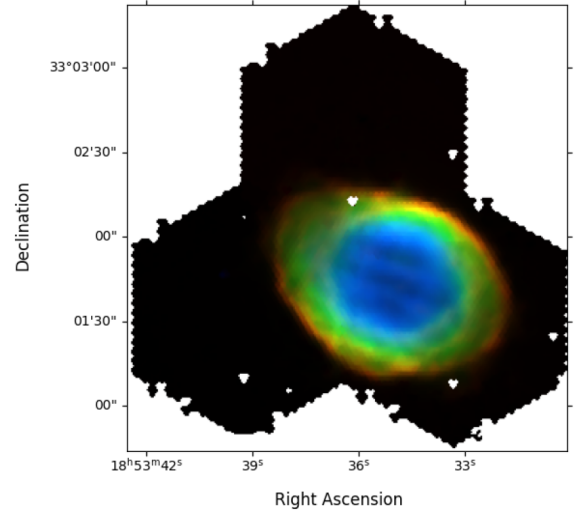
In this letter, we highlight the discovery of a peculiar feature of NGC 6720: we find a linear feature lying across the central star that presently appears to emit only in strongly ionized iron – specifically, [Fe V] 4227 Å and some weaker [Fe VI] lines. The detection of a ‘bar’ of highly ionized iron emission introduces a previously unreported structural component, with implications for the nebula’s excitation mechanisms and, potentially, for the history of mass-loss from the progenitor star. A full analysis of the whole optical nebula, as observed with WEAVE, will follow in a separate paper (Wesson et al., in preparation).

## 2 OBSERVATIONS AND DATA PROCESSING

The Ring Nebula was observed using the WEAVE LIFU mode in May and June 2023. The LIFU has 547 fibres, each 2.6 arcsec in diameter, covering a hexagonal  $78 \times 90$  arcsec<sup>2</sup> field of view with a filling factor of 0.55. The lower-resolution gratings were used for the observations, delivering a spectral resolution of  $R \sim 2500$  to both the blue and red arms of the spectrograph. The blue arm covers wavelengths from 3600 to 5955 Å, while the red covers 5800 to 9490 Å.

Observations were obtained at three pointings, with the field centres set such that the bright ring is completely covered by the mosaicked field (Fig. 1). At each position, the nebula was observed twice, with each observation consisting of three 17-min exposures, dithered in a three-point pattern to ensure complete spatial coverage of the field of view. A list of the observations used in this study is given in Table 1.

Data extraction and wavelength calibration were carried out using the WEAVE core processing system, operated by CASU in Cambridge. This results in the production of a data cube for each OB, in which the three dithers are resampled onto an equatorial coordinate grid to form a cube, with wavelengths corrected to the heliocentric reference frame. The two cubes at each pointing were then combined. We have opted to work with the data without pipeline sky subtraction, preferring to remove sky lines later as part of the customized nebular analysis. More detail on the LIFU processing may be found in M. I. Arnaudova et al. (2024). Finally, the three merged cubes were mosaicked to form a single data cube<sup>2</sup>, from which Fig. 1 and other images in this paper have been produced. As these data were obtained in the early stages of WEAVE’s operation,



**Figure 1.** Colour composite image of the Ring Nebula, reconstructed from WEAVE LIFU emission line maps of [O I] 6300 Å (red), H $\beta$  4861 Å (green), and He II 4686 Å (blue). Gaps in the reconstructed image arise from LIFU fibres which were not operational at the time of the observations.

there was no systematic collection of calibration observations and the instrument configuration was subject to change. This resulted in there being no relevant contemporaneous flux calibration data. However, the central star of the Ring Nebula, a very hot white dwarf<sup>3</sup>, is bright and well exposed, permitting its use as a flux calibrator. We compare the observed central-star spectral energy distribution (SED) with a theoretical SED (A. Csukai et al. 2025), which fits the observed SED presented by R. Sahai et al. (2025), to obtain a corrected sensitivity function that is then applied across the nebula. The theoretical SED uses a non-LTE TMAP model, optimized for PN central stars, with  $T_{\text{eff}} = 130$  kK and  $\log g = 7$  (K. Werner et al. 2003), reddened to  $E(B - V) = 0.05$  (A. Csukai et al. 2025; R. Sahai et al. 2025) using the extinction law of E. L. Fitzpatrick (1999).

## 3 EMISSION LINE MAPS

We measured both sky and nebular emission line fluxes in the data cube using ALFA, the Automated Line Fitting Algorithm (R. Wesson 2016). This publicly available code measures emission line fluxes by first subtracting an estimate of the nebular continuum. Then, on the continuum-subtracted spectrum, the code generates a population of 50 synthetic spectra with arbitrary emission line fluxes. It applies a genetic algorithm to optimize the parameters of the emission lines to achieve a good fit to the observed spectrum.

Emission lines to be fitted by ALFA are taken from an input catalogue of wavelengths. When operating on non-sky-subtracted spectra, it first fits sky emission lines using a catalogue of sky line wavelengths from R. W. Hanuschik (2003), fixing their line-of-sight velocity to zero. After fitting and subtracting sky emission, fits to the nebular emission lines are computed in a second pass. We used an input catalogue containing 168 emission lines to fit the full data cube. A full analysis of the spatially resolved emission-line spectra will be published in a forthcoming paper (Wesson et al., in preparation).

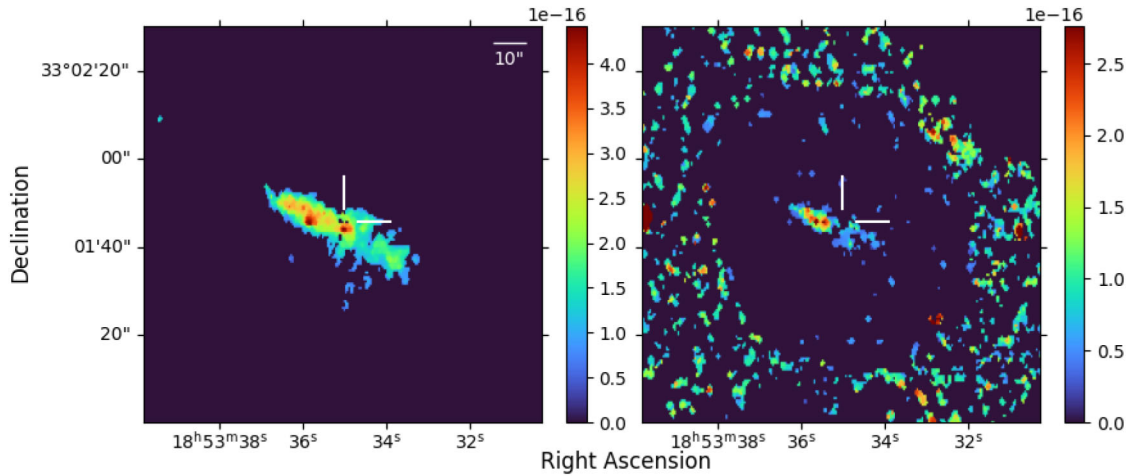
<sup>1</sup>The central star’s Gaia EDR3 good-quality parallax (Gaia Collaboration 2021) inverts to give a distance of  $787^{+29}_{-26}$  pc.

<sup>2</sup>using the PYTHON package REPROJECT (<https://reproject.readthedocs.io/>).

<sup>3</sup>Gaia DR3 2090486618786534784:  $G = 15.64$ ,  $BP = 15.59$ ,  $RP = 16.13$ .

**Table 1.** List of WEAVE science verification observations of the Ring Nebula.

OB ID	Date	Arm	RA	Dec	Airmass	Seeing (arcsec)
4876	2023-05-19	Blue	18:53:33.23	+ 33:07:11.79	1.06	1.02
	2023-05-19	Red	18:53:33.23	+ 33:07:11.79	1.06	1.02
4875	2023-05-19	Blue	18:53:36.03	+ 33:08:17.19	1.01	0.81
	2023-05-19	Red	18:53:36.03	+ 33:08:17.19	1.01	0.81
4874	2023-05-19	Blue	18:53:36.03	+ 33:08:17.19	1.01	0.81
	2023-05-19	Red	18:53:36.03	+ 33:08:17.19	1.01	0.81
4872	2023-05-20	Blue	18:53:39.13	+ 33:07:14.44	1.15	0.65
	2023-05-20	Red	18:53:39.13	+ 33:07:14.44	1.15	0.65
4873	2023-05-20	Blue	18:53:39.13	+ 33:07:14.44	1.07	0.60
	2023-05-20	Red	18:53:39.13	+ 33:07:14.44	1.07	0.60
5233	2023-06-24	Blue	18:53:33.23	+ 33:07:11.79	1.03	1.07
	2023-06-24	Red	18:53:33.23	+ 33:07:11.79	1.03	0.86

**Figure 2.** (left) [Fe V] 4227 Å emission line map, (right) summed emission of [Fe VI] emission lines at 5147, 5177, 5425, and 5678 Å. White crosshairs indicate the position of the central star. Apparent emission in the outer regions of the right-hand panel is an artefact of the low signal-to-noise ratio of the [Fe VI] lines.

#### 4 THE IRON BAR

From the spatially resolved emission line flux measurements, we constructed maps of the nebula in each emission line. An unexpected result revealed by these line maps is the presence of a narrow ‘bar’ of emission at 4227 Å (Fig. 2). We identify this line as [Fe V] 4227.20 Å. The ‘bar’ has a length of about 50 arcsec and lies along the major axis of the nebula, at a position angle measured east from north of about 70°. The morphology of this region is clearly distinct from that traced by other emission lines (Fig. 3). Previous spectroscopic studies of the Ring which employed long-slit spectroscopy (e.g. D. R. Garnett & H. L. Dinerstein 2001) would not have been able to detect this emission line if the slit was not aligned with the bar. This highlights the significant advantages of spatially resolved spectroscopy using instruments such as WEAVE to study extended objects.

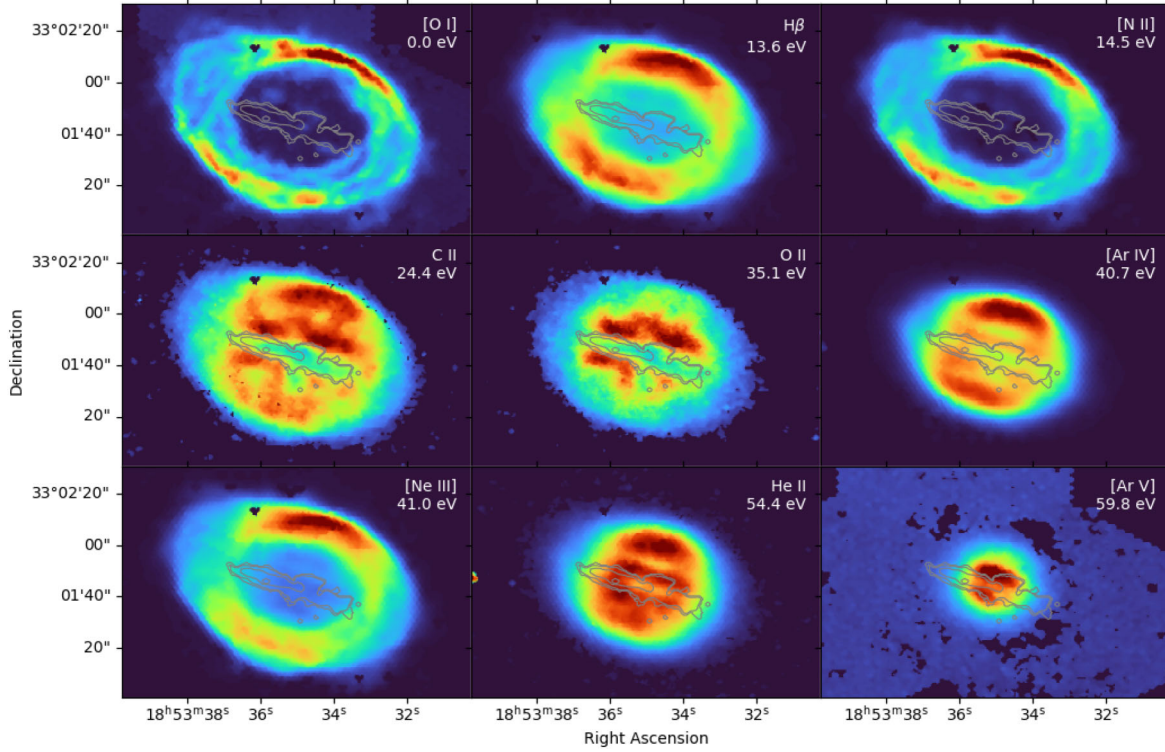
Although several other lines of [Fe V] are theoretically detectable in optical spectra, they are all weaker than the 4227 Å line, with the next strongest lines predicted to have fluxes of only  $\sim 0.2\times$  the flux of the 4227 Å line (as shown in Fig. 4). Furthermore, many of them lie close to much brighter nebular emission lines. To confirm the identification as [Fe V], we searched for other lines of highly ionized Fe. We detected several [Fe VI] lines (Fig. 4, bottom panels), but no [Fe VII] lines. The [Fe VI] lines give abundances that are broadly consistent with each other (Section 4.2), and also fall within the same ‘bar’ as the [Fe V] line (Fig. 2), thus confirming the identification of the bar’s lines as iron.

Fig. 3 shows a selection of WEAVE emission line maps, arranged in order of the ionization potentials (IPs) of the species being traced, with contours derived from the [Fe V] 4227 Å image superposed. To ionize  $\text{Fe}^{3+}$  to  $\text{Fe}^{4+}$  requires an IP of 54.8 eV. To further ionize to  $\text{Fe}^{5+}$  requires an IP of 75.0 eV, while  $\text{Fe}^{5+}$  has an IP of 100 eV. So together, the two iron ions detected require photons with energies of 54.8–100 eV to form. The distributions of the recombination lines C II 4267 Å (from  $\text{C}^{2+}$ , IP range = 24.4–47.9 eV) and O II 4649 Å (IP range 35.1–54.9 eV) are anticorrelated with that of [Fe V], with gaps, mid-nebula, just where the [Fe V] emission is located.  $\text{Ar}^{4+}$  ions span an IP range of 59.8–75.4 eV, similar to that of the  $\text{Fe}^{4+}$  ions, yet the [Ar V] 7005 Å image in Fig. 3, which is centrally concentrated as expected for such a highly ionized species, shows a gap in the middle, at the position of the [Fe V] 4227 Å emission. The He II 4686 Å image (IP range 54.4– $\infty$  eV) also does not resemble the [Fe V] and [Fe VI] images. We conclude that degree of ionization alone does not determine the observed distributions of the observed  $\text{Fe}^{4+}$  and  $\text{Fe}^{5+}$  emission.

##### 4.1 Physical conditions

To investigate the nature of the iron ‘bar’, we extracted a spectrum from a rectangular region of our mosaicked data cube, capturing the iron emission. This was centred on the central star, with dimensions  $18 \times 6.4 \text{ arcsec}^2$  and a position angle (measured east from north) of





**Figure 3.** Selected WEAVE emission-line maps, ordered by the minimum photon energy required to create the species being traced: the relevant ion is the detected ion in the case of collisionally excited lines, and the recombining ion in the case of recombination lines. The emitting species and the associated creation energy are indicated on each map. Contours derived from the [Fe V] map are overlaid in each panel. [Fe V] is produced by collisional excitation of  $\text{Fe}^{4+}$  (54.8 eV ionization potential). All images are displayed on a linear surface brightness scale with cuts at intensity percentiles of 0.1 per cent and 99.9 per cent.

72°. We determine the physical conditions and chemical abundances within the bar region using the code NEAT (R. Wesson, D. J. Stock & P. Scicluna 2012). These are given in Table 2. We measure the extinction using the four brightest hydrogen Balmer lines. The ratios of  $H\alpha/H\beta$ ,  $H\gamma/H\beta$  and  $H\delta/H\beta$  are all close to their expected Case B values<sup>4</sup>, pointing to low interstellar extinction. Previous studies have also found low interstellar extinction; R. Sahai et al. (2025) find  $A_V = 0.15$  mag, corresponding to a logarithmic extinction at  $H\beta$ ,  $c(H\beta)$ , of  $\sim 0.07$ . We proceed assuming that  $c(H\beta) = 0.0$ , but if we were to adopt a value of  $c(H\beta) = 0.1$ , the effect on our results would be small: the electron temperature  $T_e([O III])$  would be  $\sim 100$  K higher, and our derived Fe abundances would be about 2–3 per cent lower.

Electron densities and temperatures are measured from a number of standard diagnostic lines. The results are shown in Table 2. The electron densities from different diagnostics are in good agreement with each other, and we take  $460 \text{ cm}^{-3}$ , the error-weighted mean of the three measurements, to use in abundance determinations.

The temperatures calculated from [O II] and [S II] line ratios are much higher than other diagnostics. In both cases, the diagnostic lines are split between the red and blue spectral arms, and so their ratios may be more subject to systematic uncertainties. Additionally, in the latter case, the diagnostic lines at 4068 and 4076 Å are blended with several O II recombination lines, making the derived temperature unreliable. The other diagnostics are in much better agreement. The average of the four reliably determined temperature

diagnostics, weighted by their estimated uncertainties, is 10 300 K, but for calculating abundances, we use the [O III] line value only, as it is measured from well-detected lines which are all in the same spectrograph arm, and recombination of  $O^{3+}$  to  $O^{2+}$ , which could enhance the flux of the  $\lambda 4363$  line, should be less significant than recombination to singly ionized species.

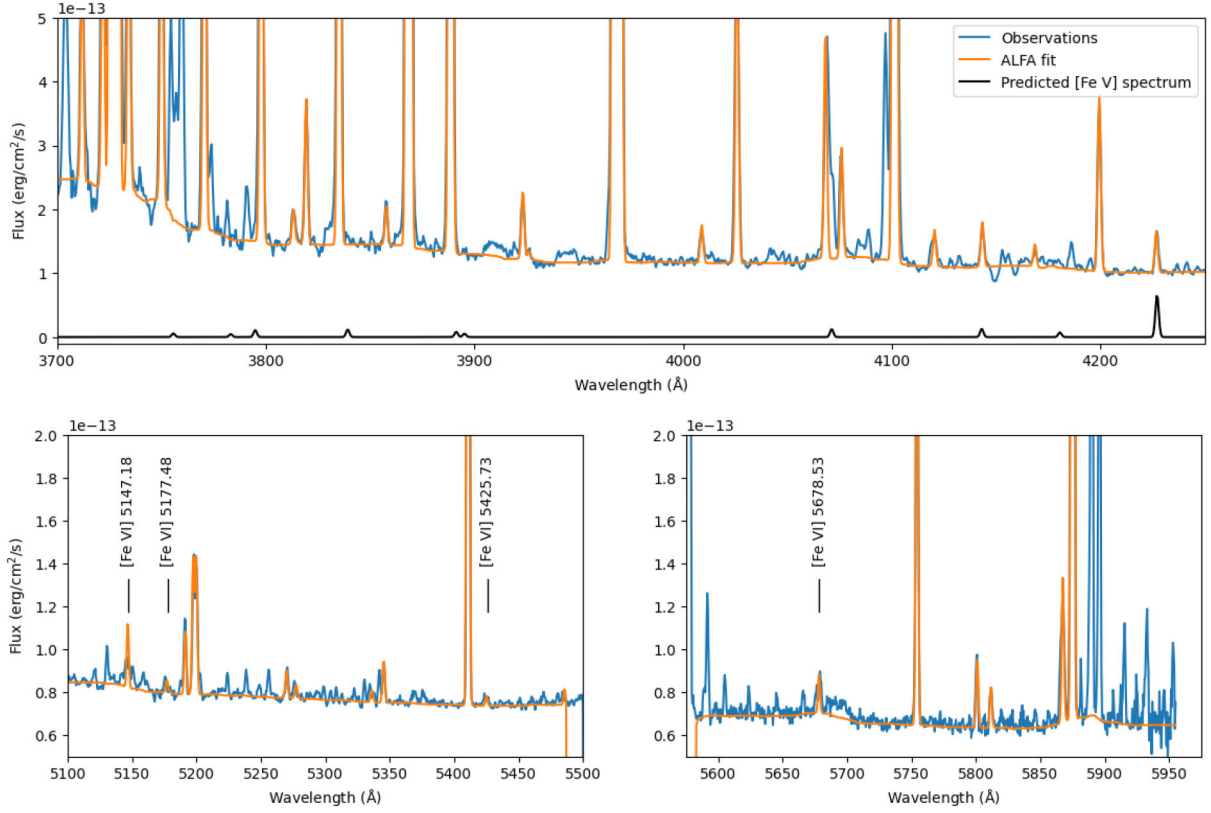
## 4.2 Iron abundance

Table 3 presents a list of the iron lines detected in the extracted ‘bar’ spectrum and the abundances derived from them. We calculate abundances using an electron temperature of 11 300 K and a density of  $460 \text{ cm}^{-3}$ , as determined in Section 4.1 and set out in Table 2. In the absence of diagnostic ratios directly pertaining to the iron bar, this is the necessary choice at this time. For [Fe V], we use collision strengths from C. P. Ballance, D. C. Griffin & B. M. McLaughlin (2007) and transition probabilities from S. N. Nahar et al. (2000). For [Fe VI], we use collision strengths from C. P. Ballance & D. C. Griffin (2008) and transition probabilities from G. X. Chen & A. K. Pradhan (2000).

The abundances we measure for  $\text{Fe}^{4+}$  and  $\text{Fe}^{5+}$  are  $7.3 \times 10^{-8}$  and  $5.4 \times 10^{-8}$ , respectively, giving an Fe/H abundance of  $> 1.3 \times 10^{-7}$  from these two ions. This is a factor of  $\sim 250$  below the solar value of  $3.24 \times 10^{-5}$  (K. Lodders et al. 2025). This depletion factor represents an upper limit, due to (i) the unobserved lower ionization stages, and (ii) the likelihood that the hydrogen line emitting volume extends further along the line of sight than the iron-emitting volume.

In NGC 6720’s bright ring, we detect the [Fe III] line at 4658.10 Å. This lies very close to C IV 4658.64 Å. Maps of this line show two spatial components, with emission interior to the bright ring

<sup>4</sup>As defined by J. G. Baker & D. H. Menzel (1938), in Case B, the nebular gas is optically thick to  $\text{Ly } \alpha$  but optically thin in all other transitions.



**Figure 4.** (top) Spectrum of the bar region, showing the clear detection of [Fe V] 4227 Å near the upper wavelength limit of the top panel. The spectrum of [Fe V] predicted for a temperature of 10 000 K and an electron density of  $1000 \text{ cm}^{-3}$  is shown in black. [Fe V] 4227 Å is by far the strongest [Fe V] line in the wavelength region observed, and a number of the already much fainter transitions lie close to other strong emission lines. (bottom) Detections of [Fe VI] lines.

**Table 2.** Electron densities and temperatures measured in the bar region of the Ring Nebula. Atomic data references can be found in R. Wesson et al. (2012). <sup>a</sup> These diagnostics are affected by line blends and recombination excitation of otherwise collisionally excited lines.

Density ( $\text{cm}^{-3}$ )	
[O II]	$480 \pm 180$
[S II]	$380 \pm 30$
[Ar IV]	$480 \pm 100$
Temperature (K)	
[O II] <sup>a</sup>	$17\,400 \pm 1500$
[S II] <sup>a</sup>	$19\,500 \pm 5100$
[N II]	$10\,100 \pm 100$
[O III]	$11\,300 \pm 100$
[Ar III]	$9100 \pm 610$
[S III]	$8500 \pm 260$

**Table 3.** Iron line intensities and ionic abundances (expressed by number) in the bar region. The observed H  $\beta$  line flux in the bar region is  $(4.96 \pm 0.09) \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ . The solar Fe/H abundance is taken from K. Lodders, M. Bergemann & H. Palme (2025).

Species	Wavelength (Å)	$I_{\lambda} (I_{H\beta} = 100)$	$X^{i+}/H^{+}$
Fe <sup>4+</sup>	4227.20	$0.25 \pm 0.02$	$(7.3 \pm 0.6) \times 10^{-8}$
Fe <sup>5+</sup>	5147.18	$0.13 \pm 0.03$	$(5.7 \pm 1.4) \times 10^{-8}$
Fe <sup>5+</sup>	5177.48	$0.02 \pm 0.01$	$(1.2 \pm 0.5) \times 10^{-8}$
Fe <sup>5+</sup>	5425.73	$0.02 \pm 0.01$	$(2.0 \pm 0.8) \times 10^{-8}$
Fe <sup>5+</sup>	5678.53	$0.08 \pm 0.03$	$(7.1 \pm 2.5) \times 10^{-8}$
Fe <sup>5+</sup>	average		$(5.4 \pm 2.5) \times 10^{-8}$
Fe <sup>4+</sup> + Fe <sup>5+</sup>			$(1.3 \pm 0.3) \times 10^{-7}$
Solar Fe/H			$3.2 \times 10^{-5}$

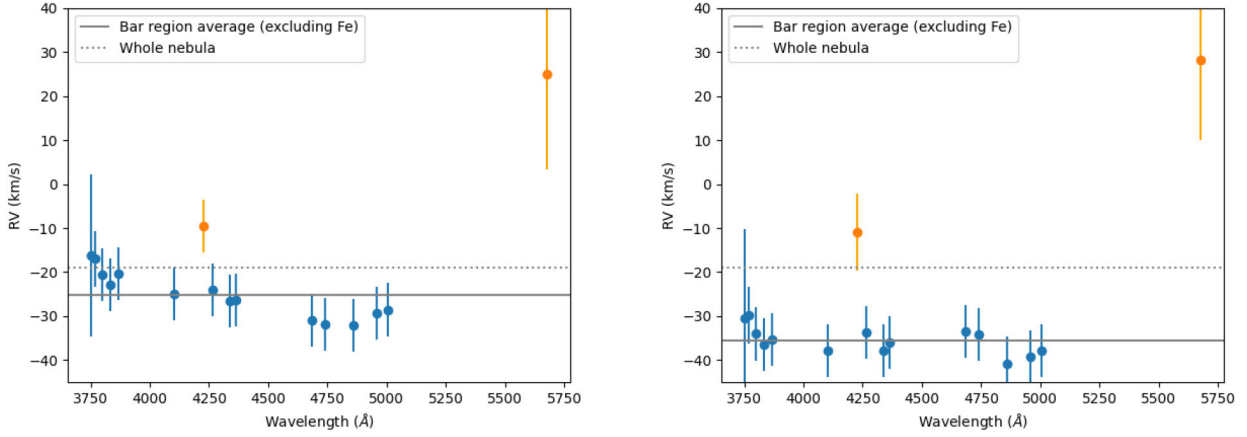
### 4.3 Kinematics

The ‘bar’ of Fe emission is brightest on the east side of the central star, but is also clearly present on the west side. Kinematic information may constrain the nature of the emission. Although our observations were taken using the LIFU’s low resolution mode, we are able to obtain some useful kinematic information.

We fit Gaussian profiles to a number of emission lines across the red and blue spectral arms, both east and west of the central star. If the ‘bar’ represented some kind of collimated bipolar outflow, one would expect to find red-shifted emission on one side of the star and blue-shifted emission on the other, unless it were to lie close to the plane of the sky.

Fig. 5 shows the radial velocities derived by fitting Gaussian

having a similar spatial pattern to He II emission, and emission in the bright ring resembling low-ionization emission lines. Attributing the emission interior to the bright ring to [C IV] and emission in the bright ring to [Fe III], we derive  $\text{Fe}^{2+}/\text{H}^{+} = (3.3 \pm 0.3) \times 10^{-8}$ , a factor of  $\sim 1000$  below the solar Fe/H abundance. The lower depletion in the bar region relative to the bright ring may suggest that iron in the bar region has been released back into the gas phase by dust destruction. The total mass of iron ions detected in the Ring Nebula’s iron bar is  $8.5 \times 10^{26} \text{ g}$  (56 per cent Fe<sup>4+</sup>; 44 per cent Fe<sup>5+</sup>), equivalent to 0.14 of an Earth mass.



**Figure 5.** Heliocentric radial velocities measured from individual spectral lines in the blue spectral region, for the bar regions east (left panel) and west (right) of the central star. The [Fe V] 4227 Å and [Fe VI] 5679 Å lines (plotted in orange) stand out as redshifted with respect to the median of the other lines.

profiles to a number of isolated and well-detected emission lines in the blue spectra. Although there is considerable scatter, most emission lines have broadly consistent radial velocities. Two outliers are apparent; these are [Fe V] 4227 Å and [Fe VI] 5678 Å, which are redshifted relative to other emission lines by  $\sim 20 \text{ km s}^{-1}$  and  $\sim 50 \text{ km s}^{-1}$ , respectively. While the average velocity of emission lines excluding Fe lines differs by about  $10 \text{ km s}^{-1}$  between the east and west sides of the central star, the Fe line velocities are the same to within their uncertainties on both sides. These observations are not consistent with a scenario in which the most-ionized Fe emission arises from a bipolar outflow.

#### 4.4 Comparison with *JWST* imagery

We compare the WEAVE map of [Fe V] 4227 Å emission with broadband images of the Ring Nebula obtained with James Webb Space Telescope (*JWST*) and presented in R. Wesson et al. (2024). Fig. 6 shows four of the *JWST* images, representative of the NIR/MIR morphologies present, with [Fe V] emission contours superposed.

The *JWST* images show two sets of features which correlate with the optical iron emission. First, a number of filters show emission on either side of a dark lane coinciding with the optical iron emission. Filters showing this pattern are F335W and F560W (shown in Fig. 6) as well as F770W, F1130W, F1280W, and F1500W (not shown). These filters all contain emission from  $\text{H}_2$  (R. Wesson et al. 2024), and so the linear features on either side of the iron bar appear to be  $\text{H}_2$  emission.

Secondly, in filters F1000W and F1800W (shown in Fig. 6) as well as F2100W (not shown), a dark lane spatially coinciding with the optical iron emission is apparent. This is clearest in the F1000W image. R. Wesson et al. (2024) estimated that this filter contains 46 per cent continuum emission, with emission-line contributions from [S IV] (26 per cent) and [Ar III] (14 per cent). However, these estimates are based on *JWST* spectroscopy of small regions close to the bright ring. WEAVE image maps show [Ar III] emission to be much brighter in the main ring than interior to it, and the smooth distribution of emission interior to the ring in the F1000W image also suggests that, in the central region, it predominantly traces continuum emission. The anticorrelation of dust continuum emission with iron emission suggests that the iron is being released into the gas phase as dust is destroyed.

## 5 DISCUSSION

The nature of the iron ‘bar’ in the Ring Nebula is unclear. While fast collimated outflows are commonly seen in planetary nebulae, and the bar appears jet-like in projection, the kinematic information shows that it is not a jet.

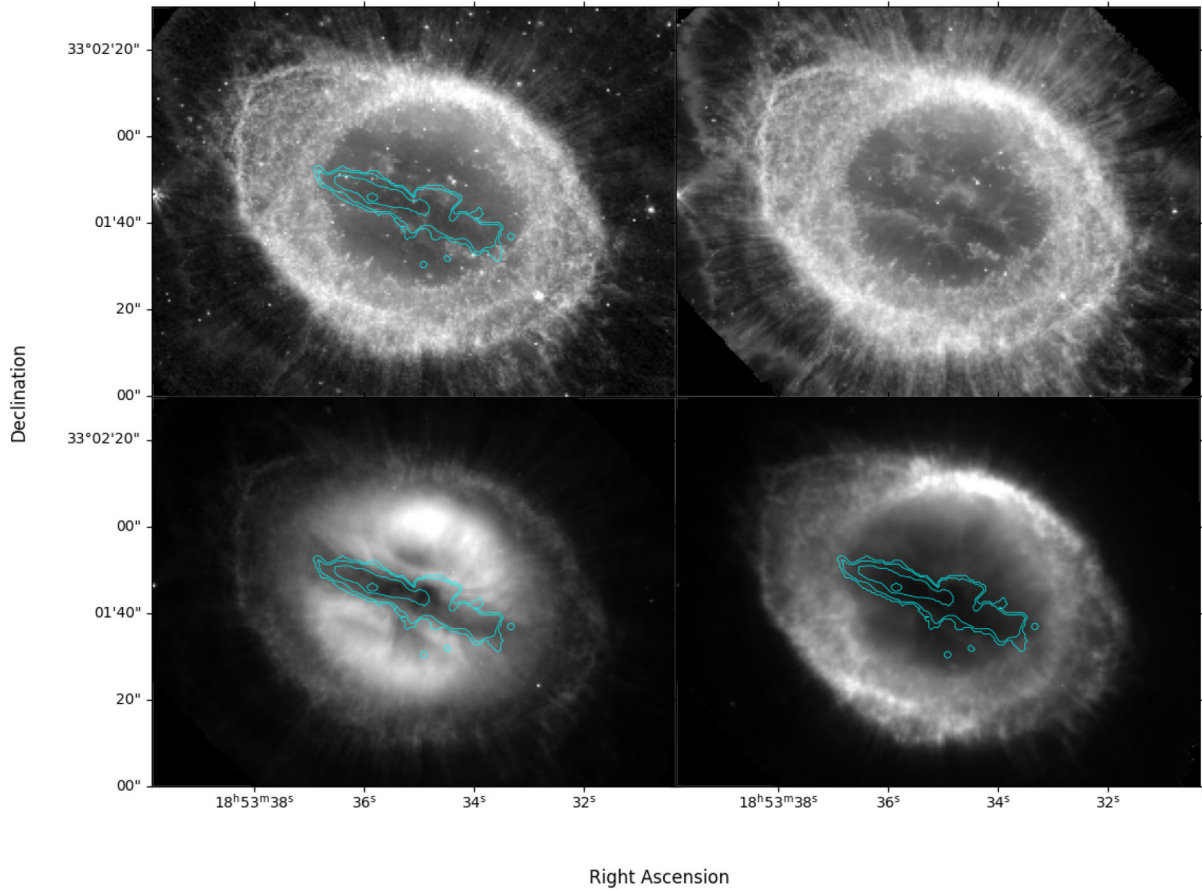
The only similar structure in another planetary nebula, to the best of our knowledge, is a recently reported detection of highly ionized Fe in spatially resolved spectroscopic observations of NGC 6818 with MUSE (Monreal-Ibero et al., in preparation<sup>5</sup>). However, in that case, the kinematic information appears to be consistent with a jet origin for the highly ionized gas.

The central star of the Ring does not bisect the iron bar, but is offset southeast of bar centre by a few arcsec (Fig. 2). R. Wesson et al. (2024) find that the central star is itself offset from the centre of the visual cavity by about 2 arcsec in the northwest direction. J. H. Kastner et al. (2025) have shown that the star is at the centre of the outer, molecular nebula, and so it is only offset within the visual cavity. The bar, on the other hand, is more nearly on the visual cavity’s geometric central axis, and is likely to be associated with the central regions of the nebula.

The  $(\text{Fe}^{4+} + \text{Fe}^{5+})/\text{H}^+$  abundance ratio of  $(1.3 \pm 0.3) \times 10^{-7}$  in Table 3 corresponds to an iron depletion factor of 250 relative to the solar value, consistent with dust formation effects: this indicates that overall, the gas in the nebula derives from the asymptotic giant branch (AGB) wind and not from the central star wind, which does not form dust. However, the hydrogen and other non-iron lines present in the iron bar spectra may very well arise mainly from a different spatial component, in front of or behind the iron emitting gas. Any line components of other species forming in the iron bar region with similar radial velocity offsets of  $\pm(15\text{--}25) \text{ km s}^{-1}$  (Fig. 5) are currently lost against the bright emission from the rest of the nebula. This prevents the true depletion in the iron bar component from being measured. A much higher spectral resolving power of at least 20 000 is needed to separate and detect what would be weak line components of other lines, in the extreme case in which the iron in the bar is undepleted, or nearly so. In this context, we note that the  $R \sim 40\,000$  long-slit spectroscopy presented by C. R. O’Dell, F. Sabbadin & W. J. Henney (2007) did provide evidence that there is some He II redshifted emission at comparable velocity (see their fig. 2).

<sup>5</sup>conference talk available at <https://www.youtube.com/watch?v=0JA2FY29zIM&t=2355s>.





**Figure 6.** Top row: *JWST* NIRCam F335W and MIRI F560W. Bottom: MIRI F1000W, MIRI F1800W. Contours derived from the WEAVE map of [Fe V] 4227 Å are displayed, except on the top right panel where they are omitted to better show the H<sub>2</sub> emission on either side of the dark lane crossing the centre.

Radiative transfer models computed using CLOUDY (C. M. Gu-nasekera et al. 2025) predict that H  $\gamma$  4340 Å in undepleted gas should be approximately as bright as [Fe V] 4227 Å. Simply multiplying the 4227/H  $\beta$  line flux ratio of  $2.5 \times 10^{-3}$  from Table 3 by the nominal iron depletion factor of 250 gives a similar result, in that a 4227/H  $\beta$  line flux ratio of 62 per cent is implied for undepleted gas.

How a component with a low or even zero gas-phase depletion of iron could be formed in the Ring Nebula is not clear. The sputtering of iron grains or iron-rich silicate grains by shock waves or by very hot gas cannot be invoked: for shock-wave destruction to occur, shock velocities significantly in excess of those materials’ sputtering threshold velocity of  $\sim 50 \text{ km s}^{-1}$  are needed (A. P. Jones et al. 1994, their figs 5 and 7). J. H. Kastner et al. (2025, their fig. 8) have detected compact CO emission components in NGC 6720 at redshifted and blueshifted radial velocities of  $\sim 45\text{--}50 \text{ km s}^{-1}$  relative to systemic, which seem too low to cause significant grain destruction. Similarly, significant thermal sputtering of such grains requires gas temperatures of  $10^6 \text{ K}$  or larger (A. G. G. M. Tielens et al. 1994, their fig. 12) – and yet, no X-ray emission has ever been reported from this low-interstellar-extinction nebula.

At present, there seem to be no obvious explanations that can account for the presence of the narrow ‘bar’ of [Fe V] and [Fe VI] emission seen in our WEAVE spectra to extend across the central regions of the Ring Nebula. Fresh observations of this newly uncovered feature at much higher spectral resolution seem essential to make progress.

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<sup>6</sup><https://weave-project.atlassian.net/wiki/display/WEAVE>

<sup>7</sup>[https://weave-project.atlassian.net/wiki/display/WEAVE/WEAVE + Acknowledgements](https://weave-project.atlassian.net/wiki/display/WEAVE/WEAVE+Acknowledgements)

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## DATA AVAILABILITY

The fully reduced and mosaicked WEAVE data cube will be available in a special data release for the ING community covering all WEAVE Science Verification data, and later as part of the first public, world-wide data release, which will take place two years after the start of full WEAVE Survey operations.

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